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SOURCES OF QUASI-PERIODIC PROPAGATING DISTURBANCES ABOVE A SOLAR POLAR CORONAL HOLE

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ABSTRACT

Quasi-periodic propagating disturbances (PDs) are ubiquitous in polar coronal holes on the Sun. It remains unclear as to what generates PDs. In this work, we investigate how the PDs are generated in the solar atmosphere by analyzing a four-hour data set taken by the Atmospheric Imaging Assembly (AIA) on board the *Solar Dynamics Observatory*. We find convincing evidence that spicular activities in the solar transition region, as seen in the AIA 304 Å passband, are responsible for PDs in the corona as revealed in the AIA 171 Å images. We conclude that spicules are an important source that triggers coronal PDs.

Key words: Sun: atmosphere – Sun: chromosphere – Sun: corona – Sun: transition region

Supporting material: animations

1. INTRODUCTION

Quasi-periodic propagating disturbances (PDs) in solar polar coronal holes have been reported since the *Skylab* era. When analyzing the Mg x 625 Å line radiance in plumes, Withbroe (1983) was the first to discover intensity disturbances with ~10% variation and propagating speeds of 100–200 km s⁻¹. Since the launch of the *Solar and Heliospheric Observatory* (*SOHO*), this phenomenon has been extensively studied. With observations taken off-limb at 1.01–1.2 R_⊙ by the Extreme-ultraviolet Imaging Telescope, DeForest & Gurman (1998) found PDs with 10%–20% intensity variations, quasi-periodicities of 10–15 minutes, and propagating speeds of 75–150 km s⁻¹ in fine structures (with a width of 3''–5'') of polar plumes. By using high spatial resolution data taken by the Atmospheric Imaging Assembly (AIA) on board the *Solar Dynamics Observatory* (*SDO*), PDs are found in both plume and inter-plume regions. In the AIA data, the periodicities of PDs are found to be 10–30 minutes, and the propagating speeds are 100–170 km s⁻¹ (Krishna Prasad et al. 2011). Spectral measurements of plumes made with, e.g., CDS and SUMER on board *SOHO* and EIS on board *Hinode*, also showed the existence of PDs with periodicities of 10–30 minutes and relative intensity variations of a few percent (Banerjee et al. 2000, 2009). Furthermore, H I Lyα measurements with UVCS/*SOHO* indicate that signatures of PDs are likely to be present up to 2.2 R_⊙ (Morgan et al. 2004).

The fact that PDs tend to propagate with a speed close to the sound speed suggests that they are likely to be signatures of propagating slow magnetic-acoustic waves (DeForest & Gurman 1998; Ofman et al. 2000; Gupta et al. 2012). Alfvén waves and fast magnetic-acoustic waves were also suggested to produce PDs in inter-plume regions (Gupta et al. 2010). On the other hand, by analyzing EIS/*Hinode* and AIA/*SDO* data, De Pontieu et al. (2009) revealed the presence of quasi-periodic outflows in the footpoints of coronal loops. Furthermore, Tian et al. (2011) discovered that high-speed quasi-periodic outflows can be found in both plume and inter-plume regions, and the outflow speeds are similar to the propagating speeds of PDs.

These outflows have been suggested to be sources of PDs (McIntosh et al. 2010).

How PDs are generated remains debated. Although an obvious candidate source, magnetoacoustic-gravity (MAG) waves in the photosphere with periods longer than ~10 minutes cannot readily penetrate the chromosphere and make it into the corona since the periods are substantially longer than some critical value P_c (Bel & Leroy 1977). For a low-beta environment, this P_c reads $4\pi c_s / (\gamma g \cos \theta)$, where $\gamma = 5/3$ is the adiabatic index, $g = 27,400 \text{ cm s}^{-2}$ is the Sun's gravitational acceleration, c_s is the sound speed, and θ is the angle between the wavevector and the background magnetic field (Yuan et al. 2014). Adopting $c_s = 7 \text{ km s}^{-1}$ as being representative of chromospheric values, one finds that $P_c = 192 / \cos \theta \text{ s}$, meaning that MAG waves with periods $\gtrsim 10$ minutes can reach the corona only along highly inclined magnetic flux tubes with $\theta \gtrsim 70^\circ$. However, if PDs have coronal sources, the restriction due to P_c is not as stringent. One possibility of this kind will be that PDs are the coronal response to spicules impinging on the corona in a repeated manner. Indeed, SUMER measurements show that spicules, the omnipresent ejections of cool chromospheric plasmas to coronal heights, can occur repeatedly at the same place (Xia et al. 2005b) with periods of 10–15 minutes (Xia et al. 2005a). Other support for this possibility comes from the recent numerical studies by Ofman et al. (2012) and Wang et al. (2013), who showed that any upflow pulse at the coronal base, regardless of its origin, leads to propagating intensity disturbances guided by magnetic flux tubes.

In the present study, by analyzing the high-resolution AIA data, we aim to examine whether or not spicules can be a source of PDs. This Letter is organized as follows. Section 2 describes the observational data, method of data analysis, and our results. Section 3 summarizes our conclusions.

2. OBSERVATIONS, DATA ANALYSIS, AND RESULTS

The data analyzed in this study were taken from 21:50 UT, 2010 August 5 to 01:50 UT, August 6 by the AIA (Lemen

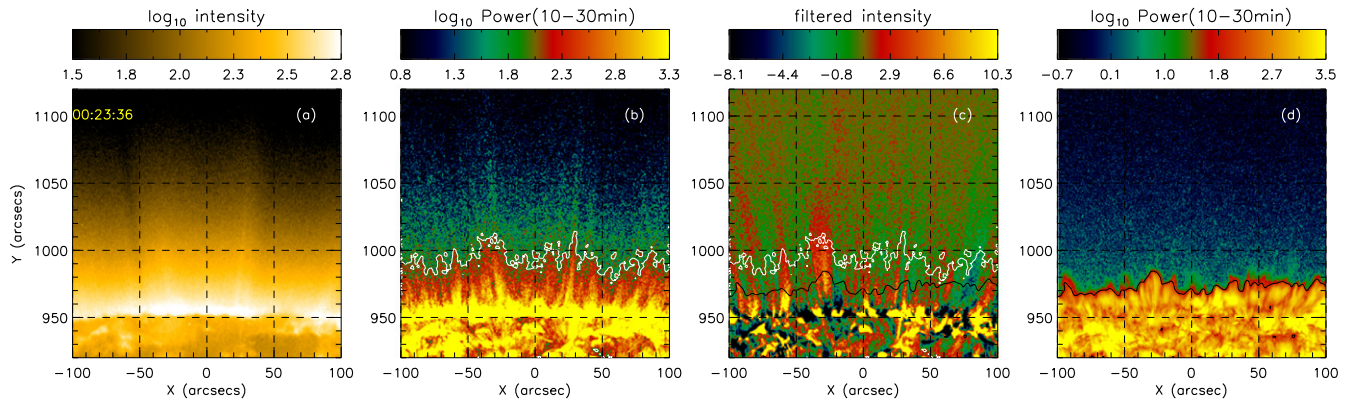


Figure 1. Field of view studied in the present work. Panel (a) gives the original AIA 171 Å image, panel (b) shows the spatial distribution of wavelet power summed over periodicities of 10–30 minutes, panel (c) presents the bandpass (10–30 minutes) filtered images, and panel (d) displays the wavelet power summed over periodicities of 10–30 minutes for the AIA 304 Å. The white contours overplotted on panels (b) and (c) outline the structures of PDs determined from panel (b), while the black lines (panels (c) and (d)) represent the contour of the AIA 304 Å power image (i.e., panel (d)).

et al. 2012). The spatial resolution of the data is about $1''.2$ and the cadence is 12 s. The AIA monitors the full disk of the Sun in 10 wavelength bands. In this work, we chose to use the passbands of 304 Å (mainly contributed by He II with a response function peaking at 0.05 MK) and 171 Å (mainly contributed by Fe IX with a response function peaking at 0.8 MK). The AIA 304 Å passband can well observe spicules in the transition region, and the 171 Å one is commonly used to identify PDs in the corona. We focus our investigation on the north polar region with solar X between $-100''$ and $100''$, where a coronal hole is clearly seen (Figure 1(a)).

We started with applying a wavelet tool kit to the time series of every pixel of the AIA 171 Å images. The wavelet power in a selected range of periodicities was then used to construct a power image. An example is given in Figure 1(b), where we show the spatial distribution of wavelet power with periodicities between 10 and 30 minutes at 00:23 UT. Note that constructing a wavelet power map involves a sequence of images over a certain time interval. However, it is possible that a power map is assigned a particular time around which this interval is centered. By following the time series of the power images, we found that PDs are an inherent structure above the polar coronal hole (see the elongated structures in Figure 1(b)). Although most of the PD structures are clearly seen below $1.1 R_{\odot}$ (i.e., solar Y being $1050''$), some PDs show clear structures reaching an altitude of $1.15 R_{\odot}$ (i.e., solar Y being $1100''$) or even higher. They are clearly present in power images at periodicities from 10 to 30 minutes, which are consistent with previous studies (DeForest & Gurman 1998; Banerjee et al. 2000, 2009; Krishna Prasad et al. 2011). To compare the PDs with spicules that were clearly visible in the AIA 304 Å channel, we further applied a simple bandpass filter to the AIA 171 Å images. The time series of intensity at each pixel was smoothed with a 50-point window to filter out signals with periods shorter than 10 minutes. The resultant time series were then subtracted by a 150-point running average of the original time series to filter out signals with periods longer than 30 minutes. The snapshot at 00:23 UT of the final product, essentially a bandpass-filtered image series, is shown in Figure 1(c). While Figures 1(b) and (c) look similar, we find that the temporal evolution of the PDs can be better followed in the bandpass-filtered images. This is crucial when comparing PDs with the rapidly evolving spicular activities.

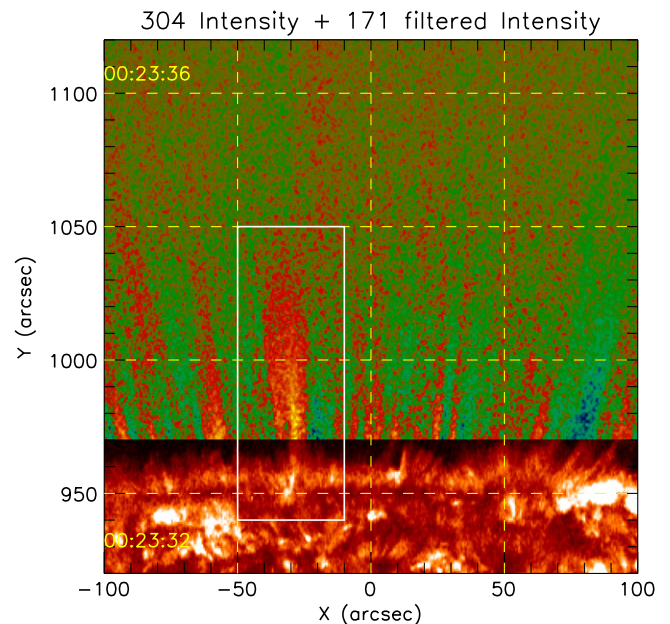


Figure 2. Composite image constructed from the original AIA 304 Å (bottom) and filtered AIA 171 Å (top) images. Spicules are visible in the AIA 304 Å image, while PDs are clear in the filtered AIA 171 Å images. The white solid lines denote the region that is studied in detail in Figures 3–5. The accompanying animation makes it easy for one to follow the evolution of the spicular activities and the PDs.

(An animation of this figure is available.)

Figure 2 and its corresponding animation show composite images constructed from the original AIA 304 Å images with solar Y ranging from $920''$ to $970''$ and bandpass-filtered AIA 171 Å images with solar Y ranging from $970''$ to $1120''$. The AIA 304 Å channel observes plasmas at around 0.05 MK, capturing the dynamics of the solar transition region and is ideal for identifying spicules. While the PDs are well identifiable in the AIA 171 Å filtered images, the series of the composite images is the best approach for investigating the correlation between the PDs in the corona and the spicules in the transition region. The animation clearly demonstrates that when a spicule is discernible in the AIA 304 Å channel, a PD is seen to be generated and/or amplified in the AIA 171 Å. Note that a one-to-one correspondence between spicules and PDs is

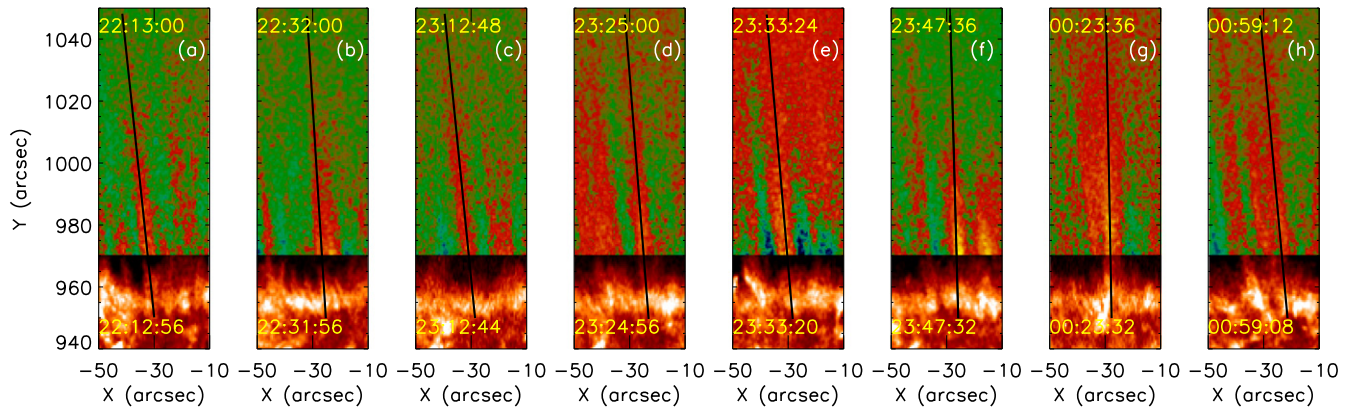


Figure 3. Snapshots showing the connection between spicules in the AIA 304 Å images (bottom) and PDs in the filtered AIA 171 Å images (top). The solid line in each image denotes the cut used to produce the time-slice plots in Figure 4.

(An animation of this figure is available.)

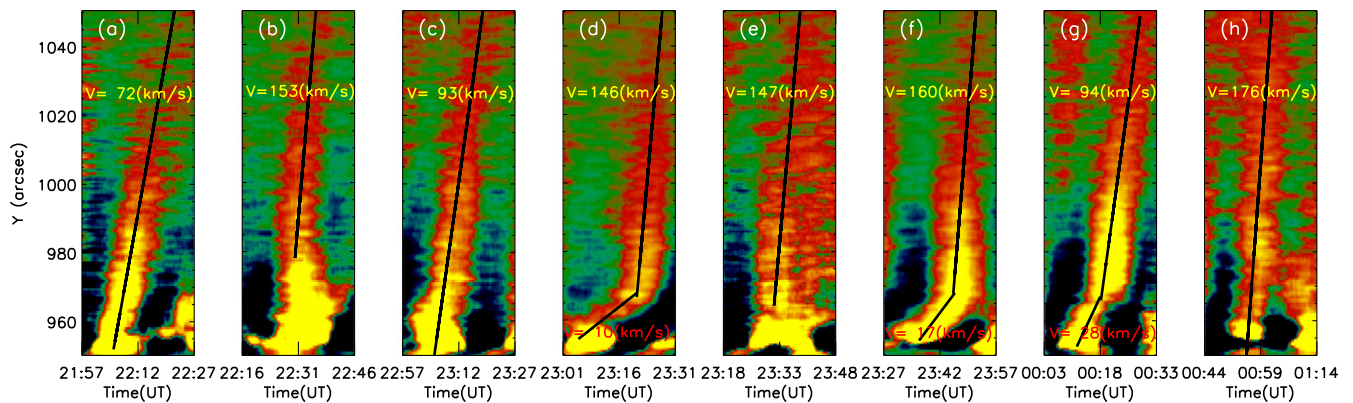


Figure 4. Time-slice images of the PDs shown in Figure 3 (here, only AIA 171 Å measurements are used). The solid lines are the slopes from which the speeds of the propagating structures are determined.

not found in some cases; see our discussion in connection to Figure 5.

We further examined a small area denoted by the white box in Figure 2. In Figure 3 and its corresponding animation, we present the evolution of the small area viewed in the composite images. Again, it demonstrates that spicular activities in the solar transition region are a source of the PDs in the corona. We also measured the speeds of the PDs by producing time-slice images along the propagation direction (the solid lines in Figure 3). The speeds measured from these PDs vary from 70 to 170 km s⁻¹, which are consistent with previous results (Withbroe 1983; DeForest & Gurman 1998; Krishna Prasad et al. 2011). In a few cases (see Figures 4(d), (f), and (g)), we found that the propagating structures have lower speeds (10–30 km s⁻¹) near the limb before gaining much higher speed above it. We suggest that the speed near the limb represents the speed of the spicular plasmas because the spicules can also be seen near the limb in the AIA 171 Å channel. The triggered PDs, on the other hand, are propagating with a much higher speed.

The PD speeds measured here are significantly larger than those of classic type I spicules (~ 20 km s⁻¹; Mariska 1992) and are close to those of type II spicules (50–150 km s⁻¹; De Pontieu et al. 2007). One may then ask whether or not type II spicules are a direct trigger of the PDs. To address this issue, we further analyzed the upward speeds of spicules that are found to trigger PDs in the corona. Because spicules viewed in

the AIA 304 Å are too crowded, only 32 relatively isolated examples were selected and their upward speeds were measured. Among these, 17 (15) spicules were found to have upward speeds larger (less) than 50 km s⁻¹. Taking a speed of 50 km s⁻¹ as the critical value separating type I from type II spicules, one would then conclude that these two types of spicules tend to be equally important in triggering PDs. However, it is not straightforward to identify the counterpart of type II spicules, the thin elongated structure originally found in *Hinode*/SOT Ca II H emissions, in the AIA 304 Å images. The reason for this is twofold. First, with a bandwidth as wide as 40 Å, the AIA 304 Å passband can actually record emissions formed at a variety of temperatures. Second, the spatial resolution in AIA 304 Å is $\sim 1''.2$, substantially less good than what is achieved by SOT ($\sim 0''.3$). An analysis combined with spectroscopic data will be crucial for associating features in the AIA 304 Å passband with type II spicules.

Our attention has been drawn to a very recent work by Pant et al. (2015), where the authors suggest that small-scale jetlike features with speeds of ~ 10 km s⁻¹ occurring in the chromospheric network (seen in IRIS 1330 Å slit-jaw images) might be responsible for the PDs in on-disk plumes. As has already been pointed out by the authors, measuring the speeds of those jetlike features is not straightforward due to the projection effect. In this sense, it remains an open question as to whether or not the jetlike features associated with PDs in on-disk plumes are spicules. Nonetheless, the present study and the

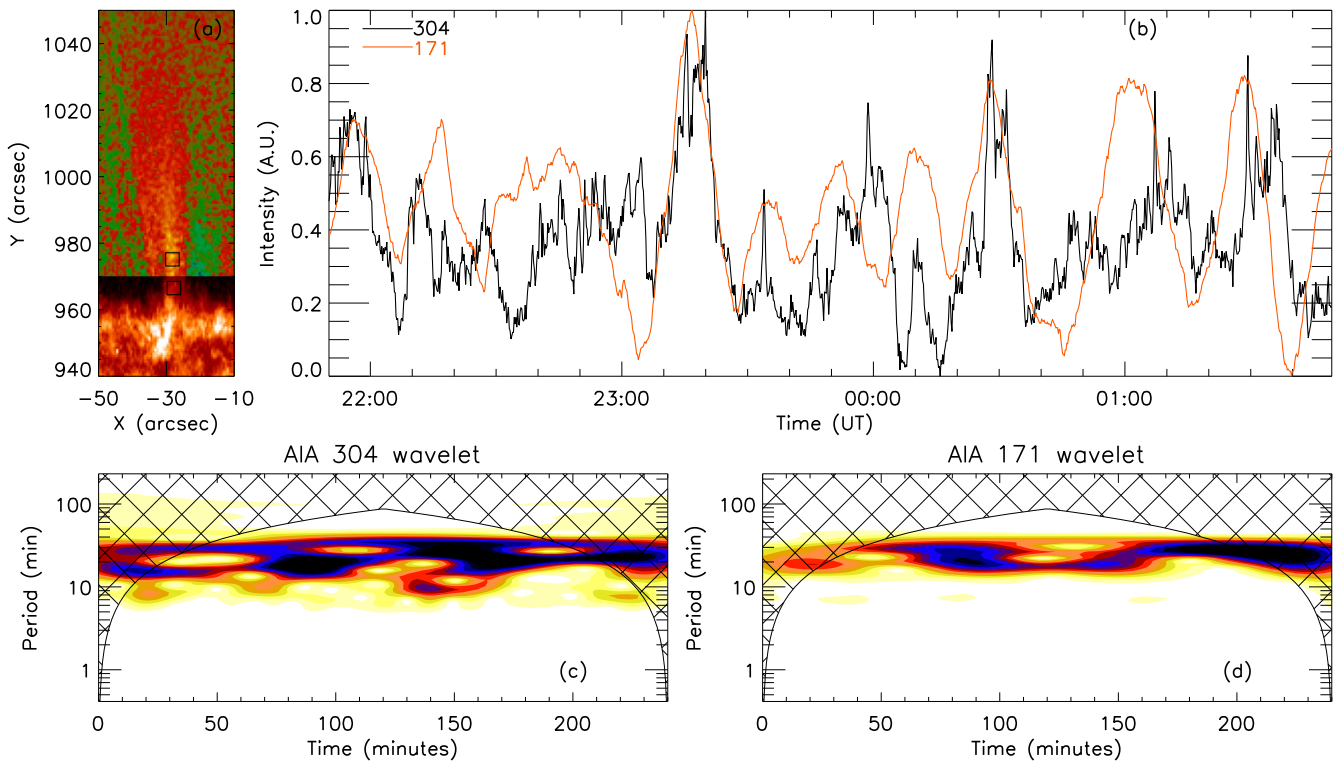


Figure 5. Light curves of the AIA 304 Å (b), the black line) and the filtered AIA 171 Å (b), the red line) taken from two small boxed regions marked in panel (a) (black solid lines). The wavelet power spectra of the two light curves are also given ((c), original 304 Å; (d), filtered 171 Å).

work by Pant et al. (2015) suggest that PDs in polar coronal holes and PDs in on-disk plumes may both have chromospheric sources. Given the possible difference in the local magnetic field topology in the two regions, it may be that this magnetic topology plays an important role in determining exactly what in the chromosphere triggers PDs in the corona. This intriguing question certainly warrants a dedicated study, which is beyond the scope of this Letter.

One may now ask why intensity disturbances triggered by spicules are quasi-periodic. An obvious possibility is that the spicules themselves possess similar periodicities, e.g., in connection to their quasi-periodic occurrence at the same location (see Section 1). This possibility is examined in Figure 5, which illustrates the quasi-periodic nature of the spicular activity and its relation to PDs. Figure 5(b) shows the light curves of the AIA 304 Å (black line) and the filtered AIA 171 Å (red line) taken from the lower and higher boxes marked in Figure 5(a), and Figures 5(c) and (d) show the corresponding wavelet spectra. We note that the two boxes are selected such that they represent the spicule in the 304 Å channel and the corresponding PDs in the 171 Å channel. At first glance, the two light curves are quite well correlated in general (see Figure 5(b)). However, some spikes in the AIA 304 Å light curve do not have a clear counterpart in the light curve of the 171 Å. The reason for this is threefold. First, the spicules are rather crowded with a variety of directions, making it difficult to distinguish their contributions to PDs in the AIA 171 Å channel. Second, spicules tend to experience substantial transverse displacement, meaning that the light curve gathered from one fixed location may correspond to a substantial number of different spicules rather than following an individual one. Third, it is likely that other non-spicular features also contribute to these spikes in the AIA 304 Å light curve. A

further investigation via wavelet analyses (Figures 5(c) and (d)) also confirms that the light curve of the AIA 304 Å owns periodicities similar to those in the AIA 171 Å light curve, i.e., 10–30 minutes. Furthermore, the consistency in the periodicities is found not only in the small region shown in Figure 5, but also in the whole field of view studied here. A power map of AIA 304 Å summed over periodicities of 10–30 minutes is displayed in Figure 1(d), which correlates well with the PDs shown in AIA 171 Å images (Figures 1(b) and (c)). This again supports the notion that the spicules seen in the transition region are a source of the PDs in the corona.

3. CONCLUSION

In this study, we have analyzed a four-hour data set of a polar coronal hole observed by the AIA. PDs with speeds of $70\text{--}170\text{ km s}^{-1}$ and periodicities of 10–30 minutes are clearly seen in the AIA 171 Å images. By following the evolution of the spicular activities in the solar transition region (seen in the AIA 304 Å channel) and the PDs in the corona (seen in the AIA 171 Å channel), we found clear evidence that spicules are one of the important sources of PDs.

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Facility: SDO.

REFERENCES

- Banerjee, D., O'Shea, E., & Doyle, J. G. 2000, *SoPh*, **196**, 63
Banerjee, D., Teriaca, L., Gupta, G. R., et al. 2009, *A&A*, **499**, L29
Bel, N., & Leroy, B. 1977, *A&A*, **55**, 239
De Pontieu, B., McIntosh, S., Hansteen, V. H., et al. 2007, *PASJ*, **59**, 655
De Pontieu, B., McIntosh, S. W., Hansteen, V. H., & Schrijver, C. J. 2009, *ApJL*, **701**, L1
DeForest, C. E., & Gurman, J. B. 1998, *ApJL*, **501**, L217
Gupta, G. R., Banerjee, D., Teriaca, L., Imada, S., & Solanki, S. 2010, *ApJ*, **718**, 11
Gupta, G. R., Teriaca, L., Marsch, E., Solanki, S. K., & Banerjee, D. 2012, *A&A*, **546**, A93
Krishna Prasad, S., Banerjee, D., & Gupta, G. R. 2011, *A&A*, **528**, L4
Lemen, J. R., Title, A. M., Akin, D. J., et al. 2012, *SoPh*, **275**, 17
Mariska, J. T. 1992, *The Solar Transition Region* (New York: Cambridge Univ. Press)
McIntosh, S. W., Innes, D. E., de Pontieu, B., & Leamon, R. J. 2010, *A&A*, **510**, L2
Morgan, H., Habbal, S. R., & Li, X. 2004, *ApJ*, **605**, 521
Ofman, L., Nakariakov, V. M., & Sehgal, N. 2000, *ApJ*, **533**, 1071
Ofman, L., Wang, T. J., & Davila, J. M. 2012, *ApJ*, **754**, 111
Pant, V., Dolla, L., Mazumder, R., et al. 2015, *ApJ*, **807**, 71
Tian, H., McIntosh, S. W., Habbal, S. R., & He, J. 2011, *ApJ*, **736**, 130
Wang, T., Ofman, L., & Davila, J. M. 2013, *ApJL*, **775**, L23
Withbroe, G. L. 1983, *SoPh*, **89**, 77
Xia, L. D., Popescu, M. D., Chen, Y., & Doyle, J. G. 2005a, in *Solar Wind 11/SOHO 16, Connecting Sun and Heliosphere*, ed. B. Fleck, T. H. Zurbuchen, & H. Lacoste (ESA SP-592; Noordwijk: ESA), 575
Xia, L. D., Popescu, M. D., Doyle, J. G., & Giannikakis, J. 2005b, *A&A*, **438**, 1115
Yuan, D., Sych, R., Reznikova, V. E., & Nakariakov, V. M. 2014, *A&A*, **561**, A19