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## ORIENTATION AND SPEED OF THE PARSEC-SCALE JET IN NGC 4261 (3C 270)

B. GLENN PINER,<sup>1,2</sup> DAYTON L. JONES,<sup>2</sup> AND ANN E. WEHRLE<sup>2</sup>

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### ABSTRACT

NGC 4261 (3C 270) is an elliptical galaxy containing a 300 pc-scale nuclear disk of gas and dust imaged by the *Hubble Space Telescope* (*HST*), around a central supermassive black hole. Previous VLBI observations of NGC 4261 revealed a gap in emission in the radio counterjet, presumably due to free-free absorption in the inner parsec of the accretion disk. Here we present three 8 GHz VLBA observations of NGC 4261 that allow us to monitor the location and depth of the gap and check for motions in the jet and counterjet. The separation between the brightest peak and the gap is stable, with an upper limit to its motion of  $0.01c$ , supporting the interpretation of the gap as absorption by an accretion disk rather than as an intrinsic jet feature. These observations span a time of order that required for orbiting material in the disk to transit the counterjet, so we are able to search for density changes (clumps) in the disk by monitoring the optical depth of the gap. The optical depth of the gap is stable to within 20% over 5 years at  $\tau = 1.1 \pm 0.1$ , corresponding to an electron density in the disk that is constant to within 10%. We measure an apparent speed in the jet of  $(0.52 \pm 0.07)c$ . An apparent speed could not be measured for the counterjet because of a lack of identifiable features. From the apparent jet speed and the jet-to-counterjet brightness ratio, we calculate the viewing angle of the jet to be  $63^\circ \pm 3^\circ$  and its intrinsic speed to be  $(0.46 \pm 0.02)c$ . From the inclination and position angles of the parsec-scale radio jet and outer *HST* disk rotation axis we calculate a difference between the parsec-scale radio jet and outer *HST* disk rotation axis of  $12^\circ \pm 2^\circ$ . Because of its well-defined *HST* disk and bright parsec-scale radio jet and counterjet, NGC 4261 is ideal for studying the combined disk-jet system, and this is the first case known to us in which both the inclination and position angles of both the disk and jet have been determined.

*Key words:* galaxies: active — galaxies: individual (NGC 4261, 3C 270) — galaxies: jets — galaxies: nuclei — radio continuum

### 1. INTRODUCTION

NGC 4261 is a nearby E2 galaxy associated with the low-luminosity (FR I) radio source 3C 270. The large-scale radio structure of NGC 4261 shows two nearly symmetric kiloparsec-scale jets, with the slightly brighter main jet lying along a position angle<sup>3</sup> of  $-92^\circ \pm 1^\circ$  (Birkinshaw & Davies 1985). The stellar rotation axis lies along position angle  $153^\circ \pm 4^\circ$  (Davies & Birkinshaw 1986), only  $6^\circ$  from the projected major axis of the galaxy (P.A. =  $159^\circ$ ). Davies & Birkinshaw conclude that the intrinsic figure of the galaxy is prolate (in which case the projected minor axis at P.A. =  $69^\circ$  is not necessarily the minor axis of the three-dimensional figure) and that the stars may rotate about either the longest or the shortest axis of the galaxy. Nolthenius (1993) gives a distance of  $27 h^{-1}$  Mpc to NGC 4261, or 40 Mpc for an assumed Hubble constant of  $67 \text{ km s}^{-1} \text{ Mpc}^{-1}$ . We use this distance to NGC 4261 throughout this paper. At a distance of 40 Mpc, 1 mas corresponds to 0.2 pc.

The nucleus of NGC 4261 is known to contain a central black hole with a mass of  $7 \times 10^8 M_\odot$  and a nuclear disk of gas and dust with a diameter of  $\sim 300$  pc, from the *Hubble Space Telescope* (*HST*) observations of Ferrarese, Ford, & Jaffe (1996; see also Jaffe et al. 1993, 1996), with values converted to our assumed distance. The rotation axis of the *HST* disk (which is interpreted as an “outer accretion disk”

by Jaffe et al. 1993) is inclined  $64^\circ$  from the line of sight, based on isophote fitting, and has a position angle of  $-107^\circ \pm 2^\circ$  (on the side toward the main jet; Ferrarese et al. 1996). Martel et al. (1999) measure a position angle of  $-112^\circ \pm 5^\circ$  for this axis from their *HST* 3CR Snapshot Survey observation of NGC 4261, and de Koff et al. (2000) measure a position angle of  $-100^\circ$  from a “model absorption map” obtained from the same data. In this paper we use the angles measured from the deeper *HST* exposures by Ferrarese et al. The lack of a relationship between the disk rotation axis and the stellar rotation axis and the fact that the disk is not centered on the nucleus or the isophotal center of the galaxy led Ferrarese et al. to consider an external origin for the *HST* disk material in a merger event.

The first VLBI images of this galaxy were made at frequencies of 1.6 and 8 GHz (Jones & Wehrle 1997, hereafter JW97). JW97 detected a parsec-scale jet and counterjet aligned with the kiloparsec-scale jets (suggesting long-term stability of the spin axis of the central black hole) and misaligned by  $14^\circ$  in position angle from the *HST* disk rotation axis (suggesting a warping of the disk in the region close to the central black hole). JW97 also detected a gap in emission on the counterjet side of the core, at a projected distance of 0.1 pc, which they interpreted as absorption by a small dense inner accretion disk. Subsequent VLBI images at 22 and 43 GHz (Jones et al. 2000) and higher resolution VLBI images at 1.6 and 5 GHz (Jones et al. 2001) confirmed free-free absorption of radio emission by a geometrically thin disk of ionized gas and allowed estimates to be made of the electron density and magnetic field of the accretion disk on subparsec scales. Absorption by neutral hydrogen has also been detected in NGC 4261 at a projected distance of  $\sim 2.5$  pc from the core by van Langevelde et al. (2000). NGC

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<sup>3</sup> Position angle is measured from north through east.

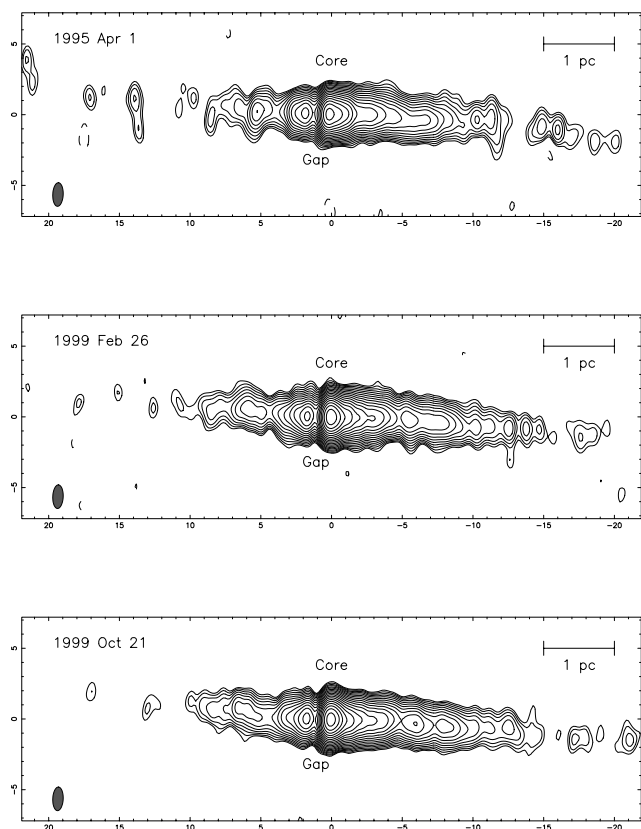


FIG. 1.—CLEAN images of NGC 4261 from the 8.4 GHz VLBA observations on 1995 April 1, 1999 February 26, and 1999 October 21. The axes are labeled in milliarcseconds (mas). All images have been restored with the uniformly weighted Gaussian beam from the 1999 February 26 epoch of  $1.7 \times 0.7$  mas FWHM at  $-2^\circ$ . The lowest contour in each image has been set to 3 times the rms noise level in the image and is equal to  $0.32$  mJy beam $^{-1}$  for the 1995 epoch and  $0.19$  mJy beam $^{-1}$  for the 1999 epochs. Successive contours are each a factor of  $\sqrt{2}$  higher. The peak flux densities are 98, 91, and 118 mJy beam $^{-1}$  for the three epochs, respectively.

4261 is not the only radio galaxy to show evidence for free-free absorption on parsec scales; such evidence has also been seen, e.g., in Centaurus A (Jones et al. 1996), Hydra A (Taylor 1996), Cygnus A (Krichbaum et al. 1998), NGC 1052 (Kellermann et al. 1999), and 3C 84 (Walker et al. 2000).

In this paper we present our multiepoch VLBI observations of NGC 4261 at 8 GHz, consisting of a single epoch from 1995 (originally presented by JW97) and two new epochs from 1999. The motivation for performing multiepoch observations was twofold. First, we wished to monitor the position and depth of the emission gap. The position of the gap should be stable if due to an accretion disk, and variations in the depth of the gap could allow detection of varying densities in the accretion disk. Second, we wanted to measure proper motions of other features in the jet and counterjet. Measurement of any two of jet proper motion, counterjet proper motion, or jet-to-counterjet brightness ratio allows calculation of the intrinsic jet speed and inclination angle. Knowledge of the radio jet inclination angle allows the complete three-dimensional orientation of the disk and jet to be found (as opposed to just the position angle offset discussed above). If all three of these quantities can be measured then even more information about the source can be obtained (including an independent measurement of the distance to the source or of the

Hubble constant, e.g., Taylor & Vermeulen 1997; Giovannini et al. 1998, or separate estimates for the fluid and pattern speeds in the jet, e.g., Taylor, Wrobel, & Vermeulen 1998; Cotton et al. 1999).

## 2. OBSERVATIONS

We observed NGC 4261 at three epochs (1995 April 1, 1999 February 26, and 1999 October 21) at 8.4 GHz with the National Radio Astronomy Observatory's Very Long Baseline Array (VLBA).<sup>4</sup> The 1995 observation has 1.5 hr of data on-source at 8.4 GHz, while the 1999 observations each recorded 8 hr of data on-source. During the 1999 October 21 observation the antenna at Brewster, WA, was not used because of a broken azimuth bearing. The 1995 observation recorded left-circular polarization, and the two 1999 observations recorded right-circular polarization. All the observations recorded a bandwidth of 64 MHz. Calibration and fringe fitting were done with the AIPS software package. Images from these data sets were produced using standard CLEAN and self-calibration procedures from the DIFMAP software package (Shepherd, Pearson, & Taylor 1994).

Images obtained from the VLBI data at these epochs are shown in Figure 1. The 1995 data set has been previously presented by JW97 but is shown again here for comparison with the more recent epochs. All images have been restored with the Gaussian beam made from the uniformly weighted data from the 1999 February 26 epoch ( $1.7 \times 0.7$  mas FWHM at  $-2^\circ$ ) to aid in comparison. The true restoring beams for the other epochs were  $1.9 \times 0.8$  mas at  $0^\circ$  for 1995 April 1 and  $2.4 \times 0.7$  mas at  $-13^\circ$  for 1999 October 21. The VLBI images clearly show a complex parsec-scale jet (to the west, or right) and fainter counterjet (to the east, or left), as well as the compact core, which has a peak flux density of about  $100$  mJy beam $^{-1}$ . The peak core flux density increased by 30% between the two 1999 observations, from 91 to 118 mJy beam $^{-1}$ . Emission from the jet and counterjet is detectable out to about 20 mas on either side of the core in these data sets. The gap in emission just to the east of the core that has been interpreted as absorption by an accretion disk by JW97 and Jones et al. (2000, 2001) is still clearly visible in the 1999 images. The gap is even more evident in the superresolved color image from 1999 October 21 that is zoomed in on the inner jet region, shown in Figure 2. The stability of this gap is discussed in the next section.

## 3. VARIABILITY OF THE JET, COUNTERJET, AND ACCRETION DISK

### 3.1. Flux Density Profiles of the Jet and Counterjet

A standard technique for analyzing VLBI images involves modeling the jet as a series of elliptical Gaussians by fitting these Gaussians either to the images or to the visibilities. This works well for a jet that can be represented by a series of discrete “blobs,” which is not the case for NGC 4261. The jets of NGC 4261, shown in Figure 1, are relatively smooth, with modest local maxima or minima superposed on an otherwise smoothly declining jet. Model-

<sup>4</sup> The National Radio Astronomy Observatory is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities for Research in Astronomy, Inc.

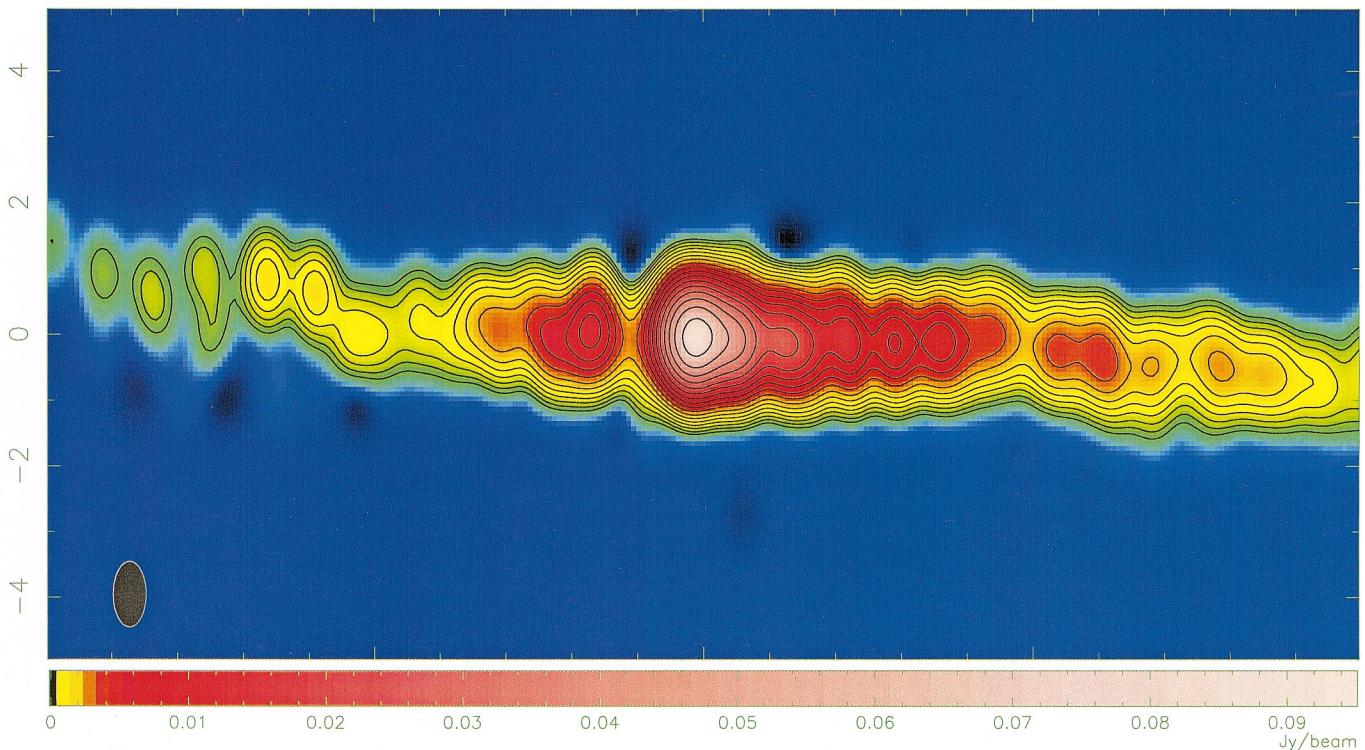


FIG. 2.—Superresolved false-color CLEAN image of NGC 4261 from the 8.4 GHz VLBA observation on 1999 October 21. The displayed region extends to  $\pm 10$  mas from the presumed core in right ascension and  $\pm 5$  mas in declination. The gap in emission is clearly visible 1 mas east of the core. The restoring beam has an FWHM of  $1.0 \times 0.5$  mas. The lowest contour represents a flux density of  $0.3 \text{ mJy beam}^{-1}$ , and successive contours are each a factor of  $\sqrt{2}$  higher.

ing the jet as a series of elliptical Gaussians would require a very large number of components and is not physically justified in this case. Instead, we study the variability of the jet and counterjet by using flux density profiles along the jets. To construct these flux density profiles we rotated the images by  $3^\circ$  (to take into account the  $-93^\circ$  position angle of the jet), restored the images with pixels one-half the size of those used in Figure 1 (but using the same beam size) to produce smoothly varying curves, and then summed the flux density along columns of pixels perpendicular to the jet. This summed flux density is expressed in janskys per mas; the total flux density is then given by integrating the area under the flux density profile curves.

### 3.2. Location and Optical Depth of the Gap

If the gap in emission to the east of the core is due to absorption by an accretion disk, then the location of the gap should be stable. In this scenario, the local maximum in emission to the east of the gap (see Fig. 1) is due to the cutoff of this absorption rather than to an enhancement of the intrinsic jet emission. This maximum should also be stable and should not propagate down the jet as do standard VLBI components.

Figure 3 shows flux density profiles along the first 4 mas of the counterjet at all three epochs, where the peak of the core emission has been used to align the profile plots. These profiles clearly show both the gap in emission and the local maximum to the east of the gap. To check the stability of the gap location, we measured the separation between the location of the maximum flux density in the core region and that of the minimum flux density in the gap region. This separation was measured to be  $1.125 \pm 0.05$ ,  $1.05 \pm 0.05$ ,

and  $1.125 \pm 0.05$  mas at 1995 April 1, 1999 February 26, and 1999 October 21, respectively, consistent with a constant position of the gap relative to the core. The upper limit to the change in the core-gap separation over these 4.6 yr is 0.075 mas, corresponding to a proper-motion upper limit of  $0.016 \text{ mas yr}^{-1}$ , or an upper limit to the apparent speed of this feature of  $0.01c$ , which is quite slow. The stability of the core-gap separation is consistent with the gap being due to absorption by a parsec-scale accretion disk. However, the stability is also consistent with the gap being a stationary feature in the jet flow. The strongest evidence for the nature

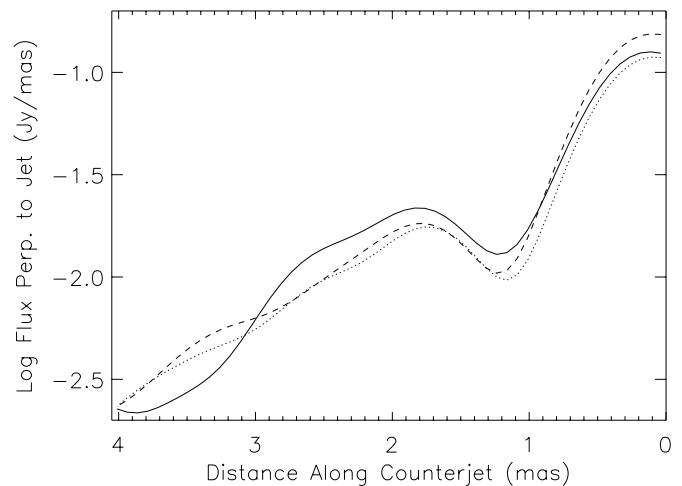


FIG. 3.—Flux density profiles along the counterjet of NGC 4261. Epoch 1995 April 1 is represented by the solid line, 1999 February 26 by the dotted line, and 1999 October 21 by the dashed line.

of the gap is the spectrum of the gap as determined from multifrequency VLBI observations (Jones et al. 2000, 2001).

If the material in the accretion disk is clumpy then we might expect to see variations in the depth of the gap as regions of the accretion disk of varying density pass in front of the counterjet. What would be the expected timescale for such variations? Given an inclination angle of  $64^\circ$  for the disk rotation axis (Ferrarese et al. 1996), the gap at a projected distance of 1.1 mas along the counterjet corresponds to a radius of 2.5 mas, or 0.5 pc, in the accretion disk. The orbital period at a radius of 0.5 pc about a black hole of  $7 \times 10^8 M_\odot$  (Ferrarese et al. 1996) is  $\approx 10^3$  yr. From Jones et al. (2000), the full apparent opening angle of the jet is between  $0.3^\circ$  and  $5^\circ$ . Using these two limits, the time for a given piece of material in the disk to pass in front of the counterjet will be of order 1–10 yr. Thus, if there are density variations, we may be able to see them over the timescale of the observations reported here. Our observations would be sensitive to clumps with a size of order the width of the counterjet at the gap location ( $\sim 10^{-2}$  to  $10^{-3}$  pc from the angular limits given above) in the direction transverse to the counterjet, but in the direction along the counterjet they would be limited by the size of the beam to clumps with a size of order 0.1 pc, which would take considerably longer (30 yr) to completely pass over the counterjet. A clump with these dimensions and with the path lengths and number densities given by Jones et al. (2001) would have a mass of order  $10^{-2}$  to  $10^{-1} M_\odot$ .

The optical depth in the region of the gap can be estimated by interpolating the jet emission over the gap region and then comparing this estimated intrinsic flux density with the observed flux density. We assume a power-law form for the intrinsic fading of the counterjet with distance from the core (see JW97; Xu et al. 2000) and find the power-law index by fitting the counterjet flux density profiles to the east of the gap, between 2 and 9 mas from the core. We find power-law indices of  $-2.0$  for the 1995 epoch and  $-1.8$  for the two 1999 epochs. The optical depths estimated in this fashion are  $1.1 \pm 0.1$ ,  $1.2 \pm 0.1$ , and  $1.0 \pm 0.1$  for 1995 April 1, 1999 February 26, and 1999 October 21, respectively. Since the optical depth is determined from flux density ratios, it should be unaffected by changes in the absolute amplitude scales, and we expect the main source of error in the optical depth measurement will come from imaging errors. Errors on the optical depth were estimated by having two of the authors independently image each data set and comparing the optical depths calculated from these independently imaged data sets. The optical depth measurements indicate that no variations in optical depth have been detected to within about 20%. Since for free-free absorption the optical depth is proportional to the square of the electron density, this corresponds to an electron density that is constant to within about 10% (assuming the path length through the disk has remained constant). Evidently, the accretion disk material in NGC 4261 is relatively smooth, at least on the scales sampled by these observations.

### 3.3. Jet Proper Motion

Since elliptical Gaussian components do not provide a good model for the NGC 4261 jet, we instead made estimates of jet proper motion from motion of local maxima on the jet profile plots and on the VLBI images. Figure 4 shows flux density profiles along the jet for the two 1999 epochs. Since even a relatively slow apparent speed would cause a

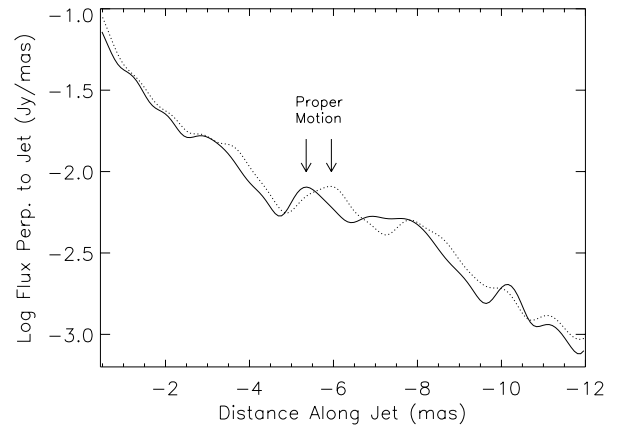


FIG. 4.—Flux density profiles along the jet of NGC 4261. Epoch, 1999 February 26 is represented by the solid line, and 1999 October 21 by the dotted line.

motion of several milliarcseconds between 1995 and 1999, connecting features between the 1995 epoch and the 1999 epochs is not feasible, and we look for proper motion only between the two 1999 epochs. For the jet, two well-defined local maxima are located at about 6 and 8 mas from the core in 1999, and they show a clear movement of about 0.5 mas between the two 1999 epochs, using the peak of the core to align the profiles. The movement of the first of these local maxima is marked on Figure 4. These local maxima are also clearly visible on the VLBI images from 1999 (in particular 1999 October 21; Fig. 1, *bottom*; 6 and 8 mas west of the core, respectively). By measuring the shift of these features between the VLBI images, as well as on the jet profile plots (using both direct measurement and an algorithm that shifts one curve with respect to the other until the sum of the squares of the difference between the two curves is minimized), we measure an apparent jet proper motion of  $0.83 \pm 0.11$  mas yr $^{-1}$ . This corresponds to an apparent speed of  $(0.52 \pm 0.07)c$ .

For the counterjet, any such local maxima that are visible during 1999 are at flux density levels much lower than those in the jet, and we conclude that an apparent speed for the counterjet cannot be reliably measured from these two epochs. Measurement of proper motion in the counterjet might be possible with additional epochs at a somewhat higher sensitivity than these and would be valuable in constraining the properties of this source.

## 4. ORIENTATION OF THE JET, COUNTERJET, AND ACCRETION DISK

### 4.1. Radio Jet Position Angle

The jets of NGC 4261 are predominantly straight, as can be seen from the plot of the jet ridgeline from the 1999 October 21 observation shown in Figure 5. Small bends of a few degrees are present, however. The jet begins at a position angle close to  $-90^\circ$ , and at 5 mas from the core bends slightly to the south. Similar (but inverted) behavior is seen on the counterjet side. In what follows we use the mean position angle of the VLBI jet over the inner 12 mas of  $-93^\circ \pm 1^\circ$ , measured from plots of the jet ridgeline, as the VLBI position angle. This agrees within the errors with the kiloparsec-scale jet position angle of  $-92^\circ \pm 1^\circ$  (Birkinshaw & Davies 1985).

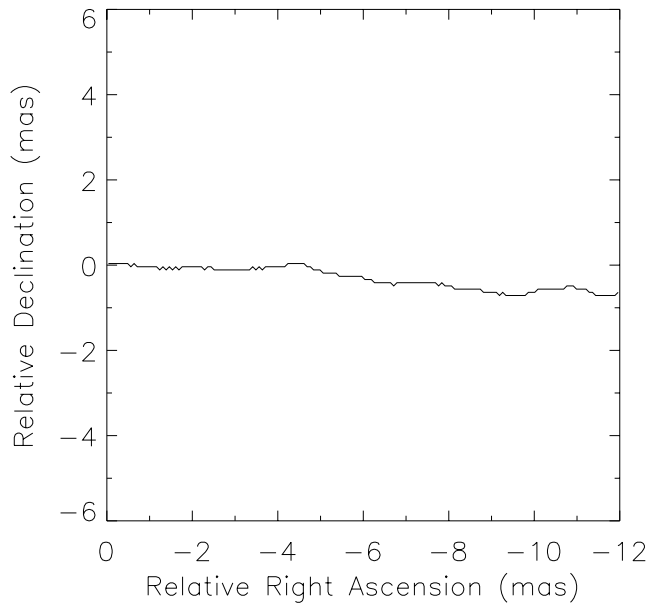


FIG. 5.—Ridgeline of the jet from the 1999 October 21 observation. The mean position angle is  $-93^\circ \pm 1^\circ$ .

#### 4.2. Radio Jet Inclination and Speed

The intrinsic speed of the jet,  $\beta$ , and the inclination angle of the jet,  $\theta$ , can be obtained from any two of the following three observable quantities: the apparent speed of the jet,  $\beta_{\text{app}}$ , the apparent speed of the counterjet, or the jet-to-counterjet brightness ratio,  $J$ . In this section we use measured values for  $\beta_{\text{app}}$  and  $J$  to calculate the intrinsic speed and inclination angle of the jet. The apparent jet speed and the jet-to-counterjet brightness ratio are given in terms of the intrinsic speed and inclination angle of the jet by

$$\beta_{\text{app}} = \frac{\beta \sin \theta}{1 - \beta \cos \theta}, \quad (1)$$

$$J = \left( \frac{1 + \beta \cos \theta}{1 - \beta \cos \theta} \right)^p, \quad (2)$$

where  $p = 2 - \alpha$  for a continuous jet and  $\alpha$  is the spectral index expressed as  $S \propto \nu^\alpha$  (see e.g., Ghisellini et al. 1993). Solving for  $\beta$  and  $\theta$  in equations (1) and (2) yields expressions for these quantities in terms of the observable quantities  $\beta_{\text{app}}$  and  $J$ :

$$\theta = \arctan \frac{2\beta_{\text{app}}}{J^{1/p} - 1}, \quad (3)$$

$$\beta = \frac{\beta_{\text{app}}}{\beta_{\text{app}} \cos \theta + \sin \theta}, \quad (4)$$

or

$$\beta = \frac{J^{1/p} - 1}{(J^{1/p} + 1) \cos \theta}. \quad (5)$$

These calculations assume that the jet and counterjet are intrinsically identical, that the bulk fluid speed in the jet is the same as the pattern speed measured on the VLBI images, and that the distance to the source is known. Measurement of the apparent speed of the counterjet would be valuable because it would allow any one of these three assumptions to be tested (while still assuming the other

two). Equation (2) also assumes that the jet's magnetic field has a random orientation; see Giovannini et al. (1994) for a discussion of this assumption.

The jet-to-counterjet brightness ratio was calculated by measuring the total flux density in the jet and the counterjet for regions beyond 2 mas from the core (to avoid the gap region where the counterjet emission is absorbed) to the visible edge of the counterjet at 9 mas. This assumes the jets are steady state, otherwise differences in light travel time would need to be taken into account. The flux density profiles support this assumption; any “components” appear to be small perturbations on an otherwise smoothly declining jet. The jet-to-counterjet brightness ratio for NGC 4261 at 8 GHz is  $J = 2.66 \pm 0.02$ , where the error was estimated from the (surely fortuitously small) scatter in the values of  $J$  measured at all three epochs. This is a small sidedness ratio compared with typical lower limits found in a sample of 17 FR I galaxies by Xu et al. (2000).

The spectral index of the jet emission is also needed. To calculate  $\alpha$ , we measured the total flux density of the jet at four frequencies (not simultaneous): 5, 8, 22, and 43 GHz. The 5 GHz data and the 22 and 43 GHz data have been previously presented by Jones et al. (2001) and Jones et al. (2000). Figure 6 shows the spectrum of the jet from data at these four frequencies, with a power-law fit to the emission between 5 and 22 GHz. At 43 GHz the jet emission is resolved out, and this point is not included in the fit. The emission between 5 and 22 GHz is well fitted by a power law with slope  $\alpha = -0.29 \pm 0.07$ . Error bars on the flux densities were estimated to be 5% at 5 and 8 GHz and 10% at 22 and 43 GHz, from typical antenna gain corrections given by amplitude self-calibration. From these measured values of apparent jet speed, jet-to-counterjet brightness ratio, and spectral index, we calculate the intrinsic speed and inclination angle of the jet from equations (3)–(5) to be  $\theta = 63^\circ \pm 3^\circ$  and  $\beta = (0.46 \pm 0.02)c$ . The systematic errors resulting from the assumptions underlying equations (3)–(5) almost surely make the true error in these derived quantities larger than the formal errors given above.

The jet speed of  $0.5c$  is only mildly relativistic and should be compared with other known speeds in FR I radio galaxies. Summaries of measured FR I speeds are given, e.g., by Tingay et al. (1998) and Cotton et al. (1999), and it is clear that a speed of  $0.5c$  is at least typical of a pattern speed in an FR I jet, although there is evidence that bulk speeds may be

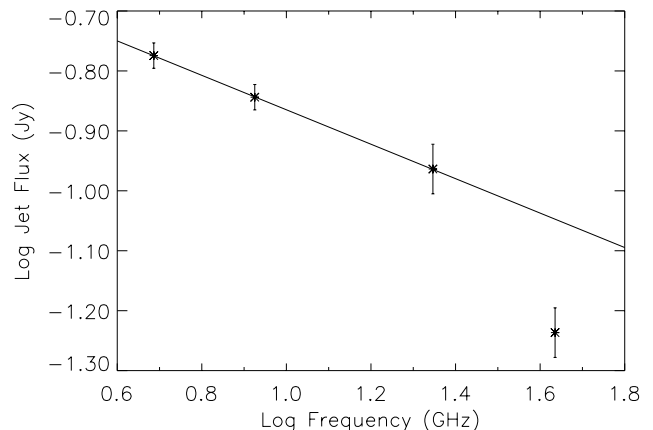


FIG. 6.—Spectrum of the NGC 4261 parsec-scale jet. The jet emission is resolved out at 43 GHz, and this point is not included in the power-law fit.

larger than observed pattern speeds in some sources (Tingay et al. 1998). Cotton et al. (1999) also find evidence that some jets may accelerate from mildly relativistic speeds as they move farther out on parsec scales. Xu et al. (2000) found that a Lorentz factor distribution with a mean of 5 reproduced the number of twin-jet sources they observed in VLBI observations of a sample of FR I galaxies, although the standard deviation of the Lorentz factor distribution was not well constrained by their observations. If the bulk flow speed in NGC 4261 were highly relativistic this would place the radio jets nearly in the plane of the sky (to reproduce the low jet-to-counterjet brightness ratio), and would make the misalignment between the VLBI jet and the *HST* disk rotation axis more severe.

#### 4.3. Relative Orientation of the Parsec-Scale Jet and Outer Accretion Disk

We can now compare the orientations of the VLBI jet and the *HST* outer disk rotation axis. The VLBI jet has an inclination angle of  $63^\circ \pm 3^\circ$  and a position angle of  $-93^\circ \pm 1^\circ$ . The *HST* outer disk rotation axis has an inclination angle of  $64^\circ \pm 5^\circ$  (from isophote fitting) and a position angle of  $-107^\circ \pm 2^\circ$  (Ferrarese et al. 1996). The angular offset between the radio jet and outer *HST* disk rotation axis is evidently only in position angle, as the inclination angles are identical within the errors. The true angular separation between the VLBI jet and the *HST* outer disk rotation axis is then  $12^\circ \pm 2^\circ$  (the apparent separation in position angle must be multiplied by the sine of the inclination angle). A sketch showing the position angles of the disk-jet system is shown in Figure 7. This angular separation is consistent with the results found by de Koff et al. (2000) from the *HST* 3CR Snapshot Survey that radio galaxies have jets nearly parallel to their dust disk axes but with some dispersion in the relation. De Koff et al. concluded their observed distribution of position angle differences could be reproduced by an intrinsically flat dis-

tribution of position angle differences of the radio jet with respect to the dust disk rotation axis between  $0^\circ$  and  $35^\circ$ .

If the VLBI jet is directed along the rotation axis of the central black hole (e.g., Rees 1978), then the rotation axes of the central black hole and outer *HST* accretion disk are misaligned by  $12^\circ \pm 2^\circ$ . The innermost regions of the accretion disk must have a rotation axis aligned with that of the central black hole, by the Bardeen-Petterson effect (Bardeen & Petterson 1975), but this happens close to the central black hole, at about 100 Schwarzschild radii ( $R_s$ ; Natarajan & Pringle 1998). For comparison, the disk observed in absorption in this paper is located at about  $10^4 R_s$ , and the size of the outer *HST* disk is about  $10^7 R_s$ . Natarajan & Pringle (1998) find that the torque that aligns the inner disk with the hole should also align the spin of the hole with the outer accretion disk on a relatively short timescale. However, more general calculations by Natarajan & Armitage (1999) show that a misalignment such as that observed here can be caused by two effects: (1) the outer disk may retain a modest warp for a long period following alignment of the hole, and (2) for holes accreting at low rates relative to the Eddington limit (which should be the case for NGC 4261, based on its low luminosity) the timescale for alignment can be much longer.

A warped accretion disk about  $5 \times 10^4 R_s$ , or 0.2 pc, in radius has been clearly observed in the peculiar spiral galaxy NGC 4258; this warp has been successfully modeled with a radiation-driven warping mechanism by Maloney, Begelman, & Pringle (1996). A possible warp in the disk of NGC 4261 is seen directly in the *HST* observations by Ferrarese et al. (1996)—with Faint Object Spectrograph data they detect an “inner” disk (about half the size of their “outer” disk) that has a rotation axis at a position angle of  $-71^\circ \pm 7^\circ$  and an inclination of  $69^\circ \pm 6^\circ$ . Although the error on the position angle of this inner disk is relatively large, the difference in position angle from the VLBI jet is significant, implying the warped disk has at least two bends. The alternative to a warped disk is that the current low accretion rate in NGC 4261 has been insufficient to perfectly align the angular momentum of the central hole. The straightness of the parsec- to kiloparsec-scale radio jets implies that the orientation of the black hole angular momentum vector has been constant for at least  $10^6$  yr (Jones et al. 2000).

#### 5. CONCLUSIONS

Our new multiepoch 8 GHz VLBA observations of NGC 4261 support the interpretation that the gap in emission in the radio counterjet is due to absorption by a parsec-scale accretion disk. The gap is stationary with respect to the core, with an upper limit to any apparent motion of 0.01c. The time required for orbiting material in the parsec-scale disk to transit the counterjet (as seen from our vantage point) is of order 1 to 10 yr, so we have been able to use the 5 year span of VLBA observations to check for density changes in the accretion disk. The optical depth ( $\tau = 1.1 \pm 0.1$ ) measured in the gap has been constant over 5 years to within our errors (20%), corresponding to an upper limit of 10% to any changes in electron density. We have measured the apparent speed of the radio jet ( $0.52 \pm 0.07$ )c and have used this and the jet-to-counterjet brightness ratio to calculate the viewing angle ( $63^\circ \pm 3^\circ$ ) and intrinsic speed ( $0.46 \pm 0.02$ )c of the radio jet. The radio jet is offset from the outer *HST* disk rotation axis by

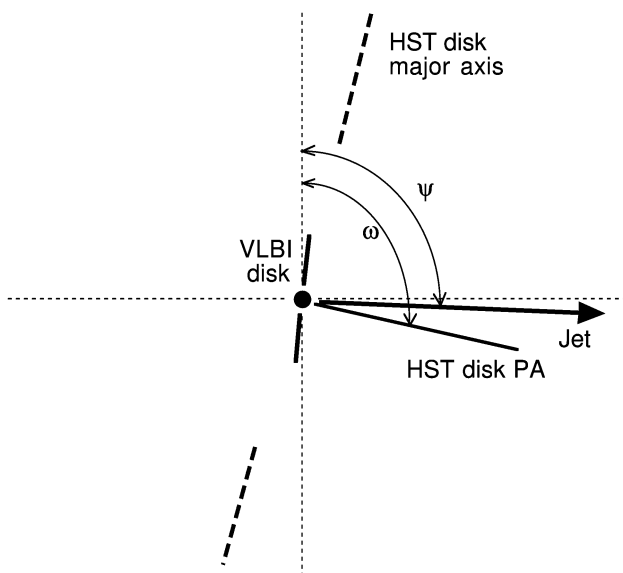


FIG. 7.—Cartoon showing the position angles of the *HST* disk and VLBI jet. Angles on this figure are not exact, and the figure is not to scale. The VLBI-scale accretion disk is assumed to be perpendicular to the VLBI jet on this figure.  $\Psi$  represents the position angle of the VLBI jet ( $-93^\circ$ ), and  $\omega$  represents the position angle of the *HST* disk rotation axis ( $-107^\circ$ ).

$12^\circ \pm 2^\circ$ , suggesting a modest warp in the outer disk or an inability of the current accretion rate to perfectly align the angular momentum of the central hole. The technique pioneered here of searching for subparsec-scale density variations in the accretion disk by monitoring the absorption of counterjet emission, while still limited by the resolution and sensitivity of the VLBI observations, holds great promise for mapping accretion disk structure on scales much smaller than those that can be directly imaged.

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