

Constraining Ion-Scale Heating and Spectral Energy Transfer in Observations of Plasma Turbulence

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We perform a statistical study of the turbulent power spectrum at inertial and kinetic scales observed during the first perihelion encounter of the *Parker Solar Probe*. We find that often there is an extremely steep scaling range of the power spectrum just above the ion-kinetic scales, similar to prior observations at 1 A.U., with a power-law index of around -4 . Based on our measurements, we demonstrate that either a significant ($>50\%$) fraction of the total turbulent energy flux is dissipated in this range of scales, or the characteristic nonlinear interaction time of the turbulence decreases dramatically from the expectation based solely on the dispersive nature of nonlinearly interacting kinetic Alfvén waves.

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Introduction.—In many astrophysical settings, the background plasma is highly turbulent and nearly collisionless. The dissipation of collisionless turbulence is important for plasma heating [1–5], but the precise mechanisms involved are a matter of debate [6]. The solar wind provides a convenient example of a collisionless plasma in which both large magnetohydrodynamic scales (MHD) and sub-ion-kinetic scales can be studied through *in situ* spacecraft measurements.

The observed ion temperature profiles in the solar wind require significant perpendicular ion heating [7], which is likely initiated at ion-kinetic scales, where particles interact efficiently with electromagnetic waves [1,8–14]. Such heating should transfer energy from the waves to the particles, generating observable electromagnetic signatures. The observation of spectral steepening in the “transition range” between MHD and ion-kinetic scales is most certainly a result of these processes [15–19].

Large-scale fluctuations appear to be nonlinearly interacting Alfvénic turbulence [20], with a power-law spectrum between $k_{\perp}^{-5/3}$ [21] and $k_{\perp}^{-3/2}$ [22,23], in rough agreement with various MHD turbulence theories [24–27]. At scales much smaller than the ion gyroradius $\rho_i = v_{\text{thi}}/\Omega_i$ (with $v_{\text{thi}} = \sqrt{2T_{0i}/m_i}$ the ion thermal speed and $\Omega_i = ZeB_0/m_i$ the ion gyrofrequency), the spectrum steepens to about $k_{\perp}^{-2.8}$ [28–31], as nondispersive Alfvénic turbulence transitions to dispersive kinetic Alfvénic turbulence [32]. This steepening occurs solely due to the dispersion relation, without any dissipation: adiabatic fluid approximation to the dynamics in this range of scales predicts a $k_{\perp}^{-7/3}$ spectrum [13,33]. Simulations in the fluid approximation obtain the slightly steeper $k_{\perp}^{-8/3}$ spectrum [34], with the difference ascribed to intermittency. Fully kinetic models recover similar scalings [35–37].

At ion scales, i.e., the transition from MHD to kinetic regimes with $k_{\perp}\rho_i \sim 1$, a spectrum significantly steeper

than the sub-ion $k_{\perp}^{-2.8}$ scaling has been reported [9,15,17,19,38–40]. Instrumental limitations initially led to interpretation of ion-scale steepening as variability in the kinetic range (i.e., $k\rho_i > 1$) spectrum, rather than a signature of processes localized to ion scales [9,38,39]. Measurements of fluctuations deep in the sub-ion scales repeatedly demonstrate a $k_{\perp}^{-2.8}$ spectra, indicating that transition range steepening disappears at scales $k_{\perp}\rho > 1$ [28–31]. Observational signatures of the transition range are intermittent, nonregular, and variable and cannot be explained by the kinetic Alfvén wave (KAW) dispersion relation [17,18,38,39,41]. Two proposed explanations for transition-range steepening are, first, strong dissipation of the turbulence around the ion scales [15–18,42]. Second, nonlinear effects that increase turbulent cascade rates: e.g., the onset of reconnection [43–45], the increasing importance of nonlinear interactions between copropagating waves [46], or the influence of coherent structures [40,47,48].

In this Letter, we report *Parker Solar Probe* (PSP) observations of strong and regular ion-scale steepening in the inner heliosphere, indicating that a dispersive transition from nonlinear Alfvén turbulence at MHD scales to kinetic Alfvén turbulence at sub-ion scales is an incomplete description of solar wind turbulence. Increased levels of ion-scale steepening in PSP observations enable a comprehensive statistical study of the phenomenon. Using continuity of turbulent energy flux, assuming local wavenumber spectral transfer [12,13,49,50], we demonstrate that the observed steepening has drastic implications for turbulent heating and/or nonlinear cascade rates. Statistical analysis shows that transition-range steepening is consistent with levels of dissipation sufficient for required bulk *in situ* heating of the solar wind [14,51–56]. As our model cannot identify heating as the unique cause of the steepening, we consider an alternative framework, in which steepening is the result of increased spectral transfer rates. Though our analysis is largely inconsistent with increased spectral transfer via imbalanced turbulence or magnetic reconnection [43,44,46], we cannot explicitly exclude other nonlinear dynamics that may anomalously accelerate turbulent energy transfer at ion scales.

Cascade model.—We use a Batchelor cascade model [13,49], which relates turbulent energy flux $\epsilon_{k_{\perp}}$ through wave-number k_{\perp} to the spectrum $E_{k_{\perp}} = b_{k_{\perp}}^2/k_{\perp}$, where $b_{k_{\perp}}$ is the turbulent amplitude, through

$$\epsilon_{k_{\perp}} = \omega_c k_{\perp} E_{k_{\perp}}. \quad (1)$$

The quantity

$$\omega_c = k_{\perp} b_{k_{\perp}} \tilde{\omega}_{k_{\perp}} = k_{\perp}^{3/2} E_{k_{\perp}}^{1/2} \tilde{\omega}_{k_{\perp}} \quad (2)$$

is the characteristic spectral transfer rate (inverse cascade time) at k_{\perp} , and $\tilde{\omega}_{k_{\perp}}$ parametrizes both dispersive (e.g., from

the dispersive KAW [33]) and/or nonlinear effects (e.g., caused by dynamic alignment [25], intermittency [34], or reconnection [43]). All unknown normalizing constants are contained in $\tilde{\omega}_{k_{\perp}}$. In statistical steady state far from the injection scales [13,57] one obtains a simple equation with solution

$$1 - \hat{Q}_{k_1, k_2} = \frac{\epsilon_{k_2}}{\epsilon_{k_1}} = \exp \left\{ - \int_{k_1}^{k_2} \frac{\gamma_{k_{\perp}} dk_{\perp}}{\omega_c k_{\perp}} \right\}, \quad (3)$$

where $\gamma_{k_{\perp}}$ is the energy dissipation rate at k_{\perp} , and \hat{Q}_{k_1, k_2} is the fractional heating rate over the range $[k_1, k_2]$. If $\gamma_{k_{\perp}} = 0$ over $[k_1, k_2]$, then $\epsilon_{k_{\perp}}$ is constant over $[k_1, k_2]$. In this Letter, we do not assume a physical mechanism for dissipation, rather we use empirical observations of spectral steepening to constrain required levels of dissipation. However, the functional form of $\gamma_{k_{\perp}}$ has significant impact on the form of the energy spectrum [12,13,17,31,58,59]. Using Eqs. (1) and (2),

$$\frac{E_{k_2}}{E_{k_1}} = \left(\frac{k_2}{k_1} \right)^{-5/3} \left(\frac{\epsilon_{k_2}}{\epsilon_{k_1}} \right)^{2/3} \left(\frac{\tilde{\omega}_{k_1}}{\tilde{\omega}_{k_2}} \right)^{2/3}. \quad (4)$$

Equation (4) relates *in situ* measurements of E_k to the turbulent heating rate (second bracket) and/or anomalous dispersive and nonlinear effects (third bracket). Surveying transition-range steepening in E_k illustrates the dramatic effect that it takes on the turbulent cascade.

Data and fitting.—The PSP mission [60] provides a set of *in situ* measurements [61,62] enabling detailed studies of turbulence in the inner heliosphere [23, 63–65]. Preliminary observations show a steep f^{-3} – f^{-4} spectrum of the magnetic field fluctuations at ion scales [66–68], but instrumental sensitivity has precluded further detailed study.

We use merged fluxgate and search coil measurements (SCM) from PSP FIELDS [61,69] operating at 293 Hz, enabling measurement of inertial, transition, and kinetic scales. We study the first PSP perihelion encounter, using the interval with maximal data rates: November 4, 2018, 09:28:19–November 7, 2018, 09:28:19. During this period, PSP was magnetically connected to a small equatorial coronal hole generating slow, but highly Alfvénic solar wind [66,70].

We partition the interval into an ensemble of 2^{16} samples (~ 223.69 s), with 50% overlap to improve statistics. Intervals were rejected for lack of finite measurements, or if the SCM was in low gain mode. Intervals with ion-scale waves, which strongly affect measurements of ion-scale turbulence, are excluded [40,66,71]. In total, 243 intervals were kept. We consider frequencies up to 100 Hz ($k_{\perp}\rho_i \sim 10$), to avoid the SCM noise floor. Average plasma properties are computed for each interval using solar wind electrons alphas and protons (SWEAP) investigation data [62]; n_i and T_i from the SWEAP-solar probe cup (SPC) [72] and n_e , T_e from SWEAP-solar probe analyzers

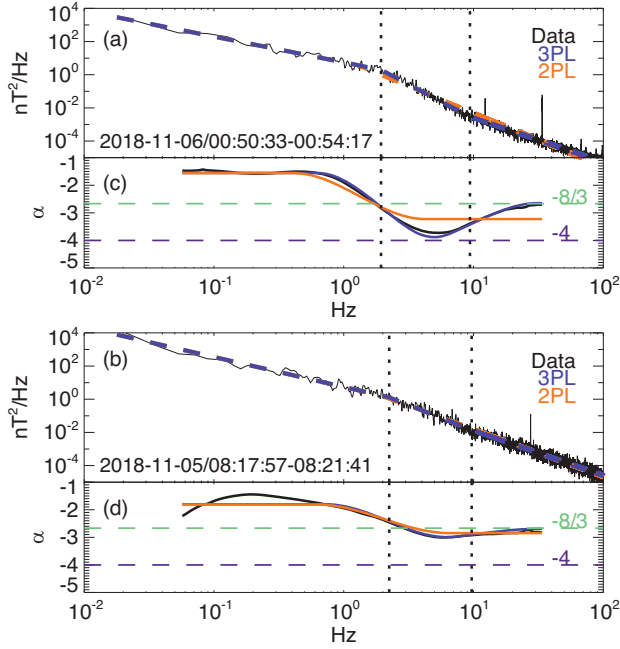


FIG. 1. (a,b) Examples of PSP/FIELDS magnetic field spectra with 3PL (blue) and 2PL fits (orange). Vertical lines show 3PL spectral breaks. (c,d) spectral indices for data (black), 3PL (blue) and 2PL fits (orange). Horizontal lines are shown corresponding to spectral indices of $-8/3$ (teal) and -4 (purple). Top interval has statistically significant spectral steepening, while the bottom interval is sufficiently fit by 2PL.

(SPAN) electron fits [73,74]. On average, β_i (β_e) is 0.26 (0.74) with a standard deviation of 0.13 (0.25), with $\beta_i = v_{\text{th}i}^2/v_A^2$ and the Alfvén speed is $v_A = B_0/\sqrt{n_{0i}m_i\mu_0}$. The average T_{0i}/T_{0e} is 0.5. The mean normalized cross helicity $\sigma_c \sim 0.75$.

KAWs have intrinsic density fluctuations [33,75–78], which at $k_{\perp}\rho_i \gtrsim 1$ provide a non-negligible contribution to the total energy. We estimate this contribution by determining δn_e from the pressure balance $\delta B_{\parallel}/B_0 = -(\beta_i/2)(1 + ZT_{0e}/T_{0i})\delta n_e/n_{0e}$, appropriate for KAWs, and estimate the total energy as

$$E_{\text{tot}} = \frac{|\delta \mathbf{B}|^2}{2\mu_0} + \frac{n_{0e}T_{0e}}{2} \left(\frac{\delta n_e}{n_0} \right)^2. \quad (5)$$

Figure 1(a) shows an example with transition-range steepening to an $\sim f^{-4}$ spectrum at ion scales, similar to observations at 1 A.U. [17,19]. At the highest frequencies, the measured spectral index is consistent with kinetic Alfvén wave scaling of $E \propto f^{-8/3}$. Transition-range steepening is not obvious in a second example interval, Fig. 1(b), but approximate $f^{-8/3}$ scaling is evident.

Continuous power-law (PL) functions with two and three spectral ranges (2PL, 3PL) are fit to each interval, optimizing χ^2 residuals. Figures 1(c) and 1(d) show the variation of spectral index over scales for the data and two models.

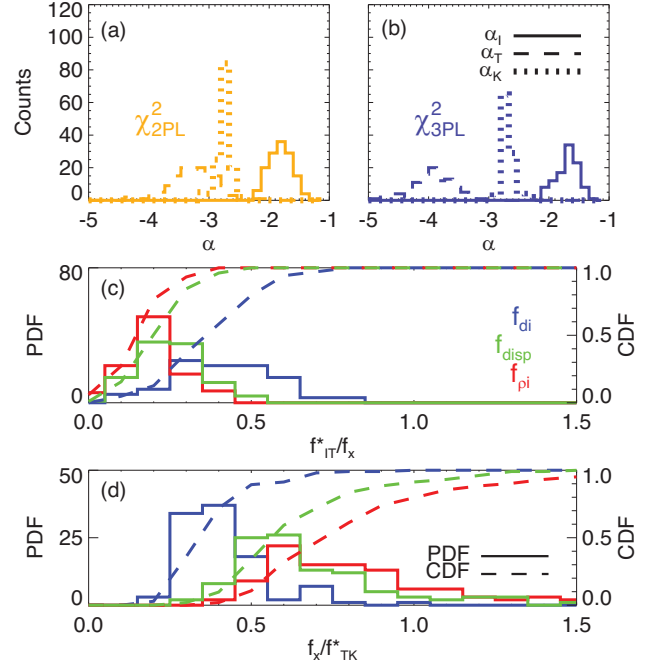


FIG. 2. (a) Probability distribution function (PDF) of fitted 3PL spectral indices for the $\chi^2_{2\text{PL}}$ (orange) population for inertial α_I (solid), transition α_T (dashed), and kinetic α_K (dotted) ranges. (b) PDF of fitted 3PL spectral indices for the $\chi^2_{3\text{PL}}$ (blue) population. (c) PDF and cumulative distribution function (CDF) of f_{IT}^* normalized to spacecraft frequencies $f_x = (f_{\rho i}, f_{di}, f_{\text{disp}})$ (red, blue, green). (d) PDF and CDF of f_{TK}^* normalized to ion scales (f_x/f_{TK}^*). Normalizations in (c) and (d) show $f_{IT}^*/f_x < 1$ and $f_{TK}^*/f_x < 1$, indicating ion scales are within the transition range.

The 3PL fit determines spectral indices of the inertial (α_I), transition, (α_T), and kinetic (α_K) ranges and the break points (f_{IT}^* and f_{TK}^*). The significance of including a third spectral range, which introduces additional degrees of freedom (f_{IT}^* and α_T), is determined by the statistical f test. Though not shown, a continuum of significance values is measured, indicating that, at times, spectral steepening is strong (i.e., the 3PL fit significantly improves χ^2), whereas some intervals have moderate or no statistically significant steepening (i.e., χ^2 is similar for both 3PL and 2PL fits). As a continuous distribution of spectral steepening is measured, intervals are separated into two equal sized populations, approximately distinguishing spectra well fit by 3PL from those sufficiently described by 2PL. The notation $\chi^2_{3\text{PL}}$ and $\chi^2_{2\text{PL}}$ refers to the populations.

Figures 2(b) and 2(c) show the histogram of measured spectral indices for 3PL fits for α_I , α_T , and α_K . Figures 2(b) and 2(c) show distributions for the $\chi^2_{2\text{PL}}$ and $\chi^2_{3\text{PL}}$ populations. Mean spectral indices for each range and population are in Table I. The total population has a mean transition-range index of $\langle \alpha_T \rangle = -3.5$, while intervals with most significant steepening have $\langle \alpha_T^{\chi^2_{3\text{PL}}} \rangle = -3.9$.

TABLE I. Measured mean spectral indices from 3PL and 2PL fits to both $\chi_{2\text{PL}}^2$ and $\chi_{3\text{PL}}^2$ populations.

Population	2PL fit		3PL fit		
	$\langle\alpha_I\rangle$	$\langle\alpha_K\rangle$	$\langle\alpha_I\rangle$	$\langle\alpha_T\rangle$	$\langle\alpha_K\rangle$
$\chi_{2\text{PL}}^2$	-1.7	-2.9	-1.7	-3.2	-2.7
$\chi_{3\text{PL}}^2$	-1.6	-3.1	-1.6	-3.9	-2.6
$\chi_{3\text{PL}}^2 \cup \chi_{2\text{PL}}^2$	-1.6	-3.0	-1.7	-3.5	-2.7

The spacecraft frequencies f_{ρ_i} , f_{d_i} , and $f_{\rho_{\text{disp}}}$ associated with the ion gyroscale ρ_i , inertial scale $d_i = \rho_i/\sqrt{\beta_i}$, and the scale at which dispersion sets in $\rho_{\text{disp}} = \rho_i \sqrt{(1+T_{0e}/T_{0i})}/2$, are computed using the Taylor hypothesis $2\pi f = v_{\text{sw}}k$. Figures 2(d) and 2(e) show distributions of the break scales of the $\chi_{3\text{PL}}^2$ population normalized to the frequencies f_{ρ_i} , f_{d_i} , and $f_{\rho_{\text{disp}}}$. Figure 2(d) shows that the break from the inertial to transition range occurs at frequencies significantly lower than ion scales.

Physical interpretation: Dissipation.—Equation (4) implies that a steep transition range is associated with either significant dissipation or nonlinear speedup of the cascade. First, we assume that the steep spectrum is due to dissipation [8,15,57,59,79–81]; i.e., $\epsilon_{k_{\perp}}$ decreases with k_{\perp} across the transition range. Using the 3PL fits, we construct a synthetic spectrum

$$E'(f) = \begin{cases} c_I f^{\alpha_I} & \text{if } f < f_{IT}^* \\ c_T f^{\alpha_K} & \text{if } f > f_{IT}^* \end{cases} \quad (6)$$

This joins the fitted inertial-range spectrum to a synthetic spectrum with the fitted kinetic-range exponent α_k at the

inertial-transition break f_{IT}^* . An example of this synthetic spectrum is shown in Fig. 3(a). We use E' to determine $\tilde{\omega}'_{k_{\perp}}$ [cf. Eqs. (1) and (2)], assuming that this synthetic spectrum is the result of constant $\epsilon'_{k_{\perp}}$. Using this synthetic $\tilde{\omega}'_{k_{\perp}}$ and the fitted 3PL spectrum $E^{3\text{PL}}$ in Eq. (4) results in an estimate of the heating in the transition range relative to the synthetic spectrum,

$$1 - \hat{Q}^* = \frac{\epsilon_{TK}}{\epsilon_{IT}} = \left(\frac{E^{3\text{PL}}(f_{TK}^*)}{E'(f_{TK}^*)} \right)^{3/2}. \quad (7)$$

Figure 3(b) shows measured ratios of $\epsilon_{TK}/\epsilon_{IT}$ as a function of transition-range spectral index α_T . Transition-range spectral indices of $\alpha_T \approx -3.5$ correspond to heating at $\hat{Q}^* > 50\%$ of the turbulent energy flux and thus may be a signature of significant ion-scale heating.

Physical interpretation: Nonlinear effects.—The steep transition-range spectrum is alternatively explained by nonlinear effects that increase nonlinear spectral transfer rates. Assuming that $\epsilon_{k_{\perp}}$ is constant, we use Eq. (4) to determine the scaling of $\tilde{\omega}_{k_{\perp}}$, and therefore ω_c , that matches each interval's measured spectrum. Figure 3(c) shows the scaling of ω_c corresponding to the interval in Fig. 1(a). In the inertial range, the wave-number power-law scalings, $\omega_c \propto k_{\perp}^{\alpha_{\omega}}$, are similar to those predicted in MHD turbulence models: between $\omega_c \propto k_{\perp}^{2/3}$ [24] and $\omega_c \propto k_{\perp}^{1/2}$ [25]. In the kinetic range, the scaling is similar to predictions of the KAW turbulence models: between $\omega_c \propto k_{\perp}^{4/3}$ [33] and $\omega_c \propto k_{\perp}^{5/3}$ [34]. The scaling of ω_c is significantly steeper in the ion-scale transition.

The fitted $E^{3\text{PL}}$ and synthetic E' spectra allow an estimate of the increase in nonlinear interactions due to transition-range steepening, using Eq. (1) and taking $\epsilon_{k_{\perp}}$ constant.

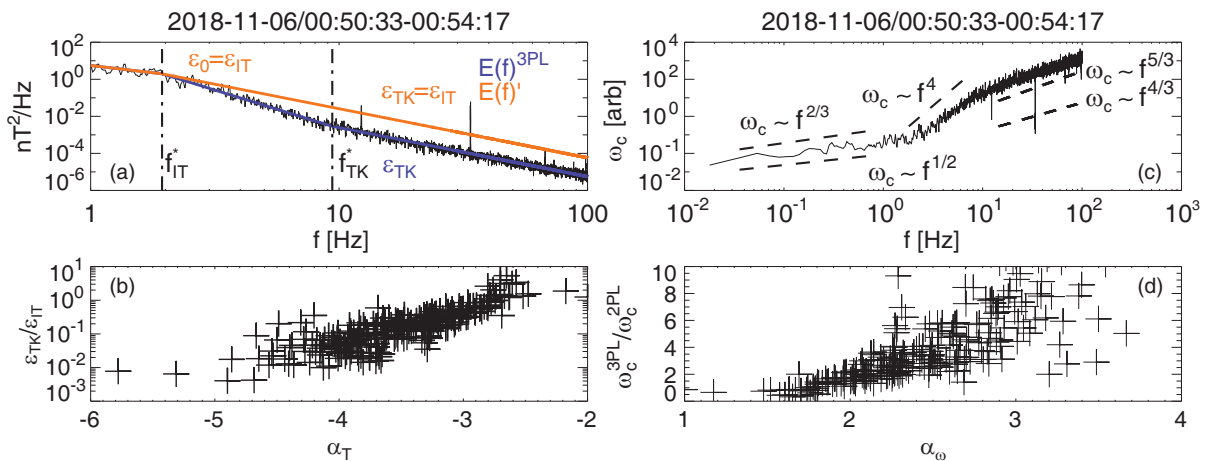


FIG. 3. (a) Measured power spectrum with transition-range steepening from Fig. 1(a) (black), with the corresponding fitted (blue) and synthetic (orange) spectra. Fit break scales are plotted (vertical lines). (b) Measured ratios of energy flux at transition-kinetic break relative to inertial-kinetic break scale $\epsilon_{TK}/\epsilon_{IT}$, plotted against transition-range spectral index α_T . (c) Nonlinear frequency ω_c as function of f , assuming $\epsilon_{k_{\perp}}$ at all frequencies. Various power-law scalings of $\omega_c \sim f^{\alpha_{\omega}}$ are shown. (d) Increase of ω_c over the transition range relative to the synthetic spectrum as a function of α_{ω} for all intervals.

Figure 3(d) shows the ratio of $\omega_c^{3\text{PL}}/\omega'_c$ evaluated at f_{TK}^* . Evidently, for nonlinear effects to explain the steeper spectra without dissipation, ω_c must anomalously increase by a large factor of $\gtrsim 5$.

Discussion.—We have performed a statistical study of the scaling properties of the turbulent fluctuation spectrum in the inner heliosphere using data from the first PSP encounter. At low frequencies, there is an “inertial range” with a spectral index ranging from $-5/3$ to $-3/2$ [21–23], while at high frequencies, there is a “kinetic range” with a spectral index of around -2.7 [28–30]. Between these, an anomalously steep transition range with a highly variable spectrum typically appears [15–19]. Even in intervals where transition-range steepening is not particularly dramatic, spectra at ion scales ($k_{\perp}\rho \sim 1$) remain steeper than the deep sub-ion-range spectrum ($k_{\perp}\rho > 1$). This indicates that three spectral ranges are regularly necessary to describe collisionless plasma turbulence throughout the inertial and ion-kinetic ranges.

Ion-scale steepening either corresponds to significant turbulent dissipation, or a dramatic change in nonlinear cascade rates. Using synthetic spectra with no ion-scale steepening, we empirically constrain the effects associated with steepening. If dissipation is the dominant effect, the typically observed transition-range spectrum of -3.5 corresponds roughly to a 50% dissipation of energy flux in this range, indicating significant ion-scale heating. The half of intervals with the strongest signatures of ion-scale steepening have a typical transition index of -3.9 , corresponding to heating rates that are $\gtrsim 90\%$ of the turbulent energy flux. This amount of heating is consistent with the turbulent dissipation required to power the solar wind [14,51–56]. Measuring the radial trends in transition-range heating may account for solar wind temperature profile and heating rates at 1 A.U.

Several competing mechanisms have been suggested for ion-scale dissipation: e.g., stochastic heating [42], Landau damping [13,17,76,77], and cyclotron resonance [10,16,18]. Correlations between signatures of the transition range and parameters associated with stochastic heating were not found [42,63]. Measurements of 3D ion distributions will enable comparison of damping rates associated with, e.g., Landau and cyclotron resonance, with our observed dissipation rates [1,8,10–14,17,18,76,82,83]. Additionally, heating processes may leave clear signatures in the distribution function: e.g., diffusive shells in the case of cyclotron resonance [84] or flattening associated with resonant damping or perpendicular stochastic heating [63,85]. Use of electric fields will enable analysis of alignment over the transition range, enabling further constraint of energy transfer, heating rates, and mechanisms [86–88].

The Batchelor cascade model (3) and (4) cannot uniquely distinguish heating from increased spectral transfer rates. If anomalous variations in cascade rates are responsible for ion-scale steepening, then the nonlinear

frequency (inverse cascade time) of the turbulence must increase by a factor of $\gtrsim 5$ relative to unsteepened intervals. One possibility for increased nonlinear transfer includes reconnection onset and loss of dynamic alignment [43–45]; however, we find no connection between the critical parameters for reconnection and the transition range. Additionally, nonlinear interactions between copropagating waves in imbalanced turbulence may cause ion-scale steepening [46]. However, the required imbalance is significantly larger than what is observed in the inner heliosphere [64,65]. No correlation between spectral steepening and cross helicity was found, throwing the influence of alignment on ion-scale steepening further into doubt. Effects of damping on highly imbalanced turbulence may provide insight into the transition range [89]. Though we exclude intervals with coherent waves, the presence of intermittency may affect nonlinear interaction at ion scales [40,47,48,50,90]. Resolving such effects in PSP likely requires electric field measurements to constrain turbulent imbalance and density fluctuations at faster cadences than currently available.

We demonstrate that ion-scale spectral steepening is a dramatic effect in collisionless plasma turbulence. While further work is needed to distinguish exact mechanisms behind observed ion-scale spectral steepening, this Letter largely constrains their effects on the turbulent cascade.

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