



# Constraining nucleosynthesis in two CEMP progenitors using fluorine

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## ABSTRACT

We present new fluorine abundance estimations in two carbon enhanced metal-poor (CEMP) stars, HE 1429–0551 and HE 1305+0007. HE 1429–0551 is also enriched in slow neutron-capture process (*s*-process) elements, a CEMP-*s*, and HE 1305+0007 is enhanced in both, slow and rapid neutron-capture process elements, a CEMP-*s/r*. The F abundances estimates are derived from the vibration–rotation transition of the HF molecule at 23358.6 Å using high-resolution infrared spectra obtained with the Immersion Grating Infrared Spectrometer (IGRINS) at the 4-m class Lowell Discovery Telescope. Our results include an F abundance measurement in HE 1429–0551 of  $A(\text{F}) = +3.93$  ( $[\text{F}/\text{Fe}] = +1.90$ ) at  $[\text{Fe}/\text{H}] = -2.53$ , and an F upper limit in HE 1305+0007 of  $A(\text{F}) < +3.28$  ( $[\text{F}/\text{Fe}] < +1.00$ ) at  $[\text{Fe}/\text{H}] = -2.28$ . Our new derived F abundance in HE 1429–0551 makes this object the most metal-poor star where F has been detected. We carefully compare these results with literature values and state-of-the-art CEMP-*s* model predictions including detailed asymptotic giant branch (AGB) nucleosynthesis and binary evolution. The modelled fluorine abundance for HE 1429–0551 is within reasonable agreement with our observed abundance, although is slightly higher than our observed value. For HE 1429–0551, our findings support the scenario via mass transfer by a primary companion during its thermally pulsing phase. Our estimated upper limit in HE 1305+0007, along with data from the literature, shows large discrepancies compared with AGB models. The discrepancy is principally due to the simultaneous *s*- and *r*-process element enhancements which the model struggles to reproduce.

**Key words:** stars: abundances – stars: AGB and post-AGB – stars: chemically peculiar.

## 1 INTRODUCTION

Very metal-poor (VMP) stars ( $[\text{Fe}/\text{H}] < -2.0$ ) observed today are long-lived stars that have witnessed the early beginnings of the Milky Way (Frebel & Norris 2015). The chemical compositions in their stellar atmospheres encode information about the first mechanisms of chemical enrichment. That is because VMP stars formed after gas in the Milky Way was enriched by only one or two generations of star formation. Furthermore, many VMP stars obtained their chemical compositions directly through mass transfer from a close companion. For these stars in particular, the stellar atmosphere reveals information about early nucleosynthesis processes in their companion progenitors. Thus, VMP stars are a key to understanding the early evolution of the Galaxy and the first generation of stars. Great efforts have been made by many groups seeking to discover and analyse metal-poor stars (see e.g. review by Beers & Christlieb 2005). An intriguing result from the HK survey by Beers, Preston &

Shectman (1985) was the discovery that a considerable fraction of VMP stars exhibit high carbon enrichment. These carbon enhanced metal-poor (CEMP) stars are defined as those with  $[\text{C}/\text{Fe}]^1 > 0.7$  (Aoki et al. 2007) and the number of CEMP stars increases with decreasing metallicities. For example, at  $[\text{Fe}/\text{H}] = -2.0$ , the CEMP fraction reaches  $\sim 20$  per cent, while at  $[\text{Fe}/\text{H}] = -4.0$ , this fraction rises up to  $\sim 80$  per cent (Placco et al. 2014). In addition to carbon, other element enhancements are found in the stellar atmospheres of these peculiar objects such as light elements and neutron-capture elements. The abundance distribution of these elements in CEMP stars normally follows a few distinctive patterns, which have a direct connection to the enrichment mechanisms that operated in the early Galaxy.

Nowadays, CEMP stars have been classified into four subclasses based on the relative abundances of barium and europium, heavy elements produced by the slow (*s*-) and the rapid (*r*-) neutron-capture

<sup>1</sup>In this work, we adopted the standard spectroscopic notation  $A(X) = \log[n(X)/n(\text{H})] + 12$  and  $[X/Y] = \log[n(X)/n(Y)]_* - \log[n(X)/n(Y)]_\odot$ .

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process, respectively. In this paper, we use the classification adopted by Abate et al. (2015a, b)<sup>2</sup> as described below:

- (i) *CEMP-s* stars are CEMP stars that satisfy the criteria  $[\text{Ba}/\text{Fe}] > +0.5$  and  $[\text{Ba}/\text{Eu}] > 0$ ;
- (ii) *CEMP-r* stars show abundance enrichments of  $[\text{Eu}/\text{Fe}] > +1.0$  and  $[\text{Ba}/\text{Eu}] < 0$ ;
- (iii) *CEMP-s/r* are CEMP-s stars,  $[\text{Ba}/\text{Fe}] > +0.5$  and  $[\text{Ba}/\text{Eu}] > 0$ , also enriched with *r*-process elements, i.e.  $[\text{Eu}/\text{Fe}] > +1.0$ ;
- (iv) *CEMP-no* stars do not exhibit enhanced abundances in *r*- and *s*-process elements, i.e.  $[\text{Ba}/\text{Fe}] < +0.5$ .

Nomoto, Kobayashi & Tominaga (2013) summarized the different channels for the formation of CEMP stars. Some of these channels require a first- or second-generation supernova to pollute the gas cloud that forms the CEMP star with carbon and other metals. Here though, we focus on the mass transfer scenario which takes place in a binary system. According to this scenario, the CEMP star receives its carbon and heavy elements from a previous asymptotic giant branch (AGB) companion.

AGB stars produce C, N, F, and neutron-capture elements in their He-shells, the products of which are mixed to the surface. Strong winds, characteristic of AGBs, expel the chemically enhanced outer layers of the star (Busso, Gallino & Wasserburg 1999; Karakas & Lattanzio 2014). When the AGB star is in a binary system, some of this wind can be accreted by a low-mass companion star (the present-day CEMP star). As the secondary star slowly evolves, it preserves in its atmosphere the chemical signature of its companion, which at present is an unseen white dwarf (Lucatello et al. 2005). In a theoretical approach, Abate et al. (2015b) performed a detailed analysis of AGB nucleosynthesis and binary evolution of some 67 CEMP-s stars from the SAGA data base (Suda et al. 2008). Their analysis provides initial parameters and final conditions of the binary system, as well as abundance predictions for several elements in those CEMP-s stars, including fluorine.

Fluorine is a particularly rare and fragile element. With just a single stable isotope, <sup>19</sup>F, fluorine is the least abundant of the elements from C to Ca. The synthesis of <sup>19</sup>F requires a series of specific conditions, and once formed, it can be easily destroyed in the stellar interiors (Woosley & Haxton 1988; Forestini et al. 1992; Jorissen, Smith & Lambert 1992; Mowlavi, Jorissen & Arnould 1996; Meynet & Arnould 2000; Lugaro et al. 2004; Palacios, Arnould & Meynet 2005; Cristallo et al. 2009). This makes fluorine difficult to produce and extremely sensitive to the environment of its formation site. Three formation sites have been proposed as principal sources of <sup>19</sup>F. AGB stars produce fluorine through complex nuclear reactions starting from <sup>14</sup>N and captures of protons, neutrons, and  $\alpha$  particles in the reaction  $^{14}\text{N}(\alpha, \gamma)^{18}\text{F}(\beta^+)^{18}\text{O}(\text{p}, \alpha)^{15}\text{N}(\alpha, \gamma)^{19}\text{F}$  (Forestini et al. 1992). The F content is then released from the star by stellar winds (Forestini et al. 1992; Mowlavi et al. 1996; Mowlavi, Jorissen & Arnould 1998; Lugaro et al. 2008). AGB stars are the only observationally confirmed site of F production (Jorissen et al. 1992; Werner, Rauch & Kruk 2005; Zhang & Liu 2005; Pandey 2006). Other formation sites are theoretically predicted to make fluorine, however, these have not yet been confirmed by observations. Wolf-Rayet stars can also produce fluorine in a similar way to AGB stars. Although production strongly depends on the metallicity and rotation of the star (Meynet & Arnould 2000; Palacios et al. 2005; Prantzos et al. 2018). Type II supernovae can synthesize F via neutrino

spallation or neutrino process,  $\nu$ -process (Kobayashi et al. 2011b). High-energy neutrinos  $\nu_\tau$  ejected during the core-collapse interact with the <sup>20</sup>Ne outer layer producing <sup>19</sup>F (Woosley & Haxton 1988; Woosley & Weaver 1995).

The chemical evolution of fluorine is highly uncertain. Most of the tests for F production are provided by observed F abundances in stars at or near solar metallicities (see e.g. Jorissen et al. 1992; Cunha, Smith & Gibson 2008; Abia et al. 2010; Recio-Blanco et al. 2012; Jönsson et al. 2014b; Guerço et al. 2019; Ryde et al. 2020). This metallicity bias is due to the difficulties to observe the HF molecular absorption line, the most used absorption feature for the diagnostic of F abundances, in VMP stars. At very low metallicities, F has been poorly studied with just two reported measurements below  $[\text{Fe}/\text{H}] < -2$ : HE 1305+0132 by Schuler et al. (2007)<sup>3</sup> and HD5223 by Lucatello et al. (2011) (some other measurements and upper limits are also available in literature, including carbon-normal halo field stars at slightly higher metallicities – e.g. Lucatello et al. 2011; Li et al. 2013). How fluorine is made in the early Universe remains unclear. Therefore, studying F in CEMP stars is a direct probe to this question.

In this paper, we present an analysis of fluorine abundances in two CEMP stars in order to test nucleosynthesis predictions from Abate et al. (2015b). Section 2 describes the target selection and observations. Section 3 includes the data reduction and analysis. We present our results and discussion in Section 4. We finalize in Section 5 summarizing our conclusions.

## 2 TARGET SELECTION AND OBSERVATIONS

### 2.1 Target selection

Abate et al. (2015a, b) modelled AGB stellar nucleosynthesis and binary evolution at low metallicity and produced a grid of theoretical binary population synthesis models of about 400 000 binary stars with different initial masses and orbital periods. The comparison of this model grid to a sample of 67 CEMP-s stars was used to determine the best-fitting models for the observed chemical abundances, surface gravity, and period (when available) of each individual star of that sample (see Abate et al. 2015a, b, for further details). We adopted those stellar parameters ( $T_{\text{eff}}$ ,  $\log g$ , and  $[\text{Fe}/\text{H}]$ ) and the F abundance predictions of each star from Abate et al. (2015b, and references therein), to generate synthetic spectra with the local thermodynamic equilibrium (LTE) stellar line analysis program MOOG (version 2014; Sneden 1973). Using the vibration–rotation transition of the HF molecule line near 2.3  $\mu\text{m}$ , we identified potential targets as those showing a synthetic HF absorption line stronger than  $\sim 5$  per cent relative to the continuum. From these potential targets, we selected two stars to perform our follow-up observations, HE 1429–0551 and HE 1305+0007.

HE 1429–0551 is a CEMP-s star with no reported abundances of Eu. We adopted the stellar parameters compiled from the literature by Abate et al. (2015b):  $T_{\text{eff}} = 4700$ ,  $\log g = 1.50$ , and  $[\text{Fe}/\text{H}] = -2.53$  from Aoki et al. (2007). A barium abundance of  $[\text{Ba}/\text{Fe}] = +1.57$  is also reported by Aoki et al. (2007). Models by Abate et al. (2015b) predict an F content in HE 1429–0551 of  $A(\text{F})_{\text{Abate}} = +4.40$ .

HE 1305+0007 is classified as a CEMP-s/r object by Goswami et al. (2006) and exhibits large enhancements in *s*- and *r*-process

<sup>2</sup>These definitions vary slightly among other authors: Beers & Christlieb (2005); Aoki et al. (2007); Masseron et al. (2010).

<sup>3</sup>This star is in fact more metal rich and a later reanalysis by Schuler et al. (2008) suggested a different result on the observed F abundance (see Section 4.1 for more details).

**Table 1.** Star ID and stellar parameters from literature data compilation by Abate et al. (2015b) in columns 1–5; our observed fluorine abundance in columns 6 and 7; modelled abundances from Abate et al. (2015b) in columns 8 and 9. Additionally, we have included stellar parameters, observed A(F) and [F/Fe] from Lucatello et al. (2011), and the predicted A(F) from Abate et al. (2015b) for in common stars. Note that F abundance in HE 1305+0007 has been also analysed by Lucatello et al. (2011; see text).

Star ID	$T_{\text{eff}}$ (K)	$\log g$	$\xi_r$ (km s $^{-1}$ )	[Fe/H]	A(F)	[F/Fe] <sup>a</sup>	A(F) <sub>Abate</sub>
HE 1429–0551	4700	1.50	1.38	–2.53	3.93	1.90	4.40
HE 1305+0007	4655	1.50	1.41	–2.28	<3.28	<1.00	4.52
Lucatello et al. (2011) sample							
CS 22942–019	5100	2.5	–	–2.5	<4.2	<2.1	5.01
CS 29497–034	4800	1.8	–	–2.9	<3.5	<1.8	–
CS 30314–067	4400	0.7	–	–2.85	<2.0	<0.3	–
CS 22948–027	4800	1.8	–	–2.5	<3.6	<1.5	4.5
CS 29502–092	4890	1.7	–	–3.18	<3.9	<2.5	–
CS 30322–023	4100	–0.3	–	–3.39	<1.8	<0.6	–
HE 1152–0355	4000	1.0	–	–1.27	3.9	0.64	–
HE 1305+0007	4560	1.0	–	–2.5	<3.5	<1.4	4.52
HD 187861	4600	1.7	–	–2.36	<3.7	<1.6	4.61
HD 5223	4500	1.0	–	–2.06	4.0	1.44	3.89
HD 122563	4600	1.1	–	–2.82	<2.5	<0.8	–

<sup>a</sup>The adopted solar abundance for F is from Grevesse, Asplund & Sauval (2007), A(F)<sub>⊙</sub> = 4.56.

**Table 2.** Observing log from infrared spectroscopy by IGRINS.

Star ID	RA (2000)	Dec. (2000)	UT date	Exposure time (s)	$K^a$
HE 1429–0551	14:32:31.56	–06:04:59.50	2019 April 09	8 × 450	10.07
			2019 April 14	4 × 480	10.07
HE 1305+0007	13:08:04.11	–00:08:39.60	2019 April 14	4 × 300	9.60

<sup>a</sup>K magnitudes from 2MASS (Cutri et al. 2003).

elements. By using high-resolution spectra, Goswami et al. (2006) derived the stellar parameters:  $T_{\text{eff}} = 4750$  K,  $\log g = 2.0$ , and  $[\text{Fe}/\text{H}] = -2.0$ , as well as a barium abundance of  $[\text{Ba}/\text{Fe}] = +2.32$ , and a remarkable high content in europium,  $[\text{Eu}/\text{Fe}] = +1.97$ . Beers et al. (2007) estimated, using  $V - K$  colour and mid-resolution optical and near-infrared spectra, the stellar parameters:  $T_{\text{eff}} = 4560$  K,  $\log g = 1.0$ , and  $[\text{Fe}/\text{H}] = -2.5$ , in addition to a Ba abundance of  $[\text{Ba}/\text{Fe}] = +2.9$ . We adopted the mean values from the stellar parameters, previously mentioned, compiled by Abate et al. (2015b). This is:  $T_{\text{eff}} = 4655$  K,  $\log g = 1.5$ , and  $[\text{Fe}/\text{H}] = -2.28$  (see Table 1). Abate et al. (2015b) predict an F abundance in HE 1305+0007 of  $A(\text{F})_{\text{Abate}} = 4.52$ .

Both objects appear to be binaries based on previous radial velocity variations. However, orbital solutions have not yet been published. HE 1429–0551 exhibits radial velocity variations above the  $3\sigma$  level, but no clear trend is observed to support orbital motion (Jorissen et al. 2016). HE 1305+0007 radial velocities were first reported by Goswami et al. (2006), and later by the Gaia Collaboration (2018), both studies have uncertainties  $>1$  km s $^{-1}$  and a velocity discrepancy  $>10$  km s $^{-1}$ . HE 1305+0007 has been flagged as a binary candidate (Arentsen et al. 2019).

## 2.2 Observations

The CEMP stars HE 1429–0551 and HE 1305+0007 were observed using the Immersion Grating Infrared Spectrometer (IGRINS; Yuk et al. 2010; Park et al. 2014; Mace et al. 2018), mounted on the Lowell Discovery Telescope at the Lowell Observatory. Observations were performed under service mode on 2019 April 9 and 14. The 0.63 arcsec slit provided spectral resolution of  $R \sim 45\,000$ , and covered

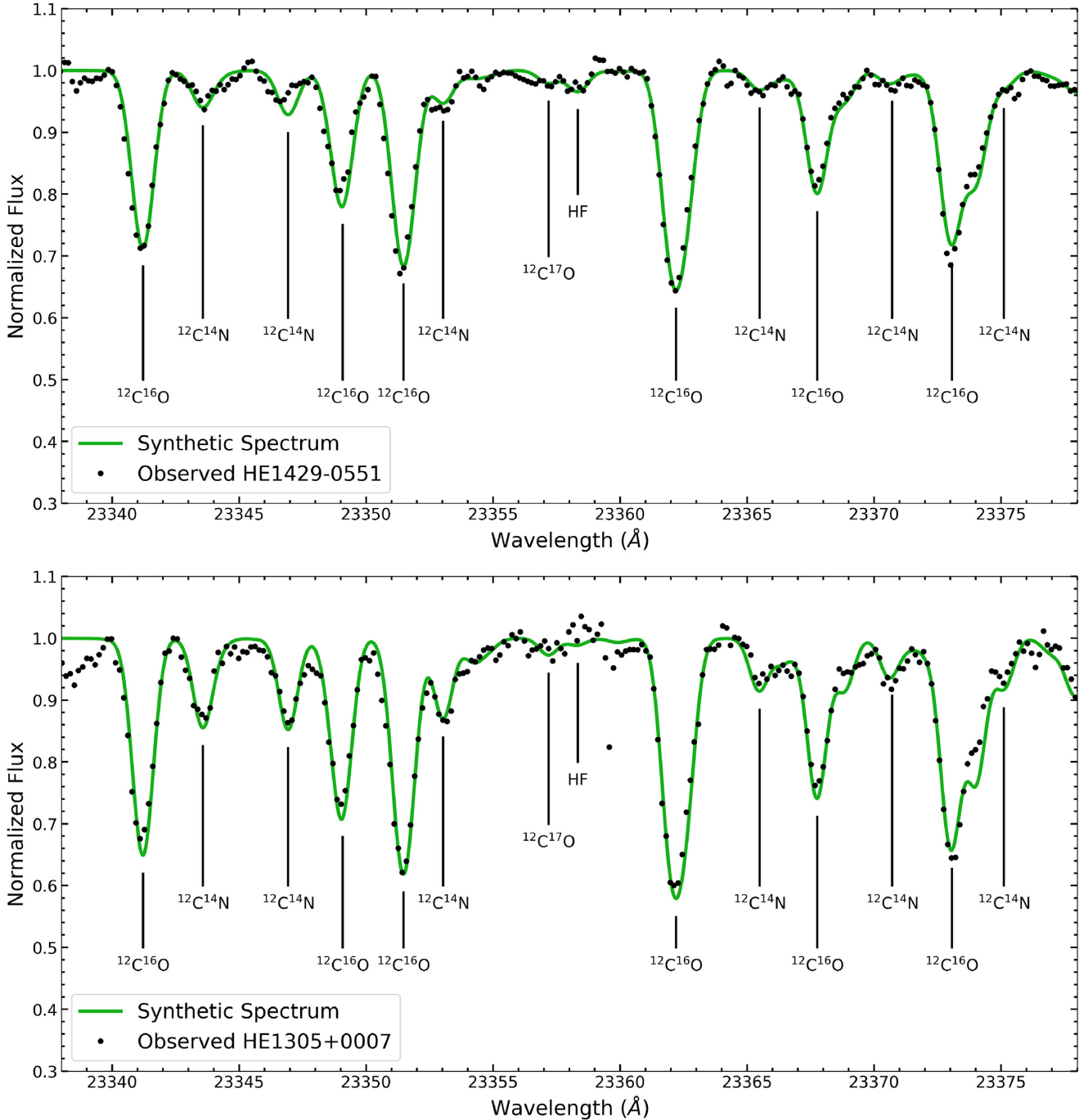
the infrared  $H$  and  $K$  bands from 1.45 to 2.45  $\mu\text{m}$ . Although the instrument has a large spectral coverage, for the purpose of this work we use a small fraction of the spectra at  $\sim 2.3$   $\mu\text{m}$ .

The entire set of observations were executed under ‘ABBA’ nod sequences. For HE 1305+0007, one ABBA sequence was programmed with an exposure time for each frame of 300 s (i.e.  $4 \times 300$ ). HE 1429–0551 was observed twice during the first night and once during the second night with exposure times of 450 and 480 s per frame, respectively. The signal-to-noise ratios (S/N), exceed 150 per resolution element in the wavelength region of interest, at 2.3  $\mu\text{m}$ , in each star. In addition to the science observations, rapidly rotating hot stars were observed each night, close in time to our program targets, in order to obtain telluric spectra. The telluric spectra are divided into the science spectra of HE 1305+0007 and HE 1429–0551 in order to remove the telluric absorption lines. The program stars and log of observations are listed in Table 2.

## 3 DATA REDUCTION AND ANALYSIS

Wavelength-calibrated spectra were produced using the IGRINS reduction pipeline package PLP2 (Lee, Gullikson & Kaplan 2017). PLP2 performs the flat-fielding, background removal, order extraction, and wavelength calibration. IRAF<sup>4</sup> tasks were used to combine the individual spectra into a single spectrum and to normalize the

<sup>4</sup>IRAF (Image Reduction and Analysis Facility) is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., under contract with the National Science Foundation.



**Figure 1.** Observed spectra (black circles) and best-fitting synthetic spectra (green line) for HE 1428–0551 (upper) and HE 1305+0007 (lower) near the HF line. Note that while the HF line is weak, most of the absorption lines are well fitted. In the lower panel, the HF line region is higher than the continuum suggesting issues with the telluric division (see Section 4).

continuum. Inspection of the spectrum of HE 1305+0007 (See Fig. 1) indicates that there was likely an issue with telluric removal in the vicinity of the HF line, i.e. the normalized flux exceeds 1.0. We refer to the effect of this last issue on our F abundance estimation in HE 1305+0007 in Section 4.

The adopted stellar parameters ( $T_{\text{eff}}$ ,  $\log g$ ,  $[\text{Fe}/\text{H}]$ ) for our sample and its literature source are described in Section 2.1 and Table 1. The microturbulent velocity  $\xi_t$ , was determined using the relation  $\xi_t = 4.2 - (6 \times 10^{-4} T_{\text{eff}})$  (Meléndez et al. 2008). The

abundance dependence due to uncertainties in the adopted stellar parameters for HE 1429–0551 are presented in Table 3. The reason that we adopted the stellar parameters from the literature instead of deriving them directly from the spectra is that (1) adopting new stellar parameters would lead to inconsistency with the values used in the Abate predictions and (2) new stellar parameters cannot be easily obtained from the IGRINS spectra alone. This is due to the lack of Fe lines in the wavelength coverage of the instrument used for the estimation of  $T_{\text{eff}}$  and  $\log g$  through excitation and ionization



**Table 3.** Fluorine abundance dependence upon model parameter uncertainties for HE 1429–0551.

Star ID	$\Delta T_{\text{eff}}$ (K)	$\Delta \log g$	$\Delta [\text{Fe}/\text{H}]$	$\Delta \xi_r$ (km s $^{-1}$ )	Total <sup>a</sup>
	–100	+0.50	+0.30	+0.30	
HE 1429–0551	+0.30	+0.17	+0.20	+0.00	+0.40

<sup>a</sup>The total value is the quadrature sum of the individual abundance dependences.

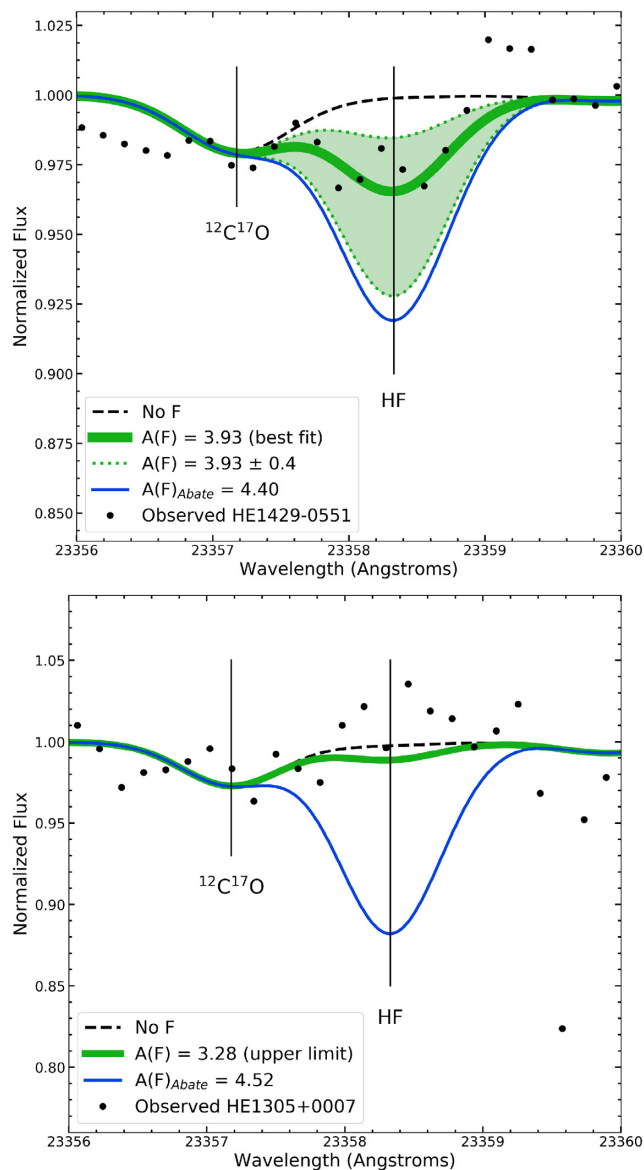
equilibrium, respectively. Alternative, photometric temperatures are feasible, but these are C-rich objects and caution should be exercised when using the standard calibrations which are generally based on C-normal objects.

The fluorine abundances were derived by comparing observed and synthetic spectra. We used the LTE program MOOG (Snedden 1973) coupled with one-dimensional LTE model atmospheres from the Kurucz (1993) grid to generate the synthetic spectra. Following previous studies of fluorine (e.g. Smith et al. 2005), we studied the vibration–rotation transition of the HF molecule absorption line, HF (1–0) R 9, at 23358.329 Å. Although this HF molecular line is not the strongest in the wavelength region, it is free of blends and is considered a reliable indicator of F abundances (Abia et al. 2009; Lucatello et al. 2011; Jönsson et al. 2014a). We adopted the most recent data for the HF(1–0) R9 molecular line from Jönsson et al. (2014a, and references therein); excitation potential  $\chi = 0.227$  and  $\log gf = -3.962$ . These HF features were incorporated in our line list and are compatible with the partition function built into MOOG. By doing this we ensure that we are using energies consistent with the partition function integrated in MOOG, avoiding overestimations in the F abundances previously reported in literature (e.g. Lucatello et al. 2011; D’Orazi et al. 2013; Jönsson et al. 2014b). We adopted additional molecular data from Goorvitch (1994) and Sneden et al. (2014) for CO and CN, respectively. For the fitting of CO and CN molecular absorption in the spectral region nearby HF, we adopted the observed C abundances from Aoki et al. (2007) for HE 1429–0551, and Goswami et al. (2006) and Beers et al. (2007) for HE 1305+0007, then we adjust O and N in order to fit our observed spectra in these stars. Our best-fitting synthetic spectra are shown in Figs 1 and 2. Note that the  $^{12}\text{C}^{17}\text{O}$  line at 23358.237 Å, is slightly blended with the blue wing of the HF line under study. However, this partial blending does not affect significantly the fluorine measurements from the HF(1–0) R9 molecular line (Lucatello et al. 2011).

#### 4 RESULTS AND DISCUSSION

Our derived fluorine abundances are:  $A(\text{F}) = +3.93$  ( $[\text{F}/\text{Fe}]^5 = +1.90$ ) for HE 1429–0551 and  $A(\text{F}) < +3.28$  ( $[\text{F}/\text{Fe}] < +1.00$ ) for HE 1305+0007. The adopted metallicity on these CEMP stars are  $[\text{Fe}/\text{H}] = -2.53$  and  $[\text{Fe}/\text{H}] = -2.28$  for HE 1429–0551 HE 1305+0007, respectively.

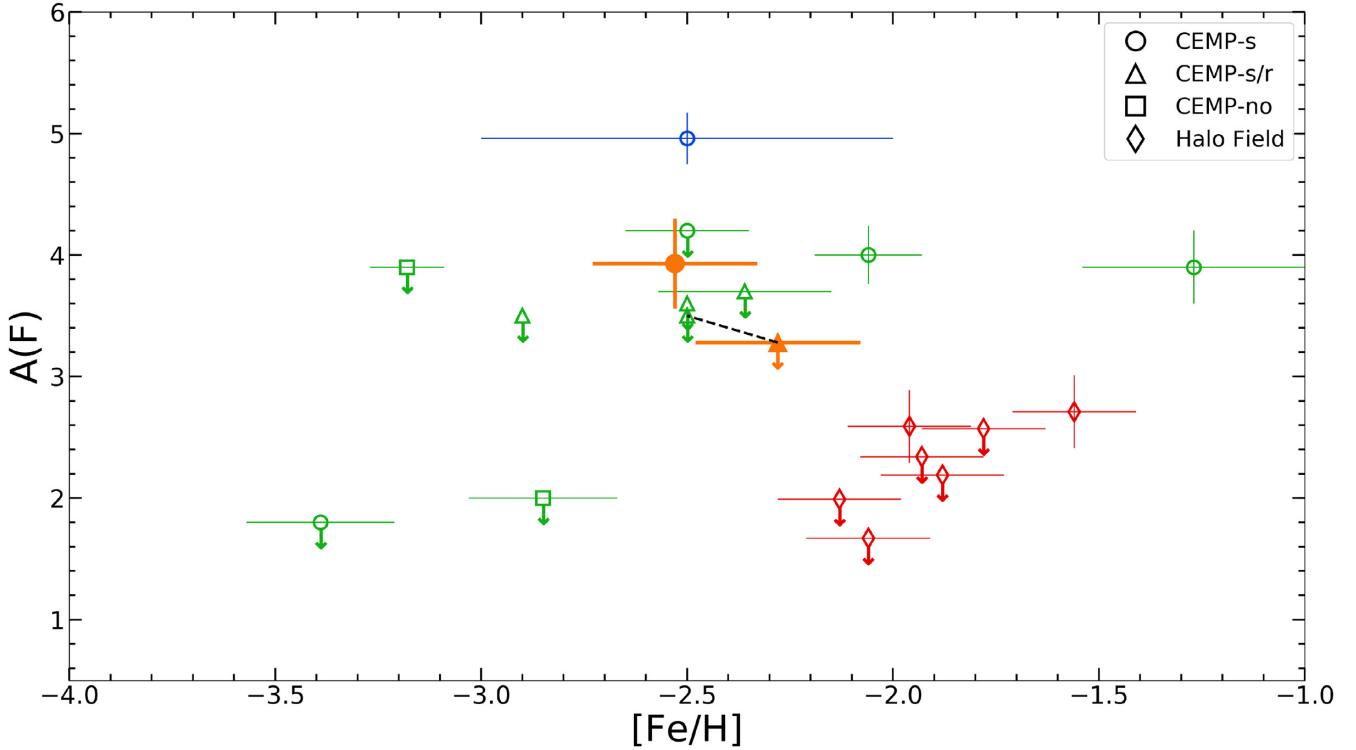
In Fig. 3, we plot our derived  $A(\text{F})$  abundances against  $[\text{Fe}/\text{H}]$  along with the literature data of F abundances on CEMP stars i.e. Schuler et al. (2007) and Lucatello et al. (2011). In this figure, we also include literature values of F abundances in carbon-normal halo field stars from Li et al. (2013) as a comparison. Fig. 4 compares our fluorine abundance estimations with those predicted by Abate et al. (2015b). In Fig. 4, additionally to our results, we include part of the Lucatello et al. (2011) sample which is in common with Abate et al. (2015b). For clarity, we connect the observed



**Figure 2.** Observed and synthetic spectra near the HF 23358.329 Å line for HE 1429–0551 (upper) and HE 1305+0007 (lower). In the upper panel, the solid green line is our best fit for HE 1429–0551 with  $A(\text{F}) = 3.93$ . The uncertainty region of our fit is shown in shaded green area and delimited by a dotted green line with  $A(\text{F}) = 3.93 \pm 0.4$ . The predicted F abundance from Abate et al. (2015b) of  $A(\text{F})_{\text{Abate}} = 4.40$  is shown in the solid blue line. In the lower panel, our upper limit for HE 1305+0007 of  $A(\text{F}) = 3.28$  is the solid green line and the predicted value from Abate et al. (2015b) of  $A(\text{F})_{\text{Abate}} = 4.52$  is the solid blue line.

F abundances with the predictions using a black dashed line. The horizontal and vertical sections of the connecting line correspond to differences in the adopted metallicities and F abundances estimates, respectively, by Lucatello et al. (2011) and Abate et al. (2015b). We

<sup>5</sup>We use the F solar abundance from Grevesse et al. (2007),  $A(\text{F})_{\odot} = 4.56$ , which is also the adopted solar abundance in the comparison literature. For an alternative F solar abundance, see Maiorca et al. (2014).



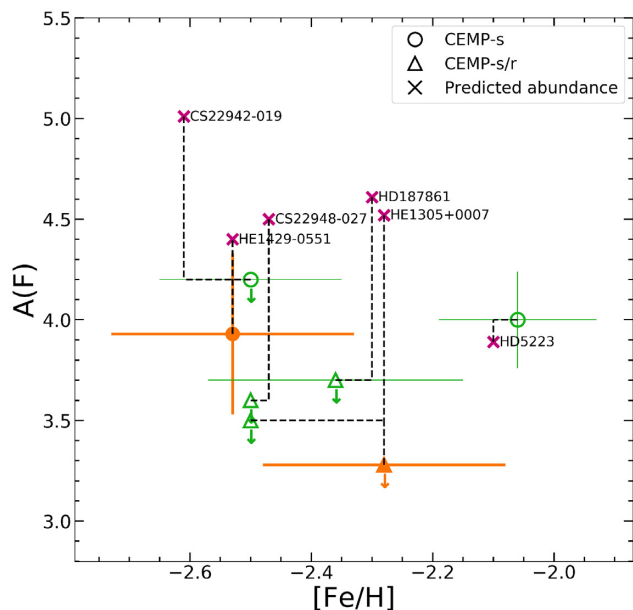
**Figure 3.** Our fluorine estimations along with literature values for  $A(\text{F})$  versus  $[\text{Fe}/\text{H}]$ . Our fluorine measurement is shown as the filled orange circle for HE 1429–0551 and our upper limit in filled orange triangle for HE 1305+0007. In blue, HE 1305+0132 from Schuler et al. (2007); in green, HE 1152–0355, HD 5223, and other upper limits from Lucatello et al. (2011); in red, HD 135148, HD 110281, and other upper limits from Li et al. (2013). The black dashed line shows the differences in the adopted  $[\text{Fe}/\text{H}]$  and the upper limit for the star in common with Lucatello et al. (2011), HE 1305+0007.

show in Figs 1 and 2 our best synthetic fit for our F measurement in HE 1429–0551 and F upper limit in HE 1305+0007. In addition to our best fit, we include in Fig. 2 a synthetic HF absorption line using the predicted F abundance by Abate et al. (2015b) for comparison. We suspect that the telluric absorption in our data for HE 1305+0007 has not been fully corrected since the normalized flux in the HF line is higher than 1.0. However, the extent of this effect in our derived upper limit is expected to be small. The predicted strength of the HF line at very low metallicities is weak, which is the principal reason for the high S/N ratio in our observations. We note that there is a  $\sim 3$  per cent excess in the normalized flux of HE 1305+0007 (see lower panel in Fig. 2). If there was an HF feature, we assume that it would have been detected if it were stronger than  $\sim 5$  per cent relative to the continuum. Moreover, the typical error in the observations of F abundances is  $\sim 0.4$  dex. This error value would account for this telluric effect if the line were observed, without substantially changing our results. Naturally, more observations in HE 1305+0007 will help to clarify the F abundance in this star. For the case of HE 1429–0551, the telluric contamination was sufficiently corrected. Despite the weakness of the line (3.5 per cent relative to the continuum), a reduced  $\chi^2$  value of 2.2 in the region of interest supports our F measurement in HE 1429–0551.

The study of F involves several sources of uncertainty. We introduce the general uncertainties in both the nucleosynthesis and observations as follows. The production of the  $^{19}\text{F}$  isotope in AGB stars takes place in the He intershell layer through complex nuclear reactions. Once formed, fluorine is dredged up to the surface during the third dredge up (TDU) along with C and *s*-process elements (Forestini et al. 1992; Mowlavi et al. 1996, 1998; Lugaro et al. 2004). As mentioned before, the principal reaction pathway for the produc-

tion of  $^{19}\text{F}$  in AGB stars is  $^{14}\text{N}(\alpha, \gamma)^{18}\text{F}(\beta^+)^{18}\text{O}(\text{p}, \alpha)^{15}\text{N}(\alpha, \gamma)^{19}\text{F}$  (Forestini et al. 1992). The protons needed in the reaction are supplied via  $^{14}\text{N}(\text{n}, \text{p})^{14}\text{C}$ , which in turn is provided with neutrons by the  $^{13}\text{C}(\alpha, \text{n})^{16}\text{O}$  reaction. This last neutron source,  $^{13}\text{C}(\alpha, \text{n})^{16}\text{O}$ , also provides neutrons for the *s*-process and its efficiency depends on the amount of  $^{13}\text{C}$  present in the He intershell. Some  $^{13}\text{C}$  existent in the intershell comes from CNO cycling in previous interpulse phases, although it is not enough to achieve the AGB *s*-process element enhancements. As an additional source of  $^{13}\text{C}$ , some protons can be partially mixed into the intershell producing  $^{13}\text{C}$  from the abundant  $^{12}\text{C}$  in the region through the reaction  $^{12}\text{C}(\text{p}, \gamma)^{13}\text{N}(\beta^+ \nu)^{13}\text{C}$  (Iben & Renzini 1982a). This reaction generates a thin layer abundant in  $^{13}\text{C}$  and  $^{14}\text{N}$  at the top of the intershell, namely the ‘ $^{13}\text{C}$  pocket’ (Iben & Renzini 1982b), which enrichment will depend on the amount of protons injected into the intershell, the partial mixing zone (PMZ). The production of F and its relation to the ‘ $^{13}\text{C}$  pocket’, suggest a positive correlation between abundances of F and *s*-process elements at the surface of the AGB star (see Section 4.3). However, some nuclear reaction rates are still not well known. For example, the  $^{19}\text{F}(\alpha, \text{p})^{22}\text{Ne}$  reaction is responsible for the destruction of  $^{19}\text{F}$ . Although this reaction is one of the most important parameters for a proper theoretical analysis of F nucleosynthesis in AGB stars, it is also one of the principal uncertainties in the nuclear reaction rates (Lugaro et al. 2004). Other sources of uncertainty in the nucleosynthesis of  $^{19}\text{F}$  include the shape of H profile of the partial mixed zone, convective mixing and mass-loss.

From an observational perspective, F abundances are difficult to measure. The formation of the HF molecule is very sensitive to  $T_{\text{eff}}$  in the atmosphere and is the main source of uncertainty in the adopted stellar parameters (Jorissen et al. 1992; Lucatello et al. 2011; Jönsson



**Figure 4.**  $A(\text{F})$  versus  $[\text{Fe}/\text{H}]$  for our data (orange symbols) and Lucatello et al. (2011) (green symbols). The Abate et al. (2015b) predictions are presented in magenta x symbols. For each star, the observed and theoretical abundances are connected by a black dashed line. Note the differences in the adopted  $[\text{Fe}/\text{H}]$  between Abate et al. (2015b) and Lucatello et al. (2011) (see text for details). HE 13005+0007 is in common between our sample and Lucatello et al. (2011), and our adopted  $[\text{Fe}/\text{H}]$  for this star also differs from their work. The fluorine abundance measured in HD 5223 by Lucatello et al. (2011) is well reproduced by the theoretical model. All the upper limits are substantially overestimated by Abate et al. (2015b), including HE 1305+0007.

et al. 2014a, b). The metallicity also affects the stellar structure in the outer layers of the star, changing its temperature profile and thus the HF absorption feature (Lucatello et al. 2011). Additionally, telluric corrections are necessary due to the high contamination in the infrared, generally blending the HF line. 3D and/or non-LTE (NLTE) effects can also play a role in the analysis of observed F abundances. Currently, 3D corrections for fluorine abundances have been explored by Li et al. (2013), but further investigation might be necessary. NLTE corrections are not available yet for the HF molecule, and the strength of its effect remains unknown. These observing issues make fluorine hard to detect. At low metallicities, previous to this work, 11 estimations of fluorine abundances have been reported in CEMP stars (Schuler et al. 2007; Lucatello et al. 2011): three measurements and eight upper limits. Carbon-normal halo field stars have been observed as well in a sample of seven stars, including two measurements and five upper limits (Li et al. 2013). With an observational success of  $\sim 28$  per cent measurements, the small observational data of F measurements in CEMP stars are not enough to be properly compared with F production in low-metallicity AGB nucleosynthesis models. However, the very presence of F (and enrichment), simultaneously with C and the characteristic  $s$ -process elements distribution in CEMP-s stars, indicate very strong evidence for AGB stellar nucleosynthesis.

#### 4.1 Comparison with literature data

Schuler et al. (2007) measured F for the first time in a CEMP-s star, HE 1305+0132. This was the most metal-poor star where an F detection has been reported with  $A(\text{F}) = +4.96$  ( $[\text{F}/\text{Fe}] =$

+2.90) at  $[\text{Fe}/\text{H}] = -2.5$ , the highest by over  $\sim 1$  dex from all other literature data at very low metallicities. Lugaro et al. (2008) showed that the observed C and F abundances by Schuler et al. (2007) in HE 1305+0132 can be justified by the binary mass transfer scenario of CEMP-s stars. Their model was able to reproduce the abundances by considering: (1) accretion of 3–11 per cent of the mass-loss from the companion star during the AGB phase and (2) dredged-up intershell material, no less than  $\sim 0.2 M_{\odot}$ , into the convective envelope by TDU. However, the analysis by Schuler et al. (2007) used older HF molecular transition data from Jorissen et al. (1992) with an excitation potential  $\chi = 0.48$  eV. This value is not compatible with the partition function in the MOOG program (Snedden 1973), leading to an offset in the reported F abundance of  $\sim +0.3$  dex (see e.g. Lucatello et al. 2011; Jönsson et al. 2014a). Moreover, the adopted metallicity for HE 1305+0132 ( $[\text{Fe}/\text{H}] = -2.5 \pm 0.5$ ) by Schuler et al. (2007) was derived by Beers et al. (2007) using mid-resolution optical and near-infrared spectra. Later, Schuler et al. (2008) reported a revised, and higher, metallicity of  $[\text{Fe}/\text{H}] = -1.92$  dex for HE 1305+0132 using high-resolution optical spectra, 0.6 dex higher than previously estimated. A more metal-rich model atmosphere contains outer layers cooler than the previous metal-poor model atmosphere used in Schuler et al. (2007). This metallicity difference can induce a lower F abundance in the measurement of HE 1305+0132 by Schuler et al. (2007), comparable to those  $[\text{F}/\text{Fe}]$  abundances found in other CEMP objects and the new values obtained in this paper. Additionally, our new F measurement in HE 1429–0551 makes this object the most metal-poor star where F has been detected. (Upper limits have been obtained at similar metallicity; Lucatello et al. 2011).

Lucatello et al. (2011) presented fluorine measurements and upper limits in 10 CEMP stars. Their sample included 8 CEMP-s and 2 CEMP-no stars. Fluorine was measured in two CEMP-s stars. For the rest of the sample, just upper limits are reported, including the two CEMP-no stars (see Fig. 3). HE 1305+0007 is in common between Lucatello et al. (2011) and our sample. In both studies, an upper limit was derived for this star. Our upper limit is 0.22 dex lower than the one obtained by Lucatello et al. (2011;  $A(\text{F}) < 3.5$ ,  $[\text{F}/\text{Fe}] < 1.4$ ). This discrepancy could be due to the differences in the adopted stellar parameters, and the different stellar atmosphere models used in the analysis of HE 1305+0007. We adopt the compiled stellar parameters from Abate et al. (2015b; see Section 2.1, and Table 1), while Lucatello et al. (2011) adopted stellar parameters from Beers et al. (2007):  $T_{\text{eff}} = 4560$ ,  $\log g = 1.0$ , and  $[\text{Fe}/\text{H}] = -2.5$ . A reanalysis of our F detection in HE 1305+0007 using the stellar parameters adopted by Lucatello et al. (2011) generates an upper limit of  $A(\text{F}) < 3.11$ ,  $[\text{F}/\text{Fe}] < 1.05$ . This differs from our original analysis by  $\Delta A(\text{F}) = +0.17$  and  $\Delta[\text{F}/\text{Fe}] = -0.05$  dex. This is principally due to the high sensitivity of the adopted absorption HF line to temperature and its moderate sensitivity to metallicity. The differences in the model atmosphere used in both analyses also have an effect in the F measurements. Lucatello et al. (2011) used spherical C-enhanced models which contribute to slightly higher F abundances for HE 1305+0007 in their study (see Section 3; Lucatello et al. 2011). Additionally, issues in the removal of telluric features near the HF line described before, could also affect our derived F upper limit in HE 1305+0007. However, our upper limit is comparable with that upper limit derived by Lucatello et al. (2011) in this star.

More recently, Li et al. (2013) derived F abundances and upper limits in seven carbon-normal metal-poor giant stars in the halo: two measurements and five upper limits (see Fig. 3). Their measurements correspond to HD110281 with  $A(\text{F}) = 2.71$  at  $[\text{Fe}/\text{H}] = -1.56$ , and HD135148 with  $A(\text{F}) = 2.59$  at  $[\text{Fe}/\text{H}] = -1.96$ . HD135148 has the highest C abundance in their sample and is a spectroscopic binary

with a companion star with the approximate mass of a white dwarf (e.g. Carney et al. 2003). Li et al. (2013) also applied 3D abundance corrections on the observed sample using 3D models from the CIFIST atmosphere grid (Ludwig et al. 2009). Since the grid of CIFIST do not have 3D models at low gravities as their star sample, four 3D models with comparable stellar parameters but higher gravity were selected to inspect the trend of the abundance corrections with changing  $T_{\text{eff}}$  and  $\log g$ . Li et al. (2013) suggested minor 3D effects for the low-gravity giants in their sample and substantial corrections ( $> 0.1$  dex) for gravities  $\log g > 2.0$ , at  $[\text{Fe}/\text{H}] = -2.0$ .

#### 4.2 Observations versus theoretical models

The F abundance in CEMP-s stars is assumed to come from an AGB companion (now an unseen white dwarf) which also synthesized C and *s*-process elements (Lugaro et al. 2004; Ivans et al. 2005; Lucatello et al. 2005; Karakas & Lattanzio 2007; Lugaro et al. 2008; Thompson et al. 2008). Within this picture, theoretical models have to consider not just AGB stellar nucleosynthesis at low metallicities, but also the binary evolution. Abate et al. (2015a, b) have undertaken an appropriate theoretical study (see Section 2.1). The initial abundances used in Abate et al. (2015a, b), including fluorine, were adopted from the Galactic Chemical Evolution (CGE) model by Kobayashi, Karakas & Umeda (2011a). In this CGE model, the contribution from massive stars is set to minimum and does not include that from WR stars. Although Kobayashi et al. (2011b) have suggested that the F production through neutrino process can be increased by a factor of 100, this is negligible once mass transfer occurs from AGB stars in the initial mass range  $[0.9, 2.5] M_{\odot}$ .

Stars HE 1429–0551 and HE 1305+0007 were originally included in the Abate et al. (2015b) sample and we can therefore compare our newly determined fluorine abundances with the predictions.

##### 4.2.1 HE 1429–0551.

Our observed F abundance,  $A(\text{F}) = +3.93$ , is slightly overpredicted but in reasonable agreement with the theoretical predictions from Abate et al. (2015b),  $A(\text{F})_{\text{Abate}} = +4.40$ . In addition to our F measurement, observed abundances of C and neutron-capture elements by Aoki et al. (2007) are well reproduced by the model for HE 1429–0551. Therefore, the observed chemical pattern of HE 1429–0551 supports the formation scenario on a binary system through mass transfer from a former AGB companion, even though its binary status has not been observationally confirmed as noted in Section 2.1. The best-fitting model from Abate et al. (2015b) for HE 1429–0551 correspond to a very long-period binary with an initial period  $P_i = 1.37 \times 10^5$  d. The initial masses of the primary star (the CEMP progenitor) and secondary star (the CEMP star) are  $M_{1,i} = 1.30 M_{\odot}$  and  $M_{2,i} = 0.71 M_{\odot}$ , respectively. The adopted PMZ mass of the primary star is  $M_{\text{PMZ}} = 10^{-3} M_{\odot}$  and the calculated accreted mass by the secondary stars is  $M_{\text{acc}} = 0.17 M_{\odot}$ . The current estimated period of the system is  $P_f = 2.3 \times 10^5$  d.

##### 4.2.2 HE 1305+0007.

Our F abundance upper limit,  $A(\text{F}) < +3.28$ , does not agree with the theoretical prediction from Abate et al. (2015b),  $A(\text{F})_{\text{Abate}} = +4.52$  dex. This is 1.24 dex higher than our estimated upper limit in this star (see Figs 2 and 4). Ours and the Lucatello et al. (2011) upper limit in this star, both disagree with the predicted abundances.

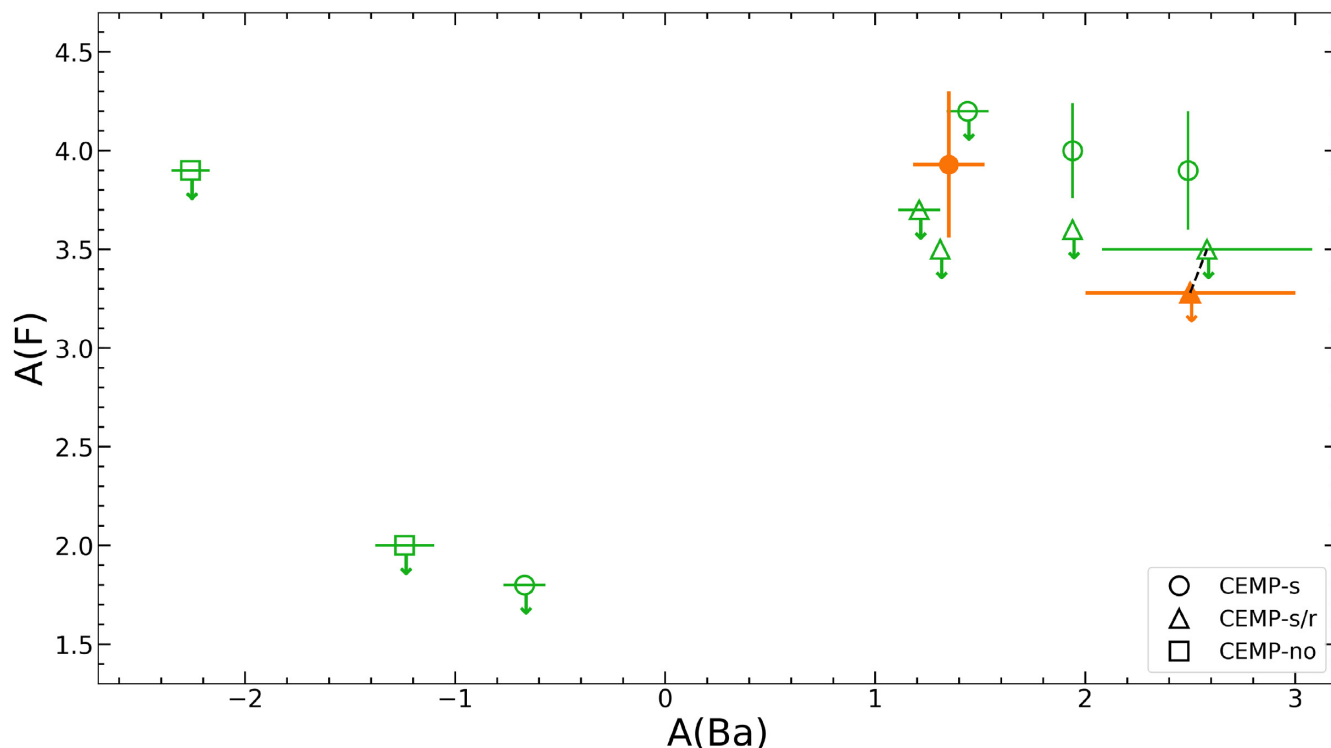
The disagreement is very likely due to the fact that HE 1305+0007 is classed as a CEMP-s/r star, which is a separate class of CEMP stars whose origins may be completely different from CEMP-s (the origin of CEMP r/s from an AGB nucleosynthesis point of view was also discussed by Lugaro et al. 2012). Our observation, as well as literature data from Lucatello et al. (2011), suggest a lower fluorine production by the HE 1305+0007 progenitor that cannot be reproduced by the nucleosynthesis model adopted by Abate et al. (2015b), while simultaneously reproducing the very enhanced abundances of the neutron-capture elements. As discussed by Abate et al. (2015b), in most CEMP-s/r stars the abundances of heavy neutron-capture elements (e.g. Ba, La, Ce, Pb, but also Nd, Sm, and Eu) are strongly enriched and to reproduce their enhancements, the algorithm determining the best-fitting model selects a solution where a large mass of very enriched material is transferred from the AGB primary star to the main-sequence secondary star of the binary system. Consequently, the best-fitting model tends to overestimate the abundances of the light elements (e.g. C, N, F, Na, Mg), which are typically enhanced to a similar extent as ‘canonical’ CEMP-s stars. This result was interpreted as an indication that CEMP-s/r star might have undergone a different nucleosynthesis history compared to CEMP-s stars, the abundances of which are generally well-reproduced by AGB nucleosynthesis models (Lugaro et al. 2012; Abate et al. 2015b).

Hampel et al. (2016, 2019) used a one-zone nucleosynthesis model with neutron captures operating at neutron densities intermediate to the *s*- and *r*-process, the *i*-process. Their adopted temperatures are consistent with the He-shell temperatures from a previous work from Stancliffe et al. (2011). The modelled abundances by Hampel et al. (2016, 2019) can satisfactorily reproduce the observed abundances of C and *s*-process elements, as well as the *r*-process elements present in HE 1305+0007. Currently, F abundances from *i*-process models are not available in the literature that we are aware of. It would be very interesting to examine and test if those models can also reproduce the observed abundances of fluorine as well as the *s/r* (or *i*) chemical pattern.

Five stars in the sample from Lucatello et al. (2011) are in common with the theoretical work of Abate et al. (2015b), including HE 1305+0007 which is also in common with this work. Fig. 4 shows the comparison between the predictions from Abate et al. (2015b) with the observations from Lucatello et al. (2011) and ours. There are some differences in the stellar parameters adopted by Lucatello et al. (2011) and Abate et al. (2015b) (i.e.  $T_{\text{eff}}$  and  $[\text{Fe}/\text{H}]$ ). The differences in the adopted  $[\text{Fe}/\text{H}]$  and the predicted-observed abundances for each star are shown in Fig. 4 using a connecting black dashed line.

The fluorine abundances observed by Lucatello et al. (2011) include CEMP-s and CEMP-s/r. Their measurement in the CEMP-s HD 5223 is in good agreement with the model. However, all their upper limits are highly overestimated by the model predictions from Abate et al. (2015b; see Fig. 4). The inconsistencies with the upper limits of the CEMP-s/r stars CS 22948–027, HD 187861, and HE 1305+0007 can be accounted for by the issues of the model on reproducing the *s/r* pattern in addition to the differences in the adopted stellar parameters mentioned before. Nevertheless, our upper limit in HE 1305+0007 was derived using the same stellar parameters as in Abate et al. (2015b), and we found similar discrepancy of 1.24 dex. Additionally, the upper limit in the CEMP-s star CS 22942–019 estimated by Lucatello et al. (2011) is highly overestimated by the model. The inconsistency with the model in this star cannot be justified by just the differences in the adopted stellar parameters.





**Figure 5.**  $A(\text{F})$  and  $A(\text{Ba})$  abundances for CEMP-s, CEMP-s/r, and CEMP-no objects from Lucatello et al. (2011) and our sample. Symbols are the same as Fig. 3 (carbon-normal field halo stars are excluded). The black dashed line shows the differences in the adopted  $A(\text{Ba})$  and the estimate upper limit in ours in common observed star with Lucatello et al. (2011), HE 1305+0007.

### 4.3 Ba and F abundances in CEMP stars

In CEMP-s stars, abundances from  $s$ -process elements are expected to correlate with those of F given that both are produced in AGB stars (Lugaro et al. 2004), the CEMP-s progenitors. However, this might not always be the case since the abundance of Ba in the envelope reaches an equilibrium value when the  $s$ -process operates at low metallicity (Lugaro et al. 2012). In contrast, F does not have an equilibrium value and its nucleosynthesis is more sensitive to the thermodynamic conditions in the intershell of the AGB star (Mowlavi et al. 1996, 1998; Lugaro et al. 2004). Progenitors of CEMP-s/r stars have to account for nucleosynthesis mechanisms able to enhance both  $s$ - and  $r$ -process elements, as well as C. Although the  $i$ -process can explain C and neutron-capture elements in the CEMP-s/r chemical pattern (e.g. Hampel et al. 2016, 2019), the effect of this process on the F production is unknown. The origin of the chemical pattern observed in CEMP-no objects still remains unclear. A variety of possible formation scenarios are proposed in literature (e.g. Umeda & Nomoto 2005; Masseron et al. 2010; Meynet et al. 2010; Chiappini 2013; Nomoto et al. 2013; Limongi & Chieffi 2018). Measured abundances of fluorine (or strong upper limits) in CEMP-no stars would give useful information on the chemical origins of CEMP-no stars.

We show in Fig. 5 the absolute abundances of Ba and F from Lucatello et al. (2011) and our sample, in green and orange, respectively. These two sample are the only data available for F and Ba in CEMP stars. Note that Ba abundances cannot be derived from the infrared  $K$  band due to the lack of features in this wavelength region. Hence, the Ba abundances shown in Fig. 5 are adopted from the literature. For our sample, we adopt the Ba abundances from the compilation by Abate et al. (2015b) – i.e. Goswami et al.

(2006) and Beers et al. (2007) for HE 1305+0007, and Aoki et al. (2007) for HE 1429–0551. The details for Lucatello’s adopted Ba abundances can be found in Lucatello et al. (2011, table 2). In addition to the inhomogeneities introduced by the Ba abundances in Fig. 5, the estimated fluorine abundances are mostly upper limits. These limitations make it difficult to quantify any trend if present. However, the CEMP-s upper limit at  $A(\text{Ba}) \sim -0.6$  and the CEMP-s measurements at higher Ba abundances suggest a possible Ba-F correlation in CEMP-s stars. We also note that the three CEMP-s stars with measured F abundances at  $A(\text{Ba}) > 1$  are higher than those upper limits from CEMP-s/r stars. This suggests that while the Ba ( $s$ -process elements proxy) production in CEMP-s and CEMP-s/r progenitors is indistinguishable, the F production might be more efficient in CEMP-s than in CEMP-s/r progenitors. Unfortunately, due to the lack of measurements and the large spread in the only two F upper limits in CEMP-no stars, no valuable information can be extracted from its F estimates.

Very recently, Ryde et al. (2020) compared the current literature data on F abundances at higher metallicities, in addition to their own sample, with O and the  $s$ -process element Ce. They showed a flat trend when studying the  $[\text{F}/\text{Ce}]$  ratio as a function of  $[\text{Fe}/\text{H}]$  for intermediate metallicities,  $-0.6 < [\text{Fe}/\text{H}] < 0.0$ . This result strengthens their conclusion on AGB stars as the dominant source of F in that metallicity range. However, their sample includes two stars below  $[\text{Fe}/\text{H}] < -0.8$  exhibiting large  $s$ -process element (Zr, La, and Ce) enhancements yet surprisingly low F abundances. (see section 5.6 and fig. 10 from Ryde et al. 2020). An inspection of radial velocities in these two stars by Ryde et al. (2020) did not show any signs of binarity, weakening the possibility of pollution from an AGB companion. A similar result was found by Cunha et al. (2003) in  $\omega$  Cen stars at similar metallicities to those from Ryde et al. (2020).

This raises multiple questions on how these objects obtained their chemical signature and what is its astronomical origin. More data are needed to observationally quantify any relation between *s*-process and F nucleosynthesis in the early Galaxy.

## 5 CONCLUSIONS

We present high-resolution fluorine abundances of two CEMP stars with  $[\text{Fe}/\text{H}] < -2$ : HE 1429–0551 and HE 1305+0007. Our new F measurement in HE 1429–0551 makes this object the lowest metallicity star where F has been detected. Despite our small sample (current literature data are also very small: three measurements, and eight upper limits; see Section 4.1, and Fig. 3), our results provide valuable information on F production in AGB stars at low metallicity and its nucleosynthesis process. We also compare our results to theoretical predictions from AGB nucleosynthesis and binary evolution model by Abate et al. (2015b) models as follow.

Our fluorine measurement in HE 1429–0551,  $A(\text{F}) = +3.93$ , is in reasonable agreement compared to the predicted abundance from Abate et al. (2015b),  $A(\text{F})_{\text{Abate}} = +4.40$  (see Figs 2 and 4). The best-fitting model of Abate et al. (2015b) also reproduces the observed abundances of C and *s*-process elements determined by Aoki et al. (2007). As noted in Section 2.1, the binary status of this object is inconclusive (Jorissen et al. 2016). Without a constraint on the orbital period, Abate et al. (2015b) found that the observed abundances were best reproduced with a very long-period ( $+10\,000$  d) binary mode with an initial period of  $P_i = 1.37 \times 10^5$  d, initial masses  $M_{1,i} = 1.30 M_{\odot}$  and  $M_{2,i} = 0.71 M_{\odot}$ . The adopted PMZ mass of the progenitor and the calculated accreted mass by HE 1429–0551 that account for the observed neutron-capture elements and F abundance is  $M_{\text{PMZ}} = 10^{-3} M_{\odot}$  and  $M_{\text{acc}} = 0.17 M_{\odot}$ , respectively. This result strongly suggests that HE 1429–0551 is indeed a binary system which has not been monitored for a sufficiently long time to constrain its orbit.

Our estimated upper limit of F abundance in HE 1305+0007,  $A(\text{F}) < +3.28$ , is significantly lower than the predicted value from Abate et al. (2015b),  $A(\text{F})_{\text{Abate}} = +4.52$  (see Figs 2 and 4). Lucatello et al. (2011) also found an F upper limit in HE 1305+0007 of  $A(\text{F})_{\text{Lucatello}} < +3.5$ , comparable to ours if we consider the differences in the analysis (see Section 4.1).

All the observed upper limits in CEMP-*s/r* stars from Lucatello et al. (2011) and this work show low F abundances compared to CEMP-*s* (see Figs 4 and 5). These observational results suggest that the CEMP-*s/r* progenitors might provide lower F abundances. The discrepancies between the observed upper limits and the predicted fluorine abundances by Abate et al. (2015b) can be due to the CEMP-*s/r* category type, and the theoretical approach using just *s*-process nucleosynthesis. Hampel et al. (2016, 2019) used *i*-process nucleosynthesis models getting a better agreement with the chemical pattern observed in CEMP-*s/r* stars, including HE 1305+0007. Although the details on how the operation of the *i*-process affects the production of F has not yet been fully explored, F abundances can be the key to understand the astronomical site of *i*-process.

We also show in Fig. 5, the  $A(\text{F})$  abundances from Lucatello et al. (2011) and this work as a function of  $A(\text{Ba})$  abundances from literature data. Although caution has to be taken in the interpretation of Fig. 5, the data may suggest a positive correlation on Ba and F abundances in CEMP-*s* stars. Additionally, we note that the only three F measurements from CEMP-*s* objects show higher F content than those in CEMP-*s/r*. This could indicate differences in the relation between *s*-process and F nucleosynthesis in CEMP-*s* and CEMP-*s/r* progenitors.

The fluorine observations presented in this paper, alongside literature data, suggest further tests and constraints on nucleosynthesis models for CEMP star progenitors. Moreover, the relation between the production of fluorine and *s*-process elements in metal-poor AGB stars is still uncertain, and thus the true nature of its production mechanism remain unclear at these metallicities. Despite the difficulty of the analysis, it is very important to extend the study to additional CEMP stars, in order to gain new insights upon nucleosynthesis and chemical enrichment in the early Galaxy.

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## DATA AVAILABILITY

Data available on request. The data underlying this article will be shared on reasonable request to the corresponding author.

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