

THE FIRST CATALOG OF ACTIVE GALACTIC NUCLEI DETECTED BY THE *FERMI* LARGE AREA TELESCOPE

A. A. ABDO^{1,2}, M. ACKERMANN³, M. AJELLO³, A. ALLAFORT³, E. ANTOLINI^{4,5}, W. B. ATWOOD⁶, M. AXELSSON^{7,8}, L. BALDINI⁹, J. BALLE¹⁰, G. BARBIELLINI^{11,12}, D. BASTIERI^{13,14}, B. M. BAUGHMAN¹⁵, K. BECHTOL³, R. BELLAZZINI⁹, B. BERENJI³, R. D. BLANDFORD³, E. D. BLOOM³, J. R. BOGART³, E. BONAMENTE^{4,5}, A. W. BORGLAND³, A. BOUVIER³, J. BREGEON⁹, A. BREZ⁹, M. BRIGIDA^{16,17}, P. BRUEL¹⁸, R. BUEHLER³, T. H. BURNETT¹⁹, S. BUSON¹³, G. A. CALIANDRO²⁰, R. A. CAMERON³, A. CANNON^{21,22}, P. A. CARAVEO²³, S. CARRIGAN¹⁴, J. M. CASANDJIAN¹⁰, E. CAVAZZUTI²⁴, C. CECCHI^{4,5}, Ö. ÇELİK^{21,25,26}, A. CELOTTI²⁷, E. CHARLES³, A. CHEKHTMAN^{1,28}, A. W. CHEN²³, C. C. CHEUNG^{1,2}, J. CHIANG³, S. CIPRINI⁵, R. CLAU³, J. COHEN-TANUGI²⁹, J. CONRAD^{8,30,68}, L. COSTAMANTE³, G. COTTER³¹, S. CUTINI²⁴, V. D'ELIA²⁴, C. D. DERMER¹, A. DE ANGELIS³², F. DE PALMA^{16,17}, A. DE ROSA³³, S. W. DIGEL³, E. DO COUTO E SILVA³, P. S. DRELL³, R. DUBOIS³, D. DUMORA^{34,35}, L. ESCANDE^{34,35}, C. FARNIER²⁹, C. FAVUZZI^{16,17}, S. J. FEGAN¹⁸, E. C. FERRARA²¹, W. B. FOCKE³, P. FORTIN¹⁸, M. FRAILIS^{32,36}, Y. FUKAZAWA³⁷, S. FUNK³, P. FUSCO^{16,17}, F. GARGANO¹⁷, D. GASPARRINI²⁴, N. GEHRELS^{21,38,39}, S. GERMANI^{4,5}, B. GIEBELS¹⁸, N. GIGLIETTO^{16,17}, P. GIOMMI²⁴, F. GIORDANO^{16,17}, M. GIROLETTI⁴⁰, T. GLANZMAN³, G. GODFREY³, P. GRANDI⁴¹, I. A. GRENIER¹⁰, M.-H. GRONDIN^{34,35}, J. E. GROVE¹, S. GUIREC⁴², D. HADASCH⁴³, A. K. HARDING²¹, M. HAYASHIDA³, E. HAYS²¹, S. E. HEALEY³, A. B. HILL^{44,69}, D. HORAN¹⁸, R. E. HUGHES¹⁵, G. IAFRATE^{11,36}, R. ITOH³⁷, G. JÓHANNESSEN³, A. S. JOHNSON³, R. P. JOHNSON⁶, T. J. JOHNSON^{21,39}, W. N. JOHNSON¹, T. KAMAE³, H. KATAGIRI³⁷, J. KATAOKA⁴⁵, N. KAWAI^{46,47}, M. KERR¹⁹, J. KNÖDLSER⁴⁸, M. KUSS⁹, J. LANDE³, L. LATRONICO⁹, C. LAVALLEY²⁹, M. LEMOINE-GOUMARD^{34,35}, M. LLENA GARDE^{8,30}, F. LONGO^{11,12}, F. LOPARCO^{16,17}, B. LOTT^{34,35}, M. N. LOVELLETTE¹, P. LUBRANO^{4,5}, G. M. MADEJSKI³, A. MAKEEV^{1,28}, G. MALAGUTI⁴¹, E. MASSARO⁴⁹, M. N. MAZZIOTTA¹⁷, W. MCCONVILLE^{21,39}, J. E. MCENERY^{21,39}, S. MCGLYNN^{8,50}, P. F. MICHELSON³, W. MITTHUMSIRI³, T. MIZUNO³⁷, A. A. MOISEEV^{25,39}, C. MONTE^{16,17}, M. E. MONZANI³, A. MORSELLI⁵¹, I. V. MOSKALENKO³, S. MURGIA³, P. L. NOLAN³, J. P. NORRIS⁵², E. NUSS²⁹, M. OHNO⁵³, T. OHSUGI⁵⁴, N. OMODEI³, E. ORLANDO⁵⁵, J. F. ORMES⁵², M. OZAKI⁵³, D. PANEQUE³, J. H. PANETTA³, D. PARENT^{1,28,34,35}, V. PELASSA²⁹, M. PEPE^{4,5}, M. PESCE-ROLLINS⁹, S. PIRANOMONTE⁵⁶, F. PIRON²⁹, T. A. PORTER³, S. RAINÒ^{16,17}, R. RANDO^{13,14}, M. RAZZANO⁹, A. REIMER^{3,57}, O. REIMER^{3,57}, T. REPOSEUR^{34,35}, J. RIPKEN^{8,30}, S. RITZ⁶, A. Y. RODRIGUEZ²⁰, R. W. ROMANI³, M. ROTH¹⁹, F. RYDE^{8,50}, H. F.-W. SADROZINSKI⁶, D. SANCHEZ¹⁸, A. SANDER¹⁵, P. M. SAZ PARKINSON⁶, J. D. SCARGLE⁵⁸, C. SGRÒ⁹, M. S. SHAW³, E. J. SISKIND⁵⁹, P. D. SMITH¹⁵, G. SPANDRE⁹, P. SPINELLI^{16,17}, J.-L. STARCK¹⁰, Ł. STAWARZ^{3,60}, M. S. STRICKMAN¹, D. J. SUSON⁶¹, H. TAJIMA³, H. TAKAHASHI⁵⁴, T. TAKAHASHI⁵³, T. TANAKA³, G. B. TAYLOR⁶², J. B. THAYER³, J. G. THAYER³, D. J. THOMPSON²¹, L. TIBALDO^{10,13,14,70}, D. F. TORRES^{20,43}, G. TOSTI^{4,5}, A. TRAMACERE^{3,63,64}, P. UBERTINI³³, Y. UCHIYAMA³, T. L. USHER³, V. VASILEIOU^{25,26}, N. VILCHEZ⁴⁸, M. VILLATA⁶⁵, V. VITALE^{51,66}, A. P. WAITE³, E. WALLACE¹⁹, P. WANG³, B. L. WINER¹⁵, K. S. WOOD¹, Z. YANG^{8,30}, T. YLINEN^{8,50,67}, AND M. ZIEGLER⁶

¹ Space Science Division, Naval Research Laboratory, Washington, DC 20375, USA

² National Research Council Research Associate, National Academy of Sciences, Washington, DC 20001, USA

³ W. W. Hansen Experimental Physics Laboratory, Kavli Institute for Particle Astrophysics and Cosmology, Department of Physics and SLAC National Accelerator Laboratory, Stanford University, Stanford, CA 94305, USA; sehealey@astro.stanford.edu

⁴ Istituto Nazionale di Fisica Nucleare, Sezione di Perugia, I-06123 Perugia, Italy; Gino.Tosti@pg.infn.it

⁵ Dipartimento di Fisica, Università degli Studi di Perugia, I-06123 Perugia, Italy

⁶ Santa Cruz Institute for Particle Physics, Department of Physics and Department of Astronomy and Astrophysics, University of California at Santa Cruz, Santa Cruz, CA 95064, USA

⁷ Department of Astronomy, Stockholm University, SE-106 91 Stockholm, Sweden

⁸ The Oskar Klein Centre for Cosmoparticle Physics, AlbaNova, SE-106 91 Stockholm, Sweden

⁹ Istituto Nazionale di Fisica Nucleare, Sezione di Pisa, I-56127 Pisa, Italy

¹⁰ Laboratoire AIM, CEA-IRFU/CNRS/Université Paris Diderot, Service d'Astrophysique, CEA Saclay, 91191 Gif sur Yvette, France

¹¹ Istituto Nazionale di Fisica Nucleare, Sezione di Trieste, I-34127 Trieste, Italy

¹² Dipartimento di Fisica, Università di Trieste, I-34127 Trieste, Italy

¹³ Istituto Nazionale di Fisica Nucleare, Sezione di Padova, I-35131 Padova, Italy

¹⁴ Dipartimento di Fisica "G. Galilei," Università di Padova, I-35131 Padova, Italy

¹⁵ Department of Physics, Center for Cosmology and Astro-Particle Physics, The Ohio State University, Columbus, OH 43210, USA

¹⁶ Dipartimento di Fisica "M. Merlin" dell'Università e del Politecnico di Bari, I-70126 Bari, Italy

¹⁷ Istituto Nazionale di Fisica Nucleare, Sezione di Bari, 70126 Bari, Italy

¹⁸ Laboratoire Leprince-Ringuet, École polytechnique, CNRS/IN2P3, Palaiseau, France

¹⁹ Department of Physics, University of Washington, Seattle, WA 98195-1560, USA

²⁰ Institut de Ciències de l'Espai (IEEC-CSIC), Campus UAB, 08193 Barcelona, Spain

²¹ NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA

²² University College Dublin, Belfield, Dublin 4, Republic of Ireland

²³ INAF-Istituto di Astrofisica Spaziale e Fisica Cosmica, I-20133 Milano, Italy

²⁴ Agenzia Spaziale Italiana (ASI) Science Data Center, I-00044 Frascati (Roma), Italy; elisabetta.cavazzuti@asdc.asi.it, gasparrini@asdc.asi.it

²⁵ Center for Research and Exploration in Space Science and Technology (CRESST) and NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA

²⁶ Department of Physics and Center for Space Sciences and Technology, University of Maryland Baltimore County, Baltimore, MD 21250, USA

²⁷ Scuola Internazionale Superiore di Studi Avanzati (SISSA), 34014 Trieste, Italy

²⁸ George Mason University, Fairfax, VA 22030, USA

²⁹ Laboratoire de Physique Théorique et Astroparticules, Université Montpellier 2, CNRS/IN2P3, Montpellier, France

³⁰ Department of Physics, Stockholm University, AlbaNova, SE-106 91 Stockholm, Sweden

³¹ Astrophysics, Oxford University, Oxford OX1 3RH, UK

³² Dipartimento di Fisica, Università di Udine and Istituto Nazionale di Fisica Nucleare, Sezione di Trieste, Gruppo Collegato di Udine, I-33100 Udine, Italy

- ³³ INAF-Istituto di Astrofisica Spaziale e Fisica Cosmica, I-00133 Roma, Italy
- ³⁴ CNRS/IN2P3, Centre d'Études Nucléaires Bordeaux Gradignan, UMR 5797, Gradignan, 33175, France; lott@cenbg.in2p3.fr
- ³⁵ Université de Bordeaux, Centre d'Études Nucléaires Bordeaux Gradignan, UMR 5797, Gradignan, 33175, France
- ³⁶ Osservatorio Astronomico di Trieste, Istituto Nazionale di Astrofisica, I-34143 Trieste, Italy
- ³⁷ Department of Physical Sciences, Hiroshima University, Higashi-Hiroshima, Hiroshima 739-8526, Japan
- ³⁸ Department of Astronomy and Astrophysics, Pennsylvania State University, University Park, PA 16802, USA
- ³⁹ Department of Physics and Department of Astronomy, University of Maryland, College Park, MD 20742, USA
- ⁴⁰ INAF Istituto di Radioastronomia, 40129 Bologna, Italy
- ⁴¹ INAF-IASF Bologna, 40129 Bologna, Italy
- ⁴² Center for Space Plasma and Aeronomic Research (CSPAR), University of Alabama in Huntsville, Huntsville, AL 35899, USA
- ⁴³ Institució Catalana de Recerca i Estudis Avançats (ICREA), Barcelona, Spain
- ⁴⁴ Université Joseph Fourier-Grenoble 1/CNRS, laboratoire d'Astrophysique de Grenoble (LAOG) UMR 5571, BP 53, 38041 Grenoble Cedex 09, France
- ⁴⁵ Research Institute for Science and Engineering, Waseda University, 3-4-1, Okubo, Shinjuku, Tokyo 169-8555, Japan
- ⁴⁶ Department of Physics, Tokyo Institute of Technology, Meguro City, Tokyo 152-8551, Japan
- ⁴⁷ Cosmic Radiation Laboratory, Institute of Physical and Chemical Research (RIKEN), Wako, Saitama 351-0198, Japan
- ⁴⁸ Centre d'Étude Spatiale des Rayonnements, CNRS/UPS, BP 44346, F-30128 Toulouse Cedex 4, France
- ⁴⁹ Università di Roma "La Sapienza," I-00185 Roma, Italy
- ⁵⁰ Department of Physics, Royal Institute of Technology (KTH), AlbaNova, SE-106 91 Stockholm, Sweden
- ⁵¹ Istituto Nazionale di Fisica Nucleare, Sezione di Roma "Tor Vergata," I-00133 Roma, Italy
- ⁵² Department of Physics and Astronomy, University of Denver, Denver, CO 80208, USA
- ⁵³ Institute of Space and Astronautical Science, JAXA, 3-1-1 Yoshinodai, Sagami-hara, Kanagawa 229-8510, Japan
- ⁵⁴ Hiroshima Astrophysical Science Center, Hiroshima University, Higashi-Hiroshima, Hiroshima 739-8526, Japan
- ⁵⁵ Max-Planck Institut für extraterrestrische Physik, 85748 Garching, Germany
- ⁵⁶ Osservatorio Astronomico di Roma, 00040 Monte Porzio Catone, Italy
- ⁵⁷ Institut für Astro- und Teilchenphysik and Institut für Theoretische Physik, Leopold-Franzens-Universität Innsbruck, A-6020 Innsbruck, Austria
- ⁵⁸ Space Sciences Division, NASA Ames Research Center, Moffett Field, CA 94035-1000, USA
- ⁵⁹ NYCB Real-Time Computing Inc., Lattingtown, NY 11560-1025, USA
- ⁶⁰ Astronomical Observatory, Jagiellonian University, 30-244 Kraków, Poland
- ⁶¹ Department of Chemistry and Physics, Purdue University Calumet, Hammond, IN 46323-2094, USA
- ⁶² University of New Mexico, MSC07 4220, Albuquerque, NM 87131, USA
- ⁶³ Consorzio Interuniversitario per la Fisica Spaziale (CIFS), I-10133 Torino, Italy
- ⁶⁴ INTEGRAL Science Data Centre, CH-1290 Versoix, Switzerland
- ⁶⁵ INAF, Osservatorio Astronomico di Torino, I-10025 Pino Torinese (TO), Italy
- ⁶⁶ Dipartimento di Fisica, Università di Roma "Tor Vergata," I-00133 Roma, Italy
- ⁶⁷ School of Pure and Applied Natural Sciences, University of Kalmar, SE-391 82 Kalmar, Sweden
- Received 2010 January 29; accepted 2010 March 27; published 2010 April 29*

ABSTRACT

We present the first catalog of active galactic nuclei (AGNs) detected by the Large Area Telescope (LAT), corresponding to 11 months of data collected in scientific operation mode. The First LAT AGN Catalog (1LAC) includes 671 γ -ray sources located at high Galactic latitudes ($|b| > 10^\circ$) that are detected with a test statistic greater than 25 and associated statistically with AGNs. Some LAT sources are associated with multiple AGNs, and consequently, the catalog includes 709 AGNs, comprising 300 BL Lacertae objects, 296 flat-spectrum radio quasars, 41 AGNs of other types, and 72 AGNs of unknown type. We also classify the blazars based on their spectral energy distributions as archival radio, optical, and X-ray data permit. In addition to the formal 1LAC sample, we provide AGN associations for 51 low-latitude LAT sources and AGN "affiliations" (unquantified counterpart candidates) for 104 high-latitude LAT sources without AGN associations. The overlap of the 1LAC with existing γ -ray AGN catalogs (LBAS, EGRET, *AGILE*, *Swift*, *INTEGRAL*, TeVCat) is briefly discussed. Various properties—such as γ -ray fluxes and photon power-law spectral indices, redshifts, γ -ray luminosities, variability, and archival radio luminosities—and their correlations are presented and discussed for the different blazar classes. We compare the 1LAC results with predictions regarding the γ -ray AGN populations, and we comment on the power of the sample to address the question of the blazar sequence.

Key words: BL Lacertae objects: general – catalogs – galaxies: active – gamma rays: galaxies

Online-only material: color figures, machine-readable tables

1. INTRODUCTION

The *Gamma-ray Large Area Space Telescope* was launched on 2008 June 11. It began its scientific operations 2 months later, and shortly thereafter, it was renamed the *Fermi Gamma-ray Space Telescope*. Its primary instrument is the Large Area Telescope (LAT; Atwood et al. 2009), the successor

to the Energetic Gamma-Ray Experiment Telescope (EGRET) on board the *Compton Gamma-Ray Observatory* (Thompson et al. 1993). The LAT offers a major increase in sensitivity over EGRET and the Italian Space Agency's *Astro-rivelatore Gamma a Immagini Leggero* (*AGILE*; Tavani et al. 2008), allowing it to study the γ -ray sky in unprecedented detail. In sky-survey mode (*Fermi*'s main observing mode), the LAT observes the entire sky every 3 hr, providing effectively uniform exposure on the timescale of days.

One of the major scientific goals of the *Fermi* mission is to investigate high-energy emission in active galactic nuclei (AGNs). Although it is generally accepted that the γ -rays

⁶⁸ Royal Swedish Academy of Sciences Research Fellow, funded by a grant from the K. A. Wallenberg Foundation.

⁶⁹ Funded by contract ERC-StG-200911 from the European Community.

⁷⁰ Partially supported by the International Doctorate on Astroparticle Physics (IDAPP) program.

detected from blazars are emitted from collimated jets of charged particles moving at relativistic speeds (Blandford & Rees 1978; Maraschi et al. 1992), open questions remain. The mechanisms by which the particles are accelerated, the precise site of the γ -ray emission, and the origin of AGN variability and the γ -ray duty cycle of blazars are not well understood. The physical reasons for the observational differences between radio-loud and radio-quiet AGNs and between flat-spectrum radio quasars (FSRQs) and BL Lacertae objects (BL Lacs) are also unclear. LAT observations of blazars and other AGNs are already helping to address these and other issues. Several in-depth spectral and/or multiwavelength studies of specific blazars (e.g., Abdo et al. 2009a, 2009b; Aharonian et al. 2009) and of non-blazar radio galaxies (Abdo et al. 2009c, 2009d, 2009e) have been performed.

The high sensitivity and nearly uniform sky coverage of the LAT make it a powerful tool for investigating the properties of large populations. The first list of bright AGNs detected by the LAT, the LAT Bright AGN Sample (LBAS; Abdo et al. 2009f) included bright AGNs at high Galactic latitude ($|b| > 10^\circ$) detected with high significance ($TS > 100$, or $\gtrsim 10\sigma$) during the first 3 months of scientific operation. This list comprised 58 FSRQs, 42 BL Lacs, two radio galaxies, and four AGNs of unknown type. Following the models used to describe the γ -ray spectra obtained with previous γ -ray observatories (e.g., Mattox et al. 1996), the early analysis reported in the LBAS was carried out by fitting the γ -ray spectra at energies above 200 MeV using a simple power-law (PL) model. This analysis revealed a fairly distinct spectral separation between FSRQs and BL Lacs, with FSRQs having significantly softer spectra. The division between the two classes was found to be at PL index $\Gamma \approx 2.2$. It has been suggested (Ghisellini et al. 2009) that this separation results from different radiative cooling of the electrons due to distinct accretion regimes in the two blazar classes. The γ -ray spectral properties and time-resolved multifrequency spectral energy distributions (SEDs) of LBAS sources were further investigated in Abdo et al. (2010a, 2010b).

Here, we report on a larger AGN sample detected after 11 months of scientific operations. The LAT first-year catalog (1FGL; Abdo et al. 2010c) contains a total of 1451 sources detected with $TS > 25$, and 1043 of these are at high Galactic latitudes ($|b| > 10^\circ$). We present a catalog of the high-latitude 1FGL sources that are associated with blazars and other AGNs. We refer to this as the First LAT AGN Catalog (1LAC). In addition to the 1LAC, we also provide, where possible, AGN associations for low-latitude LAT sources and AGN “affiliations”—candidate counterparts for which a quantitative association probability could not yet be computed—for unassociated high-latitude sources.

In Section 2, we describe the observations by the LAT and the analysis that led to the first-year catalog. In Section 3, we explain the method for associating γ -ray sources with AGN counterparts in a statistically meaningful way, present the results of this method, and describe the two schemes for classifying 1LAC AGNs. Section 4 provides a brief census of the 1LAC sample. Section 5 summarizes some of the properties of the 1LAC, including the γ -ray flux distribution, the γ -ray photon spectral index distribution, the γ -ray variability properties, the redshift distribution, and the γ -ray luminosity distribution. In Section 6, we describe the multiwavelength properties, from radio to TeV, of the 1LAC AGNs. We discuss the implications of the 1LAC results in Section 7 and conclude in Section 8.

In the following, we use a Λ CDM cosmology with values within 1σ of the *Wilkinson Microwave Anisotropy Probe*

(*WMAP*) results (Komatsu et al. 2009); in particular, we use $h = 0.71$, $\Omega_m = 0.27$, and $\Omega_\Lambda = 0.73$, where the Hubble constant $H_0 = 100 h \text{ km s}^{-1} \text{ Mpc}^{-1}$. We also define the radio spectral indices such that $S(\nu) \propto \nu^{-\alpha}$.

2. OBSERVATIONS WITH THE LARGE AREA TELESCOPE—ANALYSIS PROCEDURES

The γ -ray sources in the 1LAC are a subset of those in the 1FGL catalog; we summarize here the procedures used in producing the 1FGL catalog. The data were collected from 2008 August 4 to 2009 July 4, primarily with standard sky-survey observations. Only photons in the “Diffuse” event class⁷¹ with energies in the range 0.1–100 GeV were considered in this analysis in order to minimize contamination from misclassified cosmic rays (Atwood et al. 2009). This photon class is described in more detail (comparison to the class used in the Bright Source List (BSL; Abdo et al. 2009g) paper, systematic uncertainties) in Abdo et al. (2010c). To minimize contamination from γ -rays from the Earth’s limb, photons with incident directions greater than 105° from the local zenith were removed. In addition, time ranges during which the rocking angle of the LAT was greater than 43° were excluded from the data set because the bright limb of the Earth entered the field of view. This rocking angle limit removed only a small fraction of the data, with the excluded time intervals occurring during occasional 5 hr pointed observations at larger rocking angles (γ -ray burst (GRB) afterglow searches) and during even briefer intervals related to Sun avoidance during survey mode observations. A few minutes were excised around two bright GRBs (GRB 080916C and GRB 090510). The few time intervals with poor data quality, flagged as anything other than “Good” in the pointing/live time history (FT2) files, were also excluded. The resulting data set includes 245.6 days of live time. The standard *Fermi*-LAT *ScienceTools* software package⁷² (ver. v9r15p2) was used with the “P6_V3_DIFFUSE” set of instrument response functions.

The source-detection step made use of two wavelet algorithms, *mr_filter* (Starck & Pierre 1998) and *PGWAVE* (Ciprini et al. 2007), as well as tools that maximize a simplified likelihood function (*pointfind*; Abdo et al. 2010c) and that implement a minimum spanning tree algorithm (Campana et al. 2008). The intention in using a variety of algorithms to generate a list of “seed” positions for sources was to keep the source-detection step from being a limiting factor in the analysis. As described in the 1FGL paper, the algorithms were run independently for different energy ranges to find both soft- and hard-spectrum sources. Yet more seeds were introduced from the Roma-BZCAT (Massaro et al. 2009) and *WMAP* (Hinshaw et al. 2007; Giommi et al. 2009) catalogs if no nearby LAT seeds were present. The seeds were essentially candidate sources, and each was evaluated in detail in the subsequent steps of the 1FGL catalog analysis using the standard *gllike* tool to arrive at the final list.

The Galactic diffuse background model consistently employed throughout the analysis is the currently recommended version (*gll_iem_v02*), publicly released through the *Fermi* Science Support Center.⁷³ The isotropic background (including the γ -ray diffuse and residual instrumental backgrounds) model was derived from an overall fit of the diffuse component over

⁷¹ See http://fermi.gsfc.nasa.gov/ssc/data/analysis/documentation/Cicerone/Cicerone_Data/LAT_DP.html.

⁷² <http://fermi.gsfc.nasa.gov/ssc/data/analysis/documentation/Cicerone/>

⁷³ <http://fermi.gsfc.nasa.gov/ssc/data/access/lat/BackgroundModels.html>

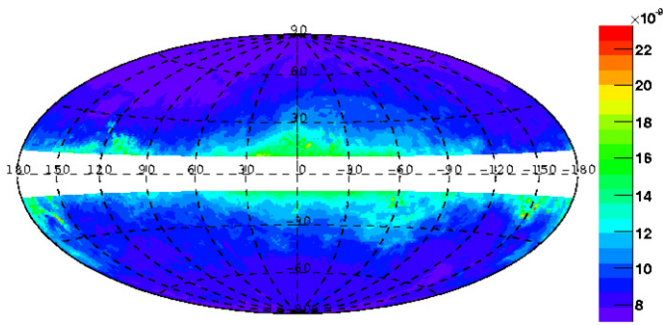


Figure 1. Point-source flux limit in units of photons $\text{cm}^{-2} \text{s}^{-1}$ for $E > 100$ MeV and photon spectral index $\Gamma = 2.2$ as a function of sky location (in Galactic coordinates).

(A color version of this figure is available in the online journal.)

the $|b| > 30^\circ$ sky. The Galactic diffuse model and corresponding isotropic spectrum are described in more detail in documentation available from the *Fermi* Science Support Center.

To evaluate the source significance, we used the maximum-likelihood algorithm implemented in *glike*. For the 1FGL catalog, a threshold of 25 was adopted for the test statistic (TS) from the *glike* likelihood analysis. Sources found to have $\text{TS} > 25$ were included in the 1FGL catalog. This corresponds approximately to a minimum significance of 4.1σ . Figure 1 displays a sky map, in Galactic coordinates, of the flux limit for photon spectral index $\Gamma = 2.2$ and $\text{TS} = 25$. The anisotropy (about a factor of 2) is due to the non-uniform Galactic diffuse background and non-uniform exposure (mostly arising from the passage of the *Fermi* satellite through the South Atlantic Anomaly). For soft sources, source confusion decreases the sensitivity to some extent (Abdo et al. 2010c).

The flux, photon spectral index (Γ), and TS of each source in the energy range 0.1–100 GeV were determined by analyzing regions of interest (ROI) typically 12° in radius. The model of the ROI used to fit the data was built taking into account all the sources detected within a given ROI. The fluxes in five bands (0.1–0.3, 0.3–1, 1–3, 3–10, and 10–100 GeV) were also evaluated, with the photon spectral index held fixed to the best fit over the whole interval. The energy flux was determined from these fluxes, resulting in better accuracy than would be obtained from the PL fitted function. For hard sources, the TSs provided by the 10–100 GeV band can be appreciable, which represents a notable difference with respect to EGRET, for which the acceptance dropped sharply in this range. The departure of the spectrum from a PL shape can be estimated via a simple χ^2 test on the five fluxes, referred to as the curvature index (Abdo et al. 2010c).

A TS map—a two-dimensional array of likelihood TS values evaluated at a finely spaced grid centered on the direction of a γ -ray source—was generated for each source using *pointfit*. TS values are determined for each grid position independently by maximum-likelihood fitting of a test point source. From the TS maps, elliptical fits to the 95% confidence source location contours were derived; for this, the decrease of the TS away from the maximum-likelihood position of a source is interpreted in terms of the χ^2 distribution with two degrees of freedom (the coordinates of the test source). The semimajor and semiminor axes of these ellipses were multiplied by 1.1 to account for systematic errors, which were evaluated by comparing the measured positions of bright sources to the known positions. The fiducial 95% error radius (the geometric mean of the semimajor and semiminor axes) is plotted against TS (derived as described

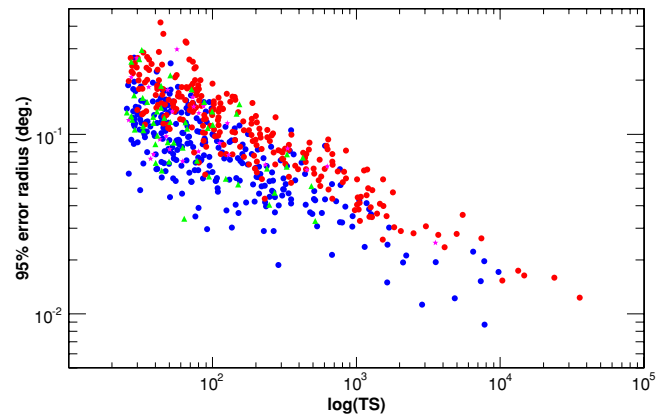


Figure 2. 95% error radius vs. TS for the 1LAC clean sample. Red circles: FSRQs, blue circles: BL Lacs, magenta stars: radio galaxies, green triangles: AGNs of unknown type.

(A color version of this figure is available in the online journal.)

above) in Figure 2. For each source, a simple variability index was derived from the χ^2 value of the monthly flux distribution with a 3% systematic uncertainty included.

3. SOURCE ASSOCIATIONS

Any procedure for associating a γ -ray source with a lower-energy counterpart necessarily relies on a spatial coincidence between the two. In the EGRET era, the problem of counterpart associations was made very difficult by the large γ -ray localization contours: the mean 95% error radius for sources in the eighteen-month EGRET sky survey (Fichtel et al. 1994) was $0^\circ.62$. The LAT offers a great improvement in source localization, resulting in a mean 95% error radius of $0^\circ.15$ for the high-latitude 1FGL sources. (For comparison to LBAS, which included only sources with $\text{TS} > 100$ and had a mean 95% error radius of $0^\circ.14$, the corresponding value for high-latitude 1FGL sources with $\text{TS} > 100$ is $0^\circ.09$.)

However, the LAT localization accuracy is not good enough to permit the determination of a lower-energy counterpart based only on positional coincidence. A firm counterpart identification is asserted only if the variability detected by the LAT corresponds with variability at longer wavelengths. In practice, such identifications are made only for a few sources (see Table 7 in Abdo et al. 2010c). For the rest, we use a method for finding associations between LAT sources and AGNs based on the calculation of association probabilities using a Bayesian approach implemented in the *gtsrcid* tool included in the *LAT ScienceTools* package and described in the BSL paper.

3.1. The Bayesian Association Method

The Bayesian method (de Ruiter et al. 1977; Sutherland & Saunders 1992), implemented by the *gtsrcid* tool in the *LAT ScienceTools*, is similar to that used by Mattox et al. (2001) to associate EGRET sources with flat-spectrum radio sources. An earlier version of the method was used and described in the BSL and LBAS papers. A more complete description is given in Abdo et al. (2010c), but we provide a basic summary here. The method uses Bayes's theorem to calculate the posterior probability that a source from a catalog of candidate counterparts is truly an emitter of γ -rays detected by the LAT. The significance of a spatial coincidence between a candidate counterpart from a catalog C and a LAT-detected γ -ray source is evaluated by examining the local density of counterparts from C in

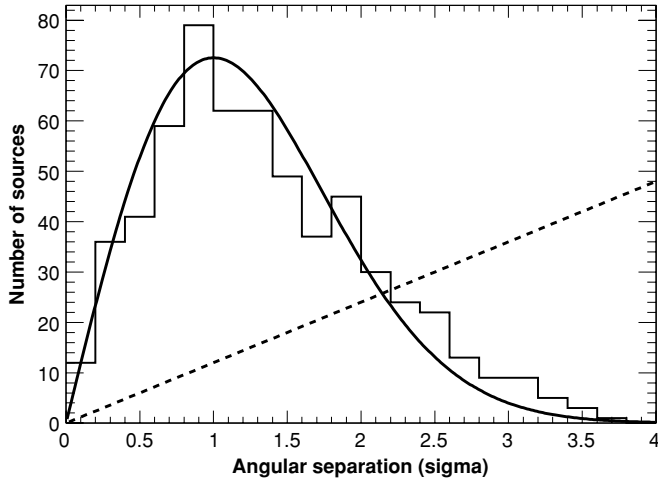


Figure 3. Normalized angular separation between the *Fermi*-LAT source position and that of the AGN counterpart. The solid curve corresponds to the expected distribution (χ^2 with two degrees of freedom) for real associations, the dashed one for purely accidental associations.

the vicinity of the LAT source. We can then estimate the likelihood that such a coincidence is due to random chance and establish whether the association is likely to be real. To each catalog C , we assign a prior probability, assumed for simplicity to be the same for all sources in C , for detection by the LAT. The prior probability for each catalog can be tuned to give the desired number of false positive associations for a given threshold on the posterior probability, above which the associations are considered reliable (see Section 3.2.1). We use a slightly different configuration of *gtsrcid* from that used for the 1FGL catalog. This allowed us to assign multiple associations to a single LAT source and to find associations with probabilities above 50% (compared with a threshold of 80% for the 1FGL associations).

Candidate counterparts were drawn from a number of source catalogs. The most important ones are the Combined Radio All-sky Targeted Eight GHz Survey (CRATES; Healey et al. 2007), the Candidate Gamma-Ray Blazar Survey (CGRaBS; Healey et al. 2008), and the Roma-BZCAT (Massaro et al. 2009). The CRATES catalog contains precise positions, 8.4 GHz flux densities, and radio spectral indices for over 11,000 flat-spectrum sources over the entire $|b| > 10^\circ$ sky. CGRaBS, a sample of the 1625 CRATES sources with radio and X-ray properties most similar to the blazars in the Third EGRET Catalog (3EG; Hartman et al. 1999), provides optical magnitudes, optical classifications, and spectroscopic redshifts. Roma-BZCAT is a list of blazars compiled based on an accurate examination of data from the literature and currently includes over 2800 sources, all observed at radio and optical frequencies and showing the observational characteristics of blazars. A complete list of the source catalogs used by *gtsrcid* can be found in Abdo et al. (2010c).

The same association method can be used at low latitudes; most of the candidate counterparts in this region are drawn from the VLBA Calibrator Survey (Beasley et al. 2002; Fomalont et al. 2003; Petrov et al. 2005, 2006, 2008; Kovalev et al. 2007). Indeed, Kovalev (2009) demonstrated the effectiveness of using VLBI data at finding radio associations for γ -ray sources, especially at low latitudes, and suggested incorporating VLBI catalogs into the *Fermi* association pipeline. The resulting associations are discussed in Section 3.2.2.

3.2. Association Results

3.2.1. The First LAT AGN Catalog (1LAC)

The 1LAC comprises all high-latitude ($|b| > 10^\circ$) sources with an association from *gtsrcid*; the full catalog includes 709 AGN associations for 671 distinct 1FGL sources and is shown in Table 1. An AGN is in the “high-confidence” sample if and only if its association probability P is at least 80%; this sample contains 663 AGNs. An AGN is in the “clean” sample if and only if it has $P \geq 80\%$, it is the sole AGN associated with the corresponding 1FGL γ -ray source (as indicated by an “S” in the last column of Table 1), and it is not “flagged” in the 1FGL catalog as exhibiting some problem or anomaly that casts doubt on its detection. This last criterion eliminates 12 sources from the clean sample: 1FGL J0217.8+7353, 1FGL J0258.0+2033, 1FGL J0407.5+0749, 1FGL J0433.5+3230, 1FGL J0539.4–0400, 1FGL J0540.9–0547, 1FGL J1424.5–7847, 1FGL J1702.7–6217, 1FGL J1727.9+5010, 1FGL J1938.2–3957, 1FGL J2212.9+0654, and 1FGL J2343.6+3437. All figures presented here are for the clean sample, which contains 599 AGNs.

The LBAS associations included one LAT source that was associated with two radio counterparts. Abdo et al. (2009f) noted that, as the number of LAT detections increased, source confusion was likely to be more of a problem, and indeed, the 1LAC includes 35 LAT sources that are associated with more than one AGN (for a total of 73 such associations). In cases of multiple associations, we list each counterpart separately in Table 1 and indicate them in the last column of the table.

The prefix “FRBA” in the column of AGN names refers to sources observed at 8.4 GHz as part of VLA program AH996 (“Finding and Rejecting Blazar Associations for *Fermi*-LAT γ -ray Sources”). The prefix “CLASS” refers to sources from the Cosmic Lens All-Sky Survey (Myers et al. 2003; Browne et al. 2003).

Figure 3 shows the distribution of normalized angular separations between the γ -ray sources and their AGN counterparts. The solid curve corresponds to the expected distribution for true associations while the dashed curve represents the expected distribution for purely random associations. These results provide confidence that most of the associations found are real.

The association probabilities can be used to estimate the number of false positive associations. In a sample of k sources with association probabilities P_i , the number of false positives is $N_{\text{false}} \approx \sum_{i=1}^k (1 - P_i)$. Among the 709 associations in the entire 1LAC, ~ 30 are false, but of the 663 sources in the high-confidence list, there are only ~ 14 false positives, and only ~ 11 of the sources in the clean sample are falsely associated. Additionally, there should be less than one false positive among the 363 most likely associations in the whole catalog.

3.2.2. Low-latitude AGNs

A simple extrapolation, based on the global density of 1LAC sources on the sky and the solid angle subtended by the Galactic plane region ($|b| < 10^\circ$), indicates that the LAT should be detecting ~ 150 AGNs at low Galactic latitudes. Diffuse radio emission, interloping Galactic point sources, and heavy optical extinction make the low-latitude sky a difficult region for AGN studies, and catalogs of AGNs and AGN candidates often avoid it partially or entirely. However, we are able to make associations with 51 low-latitude AGNs; these are presented in Table 2. Although the associations are considered valid, these sources have, in general, been studied much less uniformly and much less thoroughly than the high-latitude sources at virtually all

Table 1
First LAT AGN Catalog

1FGL Name	Associated AGN	R.A. ^a	Decl. ^a	Ang. Sep. ^b	θ_{95} ^c	Assoc. Prob.	Opt. Class	SED Class	z	F_{35} ^d	ΔF_{35} ^d	Γ ^d	$\Delta\Gamma$ ^d	σ ^d	Var. ^{?d}	Curv. ^{?d}	Note ^e	Clean ^{?f}
1FGL J0000.9–0745	CRATES J0001–0746	0.32512	–7.77417	0.089	0.153	0.96	BLL	LSP	...	1.0	0.0	2.41	0.20	5.6	N	N	S	Y
1FGL J0004.7–4737	PKS 0002–478	1.14867	–47.60517	0.032	0.153	0.99	FSRQ	LSP	0.880	0.8	0.3	2.56	0.17	6.6	N	N	S	Y
1FGL J0005.7+3815	B2 0003+38A	1.48825	38.33755	0.088	0.200	0.99	FSRQ	LSP	0.229	0.6	0.3	2.86	0.13	8.4	N	N	S	Y
1FGL J0008.3+1452	RX J0008.0+1450	2.02345	14.83982	0.072	0.143	0.70	AGN	...	0.045	0.8	0.2	2.00	0.21	4.7	N	N	S	N
1FGL J0008.9+0635	CRATES J0009+0628	2.26638	6.47256	0.119	0.117	0.93	BLL	LSP	...	0.8	0.0	2.28	0.22	5.0	N	N	S	Y
1FGL J0011.1+0050	CGRaBS J0011+0057	2.87667	0.96439	0.153	0.420	0.96	FSRQ	LSP	1.492	0.6	0.2	2.51	0.15	5.8	N	N	S	Y
1FGL J0013.1–3952	PKS 0010–401	3.24962	–39.90717	0.044	0.104	1.00	BLL	0.5	0.3	2.09	0.22	5.0	N	N	S	Y
1FGL J0013.7–5022	BZB J0014–5022	3.54675	–50.37575	0.075	0.135	1.00	BLL	HSP	...	0.6	0.2	2.23	0.22	4.4	N	N	S	Y
1FGL J0017.4–0510	CGRaBS J0017–0512	4.39925	–5.21158	0.040	0.079	1.00	FSRQ	LSP	0.227	1.5	0.3	2.60	0.07	20.2	Y	Y	S	Y
1FGL J0017.7–0019	PKS 0013–00	4.04621	–0.25347	0.390	0.315	0.57	FSRQ	LSP	1.574	0.5	0.2	2.88	0.18	5.1	N	N	S	N
1FGL J0018.6+2945	BZB J0018+2947	4.61562	29.79178	0.044	0.062	1.00	BLL	HSP	...	0.8	0.0	1.48	0.35	6.0	N	N	S	Y
1FGL J0019.3+2017	PKS 0017+200	4.90771	20.36267	0.100	0.177	0.99	BLL	LSP	...	0.7	0.2	2.38	0.15	5.9	N	N	S	Y
1FGL J0021.7–2556	CRATES J0021–2550	5.38563	–25.84703	0.110	0.116	0.86	BLL	ISP	...	0.9	0.2	1.96	0.17	7.3	N	N	S	Y
1FGL J0022.5+0607	PKS 0019+058	5.63517	6.13450	0.009	0.090	1.00	BLL	LSP	...	1.5	0.3	2.15	0.11	10.1	N	N	S	Y
1FGL J0023.0+4453	B3 0020+446	5.89767	44.94328	0.116	0.120	0.96	FSRQ	...	1.062	1.0	0.3	2.46	0.16	7.1	N	N	S	Y
1FGL J0029.9–4221	PKS 0027–426	7.57283	–42.41292	0.079	0.134	0.99	FSRQ	LSP	0.495	0.7	0.2	2.39	0.17	6.6	N	N	S	Y
1FGL J0033.5–1921	RBS 76	8.39292	–19.35944	0.006	0.053	1.00	BLL	HSP	0.610	2.8	0.4	1.89	0.08	17.5	N	N	S	Y
1FGL J0035.1+1516	RX J0035.2+1515	8.81125	15.25111	0.026	0.072	1.00	BLL	HSP	...	1.5	0.3	1.64	0.14	10.9	N	N	S	Y
1FGL J0038.4–2504	PKS 0035–252	9.56137	–24.98394	0.100	0.128	0.98	FSRQ	LSP	1.196	1.0	0.3	2.45	0.13	9.3	Y	N	S	Y
1FGL J0041.9+2318	PKS 0039+230	10.51896	23.33367	0.043	0.201	0.98	FSRQ	...	1.426	0.8	0.3	2.52	0.17	5.0	N	N	S	Y
1FGL J0045.3+2127	BZB J0045+2127	11.33042	21.46113	0.009	0.029	1.00	BLL	HSP	...	2.1	0.3	1.84	0.12	14.6	N	N	S	Y
1FGL J0047.3–2512	NGC 253	11.88806	–25.28812	0.097	0.176	1.00	AGN	...	0.001	0.7	0.2	2.15	0.17	6.2	N	N	S	Y
1FGL J0048.0–8412	PKS 0044–84	11.11192	–84.37781	0.191	0.229	0.96	FSRQ	...	1.032	1.3	0.0	2.69	0.17	5.6	N	N	S	Y
1FGL J0048.0–8412	PKS 0044–84	11.11192	–84.37781	0.191	0.229	0.96	FSRQ	...	1.032	1.3	0.0	2.69	0.17	5.6	N	N	S	Y
1FGL J0049.8–5738	PKS 0047–579	12.49775	–57.64067	0.019	0.070	1.00	FSRQ	LSP	1.797	0.6	0.2	2.42	0.16	8.3	N	N	S	Y
1FGL J0050.0–0446	PKS 0047–051	12.58971	–4.87242	0.119	0.184	0.98	FSRQ	...	0.920	0.7	0.3	2.34	0.18	5.6	N	N	S	Y
1FGL J0050.2+0235	PKS 0047+023	12.43017	2.61772	0.125	0.116	0.98	BLL	0.5	0.3	2.27	0.19	5.6	N	N	S	Y
1FGL J0050.6–0928	PKS 0048–09	12.67217	–9.48478	0.008	0.064	1.00	BLL	ISP	...	4.5	0.5	2.20	0.05	27.4	Y	N	S	Y
1FGL J0051.1–0649	PKS 0048–071	12.78421	–6.83394	0.016	0.105	1.00	FSRQ	LSP	1.975	1.9	0.3	2.36	0.09	12.8	Y	N	S	Y
1FGL J0058.0+3314	CRATES J0058+3311	14.63363	33.18811	0.104	0.110	0.95	BLL	...	1.371	1.7	0.3	2.33	0.11	10.5	Y	N	S	Y
1FGL J0058.4–3235	PKS 0055–328	14.50925	–32.57286	0.078	0.132	0.99	BLL	0.8	0.2	2.31	0.15	7.0	N	N	S	Y
1FGL J0100.2+0747	CRATES J0100+0745	15.08662	7.76428	0.048	0.047	0.97	Unknown	2.6	0.4	1.86	0.09	15.9	Y	N	S	Y
1FGL J0102.2+4223	CRATES J0102+4214	15.61313	42.23861	0.154	0.192	0.93	FSRQ	...	0.874	0.6	0.3	2.74	0.15	7.8	N	N	S	Y
1FGL J0104.4–2406	PKS 0102–245	16.24250	–24.27456	0.212	0.235	0.94	FSRQ	...	1.747	1.0	0.0	2.28	0.20	4.5	N	N	S	Y

Notes. The first page of the table is shown here.

^a J2000 coordinate of the associated AGN.

^b Angular separation, in degrees, between the 1FGL source position and the position of the associated AGN.

^c $\theta_{95} \equiv \sqrt{\theta_1\theta_2}$, where θ_1 and θ_2 are the semimajor and semiminor axes (in degrees) of the γ -ray 95% confidence region.

^d γ -ray properties from the 1FGL catalog. F_{35} : the photon flux in units of 10^{-9} photons $\text{cm}^{-2} \text{s}^{-1}$ for 1–100 GeV. ΔF_{35} : 1σ uncertainty on F_{35} in the same units. An entry of “0.0” indicates that the value of F_{35} is an upper limit. Γ : photon number spectral index. $\Delta\Gamma$: 1σ uncertainty on Γ . σ : detection significance. *Var.?*: “Y” indicates a probability < 1% that the source is steady. *Curv.?*: “Y” indicates a probability < 1% that a simple power law is a good fit to the spectrum.

^e The codes in this column provide information on multiple associations for a given 1FGL source. *S*: the 1FGL source is associated with exactly one AGN. *MM*: the 1FGL source is associated with at least two AGNs with high confidence ($P > 0.80$). *Mm*: the 1FGL source is associated with at least one AGN with high confidence ($P > 0.80$) and at least one AGN with lower confidence ($0.50 < P < 0.80$). *mm*: the 1FGL source is associated with at least two AGNs with lower confidence ($0.50 < P < 0.80$).

^f “Y” indicates that the source is in the clean sample, defined in Section 3.2.1.

(This table is available in its entirety in a machine-readable form in the online journal. A portion is shown here for guidance regarding its form and content.)

Table 2
AGN Associations for Low-latitude ($|b| < 10^\circ$) 1FGL Sources

1FGL Name	Associated AGN	R.A. ^a	Decl. ^a	Ang. Sep. ^b	θ_{95}^c	Assoc. Prob.	Opt. Class	SED Class	z	F_{35}^d	ΔF_{35}^d	Γ^d	$\Delta\Gamma^d$	σ^d	Var. ^{?d}	Curv. ^{?d}	Note ^e
1FGL J0035.9+5951	1ES 0033+595	8.96935	59.83461	0.020	0.028	1.00	BLL	3.2	0.5	1.95	0.20	13.0	N	N	S
1FGL J0046.8+5658	VCS1 J0047+5657	11.75179	56.96178	0.023	0.115	0.99	Unknown	2.3	0.4	2.27	0.17	7.5	N	N	S
1FGL J0102.8+5827	TXS 0059+581	15.69068	58.40309	0.064	0.079	0.99	FSRQ	...	0.644	4.2	0.5	2.38	0.13	16.3	Y	N	S
1FGL J0110.0+6806	4C +67.04	17.55364	68.09478	0.021	0.074	1.00	Unknown	2.0	0.5	2.35	0.21	8.5	N	N	S
1FGL J0254.2+5107	TXS 0250+508	43.49003	51.04902	0.089	0.114	0.99	Unknown	2.3	0.5	2.42	0.22	9.3	N	N	S
1FGL J0303.1+4711	4C +47.08	45.89684	47.27119	0.109	0.149	1.00	BLL	1.4	0.4	2.56	0.15	6.8	N	N	S
1FGL J0334.3+6536	RX J0333.9+6537	53.48641	65.61561	0.047	0.094	0.99	Unknown	1.3	0.5	2.19	0.22	4.1	N	N	S
1FGL J0419.0+3811	3C 111	64.58866	38.02661	0.214	0.255	0.87	AGN	...	0.049	1.5	0.5	2.61	0.22	4.3	N	N	S
1FGL J0423.8+4148	4C +41.11	65.98337	41.83409	0.025	0.031	1.00	Unknown	3.5	0.5	1.87	0.07	15.2	N	N	S
1FGL J0521.7+2114	RX J0521.7+2112	80.44152	21.21429	0.021	0.030	1.00	Unknown	5.5	0.6	1.94	0.18	21.1	N	N	S
1FGL J0533.0+4825	RX J0533.2+4823	83.31611	48.38134	0.054	0.110	1.00	FSRQ	...	1.162	2.3	0.4	2.43	0.35	10.3	Y	N	S
1FGL J0648.7-1740	TXS 0646-176	102.11874	-17.73484	0.092	0.107	0.99	FSRQ	...	1.232	2.5	0.5	2.47	0.15	10.8	Y	N	S
1FGL J0650.6-1635	PKS 0648-16	102.60242	-16.62770	0.075	0.151	0.98	Unknown	1.4	0.4	2.46	0.17	5.8	N	N	S
1FGL J0656.2-0321	OH -090	104.04634	-3.38522	0.034	0.077	1.00	Unknown	3.9	0.6	2.59	0.11	16.1	Y	N	S
1FGL J0702.2-1954	TXS 0700-197	105.67875	-19.85612	0.120	0.159	0.96	Unknown	1.7	0.4	1.92	0.16	4.2	N	N	S
1FGL J0721.4+0401	RX J0721.3+0406	110.34963	4.11228	0.083	0.125	0.98	Unknown	0.8	0.3	2.68	0.17	6.8	N	N	S
1FGL J0725.9-0053	PKS 0723-008	111.46100	-0.91571	0.043	0.077	1.00	BLL	...	0.128	0.6	0.3	2.30	0.08	6.1	N	N	S
1FGL J0730.3-1141	PKS 0727-11	112.57963	-11.68683	0.013	0.022	1.00	FSRQ	...	1.591	20.7	1.0	2.33	0.14	67.6	Y	N	S
1FGL J0754.4-1147	OI -187	118.61024	-11.78804	0.009	0.068	1.00	Unknown	1.9	0.4	2.10	0.13	10.5	N	N	S
1FGL J0825.8-2230	PKS 0823-223	126.50655	-22.50756	0.036	0.038	1.00	BLL	5.3	0.5	2.14	0.17	26.5	N	N	S
1FGL J0825.9-3216	PKS 0823-321	126.46405	-32.30645	0.036	0.110	1.00	Unknown	1.4	0.4	2.68	0.12	8.8	Y	N	S
1FGL J0827.9-3738	PKS B0826-373	127.01992	-37.51841	0.134	0.194	0.96	Unknown	1.7	0.6	2.68	0.17	5.7	Y	N	S
1FGL J0845.0-5459	PMN J0845-5458	131.26034	-54.96904	0.020	0.101	0.99	Unknown	1.9	0.4	2.24	0.17	7.7	N	N	S
1FGL J0849.6-3540	VCS2 J0849-3541	132.44010	-35.68369	0.023	0.152	0.98	Unknown	1.9	0.4	2.40	0.16	6.0	N	N	S
1FGL J0905.1-5736	PKS 0903-57	136.22158	-57.58494	0.041	0.068	1.00	FSRQ	...	0.695	1.8	0.4	2.36	0.18	7.5	Y	N	S
1FGL J1103.9-5355	PKS 1101-536	165.96759	-53.95019	0.019	0.041	1.00	Unknown	6.1	0.6	2.05	0.19	19.5	Y	N	S
1FGL J1122.9-6415	PMN J1123-6417	170.83090	-64.29339	0.059	0.088	0.99	Unknown	2.4	0.0	2.48	0.05	4.7	Y	N	S
1FGL J1307.3-6701	PKS 1304-668	197.07240	-67.11812	0.129	0.152	0.92	Unknown	1.4	0.5	2.57	0.09	5.3	Y	N	S
1FGL J1327.0-5257	PMN J1326-5256	201.70512	-52.93990	0.040	0.061	1.00	Unknown	5.2	0.6	2.33	0.11	19.2	Y	N	S
1FGL J1329.2-5605	PMN J1329-5608	202.25477	-56.13407	0.055	0.095	1.00	Unknown	4.1	0.6	2.56	0.15	15.1	Y	N	S
1FGL J1330.7-7006	PKS 1326-697	202.54615	-70.05363	0.072	0.114	0.99	Unknown	1.5	0.4	2.46	0.09	8.3	Y	N	S
1FGL J1400.9-5559	PMN J1400-5605	210.17407	-56.08210	0.092	0.122	0.99	Unknown	2.4	0.5	2.62	0.15	9.1	Y	N	S
1FGL J1514.1-4745	PMN J1514-4748	228.66677	-47.80829	0.106	0.124	0.98	Unknown	1.2	0.4	2.29	0.20	5.7	N	N	S
1FGL J1603.8-4903	PMN J1603-4904	240.96119	-49.06820	0.007	0.027	1.00	Unknown	13.4	1.1	2.12	0.14	26.8	N	Y	S
1FGL J1604.7-4443	PMN J1604-4441	241.12925	-44.69221	0.046	0.066	1.00	Unknown	7.7	0.8	2.46	0.04	22.2	Y	N	S

Notes. The first page of the table is shown here.

^a J2000 coordinate of the associated AGN.

^b Angular separation, in degrees, between the 1FGL source position and the position of the associated AGN.

^c $\theta_{95} \equiv \sqrt{\theta_1\theta_2}$, where θ_1 and θ_2 are the semimajor and semiminor axes (in degrees) of the γ -ray 95% confidence region.

^d γ -ray properties from the 1FGL catalog. F_{35} : the photon flux in units of 10^{-9} photons $\text{cm}^{-2} \text{s}^{-1}$ for 1–100 GeV. ΔF_{35} : 1σ uncertainty on F_{35} in the same units. An entry of “0.0” indicates that the value of F_{35} is an upper limit. Γ : photon number spectral index. $\Delta\Gamma$: 1σ uncertainty on Γ . σ : detection significance. *Var.?*: “Y” indicates a probability < 1% that the source is steady. *Curv.?*: “Y” indicates a probability < 1% that a simple power law is a good fit to the spectrum.

^e The codes in this column provide information on multiple associations for a given 1FGL source. *S*: the 1FGL source is associated with exactly one AGN. *MM*: the 1FGL source is associated with at least two AGNs with high confidence ($P > 0.80$). *Mm*: the 1FGL source is associated with at least one AGN with high confidence ($P > 0.80$) and at least one AGN with lower confidence ($0.50 < P < 0.80$). *mm*: the 1FGL source is associated with at least two AGNs with lower confidence ($0.50 < P < 0.80$).

(This table is available in its entirety in a machine-readable form in the online journal. A portion is shown here for guidance regarding its form and content.)

wavelengths, so we do not include them as part of the 1LAC in order to keep them from skewing any of our analyses of the overall γ -ray AGN population.

3.2.3. AGN “Affiliations”

For many of the 1FGL sources that are not formally associated with AGNs in the 1LAC, it is still possible to find nearby AGNs or AGN candidates for which reliable association probabilities cannot (yet) be computed but which show some indication that they may be the correct counterpart. For example, using the ASI Science Data Center (ASDC) multifrequency tools,⁷⁴ we performed a visual inspection of the vicinity of each unassociated 1FGL source. We considered objects inside the 95% error ellipse that showed hints of any blazar properties, such as coincident radio and X-ray emission and indications in the literature of variability, polarization, etc. For some sources, an optical spectrum, often from the Sloan Digital Sky Survey (SDSS; Adelman-McCarthy et al. 2008), was available, allowing us to classify them as BL Lacs or FSRQs. These sources have been evaluated by a method based on the known $\log N - \log S$ relationships of several types of AGNs (FSRQs, BL Lacs, flat-spectrum radio sources, etc.). We also used the figure of merit methodology developed by Sowards-Emmerd et al. (2003) and employed in the assembly of the LBAS source list; although we have discovered some problems with the calibration of the probabilities calculated by this approach, most of the resulting AGN associations seem likely to be legitimate. Both of these methods are being studied more carefully, but we list all of the unquantified correspondences (which we call “affiliations”) derived by them, along with some of the properties of the affiliated AGNs, in Table 3. In all, we find 109 AGN affiliations for 104 high-latitude LAT sources. We expect that future refinements to our association methods will allow us to turn many of these into true, quantitative associations.

3.3. Source Classification

3.3.1. Optical Classification

We classify each AGN according to its optical spectrum where available. Blazars are assigned optical classifications either as FSRQs or BL Lacs using the same scheme as for CGRaBS. In particular, following Stocke et al. (1991), Urry & Padovani (1995), and Marchã et al. (1996), we classify an object as a BL Lac if the equivalent width (EW) of the strongest optical emission line is $< 5 \text{ \AA}$, the optical spectrum shows a Ca II H/K break ratio $C < 0.4$, and the wavelength coverage of the spectrum satisfies $(\lambda_{\max} - \lambda_{\min})/\lambda_{\max} > 1.7$ in order to ensure that at least one strong emission line would have been detected if it were present. Although other definitions of BL Lac objects are sometimes applied (e.g., using the EWs of the [O II] $\lambda 3727$ and [O III] $\lambda 5007$ lines and/or different limits on C ; see, e.g., Landt et al. 2004), the definition used here can be applied over a large redshift range, with the caveat that high-redshift blazars may be classified as BL Lacs or FSRQs using different emission lines from those used for low-redshift objects. The classification of higher-redshift sources will preferentially use lines at shorter wavelengths (e.g., Ly α $\lambda 1216$ and C IV $\lambda 1549$) than for low-redshift sources (e.g., Mg II $\lambda 2798$ and H α $\lambda 6563$). A few of the 1LAC sources (e.g., radio galaxies such as Centaurus A and NGC 1275) that are not considered

blazars are listed in Table 1 simply as “AGNs”; these objects are discussed individually in greater detail in the following section. Sources for which no optical spectrum was available or for which the optical spectrum was of insufficient quality to determine the optical classification are listed as being of unknown type. Some redshifts and source types given in Table 1 may differ from previously published results (generally from the NASA Extragalactic Database⁷⁵ (NED), SDSS, and/or Véron-Cetty & Véron 2006). We thoroughly examined the data in the literature for accuracy and compatibility with our classification scheme, and our reevaluations of these results are reflected in Table 1.

A substantial number of the redshifts and optical classifications presented in Table 1 are from our own optical follow-up campaigns. Most of these come from spectroscopic observations conducted with the Marcario Low-Resolution Spectrograph (Hill et al. 1998) on the 9.2 m Hobby–Eberly Telescope (HET) at McDonald Observatory. Other facilities that have contributed optical results include the 3.6 m New Technology Telescope at La Silla, the 5 m Hale Telescope at Palomar, the 8.2 m Very Large Telescope at Paranal, and the 10 m Keck I Telescope at Mauna Kea. These spectra will be examined in greater detail in a subsequent paper (M. S. Shaw et al. 2010, in preparation). A number of the HET redshifts were confirmed, and new results were obtained, with the 3.6 m Telescopio Nazionale Galileo at La Palma. This work will also be detailed in an upcoming publication (S. Piranomonte & V. D’Elia 2010, in preparation).

3.3.2. SED Classification

The 1LAC blazars are also classified based on the peak of the synchrotron component of the broadband SED. In most cases, it is not possible to build a complete SED with simultaneous data, so we follow the scheme outlined by Abdo et al. (2010b). This scheme is an extension to all blazars of a standard classification system introduced by Padovani & Giommi (1995) for BL Lacs. We estimate the frequency of the synchrotron peak, ν_{peak}^S , using the broadband spectral indices α_{ro} (between 5 GHz and 5000 \AA) and α_{ox} (between 5000 \AA and 1 keV). This method uses an analytic relationship calibrated on the ν_{peak}^S values directly measured from the SEDs of the 48 sources studied by Abdo et al. (2010b), who confirm that these sources are representative of the global sample. This relationship is as follows:

$$\log \nu_{\text{peak}}^S = \begin{cases} 13.85 + 2.30X & \text{if } X < 0 \text{ and } Y < 0.3, \\ 13.15 + 6.58Y & \text{otherwise,} \end{cases}$$

where $X = 0.565 - 1.433\alpha_{\text{ro}} + 0.155\alpha_{\text{ox}}$ and $Y = 1.000 - 0.661\alpha_{\text{ro}} - 0.339\alpha_{\text{ox}}$. We use the estimated value of ν_{peak}^S to classify the source as a low-synchrotron-peaked, or LSP, blazar (for sources with $\nu_{\text{peak}}^S < 10^{14}$ Hz), an intermediate-synchrotron-peaked, or ISP, blazar (for 10^{14} Hz $< \nu_{\text{peak}}^S < 10^{15}$ Hz), or a high-synchrotron-peaked, or HSP, blazar (for sources with $\nu_{\text{peak}}^S > 10^{15}$ Hz). Figure 4 displays the locations of BL Lacs from the clean sample in the $(\alpha_{\text{ox}}, \alpha_{\text{ro}})$ plane. The data used for calculating the broadband spectral indices were obtained mainly from *Swift*-XRT observations, the NRAO VLA Sky Survey (NVSS; Condon et al. 1998), and the USNO-B1.0 catalog (Monet et al. 2003) and completed with data from Roma-BZCAT. It is evident from Figure 4 that γ -ray-selected BL Lacs cover the region in this plane typically occupied by blazars selected at other frequencies, such as radio and X-rays.

⁷⁴ These tools are available from the ASDC Web site at the following URLs: <http://tools.asdc.asi.it/> and <http://www.asdc.asi.it/>.

⁷⁵ <http://nedwww.ipac.caltech.edu/>

Table 3
AGN “Affiliations” for 1FGL Sources

1FGL Name	Associated AGN	R.A. ^a	Decl. ^a	Ang. Sep. ^b	θ_{95} ^c	Opt. Class	SED Class	z	F_{35} ^d	ΔF_{35} ^d	Γ ^d	$\Delta\Gamma$ ^d	σ ^d	Var.? ^d	Curv.? ^d
1FGL J0016.6+1706	CRATES J0015+1700	3.91662	17.01128	0.247	0.193	FSRQ	...	1.716	0.5	0.3	2.57	0.20	4.7	N	N
1FGL J0022.2–1850	1RXS J002209.2–185333	5.53800	–18.89248	0.054	0.111	Unknown	HSP	...	1.1	0.3	1.56	0.14	9.4	N	N
1FGL J0024.6+0346	CLASS J0024+0349	6.18841	3.81766	0.045	0.099	FSRQ	...	0.545	0.9	0.3	2.33	0.12	8.9	Y	N
1FGL J0038.0+1236	FRBA J0037+1238	9.46180	12.63863	0.060	0.098	Unknown	0.7	0.3	2.30	0.18	7.0	N	N
1FGL J0043.6+3424	CLASS J0043+3426	10.95352	34.44059	0.042	0.063	Unknown	1.7	0.3	2.09	0.10	12.4	N	N
1FGL J0048.0+2232	CLASS J0048+2235	12.01092	22.59006	0.050	0.089	FSRQ	...	1.156	1.5	0.3	2.39	0.09	13.5	Y	N
1FGL J0051.4–6242	RBS 119	12.81952	–62.70120	0.025	0.067	Unknown	1.8	0.3	1.68	0.12	12.0	N	N
1FGL J0054.9–2455	FRBA J0054–2455	13.69480	–24.92519	0.035	0.078	Unknown	HSP	...	0.7	0.2	1.95	0.22	5.2	N	N
1FGL J0110.0–4023	RBS 158	17.48483	–40.34794	0.051	0.079	Unknown	HSP	...	0.8	0.0	1.34	0.32	4.2	N	N
1FGL J0115.7+0357	CLASS J0115+0356	18.91880	3.94536	0.021	0.076	BLL	1.4	0.3	2.07	0.15	8.7	N	N
1FGL J0124.6–0616	AT20G J0124–0625	21.21033	–6.41722	0.146	0.163	BLL	0.6	0.3	2.20	0.17	5.5	N	N
1FGL J0134.4+2632	RX J0134.4+2638	23.61779	26.64588	0.103	0.155	Unknown	HSP	...	0.9	0.3	2.26	0.16	6.6	N	N
1FGL J0157.0–5259	RBS 259	29.23839	–53.03285	0.049	0.100	Unknown	0.8	0.2	1.85	0.20	7.6	N	N
1FGL J0217.9–6630	CRATES J0216–6636	34.21187	–66.61169	0.153	0.108	BLL	0.9	0.3	1.94	0.17	7.0	N	N
1FGL J0226.3+0937	FRBA J0226+0937	36.55704	9.62398	0.034	0.133	Unknown	0.9	0.3	1.99	0.17	6.4	N	N
1FGL J0256.9+2920	FRBA J0256+2924	44.22623	29.41500	0.066	0.105	AGN	...	0.190	0.6	0.3	1.69	0.30	4.2	N	N
1FGL J0315.6–5109	AT20G J0314–5104	48.60733	–51.07547	0.218	0.314	BLL	0.7	0.3	2.60	0.20	6.0	N	N
1FGL J0318.1+0254	CLASS J0317+0248	49.44965	2.81183	0.135	0.165	FSRQ	...	0.748	0.8	0.3	2.20	0.18	4.1	N	N
1FGL J0333.7+2919	FRBA J0333+2916	53.45420	29.27542	0.057	0.044	Unknown	ISP	...	1.0	0.3	1.57	0.21	6.5	N	N
1FGL J0342.2+3859	CLASS J0342+3859	55.56779	38.98507	0.012	0.102	FSRQ	...	0.945	1.1	0.0	2.17	0.29	4.7	N	N
1FGL J0401.3–3152	PKS 0400–319	60.58863	–31.79053	0.232	0.155	FSRQ	...	1.288	0.6	0.3	2.37	0.17	6.0	N	N
1FGL J0445.2–6008	AT20G J0445–6015	71.25667	–60.25006	0.105	0.169	AGN	...	0.097	0.8	0.3	1.98	0.29	4.6	N	N
1FGL J0506.9–5435	RBS 621	76.74127	–54.58422	0.014	0.058	Unknown	1.0	0.0	1.42	0.31	6.7	N	N
1FGL J0515.2+7355	CLASS J0516+7351	79.12993	73.85240	0.115	0.167	BLL	...	0.249	0.7	0.3	2.04	0.25	5.0	N	N
1FGL J0515.9+1528	CLASS J0515+1527	78.94732	15.45461	0.032	0.069	BLL	1.2	0.4	2.01	0.18	8.4	N	N
1FGL J0517.6+0857	CLASS J0517+0858	79.41693	8.97661	0.025	0.109	FSRQ	...	0.328	1.1	0.4	2.58	0.12	8.5	N	N
1FGL J0537.7–5717	PKS 0541–834	84.45325	–57.30806	0.014	0.073	FSRQ	1.0	0.0	1.81	0.42	6.1	N	N
1FGL J0538.4–3910	1RXS J053810.0–390839	84.54304	–39.14562	0.059	0.077	Unknown	HSP	...	1.3	0.3	2.16	0.12	9.0	N	N
1FGL J0541.9–0204	CRATES J0541–0154	85.47800	–2.07800	0.011	0.075	Unknown	1.5	0.6	2.30	0.12	6.0	N	N
1FGL J0600.5–2006	CRATES J0601–2004	90.47012	–20.07922	0.303	0.260	FSRQ	...	1.216	0.7	0.3	2.41	0.15	6.2	N	N
1FGL J0603.0–4012	1WGA J0602.8–4018	90.71162	–40.31417	0.109	0.143	Unknown	1.2	0.3	2.25	0.16	5.6	N	N
1FGL J0604.2–4817	IES 0602–482	91.03918	–48.29059	0.011	0.062	Unknown	1.1	0.3	2.12	0.16	7.2	N	N
1FGL J0605.1+0005	CLASS J0604+0000	91.24341	0.01204	0.100	0.129	Unknown	1.1	0.3	1.94	0.23	5.4	N	N
1FGL J0608.1–0630	CRATES J0609–0615	92.41654	–6.25161	0.457	0.306	Unknown	1.8	0.5	2.54	0.10	8.1	N	Y
1FGL J0609.3–0244	NVSS J060915–024754	92.31252	–2.79840	0.065	0.096	Unknown	HSP	...	1.2	0.3	2.02	0.20	5.9	N	N
1FGL J0622.2+3751	CLASS J0621+3750	95.49016	37.84916	0.058	0.131	Unknown	2.0	0.4	2.36	0.09	10.2	N	Y
1FGL J0706.5+3744	CLASS J0706+3744	106.63207	37.74344	0.007	0.063	BLL	HSP	...	1.1	0.3	2.01	0.15	8.2	N	N
1FGL J0707.3+7742	FRBA J0706+7741	106.71367	77.69349	0.027	0.088	Unknown	1.4	0.3	2.30	0.13	10.5	N	N
1FGL J0723.6+2908	CLASS J0723+2859	110.97850	28.99163	0.165	0.215	Unknown	0.9	0.3	2.10	0.17	5.8	Y	N

Notes. The first page of the table is shown here.

^a J2000 coordinate of the associated AGN.

^b Angular separation, in degrees, between the 1FGL source position and the position of the associated AGN.

^c $\theta_{95} \equiv \sqrt{\theta_1 \theta_2}$, where θ_1 and θ_2 are the semimajor and semiminor axes (in degrees) of the γ -ray 95% confidence region.

^d γ -ray properties from the 1FGL catalog. F_{35} : the photon flux in units of 10^{-9} photons $\text{cm}^{-2} \text{s}^{-1}$ for 1–100 GeV. ΔF_{35} : 1σ uncertainty on F_{35} in the same units. An entry of “0.0” indicates that the value of F_{35} is an upper limit. Γ : photon number spectral index. $\Delta\Gamma$: 1σ uncertainty on Γ . σ : Detection significance. *Var.?*: “Y” indicates a probability < 1% that the source is steady. *Curv.?*: “Y” indicates a probability < 1% that a simple power law is a good fit to the spectrum.

(This table is available in its entirety in a machine-readable form in the online journal. A portion is shown here for guidance regarding its form and content.)

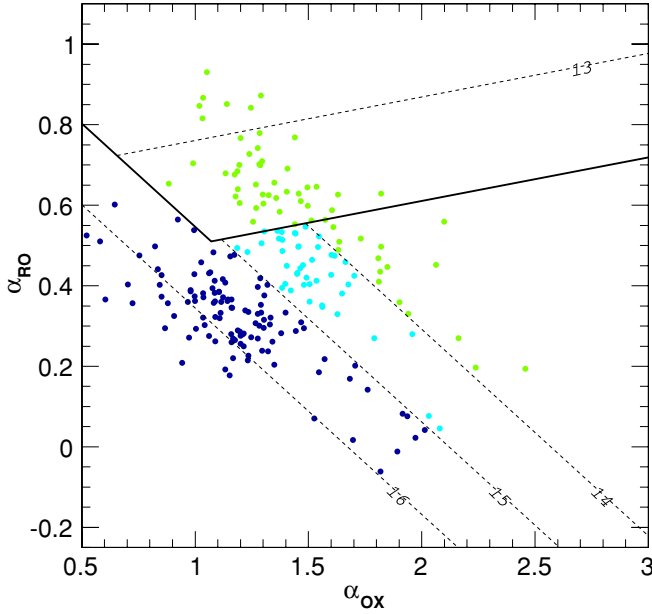


Figure 4. Loci of the BL Lacs in the clean sample in the $(\alpha_{\text{ox}}, \alpha_{\text{ro}})$ plane (green: LSP-BL Lacs, cyan: ISP-BL Lacs, blue: HSP-BL Lacs). The $\{X < 0\} \cap \{Y < 0.3\}$ region (see Section 3.3.2) is delineated in solid black. The dashed lines correspond to $\log(\nu_{\text{peak}}^{\text{S}} \text{ (Hz)}) = 13, 14, 15,$ and 16 .

(A color version of this figure is available in the online journal.)

Figure 5 shows the soft X-ray flux versus the radio flux density at 1.4 GHz for the blazars in the clean sample. As in LBAS, we find that the broadband properties of the 1LAC blazars are consistent with those of the parent populations of BL Lacs and FSRQs.

We note that the SED classification method assumes that the optical and X-ray fluxes come from the nonthermal emission without contamination from the disk or accretion. For blazars in which the thermal components are non-negligible, this method may lead to a significant overestimation of the position of $\nu_{\text{peak}}^{\text{S}}$.

4. 1LAC POPULATION CENSUS

Table 4 summarizes the breakdown of 1LAC sources by type for the full catalog, the high-confidence sample, and the clean sample. The fraction of blazars (BL Lacs and FSRQs) that are BL Lacs (51% for the high-confidence sample) is higher than in LBAS (39%) and much higher than in 3EG (21%). As discussed in Section 5.1, for a given significance, BL Lacs can typically be detected by the LAT at a lower flux than can FSRQs, which have softer spectra. The presence of one ISP-FSRQ and one HSP-FSRQ warrants comment. The existence of higher-peaked FSRQs has not been definitively proven yet. We stress that the SED classifications given here can be improved with simultaneous data, and for blazars in which the thermal components are non-negligible, it may lead to a significant overestimation of $\nu_{\text{peak}}^{\text{S}}$. Moreover, the average error on the γ -ray spectral indices is 0.14, so it cannot be ruled out that the candidate ISP- and HSP-FSRQs actually have lower synchrotron peak frequencies and are consistent with the rest of the sample.

The distributions of synchrotron peak positions $\nu_{\text{peak}}^{\text{S}}$ are shown in Figure 6. The large fraction of HSPs in the 1LAC sample contrasts sharply with the results from the 3EG sample, in which most BL Lacs (>80%) were of the LSP type. This is also a consequence of the dependence of the limiting flux on the spectral index (see Section 5.1 and Figure 9).

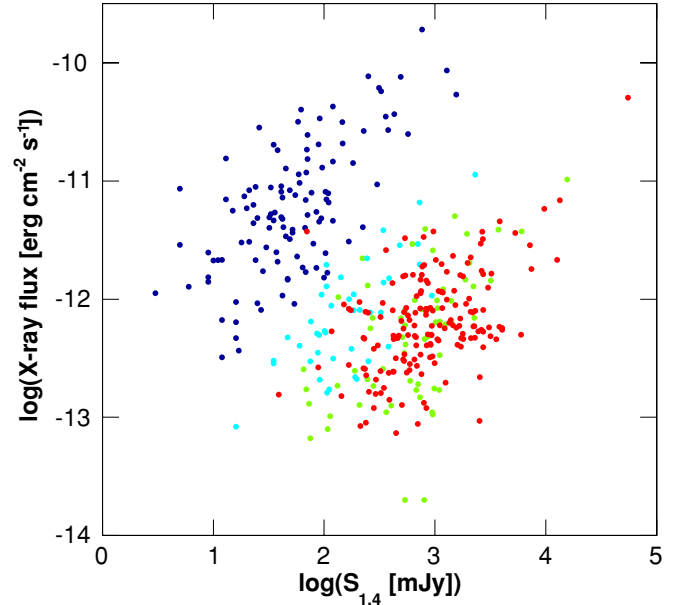


Figure 5. Soft X-ray (0.1–2.4 keV) flux vs. radio flux density at 1.4 GHz for the clean sample (red: FSRQs, green: LSP-BL Lacs, cyan: ISP-BL Lacs, blue: HSP-BL Lacs).

(A color version of this figure is available in the online journal.)

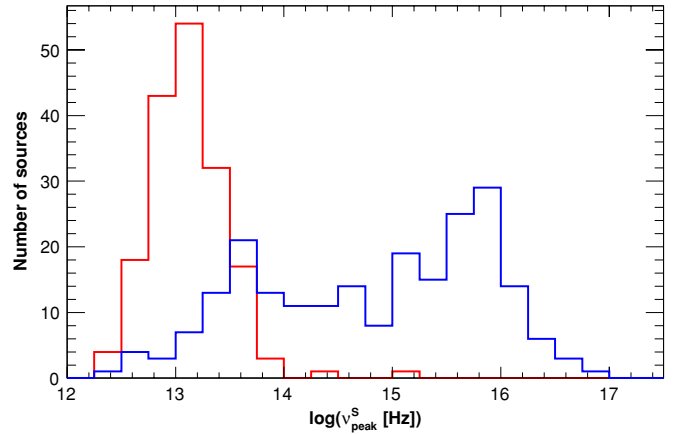


Figure 6. Distributions of $\nu_{\text{peak}}^{\text{S}}$ for FSRQs (red) and BL Lacs (blue) in the clean sample.

(A color version of this figure is available in the online journal.)

Table 4
Census of 1LAC Sources

AGN Type	Number of AGNs in		
	Entire 1LAC Sample	High-confidence Sample ^a	Clean Sample ^a
All	709	663	599
FSRQ	296	281	248
...LSP	189	185	171
...ISP	3	2	1
...HSP	2	2	1
BL Lac	300	291	275
...LSP	69	67	62
...ISP	46	44	44
...HSP	118	117	113
Other AGN	41	30	26
Unknown	72	61	50

Note.

^a See Section 3.2.1 for the definitions of these samples.

4.1. Misaligned AGNs

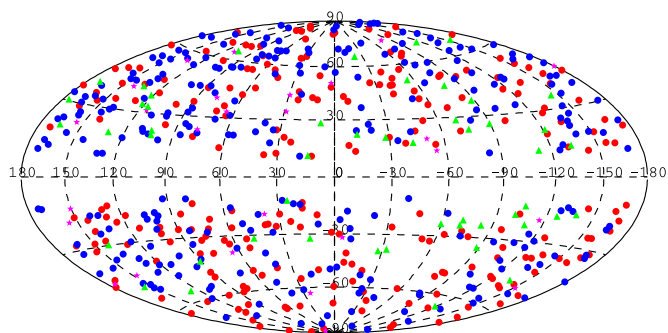


Figure 7. Locations of the sources in the clean sample. Red circles: FSRQs, blue circles: BL Lacs, magenta stars: radio galaxies, green triangles: AGNs of unknown type.

(A color version of this figure is available in the online journal.)

Figure 7 shows the sky locations of the sources in the clean sample. It is clear that the distribution is not isotropic; there are more sources in the northern galactic hemisphere than in the southern one. The Galactic latitude distributions for FSRQs and BL Lacs are shown in Figure 8 along with the corresponding distributions for the Roma-BZCAT and 1FGL. The anisotropy, defined as $(N_+ - N_-)/(N_+ + N_-)$, where N_+ and N_- are the number of sources at $b > 10^\circ$ and $b < -10^\circ$, respectively, amounts to -4% ($N_+ = 135$, $N_- = 146$) for high-confidence FSRQs and 18% ($N_+ = 171$, $N_- = 120$) for high-confidence BL Lacs. For comparison, the sources in the 1FGL catalog are very evenly distributed over the two hemispheres: 559 sources with $b > 10^\circ$ and 550 sources with $b < -10^\circ$.

We have performed a search for possible associations between LAT detections and non-blazar radio-loud sources (i.e., radio galaxies and steep-spectrum radio quasars) using three main low-frequency surveys: the 3CR catalog (Bennett 1962; Spinrad et al. 1985); its revised version, the 3CRR catalog (Laing et al. 1983); and the Molonglo Southern 4 Jy Sample (MS4; Burgess & Hunstead 2006). These catalogs are flux limited (3CR: 9 Jy; 3CRR: 10.9 Jy; MS4: 4 Jy) and cover large portions of the northern (3CR, 3CRR) and southern (MS4) sky. In addition, because these surveys were conducted at low frequencies (3CR and 3CRR: 178 MHz; MS4: 408 MHz), they detect radio sources primarily in the relatively steep part of the synchrotron emission spectrum, which is generally associated with the extended lobes rather than the compact cores.

As a result, these catalogs are particularly appropriate for selecting AGNs with jets misaligned with respect to the line of sight. The sources in these surveys exhibit radio maps with resolved and possibly symmetrical structures and steep radio spectra ($\alpha_r > 0.5$). At higher radio frequencies, the distinction between blazars and misaligned radio sources can be less sharp as the compact core emission (characterized by a flat or inverted spectrum) can emerge and dominate the optically thin synchrotron emission from extended regions. In addition to Cen A (Abdo et al. 2009f), NGC 1275 (Abdo et al. 2009c, 2009f), and M87 (Abdo et al. 2009d), this search has resulted in the classification of five LAT-detected sources as misaligned AGNs: three radio galaxies (NGC 1218 = 3C 78, PKS 0625–35, and NGC 6251) and two steep-spectrum radio quasars (3C 207

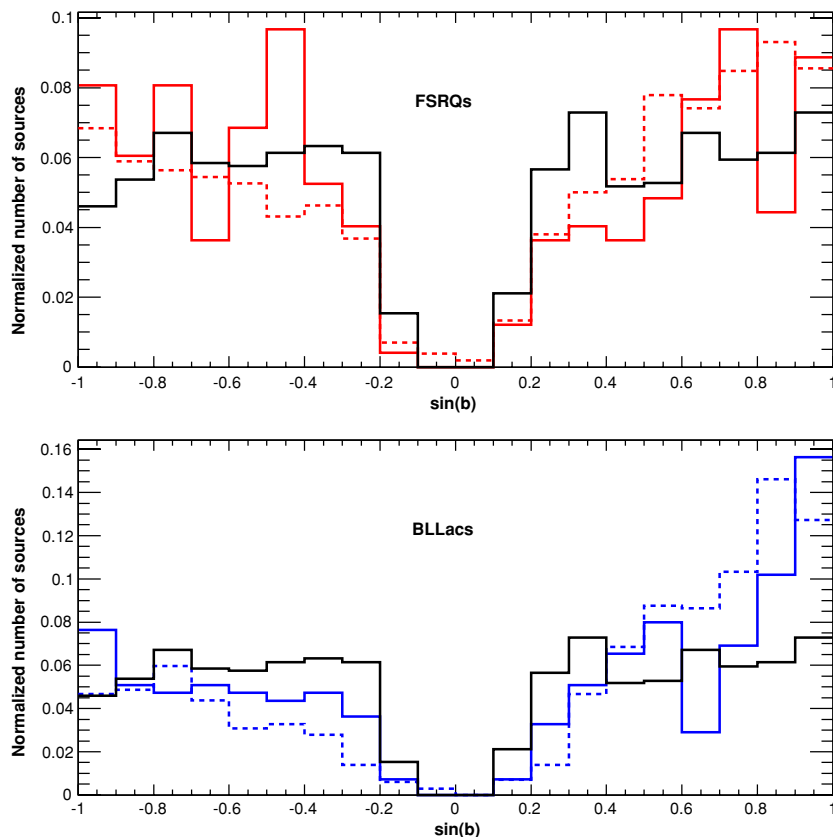


Figure 8. Galactic latitude distributions of FSRQs (top) and BL Lacs (bottom) in the clean sample (solid red). The corresponding distributions for the Roma-BZCAT blazars (dashed red) and the 1FGL sources with $|b| > 10^\circ$ (solid black) are shown for comparison.

(A color version of this figure is available in the online journal.)

and 3C 380). The properties of steep-spectrum radio sources detected by the LAT will be examined further in a future publication (A. A. Abdo et al. 2010a, in preparation).

Five other 1LAC sources have flat or nearly flat radio spectra between ~ 1 GHz and 8.4 GHz and strong emission lines in their optical spectra, but these lines are narrow, in contrast to the broad emission features seen in FSRQs. Under the standard AGN unification paradigm (Antonucci 1993; Urry & Padovani 1995), this is interpreted as an indication that the jet axis is at a larger angle to the line of sight than for a typical FSRQ. This causes the broad-line region to be obscured by the dusty torus surrounding the central black hole, leaving only narrow lines in the resulting spectrum, and such sources are commonly described as narrow-line radio galaxies (NLRGs). The NLRGs appearing in the 1LAC are 4C +15.05, PKS 1106+023, CGRaBS J1330+5202, 4C +15.54, and CGRaBS J2250–2806.

4.2. Radio-quiet AGNs

Radio-loud narrow-line Seyfert 1 galaxies have previously been shown to be emitters of γ -rays (PMN J0948+0022, Abdo et al. 2009e; B2 0321+33B, PKS 1502+036, and PKS 2004–447, Abdo et al. 2009h). There is a tentative indication that the LAT may have detected radio-quiet Seyfert galaxies or radio-quiet quasars. Indeed, the 1LAC includes 10 radio-quiet AGN associations for LAT sources (RX J0008.0+1450, 1WGA 0405.6–1313, CXOMP J084045.2+131617, CXOMP J084054.3+131456, SDSS J112042.47+071311.5, CXOCY J113008.8–144737, SDSS J125257.95+525925.6, SDSS J143847.94+371342.7, QSO J150255.20–415430.2, and SDSS J155140.52+085226.1). In seven of these cases, however, the LAT source is also associated with at least one radio-loud AGN (a blazar or a radio galaxy) with higher probability, so the γ -ray emission is very likely attributable to the radio-loud source. 1FGL J1120.4+0710 is associated both with SDSS J112042.47+071311.5, a Seyfert 1 galaxy, and CRATES J1120+0704; although the latter is of unknown type, it is a radio-loud source (probably a blazar) and is the likely source of the γ -ray emission. The remaining two radio-quiet 1LAC associations, RX J0008.0+1450 and QSO J150255.20–415430.2, have association probabilities under 70%. More data and further study will be necessary to establish whether these objects are indeed emitters of γ -rays.

4.3. Notes on Individual Sources

Here, we provide additional information on sources with unusual or interesting properties. Among these are the sources classified optically as “AGNs of other types” (except those mentioned as misaligned AGNs or radio-quiet AGNs in the previous two subsections).

1FGL J0047.3–2512. This source is associated with NGC 253, a starburst galaxy previously detected by the LAT (Abdo et al. 2010d). It has for some time been designated a Seyfert galaxy in the Véron catalog (most recently in Véron-Cetty & Véron 2006), but this classification is questionable (see, e.g., Förster Schreiber et al. 2004). Engelbracht et al. (1998) determine that the LINER-like characteristics of NGC 253 can be attributed entirely to starburst activity.

1FGL J0339.1–1734. This source is associated with PKS 0336–177. The optical spectrum from the final release of 6dFGS (Jones et al. 2004, 2009) shows no strong emission lines, but the Ca II H/K break is too large to satisfy the criterion $C < 0.4$ for classification as a BL Lac.

1FGL J0405.6–1309. This source is associated with PKS 0403–13 (an FSRQ) and with 1WGA J0405.6–1313. Mason et al. (2000) classify the latter simply as an “AGN” but provide the redshift ($z = 0.226$) based on the detection of three emission lines, including broad H β . The source is radio-faint (undetected in NVSS) and likely a Seyfert galaxy.

1FGL J0627.3–3530. This source is associated with PKS 0625–35, which shows an FR-I radio morphology. However, Wills et al. (2004) suggest a BL Lac classification for this object. They also note that, in contrast with the other genuine FR-I sources in their sample, an adequate fit to the optical continuum of PKS 0625–35 required a nonthermal PL in addition to emission from the host galaxy.

1FGL J0645.5+6033. This source is associated with BZU J0645+6024, a source with broad emission lines but with a steep radio spectrum.

1FGL J0923.2+4121. This source is associated with B3 0920+416, which Falco et al. (1998) characterize as a late-type galaxy.

1FGL J0956.5+6938. This source is associated with M82, also known as NGC 3034 and 3C 231, a starburst galaxy previously detected by the LAT (Abdo et al. 2010d). There are indications that it may host a weak AGN (Wills et al. 1999), though this is not confirmed.

1FGL J1202.9+6032. This source is associated with CRATES J1203+6031. Falco et al. (1998) describes it as an early-type galaxy; it is listed in NED as a LINER.

1FGL J1305.4–4928. This source is associated with NGC 4945, which is both a Seyfert 2 galaxy and a starburst galaxy. It is the third starburst galaxy seen by the LAT, and its detection is reported here for the first time. A more careful analysis will be required to determine whether the γ -ray emission comes from the starburst activity or the AGN or both.

1FGL J1307.0–4030. This source is associated with ESO 323-G77, a nearby ($z = 0.015$) Seyfert 1.2 galaxy detected also by *Swift*-BAT (see, e.g., Ajello et al. 2009 and Section 6.3). It exhibits a hard X-ray spectrum typical of radio-quiet AGN. There is little indication that the source is a starburst galaxy. If the γ -ray emission from 1FGL J1307.0–4030 truly comes from ESO 323-G77, then its origin might be connected to the central AGN; this hypothesis is being examined more closely.

1FGL J1641.0+1143. This source is associated with CRATES J1640+1144, described by Mitton et al. (1977) only as a “galaxy.” Since the authors do classify other sources as BL Lacs, it seems likely that CRATES J1640+1144 is not a BL Lac.

1FGL J1642.5+3947. In LBAS, the brightest LAT source in this region, 0FGL J1641.4+3939, was associated with the faint blazar CLASS J1641+3935 with low probability. However, with an improved LAT localization, the γ -ray source is now affiliated with the well-known, strongly variable superluminal blazar 3C 345 (Unwin et al. 1983; Zensus et al. 1995). This is generally a difficult region for a confident association (see Reyes & Cheung 2009) due to the known concentration of AGNs in the field (Arp 1997), including the $z = 1.666$ superluminal quasar NRAO 512 (Lister et al. 2009). All the sources mentioned are within the localization of EGR J1642+3940 (Casandjian & Grenier 2008).

1FGL J1647.4+4948. This source is associated with CGRaBS J1647+4950, which Falco et al. (1998) characterize as a late-type galaxy.

1FGL J1724.0+4002. This source is associated with B2 1722+40. The optical spectrum from Vermeulen et al. (1996)

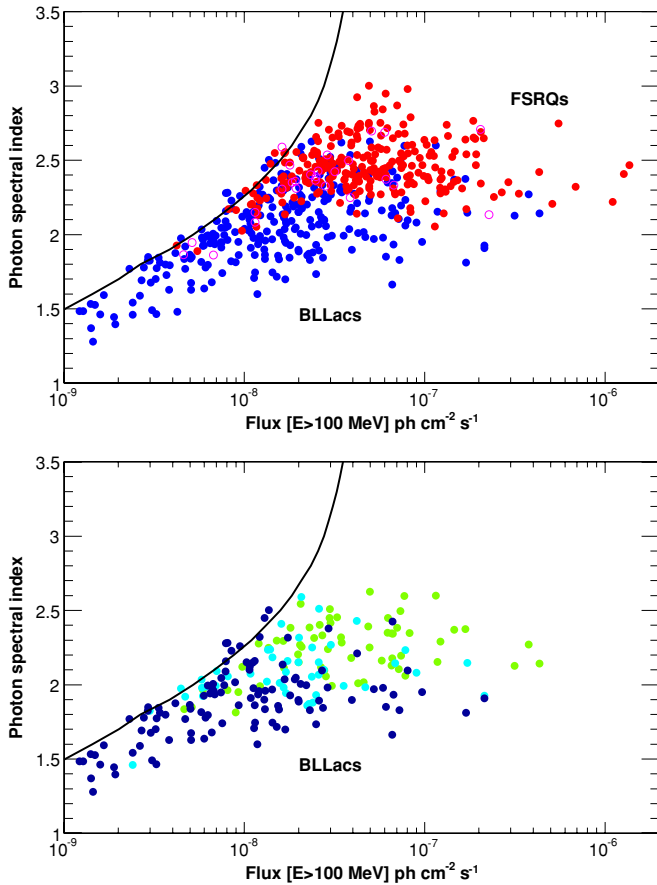


Figure 9. Top: flux ($E > 100$ MeV) vs. photon spectral index for the sources in the clean sample. The solid curves represent the $TS = 25$ flux limit estimated for a Galactic position $(l, b) = (40^\circ, 40^\circ)$. Bottom: the same for the different BL Lac subclasses. Green: LSPs, cyan: ISPs, blue: HSPs.

(A color version of this figure is available in the online journal.)

shows narrow forbidden emission lines and absorption features from the host galaxy.

1FGL J1756.6+5524. This source is associated with BZB J1756+5522 (a BL Lac) and with CRATES J1757+5523. The optical spectrum of the latter from Caccianiga et al. (2002) shows strong absorption features from the host galaxy and a large Ca II H/K break. The authors describe it as a passive elliptical galaxy.

1FGL J2008.6–0419. This source is associated with 3C 407, a source with broad emission lines but with a fairly steep radio spectrum.

1FGL J2038.1+6552. This source is associated with NGC 6951, which has been classified as a Seyfert 2 galaxy and a LINER. It also has a circumnuclear ring of starburst activity.

1FGL J2204.6+0442. This source is associated with 4C +04.77, which is often called a BL Lac in the literature. A spectrum from Véron-Cetty & Véron (1993), however, shows that the object exhibits strong broad emission lines characteristic of Seyfert 1 galaxies.

5. PROPERTIES OF THE 1LAC SOURCES

5.1. Flux and Photon Spectral Index Distributions

As demonstrated in Abdo et al. (2010a), many bright LAT blazars (most notably FSRQs and some LSP-BL Lacs) exhibit breaks in their γ -ray spectra. Despite this caveat, determining a photon spectral index fitted over the whole band is useful as it

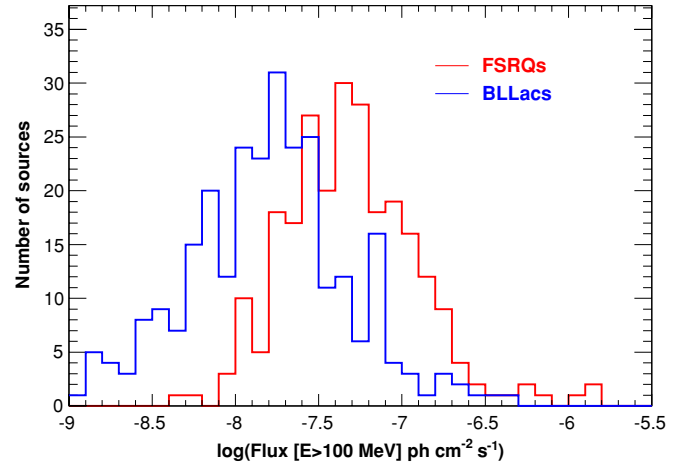


Figure 10. Flux ($E > 100$ MeV) distributions for the FSRQs (red) and BL Lacs (blue) in the clean sample.

(A color version of this figure is available in the online journal.)

does reflect the source's spectral hardness and can be obtained with reasonable accuracy even for fairly faint sources.

The eleven-month average photon spectral index is plotted against the flux ($E > 100$ MeV) estimated from a PL fit in Figure 9 for blazars in the clean sample. As the figure shows, the limiting flux corresponding to $TS = 25$ depends fairly strongly on the photon spectral index; the solid curve corresponds to a simple analytical estimate (Lott et al. 2007). This effect is discussed in greater detail in Abdo et al. (2010c). FSRQs (red circles in Figure 9) mostly cluster in the soft spectral index region, while BL Lacs (blue circles) primarily occupy the hard spectral index region, confirming the trend seen for the LBAS sources (Abdo et al. 2009f). This implies that the limiting flux is different for FSRQs and BL Lacs. The respective flux distributions are compared in Figure 10. The mean fluxes are 8.5×10^{-8} photons $\text{cm}^{-2} \text{s}^{-1}$ and 2.9×10^{-8} photons $\text{cm}^{-2} \text{s}^{-1}$ for FSRQs and BL Lacs, respectively. The faintest BL Lacs are about a factor of 3 fainter than the faintest FSRQs.

The overall flux distribution (of FSRQs and BL Lacs combined) is compared to that measured by EGRET in Figure 11(a). The high-flux ends of the two distributions are in reasonable agreement. The peak flux distributions (maximum flux in a \approx fifteen-day viewing period for EGRET versus maximum monthly flux for the LAT) are compared in Figure 11(b). The peak fluxes observed by EGRET are substantially higher than those observed for the brightest 1LAC sources; this is illustrated in Figures 11(c) and (d), which show the peak flux versus mean flux and the peak flux/mean flux ratio, respectively. The effect of using two different time binnings to determine the peak flux (15 days for EGRET versus 1 month for the LAT) has been studied for the bright LBAS sources; the fifteen-day peak flux is only 13% higher, on average, than the one-month peak flux. Hence, the main reason for the different peak fluxes observed by EGRET and the LAT is probably the different time spans over which the observations were conducted (4.5 years for EGRET versus 11 months for the LAT), enabling the sources to explore a wider range of different states during the EGRET era. Note also that the EGRET observations were often triggered by flaring alerts provided by other facilities, causing a bias that is not present for the LAT.

The photon spectral index distributions for the clean sample, shown in Figure 12, confirm the trend found for the LBAS

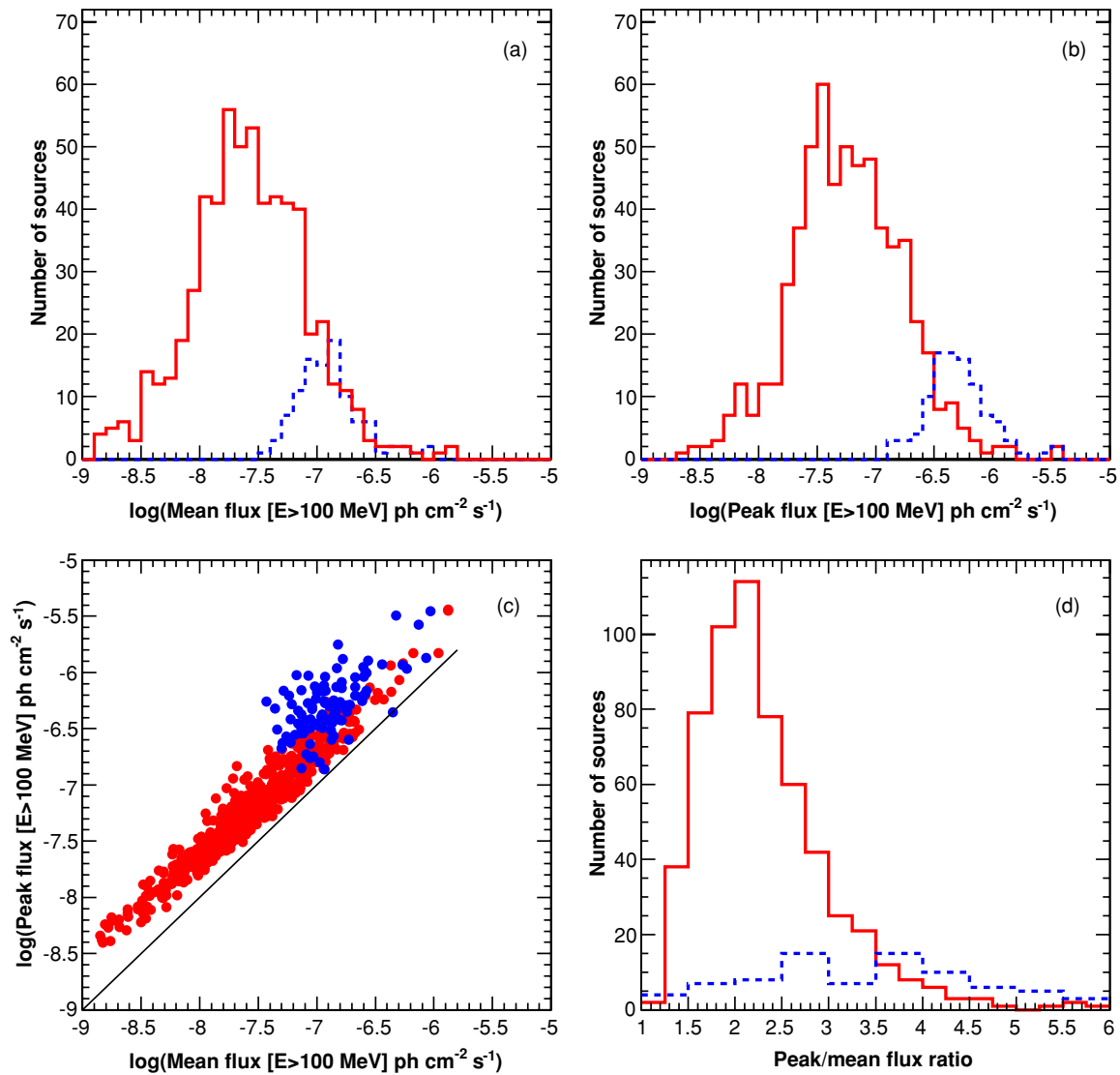


Figure 11. (a) Comparison of mean flux distributions for blazars detected by the *Fermi*-LAT (solid) and by EGRET (dashed). (b) Same as (a) but for the peak flux (i.e., maximum flux in a \sim fifteen-day viewing period for EGRET, maximum monthly flux for the LAT) distributions. (c) Peak flux vs. mean flux for 1LAC AGNs (filled circles) and for EGRET AGNs (open circles). (d) Same as (a), but for the peak flux/mean flux ratio.

(A color version of this figure is available in the online journal.)

sources (Abdo et al. 2009f) and already mentioned above, with a clear difference between FSRQs and BL Lacs. Only 7% of FSRQs (17/231) have a photon spectral index $\Gamma < 2.2$, and only two FSRQs have $\Gamma < 2$ (with uncertainties of ~ 0.2). Thus, it is quite clearly established that LAT-detected FSRQs exhibit soft γ -ray spectra. The 1LAC BL Lac distribution is broader than that observed for LBAS; it shows a larger overlap with the FSRQ distribution. In Abdo et al. (2010a), the BL Lacs with $\Gamma > 2.2$ were found to be mostly LSP sources. This trend is also confirmed for the 1LAC sources, as illustrated in Figure 13, which shows Γ versus ν_{peak}^S . Abdo et al. (2010a) also point out that the BL Lac photon spectral index distribution suffers from a bias since the high-spectral-index (i.e., soft) end is cut off somewhat due to the TS threshold, as seen in Figure 9. Figure 14 shows the photon spectral index distributions for the different blazar classes for sources with $F[E > 100 \text{ MeV}] > 3 \times 10^{-8} \text{ photons cm}^{-2} \text{ s}^{-1}$, above which the sample is essentially complete (see Figure 9). The completeness of the sample will be examined more carefully in an upcoming paper on the LAT blazar population (Abdo et al. 2010e).

The presence of a few very hard ($\Gamma < 1.7$) sources warrants comment. These are fairly low-significance sources (TS < 150), as is evident from Figure 15. Note that typical statistical error bars reach 0.1–0.2 for all sources with TS in that range. Figure 12 also displays the photon spectral index distribution for sources with unknown types. This distribution overlaps with both those of FSRQs and BL Lacs, which seems to rule out that these sources belong preferentially to either class.

Since the flux of each source is evaluated in five different bands, the departure of the spectrum from a PL shape can be estimated via a simple χ^2 test, referred to as the curvature index (Abdo et al. 2010c). Figure 16 displays the curvature index versus total flux for FSRQs and for the three BL Lac subclasses. The trend corroborates the findings reported in Abdo et al. (2010a), namely, that the spectra of bright FSRQs and LSP-BL Lacs show strong departures from a PL shape while those of HSP-BL Lacs are essentially compatible with PLs. Note that due to their harder spectra, any deviation from a PL behavior would be more easily visible for HSP-BL Lacs. Because of larger statistical uncertainties, no conclusion can be drawn from

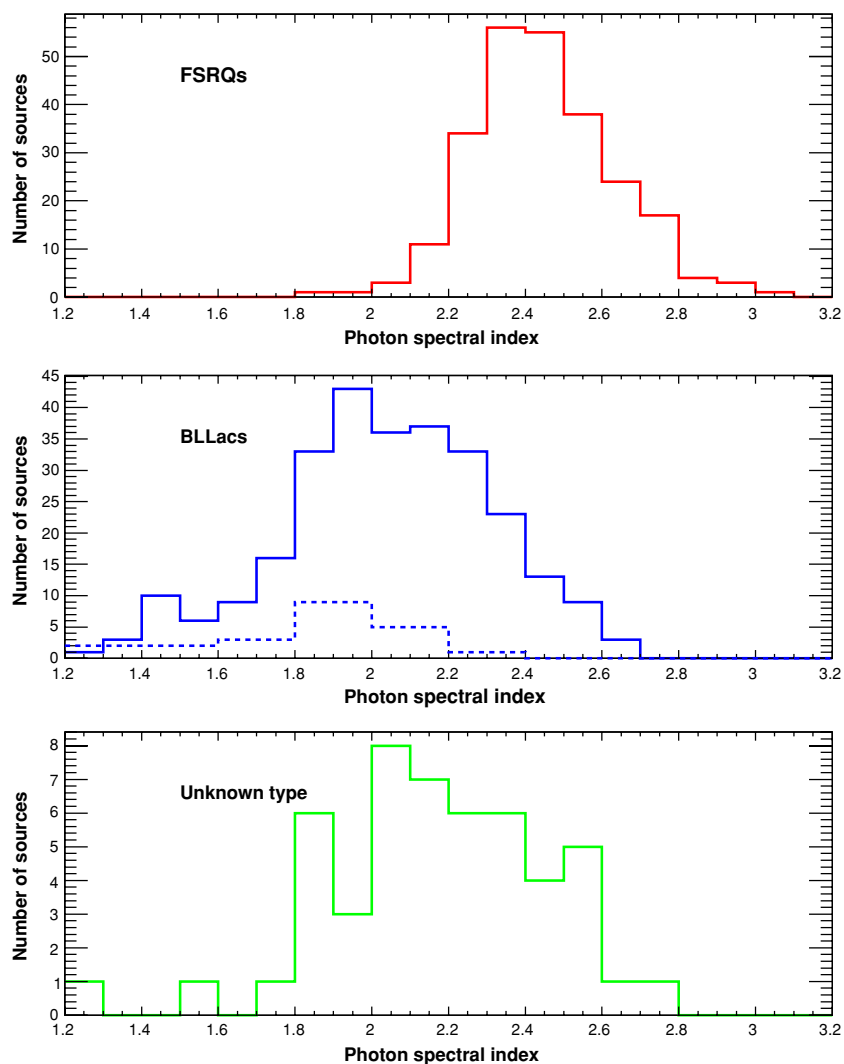


Figure 12. Photon spectral index distributions for the FSRQs (top), BL Lacs (middle), and AGNs of unknown type (bottom) in the clean sample. For BL Lacs, the dashed histogram corresponds to the TeV-detected sources.

(A color version of this figure is available in the online journal.)

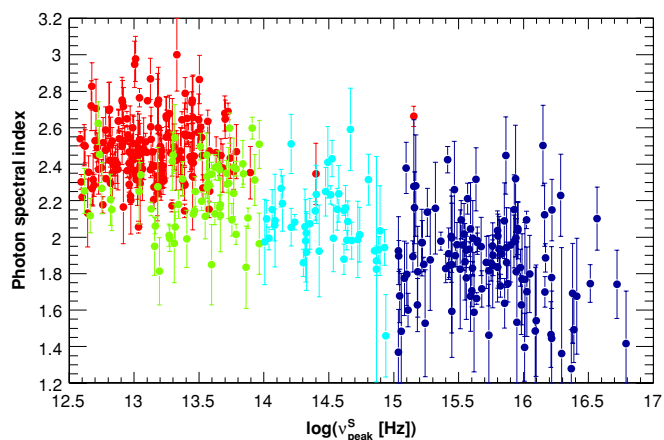


Figure 13. Photon spectral index vs. peak frequency of the synchrotron component of the SED for FSRQs (red) and BL Lacs (green: LSPs, cyan: ISPs, blue: HSPs) in the clean sample.

(A color version of this figure is available in the online journal.)

the χ^2 test for fainter sources (a similar limitation affects the variability index and is described below).

5.2. Variability

One of the defining characteristics of AGNs is their variability, measured at all timescales and at all wavelengths. In the LBAS sample, 46 sources (35 FSRQs, 10 BL Lacs, and one blazar of unknown type) were flagged as variable based on the results of a χ^2 test applied to weekly light curves covering the first 3 months of the LAT sky survey. Recently, Abdo et al. (2010f) detected variability in 64 LBAS sources by analyzing weekly light curves over a period of 11 months.

The electronic version of the 1FGL catalog provides the light curves for all sources over 11 time intervals of 30.37 days each in the energy range 100 MeV–100 GeV. The procedure used to make these light curves is described in Abdo et al. (2010c). Our analysis of the light curves of the sources in the clean sample shows that only six sources (3C 454.3, PKS 1510–08, 3C 279, PKS 1502+106, 3C 273, and PKS 0235+164) have a maximum monthly flux ($E > 100$ MeV) greater than 10^{-6} photons $\text{cm}^{-2} \text{s}^{-1}$ (see also Figure 11(b)). The number of sources with peak flux values above 10^{-6} photons $\text{cm}^{-2} \text{s}^{-1}$ increases if we consider shorter time intervals, as illustrated by the 41 Astronomer’s Telegrams issued by the LAT Collaboration (see below), mostly related to sources that showed short flares that reached this flux level on the timescale of days.

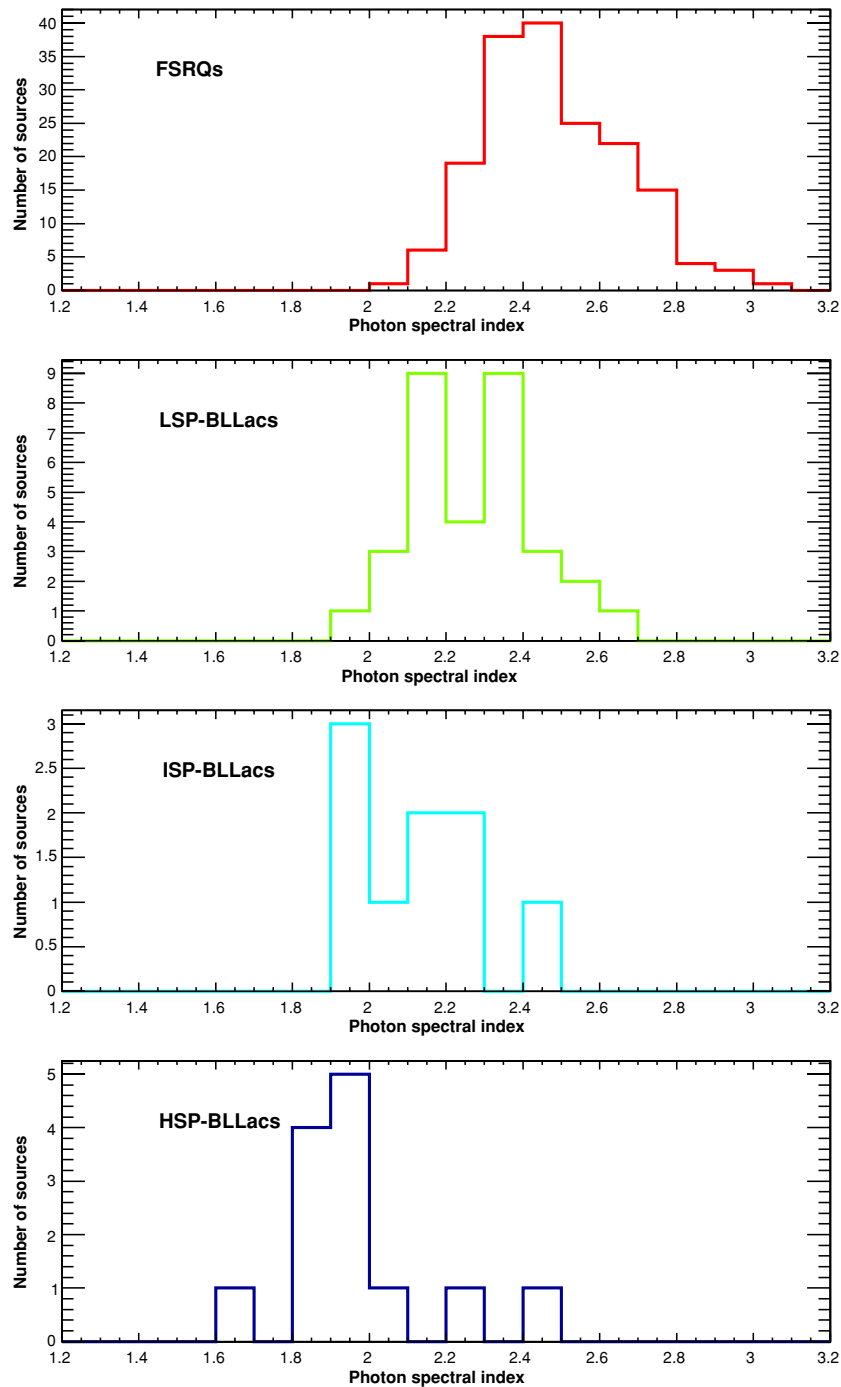


Figure 14. Photon spectral index distributions for the different blazar classes for sources in the clean sample with $F [E > 100 \text{ MeV}] > 3 \times 10^{-8} \text{ photons cm}^{-2} \text{ s}^{-1}$. (A color version of this figure is available in the online journal.)

The 1FGL catalog also lists a variability index V , obtained using a simple χ^2 test, that can be used to determine the probability that a source is variable. Sources for which $V > 23.21$ have a 99% probability (for 10 degrees of freedom) of being variable. Figure 17 shows the distribution of the variability index for blazars (BL Lacs and FSRQs) in the clean sample. 189 blazars in the clean sample are found to be variable (to the right of the vertical line in Figure 17); they comprise 129 FSRQs (68%), 46 BL Lacs (24%), five AGNs of other types (NGC 1275, B2 0321+33B, 4C +15.54, B2 1722+40, and CGRaBS J2250–2806), and nine blazars of unknown type. The 46 variable BL Lacs include 24 LSPs, nine ISPs, 10 HSPs, and

three BL Lacs (CRATES J0058+3311, CGRaBS J0211+1051, and CRATES J1303+2433) for which there are not enough multiwavelength data to permit an SED classification. Figure 18 shows the distribution of the photon spectral index for the variable blazars. This distribution is dominated by sources with $\Gamma > 2.2$. These are quite bright sources observed at energies higher than the peak energies of their SEDs, where the amplitude of the variability is generally larger.

It is important to bear in mind that for a source to be labeled as variable on the basis of its variability index, it must be both intrinsically variable and sufficiently bright. Larger statistical flux uncertainties obtained for fainter sources lead to a reduction

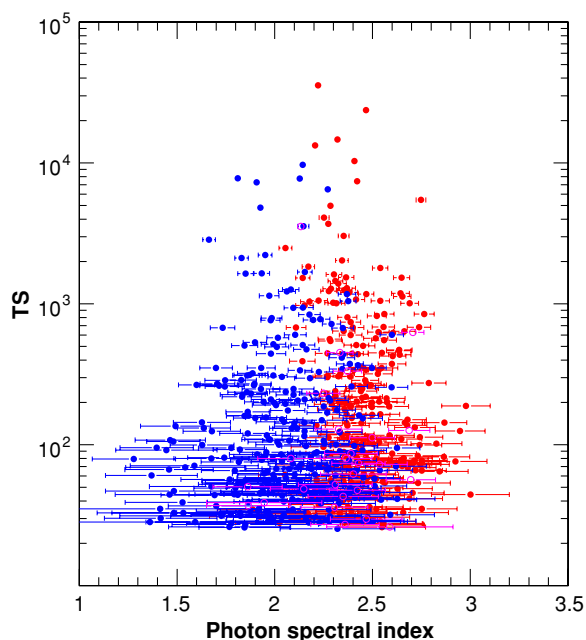


Figure 15. TS vs. photon spectral index for FSRQs (red) and BL Lacs (blue) in the clean sample.

(A color version of this figure is available in the online journal.)

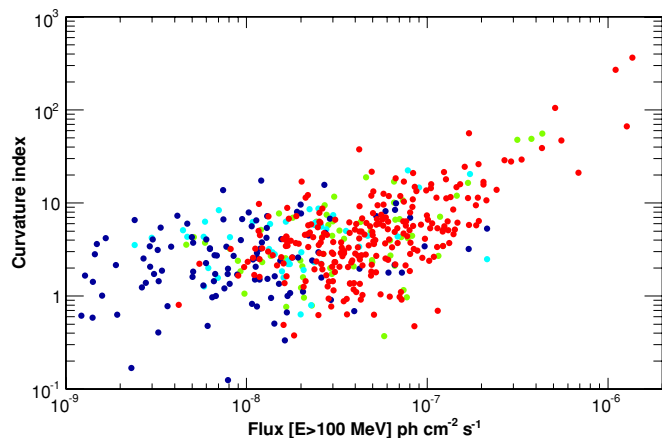


Figure 16. Spectral curvature index vs. flux for FSRQs (red) and BL Lacs (green: LSPs, cyan: ISPs, blue: HSPs) in the clean sample.

(A color version of this figure is available in the online journal.)

of the variability index for a given fractional flux variation. To illustrate this effect, the variability index is plotted against the flux in Figure 19 for FSRQs and BL Lacs. The curves display the evolution of the variability index for two sample sources (the FSRQ 3C 454.3 and the BL Lac object AO 0235+164) that would be observed for the same temporal variation but lower mean fluxes. As a result of this effect, fainter sources appear less variable than brighter sources simply because we cannot measure their variability as well. This may also explain why only $\sim 17\%$ of the 1LAC BL Lacs, which are generally fainter and harder than FSRQs, are found to be variable by this criterion. The histograms in Figure 17 are artificially broadened by the flux dependence of the variability index shown in Figure 19. Allowing for the behavior shown by the trends for 3C 454.3 and AO 0235+164, there is no strong evidence for a significant difference in the intrinsic variability of FSRQs and BL Lacs.

The LAT Collaboration routinely issues Astronomer's Telegrams (ATels) to alert the community to transient sources or

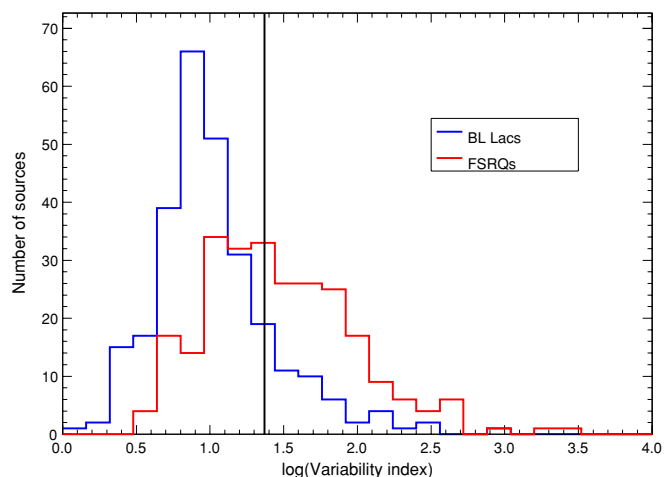


Figure 17. Distribution of the variability index for all BL Lacs and FSRQs in the clean sample. The vertical line is drawn at $V = 23.21$, which corresponds to a 99% probability that a source is variable.

(A color version of this figure is available in the online journal.)

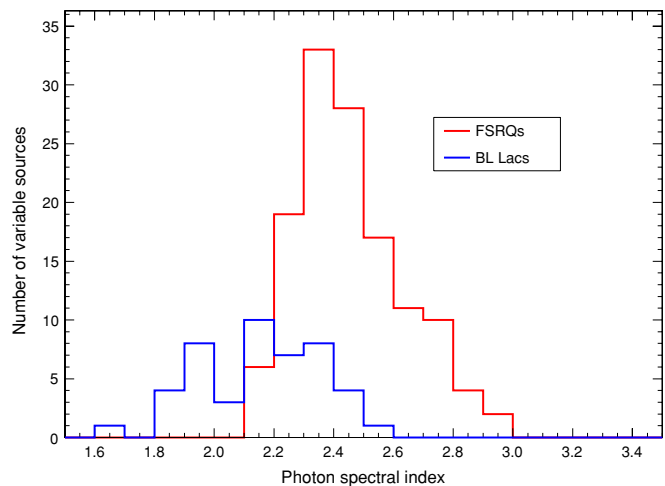


Figure 18. Distribution of the photon spectral index for the variable BL Lacs and FSRQs in the clean sample.

(A color version of this figure is available in the online journal.)

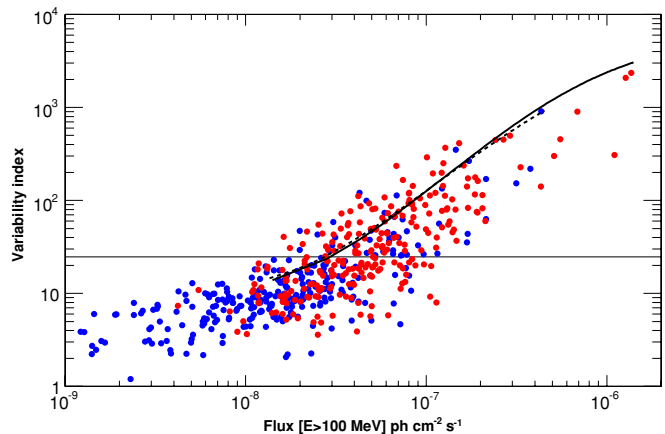


Figure 19. Variability index vs. flux for FSRQs (red) and BL Lacs (blue) in the clean sample. The curves display the evolution of the variability index for the FSRQ 3C 454.3 (solid) and the BL Lac object AO 0235+164 (dashed) that would be observed for the same temporal variation but lower mean fluxes.

(A color version of this figure is available in the online journal.)

sources that are exhibiting flaring states in order to encourage simultaneous multiwavelength follow-up. Through 2010

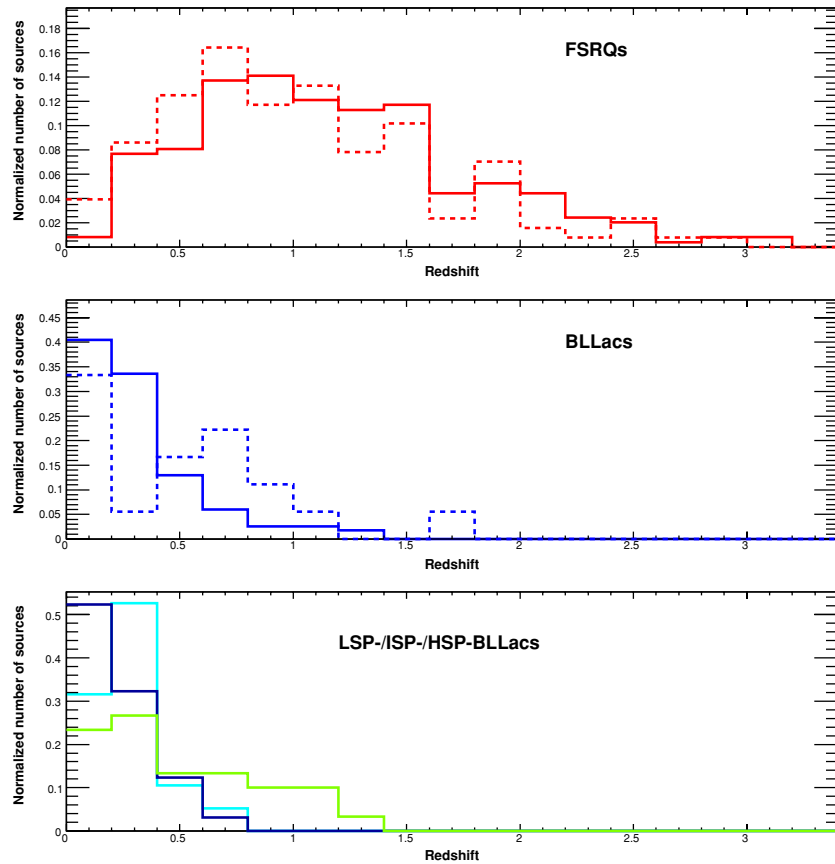


Figure 20. Top: normalized redshift distribution for the FSRQs in the clean sample (solid histogram). The redshift distribution for FSRQs in the *WMAP* catalog is also shown for comparison (dashed histogram). Middle: the same for BL Lacs. Bottom: redshift distributions for LSP-BL Lacs (green), ISP-BL Lacs (cyan), HSP-BL Lacs (blue) in the clean sample.

(A color version of this figure is available in the online journal.)

January 20, 43 1LAC sources (and one at low latitude) have been the subjects of ATels from the LAT team. They comprise 32 FSRQs (with redshifts as high as $z = 2.534$, for B3 1343+451), 10 BL Lacs, and one AGN of another type (CGRaBS J2250–2806, an NLRG).

5.3. Redshift Distributions

The redshift distributions for FSRQs and BL Lacs in the clean sample are presented in Figure 20 along with the corresponding ones for the *WMAP* blazars (Hinshaw et al. 2007; Giommi et al. 2009), which constitute a flux-limited all-sky sample above 1 Jy in the central *WMAP* observing band (*Q* band, 41 GHz). Note that only 121 out of 291 (42%) high-confidence 1LAC BL Lacs have measured redshifts. This is notably worse than for LBAS, in which 29 out of 42 (69%) BL Lacs had measured redshifts. The redshift distributions of γ -ray-detected blazars—both for FSRQs and for BL Lacs—are fairly similar to those for *WMAP*, peaking around $z = 1$ for FSRQs and at a lower redshift for BL Lacs (the lowest-redshift 1LAC BL Lac has $z = 0.030$). The highest redshift for a high-confidence 1LAC FSRQ is $z = 3.10$.

The photon spectral index is plotted against redshift in Figure 21. For FSRQs, no significant evolution is visible. This behavior is compatible with what was previously observed for LBAS. The attenuation effect of the extragalactic background light (EBL) would tend to introduce spurious evidence of evolution (Chen et al. 2004), but the soft spectra of FSRQs and the common presence of spectral breaks at a few GeV (Abdo et al. 2010a) both minimize this effect. A stronger

evolution is seen for BL Lacs: hard sources are mostly located at low redshifts, while most high-redshift sources are softer than average (though it is important to bear in mind that most BL Lacs do not have measured redshifts). This trend was less clear for the LBAS BL Lacs (Abdo et al. 2009f) due to lower statistics. Figure 22 displays the photon spectral index distributions for sources with $z < 0.5$ (top) and $z > 0.5$ (middle), which are clearly different. This provides some insight into the properties of BL Lacs without measured redshifts. The photon spectral index distributions of BL Lacs with measured redshifts and with unknown redshifts are shown in the bottom panel of Figure 22. The distribution of BL Lacs with unknown redshifts includes notably fewer hard sources than that for BL Lacs with known redshifts. A Kolmogorov–Smirnov (K–S) test yields a probability of 6×10^{-3} that the two distributions are drawn from the same underlying population. The similarity of the distribution of BL Lacs with unknown redshifts to that of BL Lacs with $z > 0.5$ (K–S probability of 0.54 that the two distributions are drawn from the same underlying population) supports the idea of some bias toward higher redshifts for this class of objects.

5.4. Luminosity Distributions

The γ -ray luminosity L_γ is calculated as in Ghisellini et al. (2009):

$$L_\gamma = 4\pi d_L^2 \frac{S(E_1, E_2)}{(1+z)^{2-\Gamma}},$$

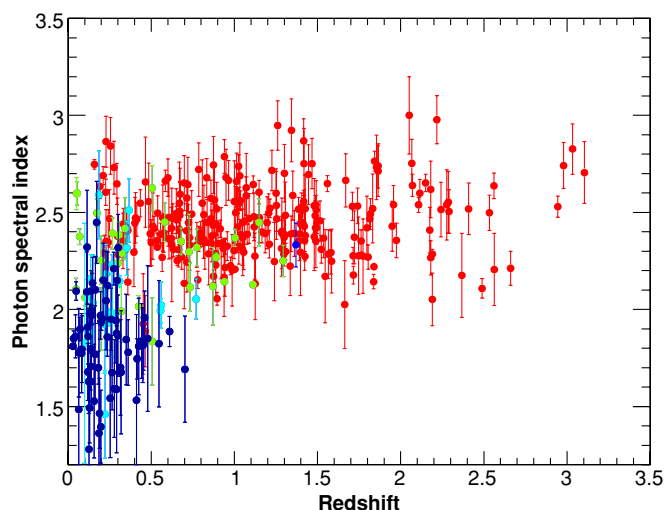


Figure 21. Photon spectral index vs. redshift for the FSRQs (red) and BL Lacs (green: LSPs, cyan: ISPs, blue: HSPs) in the clean sample. (A color version of this figure is available in the online journal.)

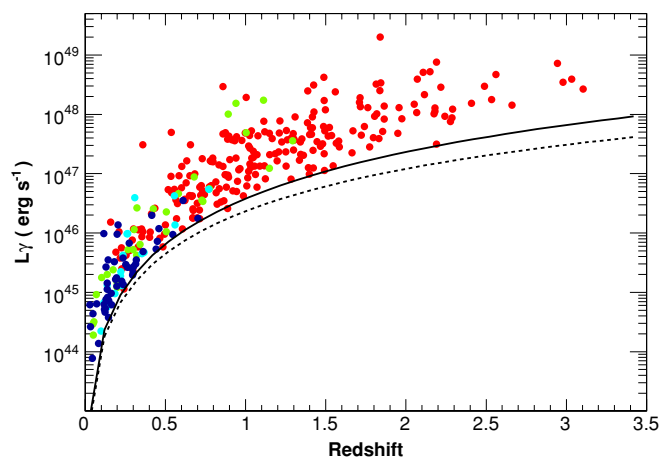


Figure 23. γ -ray luminosity vs. redshift for the different AGN classes (red: FSRQs, green: LSP-BL Lacs, cyan: ISP-BL Lacs, blue: HSP-BL Lacs, magenta: radio galaxies) in the clean sample. The curves correspond to approximate experimental limits calculated for a photon spectral index $\Gamma = 2.2$ (solid) and $\Gamma = 1.8$ (dashed). (A color version of this figure is available in the online journal.)

where d_L is the luminosity distance, Γ is the photon spectral index, S is the energy flux, and E_1 and E_2 are the lower and upper energy bounds (taken here to be 100 MeV and 100 GeV), respectively. Implicit in this derivation is the assumption that $E_2 \gg E_1$, which is satisfied for our calculations. Figure 23 shows L_γ plotted against the source redshift. The curves

represent approximate instrumental limits calculated for two photon indices, $\Gamma = 1.8$ and $\Gamma = 2.2$. The low-redshift FSRQs with $L_\gamma > 10^{48}$ erg s $^{-1}$ could potentially still be detected at redshifts $z > 3.1$.

Figure 24 shows the photon spectral index plotted against the γ -ray luminosity. The Pearson correlation coefficient

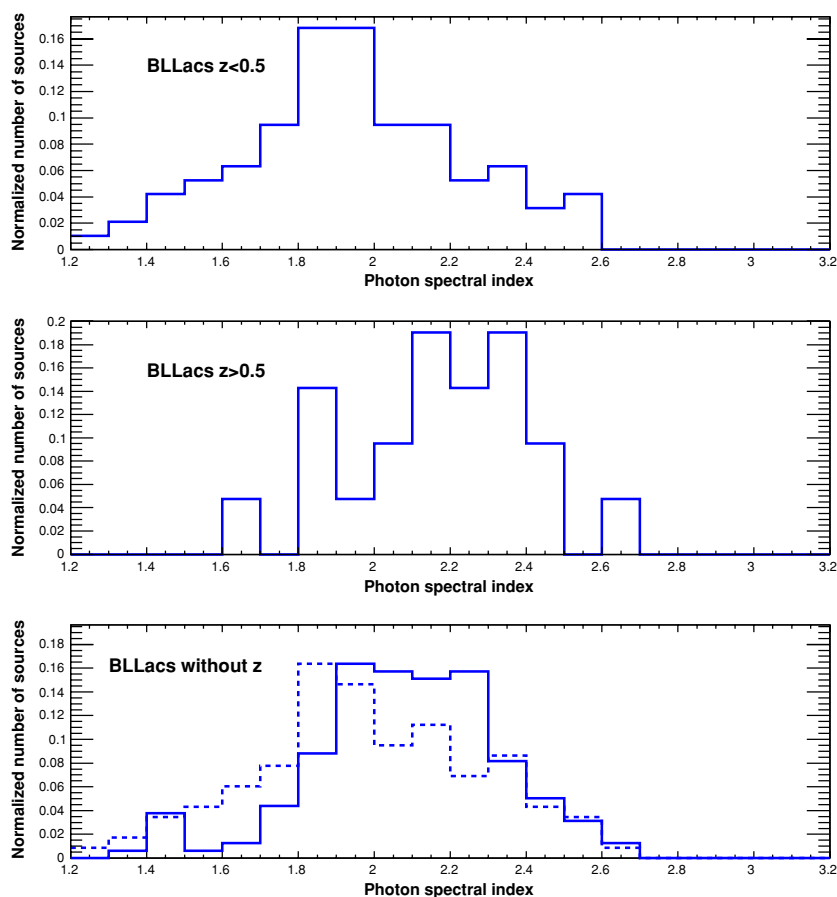


Figure 22. Photon spectral index distributions for BL Lacs in the clean sample with $z < 0.5$ (top), with $z > 0.5$ (middle), and with unknown redshifts (bottom, solid histogram). The total distribution for BL Lacs with known redshifts is shown for comparison as a dashed histogram in the bottom panel. (A color version of this figure is available in the online journal.)

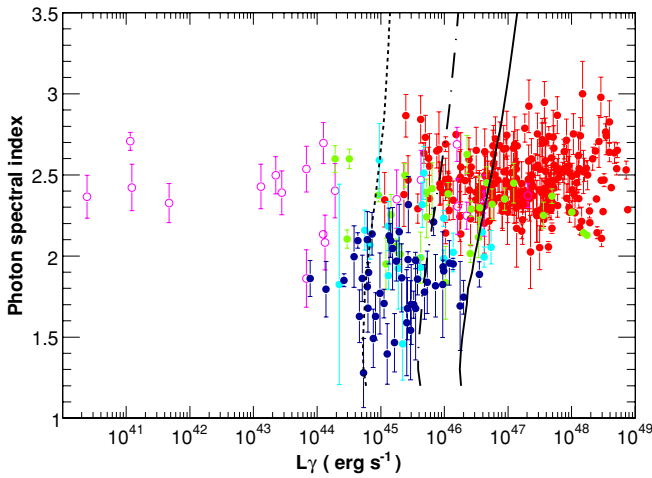


Figure 24. Photon spectral index vs. γ -ray luminosity for the different AGN classes (red: FSRQs, green: LSP-BL Lacs, cyan: ISP-BL Lacs, blue: HSP-BL Lacs, magenta: radio galaxies) in the clean sample. The curves represent approximate instrumental limits for $z = 0.2$ (dashed), $z = 0.5$ (dot-dashed), and $z = 1$ (solid).

(A color version of this figure is available in the online journal.)

for the two parameters is 0.17. It is important to bear in mind two issues when interpreting this correlation: the higher-flux limit for soft sources (see Figures 9 and 10) and the difference in redshift distributions between FSRQs and BL Lacs (see Figure 20). Given their relative softness (and thus, high-flux limit), FSRQs are excluded from the soft-faint region. BL Lacs are also partly excluded because of their relatively low fluxes. Luminosity limits calculated for different redshifts (with the same approach as for the flux limit shown in Figure 9) are displayed for reference in Figure 24 and illustrate somewhat the effect of the Malmquist bias.⁷⁶

6. MULTIWAVELENGTH PROPERTIES OF ILAC SOURCES

6.1. Sources Detected at TeV Energies

Over the last two decades, ground-based γ -ray instruments operating in the “TeV” or very high energy (VHE; $E \gtrsim 100$ GeV) regime have detected 32 AGNs, with the pace of discovery increasing significantly⁷⁷ as the latest generation of instruments—CANGAROO, H.E.S.S., MAGIC, and VERITAS—has been commissioned. Of these 32 TeV AGNs, the majority are BL Lacs (23 HSPs, three ISPs, and two LSPs), with the remainder comprising one FSRQ, two FR-I radio galaxies, and one AGN of unknown type. A detailed analysis of the TeV AGNs detected with *Fermi* during the first 5.5 months of operation is given by Abdo et al. (2009i, and see references to the TeV detections therein). Five new TeV AGNs have since been detected (1ES 0414+009, Hofmann & Fegan 2009; PKS 0447–437, Raue et al. 2009; RBS 0413, Ong & Fortin 2009; 1ES 0502+675, Ong 2009a; VER J0521+211, Ong 2009b). Two of these detections and one previous detection (PKS 1424+240; Acciari et al. 2010) were motivated directly by the detection of GeV emission with *Fermi*.

⁷⁶ In a flux-limited sample, sources located at larger distances appear more luminous than closer ones since fainter sources are below the detection threshold (Malmquist 1920).

⁷⁷ TeVCat (<http://tevcat.uchicago.edu/>) presents an up-to-date catalog of TeV sources.

Table 5
Positional Coincidences of ILAC Sources with TeV Sources

LAT Source	TeV Association	R.A. ^a	Decl. ^a
1FGL J0222.6+4302	3C 66A	35.8000	43.0117
1FGL J0319.7+1847	RBS 0413	49.9658	18.7594
1FGL J0416.8+0107	1ES 0414+009	64.2184	1.0901
1FGL J0449.5–4350	PKS 0447–437	72.3529	–43.8358
1FGL J0507.9+6738	1ES 0502+675	76.9842	67.6233
1FGL J0521.7+2114 ^b	VER J0521+211 ^b	80.4792	21.1900
1FGL J0710.6+5911	RGB J0710+591	107.6254	59.1390
1FGL J0721.9+7120	S5 0716+714	110.4725	71.3433
1FGL J0809.5+5219	1ES 0806+524	122.4550	52.3161
1FGL J1015.1+4927	1ES 1011+496	153.7671	49.4336
1FGL J1103.7–2329	1ES 1101–232	165.9083	–23.4919
1FGL J1104.4+3812	Mkn 421	166.1138	38.2089
1FGL J1136.6+7009	Mkn 180	174.1100	70.1575
1FGL J1221.3+3008	1ES 1218+304	185.3413	30.1769
1FGL J1230.8+1223	M87	187.7058	12.3911
1FGL J1221.5+2814	W Com	185.3821	28.2331
1FGL J1256.2–0547	3C 279	194.0463	–5.7894
1FGL J1325.6–4300	Cen A	201.3667	–43.0183
1FGL J1426.9+2347	PKS 1424+240	216.7516	23.8000
1FGL J1428.7+4239	H 1426+428	217.1358	42.6725
1FGL J1555.7+1111	PG 1553+113	238.9292	11.1900
1FGL J1653.9+3945	Mkn 501	253.4675	39.7600
1FGL J2000.0+6508	1ES 1959+650	299.9996	65.1486
1FGL J2009.5–4849	PKS 2005–489	302.3721	–48.8219
1FGL J2158.8–3013	PKS 2155–304	329.7196	–30.2217
1FGL J2202.8+4216	BL Lac	330.6804	42.2778
1FGL J2347.1+5142 ^b	1ES 2344+514 ^b	356.7700	51.7050
1FGL J2359.0–3035	H 2356–309	359.7875	–30.6228

Notes.

^a J2000 coordinate, in degrees, from TeVCat (<http://tevcat.uchicago.edu/>).

^b This source is at low Galactic latitude ($|b| < 10^\circ$) and is thus not formally in the ILAC but appears in Table 2.

Of the TeV AGNs, the 28 listed in Table 5 are detected by *Fermi* as 1FGL sources, with a mean photon spectral index of $\langle \Gamma_{\text{GeV}} \rangle = 2.02 \pm 0.01$. Taking only the subsample of 25 GeV–TeV BL Lacs, the mean index is $\langle \Gamma_{\text{GeV}} \rangle = 1.92 \pm 0.01$ ($\sigma_\Gamma = 0.26$), compared with 2.07 ± 0.01 ($\sigma_\Gamma = 0.28$) for the larger sample of *Fermi* BL Lacs (see Figure 12(b)), illustrating that the TeV sources are among the hardest BL Lacs in the GeV regime. The measured spectra of the majority of the GeV–TeV BL Lacs are well described by PLs in both regimes. For many of the sources, the photon spectral index of the GeV emission differs significantly from that of the TeV emission, which may be an indication of the presence of a break in the γ -ray spectrum between the two regimes. However, most of the TeV spectra have not been measured simultaneously with the GeV spectra, and caution is advised when comparing the spectra in detail. The largest such break, consistent with $\Delta\Gamma \equiv \Gamma_{\text{TeV}} - \Gamma_{\text{GeV}} \sim 2$, is evident in the spectra of 1ES 1101+496, H 1426+428, and PG 1553+113. By contrast, the spectra of the nearby radio galaxies M87 and Cen A show no evidence of a spectral break. The mean break index is $\langle \Delta\Gamma \rangle = 1.3$. The values of $\Delta\Gamma$ for the GeV–TeV sources are shown in Figure 25, plotted against the redshift of the source. Among the possible explanations for the apparent deficit of sources with small $\Delta\Gamma$ at high redshift is the effects of pair production with the EBL, which is expected to introduce a redshift-dependent steepening into the TeV spectra of extragalactic objects.

The four TeV AGNs not detected thus far by *Fermi* are RGB J0152+017, 1ES 0229+200, 1ES 0347–121, and PKS

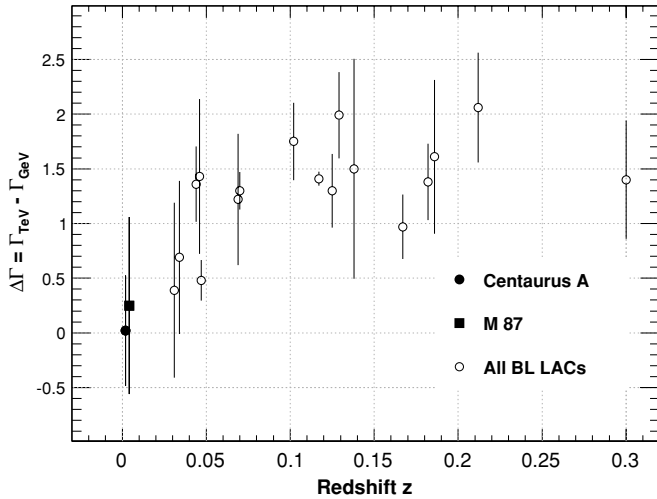


Figure 25. GeV/TeV spectral break $\Delta\Gamma \equiv \Gamma_{\text{TeV}} - \Gamma_{\text{GeV}}$ vs. redshift. Only sources with well measured redshifts and TeV photon spectral indices are shown. The TeV measurements come from the references given in Abdo et al. (2009i); new measurements for PKS 2005–489 (Acero et al. 2010) and S5 0716+714 (Mazin et al. 2009) are used. 3C 279, an FSRQ whose TeV spectrum was measured during an extreme flaring episode, is also excluded.

0548–322, all HSP-BL Lacs. Extrapolating the measured PL TeV spectra from all the TeV AGNs down to 200 GeV, it is evident that these four have among the smallest fluxes of all TeV AGN at this energy.

6.2. Sources Detected Previously at GeV Energies

Of the 709 sources in the 1LAC, 114 sources were included in the BSL and were thus significantly detected (at 10σ or greater) during the first 3 months of LAT observations. All of those sources were previously associated with AGNs as part of LBAS. Three low-confidence LBAS sources (0FGL J0909.7+0145, 0FGL J1248.7+5811, and 0FGL J1641.4+3939) are not confirmed in the 1LAC sample.

Ten years after EGRET, it is interesting to look at the fraction of the AGNs that were active in the EGRET era and are detected again by the LAT with a comparable flux. We consider two sources to be “positionally coincident” when the separation between their positions is less than the quadratic sum of their 95% error radii. In the 1LAC sample, 63 AGNs are positionally coincident with 3EG sources. Of these, the 3EG catalog lists 51 sources as AGN identifications and four as AGN associations. The 3EG catalog listed a total of 62 sources as AGN identifications, so there are 11 identified EGRET AGNs that are not positionally coincident with 1LAC sources. Forty-five sources from the 1LAC have positions compatible with the revised EGRET catalog (EGR; Casandjian & Grenier 2008), while 22 sources are positionally coincident with sources from the high-energy EGRET catalog (GEV; Lamb & Macomb 1997). In all, 75 1LAC sources have coincident detections from one or more of these three EGRET catalogs.

Eleven of the sources in this catalog are included in the first-year *AGILE* catalog (1AGL; Pittori et al. 2009). All 11 are identified as blazars by *AGILE*, and indeed, these 11 comprise the total sample of blazars in the 1AGL catalog. Only two of these *AGILE*-detected sources do not have similar EGRET detections. In all, 77 1LAC sources have been cataloged by other GeV instruments.

These 77 sources are listed in Table 6 along with the mean fluxes and photon indices measured by the LAT and by EGRET.

These AGNs are composed of 49 FSRQs, 21 BL Lacs, four AGNs of other types, and three AGNs of unknown types.

During its 4.5 year mission, EGRET found few sources with flux ($E > 100$ MeV) less than 10×10^{-8} photons $\text{cm}^{-2} \text{s}^{-1}$. A large number of the 1LAC sources have fluxes well below this value; such sources would not have been visible to EGRET. Some sources, such as 1FGL J0428.6–3756 (associated with PKS 0426–380) and 1FGL J2229.7–0832 (associated with PKS 2227–08), have flux values well above the EGRET threshold but were not seen by EGRET and yet are not noted as being variable in the 1FGL data. These examples continue to demonstrate that long-duration variability is evident in AGN γ -ray flux.

All of the EGRET sources seen at 10σ significance and associated with flaring blazars have been detected by the LAT and appear with blazar associations.

6.3. Sources Detected in the Hard X-ray Band

In recent years, a new generation of hard X-ray telescopes has drastically improved our knowledge of the source populations in this energy band. The *Swift*-BAT and *INTEGRAL*-IBIS instruments have been conducting surveys at hard X-ray energies; the most recent catalogs are the fourth IBIS catalog (Bird et al. 2010) of 723 sources detected at 17–100 keV and the 54 month Palermo BAT catalog (Cusumano et al. 2010) of 1049 sources⁷⁸ detected at 14–150 keV. These two telescopes perform in a complementary fashion; while IBIS is the more sensitive instrument with better angular resolution, it has a smaller field of view and has concentrated exposure in the Galactic plane. Conversely, the BAT instrument has lower instantaneous sensitivity but has a much larger field of view and a much more uniform exposure across the whole sky. An initial comparison between the LBAS sources and the fourth IBIS catalog by Ubertini et al. (2009) showed that only a small subset of the >250 *INTEGRAL* AGNs, of which 19 are blazars, was detected by the LAT.

Restricting the catalogs to $|b| > 10^\circ$, there are 291 IBIS sources and 736 BAT sources. If we consider all of the 1LAC sources with associated counterparts that fall within the error circles of the hard X-ray catalogs, we find that 50 of the 1LAC sources can be associated with known hard X-ray sources. Of these 50 sources, one appears only in the fourth IBIS catalog, 16 appear in both the fourth IBIS catalog and the 54 month Palermo BAT catalog, and the remaining 35 are detected only in the 54 month Palermo BAT catalog. This is still a very small subset of AGNs despite the depth of the LAT catalog and the sky coverage of the BAT catalog.

Table 7 lists the high-confidence 1LAC AGNs detected in hard X-rays. This sample contains 27 FSRQs, 16 BL Lacs, and seven AGNs of other types. Most of the AGNs (54%) detected in hard X-rays are FSRQs; this is expected since the external Compton peak of these high-luminosity objects reaches a maximum at MeV energies and declines in the hard X-rays, maintaining a substantial fraction of the emitting power. The difference in the spectral shapes of BL Lacs and FSRQs observed in hard X-rays indicates that this energy range probes the high-energy tail of the synchrotron peak in BL Lacs and the ascending part of the Compton peak in FSRQs.

Finally, there is evidence that the hard X-ray emission from most bright LAT blazars is, apparently, missed in spite of the

⁷⁸ Cusumano et al. (2010) describe the 39 month catalog; the 54 month catalog is available at: <http://www.ifa.inaf.it/cgi-bin/INAF/pub.cgi?href=activities/bat/index.html>

Table 6
Positional Coincidences of 1LAC Sources with EGRET/AGILE Sources

1FGL Name	3EG Name	EGR Name	GEV Name	1AGL Name	LAT Flux ^a	LAT Ph. Index	3EG Flux ^a	EGRET Ph. Index	Probable Class
1FGL J0204.5+1516	3EG J0204+1458	EGR J0204+1505			1.79 ± 0.58	2.47 ± 0.18	23.6 ± 5.6	2.23 ± 0.28	AGN
1FGL J0210.6-5101	3EG J0210-5055	EGR J0210-5058	GEV J0210-5053		14.59 ± 0.96	2.37 ± 0.04	85.5 ± 4.5	1.99 ± 0.05	BL Lac
1FGL J0222.6+4302	3EG J0222+4253	EGR J0223+4300	GEV J0223+4254		21.45 ± 0.97	1.93 ± 0.02	18.7 ± 2.9	2.01 ± 0.14	BL Lac
1FGL J0238.6+1637	3EG J0237+1635		GEV J0237+1648		43.37 ± 1.07	2.14 ± 0.02	65.1 ± 8.8	1.85 ± 0.12	BL Lac
1FGL J0339.2-0143	3EG J0340-0201	EGR J0338-0203			4.00 ± 0.65	2.50 ± 0.10	118.8 ± 22.0	1.84 ± 0.25	FSRQ
1FGL J0416.5-1851	3EG J0412-1853	EGR J0413-1851			3.59 ± 0.58	2.37 ± 0.10	49.5 ± 16.1	3.25 ± 0.68	FSRQ
1FGL J0423.2-0118	3EG J0422-0102				13.99 ± 0.90	2.42 ± 0.04	50.2 ± 10.4	2.44 ± 0.19	FSRQ
1FGL J0433.5+2905	3EG J0433+2908	EGR J0433+2906	GEV J0433+2907		6.57 ± 0.83	2.13 ± 0.06	22.0 ± 2.8	1.90 ± 0.10	BL Lac
1FGL J0442.7-0019	3EG J0442-0033	EGR J0442-0027	GEV J0441-0044		17.27 ± 0.91	2.44 ± 0.04	79.0 ± 10.1	2.37 ± 0.18	FSRQ
1FGL J0448.6+1118	3EG J0450+1105				5.90 ± 0.88	2.51 ± 0.09	109.5 ± 19.4	2.27 ± 0.16	BL Lac
1FGL J0455.6-4618	3EG J0458-4635				6.15 ± 0.83	2.57 ± 0.09	7.7 ± 2.1	2.75 ± 0.35	FSRQ
1FGL J0457.0-2325	3EG J0456-2338	EGR J0456-2334			51.14 ± 1.07	2.21 ± 0.02	14.7 ± 4.2	3.14 ± 0.47	FSRQ
1FGL J0501.0-0200	3EG J0500-0159				3.74 ± 0.70	2.50 ± 0.11	11.2 ± 2.3	2.45 ± 0.27	FSRQ
1FGL J0509.3+0540		EGR J0509+0550	GEV J0508+0540		5.77 ± 0.73	2.16 ± 0.06	BL Lac
1FGL J0516.7-6207	3EG J0512-6150				5.76 ± 0.95	2.28 ± 0.09	7.2 ± 1.7	2.40 ± 0.26	Unknown
1FGL J0531.0+1331	3EG J0530+1323	EGR J0530+1331	GEV J0530+1340		16.85 ± 1.28	2.64 ± 0.06	93.5 ± 3.6	2.46 ± 0.04	FSRQ
1FGL J0538.8-4404	3EG J0540-4402	EGR J0540-4358	GEV J0540-4359	1AGL J0538-4424	37.77 ± 1.06	2.27 ± 0.02	25.3 ± 3.1	2.41 ± 0.12	BL Lac
1FGL J0540.9-0547	3EG J0542-0655				3.73 ± 1.01	2.37 ± 0.12	66.5 ± 19.5	...	FSRQ
1FGL J0721.9+7120	3EG J0721+7120	EGR J0723+7134	GEV J0719+7133	1AGL J0722+7125	17.26 ± 0.75	2.15 ± 0.03	17.8 ± 2.0	2.19 ± 0.11	BL Lac
1FGL J0738.2+1741	3EG J0737+1721	EGR J0737+1720			4.60 ± 0.48	2.02 ± 0.06	16.4 ± 3.3	2.60 ± 0.28	BL Lac
1FGL J0742.2+5443	3EG J0743+5447	EGR J0743+5438			5.22 ± 0.75	2.45 ± 0.09	30.3 ± 5.0	2.03 ± 0.20	FSRQ
1FGL J0830.5+2407	3EG J0829+2413	EGR J0829+2415			8.35 ± 0.84	2.79 ± 0.09	24.9 ± 3.9	2.42 ± 0.21	FSRQ
1FGL J0831.6+0429		EGR J0829+0510			7.35 ± 0.76	2.50 ± 0.07	BL Lac
1FGL J0842.2+7054	3EG J0845+7049				8.06 ± 1.04	2.98 ± 0.12	10.2 ± 1.8	2.62 ± 0.16	FSRQ
1FGL J0850.0-1213	3EG J0852-1216	EGR J0852-1224			5.02 ± 0.64	2.27 ± 0.07	44.4 ± 11.6	1.58 ± 0.58	FSRQ
1FGL J0854.8+2006	3EG J0853+1941	EGR J0853+2015			7.03 ± 0.82	2.38 ± 0.07	10.6 ± 3.0	2.03 ± 0.35	BL Lac
1FGL J0957.7+5523		EGR J0957+5513	GEV J0956+5508		11.39 ± 0.62	2.05 ± 0.03	FSRQ
1FGL J1000.1+6539	3EG J0958+6533	EGR J0956+6524			2.59 ± 0.67	2.51 ± 0.16	15.4 ± 3.0	2.08 ± 0.24	BL Lac
1FGL J1104.4+3812	3EG J1104+3809	EGR J1104+3813	GEV J1104+3809	1AGL J1104+3754	16.93 ± 0.58	1.81 ± 0.02	13.9 ± 1.8	1.57 ± 0.15	BL Lac
1FGL J1133.1+0033	3EG J1133+0033				2.71 ± 0.53	2.18 ± 0.11	10.6 ± 3.0	2.73 ± 0.63	BL Lac
1FGL J1159.4+2914	3EG J1200+2847		GEV J1201+2906		12.24 ± 0.74	2.37 ± 0.04	50.9 ± 11.9	1.98 ± 0.22	FSRQ
1FGL J1221.5+2814				1AGL J1222+2851	7.82 ± 0.61	2.06 ± 0.04	BL Lac
1FGL J1224.7+2121	3EG J1224+2118				8.07 ± 0.75	2.55 ± 0.07	13.9 ± 1.8	2.28 ± 0.13	FSRQ
1FGL J1225.8+4336	3EG J1227+4302				2.82 ± 0.69	2.81 ± 0.18	21.7 ± 7.1	...	Unknown
1FGL J1229.1+0203	3EG J1229+0210	EGR J1229+0203		1AGL J1228+0142	55.30 ± 1.48	2.75 ± 0.03	15.4 ± 1.8	2.58 ± 0.09	FSRQ
1FGL J1239.5+0443		EGR J1237+0434		1AGL J1238+0406	8.41 ± 0.75	2.35 ± 0.06	FSRQ
1FGL J1256.2-0547	3EG J1255-0549	EGR J1256-0552	GEV J1256-0546	1AGL J1256-0549	68.84 ± 1.37	2.32 ± 0.02	179.7 ± 6.7	1.96 ± 0.04	FSRQ
1FGL J1258.7-2221		EGR J1259-2209			5.66 ± 0.80	2.39 ± 0.08	FSRQ
1FGL J1321.1+2214	3EG J1323+2200				1.62 ± 0.50	2.21 ± 0.15	18.1 ± 4.0	1.86 ± 0.35	FSRQ
1FGL J1325.6-4300	3EG J1324-4314				20.40 ± 1.47	2.71 ± 0.06	13.6 ± 2.5	2.58 ± 0.26	AGN
1FGL J1337.7-1255		EGR J1337-1310			7.21 ± 1.12	2.50 ± 0.09	FSRQ
1FGL J1408.9-0751	3EG J1409-0745	EGR J1409-0736	GEV J1409-0741		5.21 ± 0.71	2.42 ± 0.08	97.6 ± 9.1	2.29 ± 0.11	FSRQ
1FGL J1422.7+3743	3EG J1424+3734	EGR J1424+3730			2.36 ± 0.72	2.63 ± 0.19	16.3 ± 4.9	3.25 ± 0.46	BL Lac
1FGL J1428.2-4204	3EG J1429-4217	EGR J1428-4240			6.70 ± 0.82	2.31 ± 0.07	29.5 ± 5.3	2.13 ± 0.21	FSRQ
1FGL J1457.5-3540	3EG J1500-3509				33.27 ± 1.18	2.27 ± 0.02	10.9 ± 2.8	2.99 ± 0.37	FSRQ
1FGL J1503.5-1544	3EG J1504-1537	EGR J1504-1539			0.76 ± 0.38	1.74 ± 0.19	33.2 ± 10.3	...	BL Lac
1FGL J1505.1-3435	3EG J1500-3509				1.11 ± 0.54	2.02 ± 0.19	10.9 ± 2.8	2.99 ± 0.37	Unknown
1FGL J1512.8-0906	3EG J1512-0849	EGR J1512-0857		1AGL J1511-0908	127.10 ± 1.85	2.41 ± 0.01	18.0 ± 3.8	2.47 ± 0.21	FSRQ

Table 6
(Continued)

1FGL Name	3EG Name	EGR Name	GEV Name	1AGL Name	LAT Flux ^a	LAT Ph. Index	3EG Flux ^a	EGRET Ph. Index	Probable Class
1FGL J1607.1+1552	3EG J1605+1553	EGR J1607+1533			3.86 ± 0.59	2.25 ± 0.08	42.0 ± 12.3	2.06 ± 0.41	AGN
1FGL J1609.0+1031	3EG J1608+1055	EGR J1608+1051			6.19 ± 0.87	2.72 ± 0.10	34.9 ± 5.6	2.63 ± 0.24	FSRQ
1FGL J1613.5+3411	3EG J1614+3424		GEV J1613+3432		1.02 ± 0.48	2.29 ± 0.22	26.5 ± 4.0	2.42 ± 0.15	FSRQ
1FGL J1625.7–2524	3EG J1626–2519		GEV J1626–2502		10.53 ± 1.42	2.36 ± 0.06	42.6 ± 6.6	2.21 ± 0.13	FSRQ
1FGL J1626.2–2956	3EG J1625–2955	EGR J1625–2958	GEV J1626–2955		5.02 ± 0.94	2.36 ± 0.10	258.9 ± 15.3	2.07 ± 0.07	FSRQ
1FGL J1635.0+3808	3EG J1635+3813		GEV J1636+3812		19.21 ± 1.17	2.47 ± 0.04	107.5 ± 9.6	2.15 ± 0.09	FSRQ
1FGL J1635.4+8228	3EG J1621+8203				3.78 ± 0.71	2.50 ± 0.12	10.4 ± 3.0	2.29 ± 0.49	AGN
1FGL J1702.7–6217	3EG J1659–6251				4.74 ± 1.04	2.54 ± 0.13	47.0 ± 13.1	2.54 ± 0.37	FSRQ
1FGL J1728.2+0431	3EG J1727+0429	EGR J1727+0416			6.66 ± 1.05	2.65 ± 0.11	17.9 ± 4.1	2.67 ± 0.26	FSRQ
1FGL J1733.0–1308	3EG J1733–1313				8.73 ± 1.12	2.34 ± 0.07	36.1 ± 3.4	2.23 ± 0.10	FSRQ
1FGL J1740.0+5209	3EG J1738+5203	EGR J1740+5213			17.20 ± 1.03	2.71 ± 0.05	18.2 ± 3.5	2.42 ± 0.23	FSRQ
1FGL J1849.3+6705				1AGL J1846+6714	24.24 ± 0.89	2.25 ± 0.03	FSRQ
1FGL J1911.2–2007	3EG J1911–2000	EGR J1912–2000			12.37 ± 1.07	2.42 ± 0.05	17.5 ± 2.7	2.39 ± 0.18	FSRQ
1FGL J2006.6–2302	3EG J2006–2321				7.16 ± 0.95	2.68 ± 0.10	19.8 ± 4.4	2.33 ± 0.36	FSRQ
1FGL J2009.5–4849			GEV J2009–4827		3.94 ± 0.49	1.90 ± 0.06	BL Lac
1FGL J2025.6–0735	3EG J2025–0744		GEV J2024–0812	1AGL J2026–0732	29.15 ± 1.18	2.35 ± 0.03	74.5 ± 13.4	2.38 ± 0.17	FSRQ
1FGL J2031.5+1219		EGR J2032+1226			4.06 ± 0.90	2.41 ± 0.12	BL Lac
1FGL J2035.4+1100	3EG J2036+1132				7.74 ± 1.11	2.68 ± 0.10	13.3 ± 3.1	2.83 ± 0.26	FSRQ
1FGL J2056.3–4714	3EG J2055–4716				17.00 ± 1.03	2.54 ± 0.05	23.6 ± 6.0	2.04 ± 0.35	FSRQ
1FGL J2158.8–3013	3EG J2158–3023	EGR J2200–3015			21.44 ± 0.70	1.91 ± 0.02	30.4 ± 7.7	2.35 ± 0.26	BL Lac
1FGL J2202.8+4216	3EG J2202+4217	EGR J2204+4225			16.81 ± 1.01	2.38 ± 0.04	39.9 ± 11.6	2.60 ± 0.28	BL Lac
1FGL J2212.1+2358	3EG J2209+2401				1.19 ± 0.50	2.13 ± 0.19	14.6 ± 4.2	2.48 ± 0.50	FSRQ
1FGL J2212.9+0654		EGRc J2215+0653			3.01 ± 0.61	2.33 ± 0.11	FSRQ
1FGL J2232.5+1144	3EG J2232+1147				14.70 ± 0.97	2.56 ± 0.05	19.2 ± 2.8	2.45 ± 0.14	FSRQ
1FGL J2235.7–4817		EGR J2233–4812			2.02 ± 0.56	2.43 ± 0.17	FSRQ
1FGL J2253.9+1608	3EG J2254+1601	EGR J2253+1606	GEV J2253+1622	1AGL J2254+1602	136.81 ± 1.74	2.47 ± 0.01	53.7 ± 4.0	2.21 ± 0.06	FSRQ
1FGL J2258.0–2757		EGR J2258–2745			5.61 ± 0.85	2.67 ± 0.10	FSRQ
1FGL J2323.5–0315	3EG J2321–0328				6.05 ± 0.80	2.45 ± 0.08	38.2 ± 10.1	...	FSRQ

Notes. Associated sources from the 3rd EGRET (3EG; Hartman et al. 1999), Revised EGRET (EGR; Casandjian & Grenier 2008), high-energy EGRET (GEV; Lamb & Macomb 1997), and one-year *AGILE* (1AGL; Pittori et al. 2009) catalogs. The LAT and EGRET fluxes and spectral indices are also provided.

^a In units of 10^{-8} photons $\text{cm}^{-2} \text{s}^{-1}$.

Table 7
Positional Associations with Hard X-ray Sources

LAT Source	Hard X-ray Source	R.A. ^a	Decl. ^a	Type
1FGL J0217.8+7353	1ES 0212+735	34.32489	73.822807	FSRQ
1FGL J0319.7+4130	NGC 1275	49.950871	41.501099	AGN
1FGL J0325.0+3403	B2 0321+33B	51.18626	34.176857	AGN
1FGL J0334.2+3233	NRAO 140	54.128605	32.305359	FSRQ
1FGL J0405.6–1309	PKS 0403–13	61.39616	–13.142807	FSRQ
1FGL J0522.8–3632	PKS 0521–36	80.746483	–36.458878	BLL
1FGL J0531.0+1331	PKS 0528+134	82.754601	13.562002	FSRQ
1FGL J0538.8–4404	PKS 0537–441	84.750519	–44.104771	BLL
1FGL J0539.1–2847	PKS 0537–286	84.951698	–28.645336	FSRQ
1FGL J0636.1–7521	PKS 0637–75	99.080544	–75.244789	FSRQ
1FGL J0710.6+5911	BZB J0710+5908	107.664192	59.14875	BLL
1FGL J0746.6+2548	B2 0743+25	116.609474	25.811897	FSRQ
1FGL J0750.6+1235	PKS 0748+126	117.666161	12.520621	FSRQ
1FGL J0806.2+6148	CGRaBS J0805+6144	121.314117	61.73798	FSRQ
1FGL J0842.2+7054	4C +71.07	130.395554	70.889877	FSRQ
1FGL J0949.0+0021	CGRaBS J0948+0022	147.228745	0.350572	FSRQ
1FGL J0956.5+6938	M82	148.94429	69.694511	AGN
1FGL J1048.7+8054	CGRaBS J1044+8054	161.060471	80.926788	FSRQ
1FGL J1103.7–2329	CRATES J1103–2329	165.899918	–23.477991	BLL
1FGL J1104.4+3812	Mkn 421	166.116058	38.208893	BLL
1FGL J1130.2–1447	PKS 1127–14	172.528671	–14.813811	FSRQ
1FGL J1136.2+6739	BZB J1136+6737	174.020081	67.657433	BLL
1FGL J1221.3+3008	B2 1218+30	185.328064	30.148579	BLL
1FGL J1222.5+0415	4C +04.42	185.589111	4.235839	FSRQ
1FGL J1224.7+2121	4C +21.35	186.224121	21.383167	FSRQ
1FGL J1229.1+0203	3C 273	187.27565	2.043433	FSRQ
1FGL J1256.2–0547	3C 279	194.041	–5.803322	FSRQ
1FGL J1305.4–4928	NGC 4945	196.353302	–49.473759	AGN
1FGL J1307.0–4030	ESO 323-G77	196.643127	–40.422039	AGN
1FGL J1325.6–4300	Cen A	201.362106	–43.026054	AGN
1FGL J1331.9–0506	PKS 1329–049	203.010651	–5.174589	FSRQ
1FGL J1417.8+2541	2E 1415+2557	214.479584	25.733086	BLL
1FGL J1428.7+4239	1ES 1426+428	217.147217	42.666458	BLL
1FGL J1442.8+1158	1ES 1440+122	220.704819	12.045237	BLL
1FGL J1512.8–0906	PKS 1510–08	228.205551	–9.086402	FSRQ
1FGL J1517.8–2423	AP Lib	229.457718	–24.370974	BLL
1FGL J1555.7+1111	PG 1553+113	238.886841	11.209181	BLL
1FGL J1626.2–2956	PKS 1622–29	246.539917	–29.818855	FSRQ
1FGL J1653.9+3945	Mkn 501	253.446915	39.768932	BLL
1FGL J1829.8+4845	3C 380	277.403839	48.753803	AGN
1FGL J1925.2–2919	PKS B1921–293	291.157959	–29.235388	FSRQ
1FGL J1938.2–3957	PKS 1933–400 ^b	294.294	–39.932	FSRQ
1FGL J2000.0+6508	1ES 1959+650	299.94632	65.150742	BLL
1FGL J2148.5+0654	4C +06.69	327.031952	6.948918	FSRQ
1FGL J2202.8+4216	BL Lac	330.72821	42.273628	BLL
1FGL J2229.7–0832	PKS 2227–08	337.407898	–8.519718	FSRQ
1FGL J2232.5+1144	CTA 102	338.135529	11.73317	FSRQ
1FGL J2253.9+1608	3C 454.3	343.483856	16.153084	FSRQ
1FGL J2327.7+0943	PKS 2325+093	351.892822	9.633142	FSRQ
1FGL J2359.0–3035	1H 2351–315	359.779633	–30.594545	BLL

Notes.

^a J2000 coordinate, in degrees, from the 54 month Palermo BAT catalog where available or from the fourth IBIS catalog otherwise.

^b Hard X-ray source found only in the fourth IBIS catalog.

sub-milliCrab sensitivity reached with *INTEGRAL* and *Swift* in the deepest fields.

6.4. Radio Properties

The high-latitude 1FGL sources associated with AGNs and presented here in the 1LAC are all radio sources at some level. The large number of sources found in CRATES can, by definition, be characterized as relatively bright, flat-spectrum radio

sources. However, since a significant number of associations are found from other catalogs, it is interesting to contemplate the distributions of the flux densities, luminosities, and spectral indices.

In Figure 26, we plot the distribution of the radio flux density at 8.4 GHz for the clean sample and separately for the FSRQs and the BL Lacs in this sample. The overall distribution, including also the blazars of unknown type and the other AGNs,

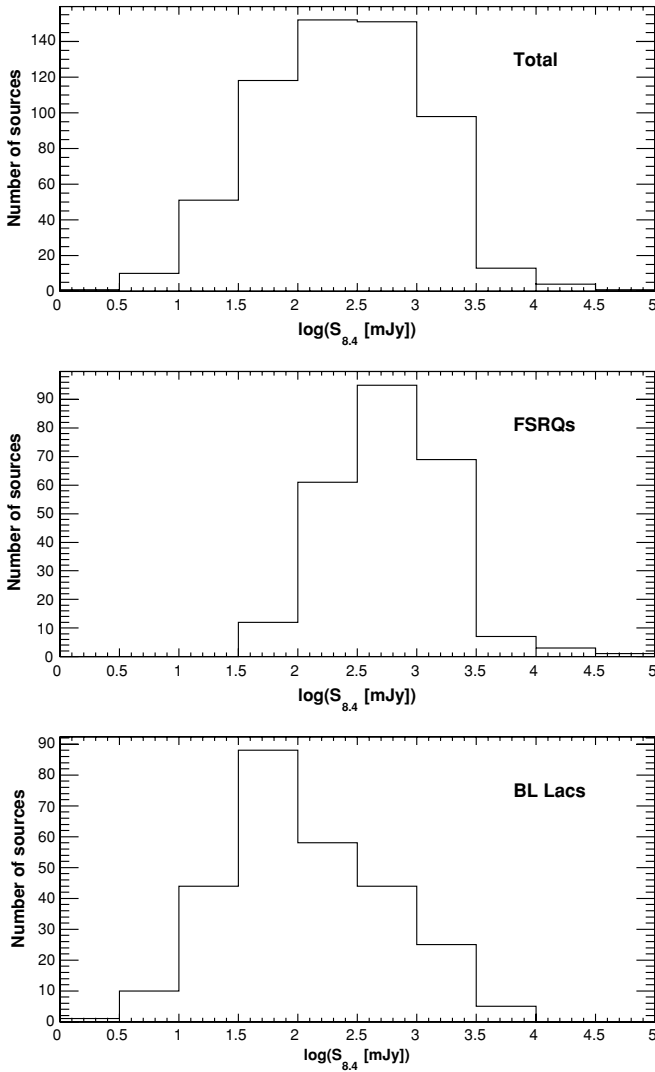


Figure 26. Distributions of radio flux densities at 8.4 GHz for all sources in the clean sample (top), FSRQs (middle), and BL Lacs (bottom).

has one main peak (corresponding to $\langle S_{8\text{ GHz, ILAC}} \rangle \sim 800$ mJy). However, the FSRQs are on average brighter than the BL Lacs, with $\langle S_{8\text{ GHz, FSRQ}} \rangle \sim 1300$ mJy and $\langle S_{8\text{ GHz, BLL}} \rangle \sim 400$ mJy.

In Figure 27, we plot the corresponding radio luminosities, calculated and K -corrected for the sources with measured redshifts. The distributions for FSRQs and BL Lacs are clearly different, with the FSRQs concentrated at higher radio luminosities ($\log(L_{r, \text{FSRQ}} (\text{erg s}^{-1})) = 44.1 \pm 0.7 \text{ erg s}^{-1}$) and the BL Lacs distributed over a broad range ($\log(L_{r, \text{BLL}} (\text{erg s}^{-1})) = 42.2 \pm 1.1 \text{ erg s}^{-1}$). This is similar to the distribution found in the LBAS, and the only sources with radio luminosities below $L_r = 10^{40} \text{ erg s}^{-1}$ are a few nearby AGNs (Cen A, NGC 253, and NGC 4945).

Finally, in Figure 28, we plot the spectral index distribution, calculated between either 1.4 GHz from NVSS for sources with $\delta > -40^\circ$ or 0.84 GHz from the Sydney University Molonglo Sky Survey (SUMSS; Mauch et al. 2003) for sources with $\delta < -40^\circ$ and 8.4 GHz from CRATES (or other measurements from NED). The overall distribution is consistent with a flat spectral index ($\alpha = 0.10 \pm 0.33$). No difference is found between FSRQs and BL Lacs. Interestingly, a tail of sources with steeper spectra is found, and it is mostly composed of non-blazars, such as radio galaxies. This suggests that γ -ray sources can also be

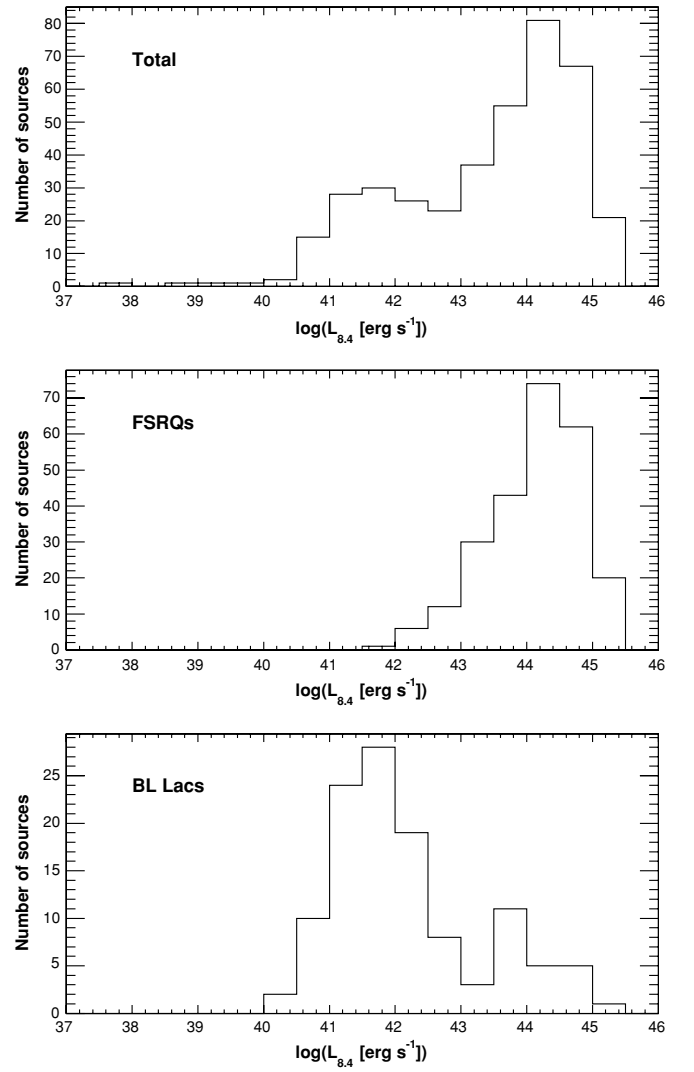


Figure 27. Distributions of radio luminosities at 8.4 GHz for all sources in the clean sample (top), FSRQs (middle), and BL Lacs (bottom).

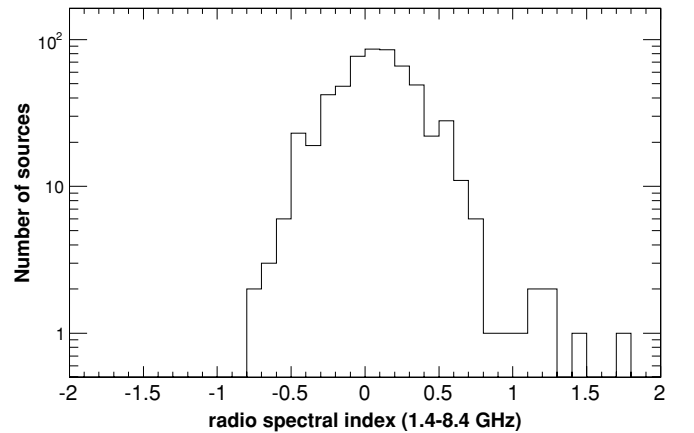


Figure 28. Distribution of radio spectral indices between 1.4 GHz (0.8 GHz for southern sources) and 8.4 GHz for sources in the clean sample.

associated with radio sources that are less closely aligned with the line of sight than blazars (see Section 4.1).

The relationship between the radio emission and the high-energy emission merits a deeper discussion, as highlighted by recent works discussing the apparent VLBA jet speed (Lister et al. 2009), opening angle (Pushkarev et al. 2009), and Doppler

factors (Savolainen et al. 2010). In particular, since the different AGN classes have different properties in the γ -ray band, particular caution is necessary as far as the correlation between radio and γ -ray fluxes for LAT sources is concerned (Kovalev et al. 2009; Giroletti et al. 2010). A paper dedicated to this subject is in preparation.

7. DISCUSSION

Of the 1043 high-latitude ($|b| > 10^\circ$) γ -ray sources in the 1FGL catalog, 671 are associated with 709 AGNs; these constitute the 1LAC. The 663 high-confidence AGN sources detected by *Fermi* in 11 months of data can be compared with the 66 high-confidence AGNs detected by EGRET at $>5\sigma$ significance over its lifetime (Hartman et al. 1999). We estimate that a large fraction of the other 372 high-latitude sources are likely AGNs on the basis of their spectral and variability properties; indeed, we have found plausible AGN candidates for ~ 100 of them (see Section 3.2.3). In addition, ~ 20 known millisecond pulsars, and possibly a comparable number of yet unidentified millisecond pulsars, are among the remaining high-latitude 1FGL sources.

The 1LAC represents a \sim five-fold increase in the number of sources associated with γ -ray blazars over previously published source lists (in particular, the 106 high-confidence blazars in the LBAS catalog). It is likely that other blazars remain unidentified, as implied by the observed north–south anisotropy of the associated sources, which reflects the incompleteness of the counterpart AGN catalogs. This is partly supported by the similarity of the γ -ray properties, including flux and photon spectral index, of the unidentified sources to those of the 1LAC blazars (A. A. Abdo et al. 2010b, in preparation). There are also indications from our early radio follow-up of the unidentified sources that the LAT is detecting blazars that are fainter in the radio than those appearing in the catalogs of flat-spectrum radio sources from which we draw candidate associations.

The overall properties of the 1LAC sources are generally consistent with those of the LBAS (Abdo et al. 2009f, 2010a, 2010b), including observations of:

1. *A small number of non-blazar sources.* The misaligned AGNs include six radio galaxies, of which only three (NGC 1275, Cen A, and M87) were previously reported. Five of the six radio galaxies are low-power FR-I galaxies associated with LAT sources with very high probabilities ($P > 98\%$). The sixth, NGC 6251, is a moderately powerful radio galaxy (Perley et al. 1984) and has a lower association probability. The misaligned FR-II radio galaxy 3C 111 is associated with a LAT source (see Table 2), but it is not a member of the 1LAC since it resides at low Galactic latitude. In addition, two steep-spectrum radio quasars and five NLRGs are among the other non-blazar sources, though many of these have prominent radio cores and some lack extended radio emission.
2. *Redshift distributions peaking at $z \approx 1$ for 1LAC FSRQs and at low redshift for 1LAC BL Lacs with known redshifts.* The maximum redshift of the high-confidence 1LAC FSRQs is $z = 3.10$, larger than the maximum redshift ($z = 2.286$) for an EGRET blazar. For comparison, the maximum redshift in both CGRaBS and Roma-BZCAT is for an FSRQ with $z > 5$. The redshift distribution for BL Lacs in the 1LAC with known redshift displays a broad low-redshift peak from the lowest known redshift (for Mkn 421) at $z = 0.030$ to $z \approx 0.4$. The good agreement with the *WMAP*

redshift distribution for FSRQs, with a comparable number of sources, supports the idea that similar populations are being sampled by *WMAP* and the LAT. Both the FSRQ and BL Lac redshift distributions are compatible with the corresponding distributions from LBAS; the K–S probabilities that they derive from the same parent distributions are 0.99 and 0.89 for FSRQs and BL Lacs, respectively.

No strong evidence for a new population of misaligned FSRQs emerging at lower redshifts is found. Other than NGC 6251, just mentioned, notably absent are detections, at least at the 1LAC significance level of $TS > 25$, of nearby powerful FR-II radio galaxies such as Cyg A and Pic A, though the low-latitude FR-II radio galaxy 3C 111 is detected by the LAT.

3. *A high BL Lac/FSRQ ratio, close to unity.* This ratio is even higher than that found for the LBAS sources, which comprise 42 BL Lacs and 57 FSRQs. It stands in sharp contrast to the population of 3EG blazars, in which FSRQs outnumber BL Lacs by a factor of 3.
4. *A high HSP/LSP ratio among BL Lacs.* The large HSP/LSP ratio for BL Lacs is a result of the fact that the sample is significance limited. The LAT detects hard-spectrum sources at higher significance than soft-spectrum sources with comparable photon number fluxes. Moreover, the LAT is far more sensitive to multi-GeV photons than was EGRET, which lost sensitivity above several GeV due to self-vetoing effects. Consequently, HSP-BL Lacs, with harder spectra than LSP-BL Lacs, are more prevalent in the 1LAC than in the 3EG catalog.
5. *Little evidence for different variability properties for FSRQs and BL Lacs.* No significant differences in the variability of FSRQs and BL Lacs are observed for sources at comparable fluxes. Because HSP blazars are generally observed at low-flux levels, however, the comparison of variability properties primarily concerns LSP/ISP-BL Lacs and FSRQs.
6. *A fairly strong correlation between photon spectral index and blazar class for the detected sources.* FSRQs, which are almost all LSP blazars, are found to be soft, while BL Lacs, whether of the LSP, ISP, or HSP type, represent on average a population of harder-spectrum sources. The average photon spectral index (Γ) continuously shifts to lower values (i.e., harder spectra) as the class varies from FSRQs ($\langle \Gamma \rangle = 2.48$) to LSP-BL Lacs ($\langle \Gamma \rangle = 2.28$), ISP-BL Lacs ($\langle \Gamma \rangle = 2.13$), and HSP-BL Lacs ($\langle \Gamma \rangle = 1.96$).

Furthermore, as noted by Ghisellini et al. (2009) and seen in Figure 24, the γ -ray blazars detected with the *Fermi*-LAT separate into hard-spectrum, low-luminosity sources, primarily consisting of BL Lac objects, and high-luminosity, soft-spectrum sources, primarily consisting of FSRQs and LSP-BL Lacs. Compared with the LBAS sample, the blazar divide now displays an extension of the sample toward low-luminosity objects in each of the blazar subclasses. In addition, we identify a third branch consisting of radio galaxies, which are distinguished from the aligned blazars by their soft spectra and very low γ -ray luminosities.

The finding that FSRQs in the 1LAC almost all have soft ($\Gamma \gtrsim 2.2$) γ -ray spectra suggests a connection between the presence of strong emission lines and the nonthermal electron maximum Lorentz factor or spectral shape. It is consistent with scenarios (Ghisellini et al. 1998; Böttcher & Dermer 2002) in which the strong ambient radiation fields, as revealed by the presence of emission lines, affects the acceleration and cooling

of particles and controls the formation of the blazar SED. The luminosity dependence of the γ -ray spectral hardness is also consistent with the blazar sequence (Ghisellini et al. 1998; Fossati et al. 1998; Ghisellini & Tavecchio 2008), in which the frequencies of the synchrotron and Compton peaks are inversely correlated with the apparent isotropic luminosities of the blazars. This behavior has been interpreted in terms of an evolution of FSRQs into BL Lac objects as the supply of matter fueling the accretion disk declines, the surrounding external radiation fields become less intense, and the jet weakens (Böttcher & Dermer 2002; Cavaliere & D’Elia 2002; Ghisellini & Tavecchio 2008).

Important caveats must, however, be considered before drawing firm conclusions about the blazar sequence from the 1LAC sources. First, identification of the blazars used to construct the blazar sequence is biased by the flux limits and frequency ranges of the underlying radio and X-ray samples (Giommi et al. 1999; Padovani et al. 2003), and different conclusions about its validity can be drawn depending on the catalogs used to examine the sequence (Nieppola et al. 2006). Second, specific sources that do not conform to the sequence have been identified (Caccianiga & Marchã 2004; Giommi et al. 2007). Third, the 1LAC is significance limited rather than flux limited, which results in preferential detection of lower-flux, harder-spectrum sources. This means that there is a strong bias against detecting weak, soft-spectrum γ -ray sources from low-luminosity LSP blazars that might violate the sequence. On the other hand, hard-spectrum HSP sources should be more easily detected, but there are no high-luminosity HSP sources detected in the LAT band (see Figure 24), as all high-luminosity objects have SEDs of the LSP type. This may, however, reflect the difficulty of measuring the redshifts of BL Lacs, despite efforts even with 10 m class telescopes (M. S. Shaw et al. 2010, in preparation). Even for sources associated with AGNs with high probabilities, a large fraction lack redshift measurements, including 58% of the high-confidence BL Lacs and ~ 30 BL Lacs with $\Gamma \lesssim 1.9$. The recent work of Plotkin et al. (2010) shows that many BL Lacs might have redshifts larger than 2, up to $z = 4.9$ (though these tend to be radio-faint, X-ray-faint objects). Considering this result, and the indication that BL Lacs with unknown redshifts may be found preferentially at higher redshift (see Section 5.3, Plotkin et al. 2010, and Meisner & Romani 2010), it is possible that a number of distant, high-luminosity HSP-BL Lacs are included in the 1LAC sample. According to our SED classification, ~ 20 BL Lacs with unknown redshifts are already classified as HSPs. If they were all at high redshift (e.g., $z = 2$), and excluding those already detected at VHE (which indicates a low redshift due to a lack of EBL absorption), ~ 7 objects could have γ -ray luminosities $L_\gamma \gtrsim 10^{48}$ erg s $^{-1}$.

The existence of such sources, though quite rare, would be contrary to expectations from the blazar sequence (Ghisellini et al. 1998; Fossati et al. 1998; Ghisellini & Tavecchio 2008). In the absence of redshift measurements of possible high-luminosity, hard-spectrum sources, we can only conclude that the *Fermi* results seem compatible with the sequence. Identifying hard-spectrum blazar sources at high redshifts is also important for studies of the cosmological evolution of the EBL (A. A. Abdo et al. 2010c, in preparation), which are hampered by their absence.

Related to this, it is interesting to compare the redshift distribution of the LAT blazars (in particular, the FSRQs) with that of the objects detected by BAT on *Swift* (Ajello et al. 2009). Despite a fourfold increase in exposure time with respect to the LBAS sample and a consideration of γ -ray sources with

TS > 25 in the 1LAC compared with TS > 100 in LBAS, the number of LAT-detected blazars still drops sharply at $z \sim 2.5$, and only two of the high-confidence 1LAC blazars have $z > 3$ (keeping in mind biases from the lack of redshift measurements). This is very different from the BAT survey, in which half of all FSRQs have $z > 2$, and the distribution extends to $z \sim 4$. This behavior may be indicative of a shift in the SED peak frequencies toward lower values (i.e., a “redder” SED) for blazars at high redshifts. Indeed, the jets of the high-redshift BAT blazars are found to be more powerful than those of the LAT blazars and are among the most powerful known (Ghisellini et al. 2010). The peak of the inverse Compton flux for these objects, estimated to be in the MeV or even sub-MeV range, is located closer to the BAT band than to the LAT band, and the LAT instead samples the cutoff region of the inverse Compton spectrum.

Comparison of *Fermi* results with pre-launch *GLAST* expectations depends on how the 1LAC results are characterized. Except for a few dozen millisecond pulsars, nearly all of the 1043 high-latitude γ -ray sources in the 1FGL are likely to be AGNs, but only 671 high-latitude 1FGL sources are associated with AGNs. Adopting a radio/ γ -ray correlation, Stecker & Salamon (1996) and Chiang & Mukherjee (1998) predicted many thousands of blazars to a flux limit of 2×10^{-9} photons cm $^{-2}$ s $^{-1}$ (for $E > 100$ MeV) reached after five years for a source with $\Gamma = 2.1$, as did Narumoto & Totani (2006), who treated the problem using radio and X-ray luminosity functions. Mücke & Pohl (2000), using a physical blazar model matched to the EGRET data, predicted an increased fraction of BL Lac objects at fainter flux levels, with nearly 2000 BL Lacs and 1000 FSRQs at the one-year flux limit of 4×10^{-9} photons cm $^{-2}$ s $^{-1}$ (for $E > 100$ MeV). Dermer (2007) predicted that ~ 800 FSRQs and ~ 200 BL Lacs would be detected after one year, also employing a physical blazar model. Assuming the validity of the blazar sequence, Inoue & Totani (2009) predicted ~ 600 – 1200 blazars to a flux limit of 2×10^{-9} photons cm $^{-2}$ s $^{-1}$ (for $E > 100$ MeV). However, predictions based on a minimum flux depend sensitively on the source spectral index, as shown by Figure 9, so more reliable predictions needed to take this effect into account.

The populations of AGNs that can be studied jointly at GeV and at TeV energies have increased dramatically beyond what could have been expected based on the experience of the EGRET era. There is now a large and increasing population of GeV–TeV AGNs that can be best studied through observations with both *Fermi* and the ground-based γ -ray observatories. Since its launch, the number of TeV AGNs detected also by *Fermi* has steadily increased as its exposure on these sources grows. Complementing this, the TeV observatories have started to observe the more promising *Fermi* AGNs, and a number of detections have resulted. One promising feature of joint GeV–TeV observations is the power to infer the redshifts of BL Lacs from spectral modeling, as in the case of PG 1553+113 (Abdo et al. 2010g).

8. CONCLUSIONS

We have presented the 1LAC, consisting of 671 high-latitude γ -ray sources with TS > 25 measured in the first 11 months of the *Fermi*-LAT sky survey. Due to multiple associations, these sources are associated with 709 AGNs. Besides a growing number of blazars and six radio galaxies, the extragalactic γ -ray sky now includes a small number of steep-spectrum quasars and NLRGs. Confirming many of the LBAS results, we find a strong correlation between the 0.1–100 GeV γ -ray spectral index and the blazar type, whether defined in terms of optical emission line

strength or peak frequency of the synchrotron component of the SED. The relationship between the γ -ray spectral properties and the spectral luminosity is complicated by selection biases for source identification and redshift determination, but the *Fermi* results seem compatible with the blazar sequence.

The 1LAC represents significant progress in terms of the number of detected sources, the diversity of source types, and the accuracy of measured γ -ray properties. This sample is likely to evolve in the future as more unidentified sources can be associated with AGNs thanks to better counterpart catalogs and continued correlated variability studies with more LAT observations. Bearing in mind the instrumental limitations described above, the 1LAC provides an essential foundation on which to build a much better understanding of the population of γ -ray-emitting AGNs.

The *Fermi* LAT Collaboration acknowledges generous ongoing support from a number of agencies and institutes that have supported both the development and the operation of the LAT as well as scientific data analysis. These include the National Aeronautics and Space Administration and the Department of Energy in the United States; the Commissariat à l'Énergie Atomique and the Centre National de la Recherche Scientifique/Institut National de Physique Nucléaire et de Physique des Particules in France; the Agenzia Spaziale Italiana and the Istituto Nazionale di Fisica Nucleare in Italy; the Ministry of Education, Culture, Sports, Science and Technology (MEXT), High Energy Accelerator Research Organization (KEK) and Japan Aerospace Exploration Agency (JAXA) in Japan; and the K. A. Wallenberg Foundation, the Swedish Research Council and the Swedish National Space Board in Sweden.

Additional support for science analysis during the operations phase is gratefully acknowledged from the Istituto Nazionale di Astrofisica in Italy and the Centre National d'Études Spatiales in France.

This research has made use of the NASA/IPAC Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration. Part of this work is based on archival data, software, or online services provided by the ASI Science Data Center (ASDC).

Some of the results presented in this paper are based on observations obtained with the Hobby–Eberly Telescope (HET), a joint project of the University of Texas at Austin, the Pennsylvania State University, Stanford University, Ludwig-Maximilians-Universität München, and Georg-August-Universität Göttingen. The HET is named in honor of its principal benefactors, William P. Hobby and Robert E. Eberly. The Marcario Low-Resolution Spectrograph (LRS) is named for Mike Marcario of High-Low-Resolution Optics, who fabricated several optics for the instrument but died before its completion. The LRS is a joint project of the Hobby–Eberly Telescope partnership and the Instituto de Astronomía de la Universidad Nacional Autónoma de México.

This work is also partly based on optical spectroscopy performed at the Telescopio Nazionale Galileo (TNG), La Palma, Canary Islands (proposal AOT20/09B). These observations confirm some of the redshifts and classifications found with the HET. We thank the HET and TNG personnel for their assistance during the observing runs.

The data in this paper are based partly on observations obtained at the Hale Telescope, Palomar Observatory, as part of a collaborative agreement between the California Institute of Technology, its divisions Caltech Optical Observatories and the

Jet Propulsion Laboratory (operated for NASA), and Cornell University.

Some of the data presented herein were obtained at the W. M. Keck Observatory, which is operated as a scientific partnership among the California Institute of Technology, the University of California, and the National Aeronautics and Space Administration. The Observatory was made possible by the generous financial support of the W. M. Keck Foundation.

The authors wish to recognize and acknowledge the very significant cultural role and reverence that the summit of Mauna Kea has always had within the indigenous Hawaiian community. We are most fortunate to have the opportunity to conduct observations from this mountain.

Some of the data in this paper are based on observations made with ESO telescopes at the La Silla Observatory under programs 083.B-0460(B) and 084.B-0711(B).

The National Radio Astronomy Observatory is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc.

Funding for the Sloan Digital Sky Survey (SDSS) and SDSS-II has been provided by the Alfred P. Sloan Foundation, the Participating Institutions, the National Science Foundation, the U. S. Department of Energy, the National Aeronautics and Space Administration, the Japanese Monbukagakusho, the Max Planck Society, and the Higher Education Funding Council for England. The SDSS Web site is <http://www.sdss.org/>.

The SDSS is managed by the Astrophysical Research Consortium (ARC) for the Participating Institutions. The Participating Institutions are the American Museum of Natural History, Astrophysical Institute Potsdam, University of Basel, University of Cambridge, Case Western Reserve University, The University of Chicago, Drexel University, Fermilab, the Institute for Advanced Study, the Japan Participation Group, The Johns Hopkins University, the Joint Institute for Nuclear Astrophysics, the Kavli Institute for Particle Astrophysics and Cosmology, the Korean Scientist Group, the Chinese Academy of Sciences (LAMOST), Los Alamos National Laboratory, the Max-Planck-Institute for Astronomy (MPIA), the Max-Planck-Institute for Astrophysics (MPA), New Mexico State University, Ohio State University, University of Pittsburgh, University of Portsmouth, Princeton University, the United States Naval Observatory, and the University of Washington.

Facilities: *Fermi*(LAT)

REFERENCES

- Abdo, A. A., et al. 2009a, *ApJ*, **699**, 817 (3C 454.3)
 Abdo, A. A., et al. 2009b, *ApJ*, **697**, 934 (PKS 1454–354)
 Abdo, A. A., et al. 2009c, *ApJ*, **699**, 31 (NGC 1275)
 Abdo, A. A., et al. 2009d, *ApJ*, **707**, 55 (M87)
 Abdo, A. A., et al. 2009e, *ApJ*, **699**, 976 (PMN J0948+0022)
 Abdo, A. A., et al. 2009f, *ApJ*, **700**, 597 (LBAS)
 Abdo, A. A., et al. 2009g, *ApJS*, **183**, 46 (BSL)
 Abdo, A. A., et al. 2009h, *ApJ*, **707**, L142 (γ -ray emission from radio-loud narrow-line Seyfert 1s)
 Abdo, A. A., et al. 2009i, *ApJ*, **707**, 1310 (*Fermi* observations of TeV AGNs)
 Abdo, A. A., et al. 2010a, *ApJ*, **710**, 1271 (GeV spectral properties of LBAS blazars)
 Abdo, A. A., et al. 2010b, *ApJ*, submitted (arXiv:0912.2040; SEDs of *Fermi* bright blazars)
 Abdo, A. A., et al. 2010c, *ApJ*, submitted (arXiv:1002.2280; LAT first-year source catalog)
 Abdo, A. A., et al. 2010d, *ApJ*, **709**, L152 (starburst galaxies)
 Abdo, A. A., et al. 2010e, *ApJ*, submitted (arXiv:1003.0895; blazar population studies)

- Abdo, A. A. 2010f, *ApJ*, submitted (arXiv:1004.0348; variability properties of LBAS blazars)
- Abdo, A. A., et al. 2010g, *ApJ*, 708, 1310 (PG 1553+113)
- Acciari, V. A., et al. 2010, *ApJ*, 708, L100
- Acleo, F., et al. 2010, *A&A*, 511, A52
- Adelman-McCarthy, J. K., et al. 2008, *ApJS*, 175, 297
- Aharonian, F., et al. 2009, *ApJ*, 696, L150
- Ajello, M., et al. 2009, *ApJ*, 699, 603
- Antonucci, R. 1993, *ARA&A*, 31, 473
- Arp, H. 1997, *A&A*, 327, 479
- Atwood, W. B., et al. 2009, *ApJ*, 697, 1071
- Beasley, A. J., Gordon, D., Peck, A. B., Petrov, L., MacMillan, D. S., Fomalont, E. B., & Ma, C. 2002, *ApJS*, 141, 13
- Bennett, A. S. 1962, *MNRAS*, 125, 75
- Bird, A. J., et al. 2010, *ApJS*, 186, 1
- Blandford, R. D., & Rees, M. J. 1978, in Proc. Pittsburgh Conf. on BL Lac Objects, ed. A. M. Wolfe (Pittsburgh, PA: Univ. Pittsburgh), 328
- Böttcher, M., & Dermer, C. D. 2002, *ApJ*, 564, 86
- Browne, I. W. A., et al. 2003, *MNRAS*, 341, 13
- Burgess, A. M., & Hunstead, R. W. 2006, *AJ*, 131, 100
- Caccianiga, A., & Marchã, M. J. M. 2004, *MNRAS*, 348, 937
- Caccianiga, A., Marchã, M. J., Antón, S., Mack, K., & Neeser, M. J. 2002, *MNRAS*, 329, 877
- Campana, R., Massaro, E., Gasparini, D., Cutini, S., & Tramacere, A. 2008, *MNRAS*, 383, 1166
- Casandjian, J., & Grenier, I. A. 2008, *A&A*, 489, 849
- Cavaliere, A., & D'Elia, V. 2002, *ApJ*, 571, 226
- Chen, A., Reyes, L. C., & Ritz, S. 2004, *ApJ*, 608, 686
- Chiang, J., & Mukherjee, R. 1998, *ApJ*, 496, 752
- Ciprini, S., et al. 2007, in AIP Conf. Ser. 921, The First GLAST Symposium, ed. S. Ritz, P. Michelson, & C. A. Meegan (Melville, NY: AIP), 546
- Condon, J. J., Cotton, W. D., Greisen, E. W., Yin, Q. F., Perley, R. A., Taylor, G. B., & Broderick, J. J. 1998, *AJ*, 115, 1693
- Cusumano, G., et al. 2010, *A&A*, 510, A48
- de Ruiter, H. R., Arp, H. C., & Willis, A. G. 1977, *A&AS*, 28, 211
- Dermer, C. D. 2007, *ApJ*, 659, 958
- Engelbracht, C. W., Rieke, M. J., Rieke, G. H., Kelly, D. M., & Achtermann, J. M. 1998, *ApJ*, 505, 639
- Falco, E. E., Kochanek, C. S., & Munoz, J. A. 1998, *ApJ*, 494, 47
- Fichtel, C. E., et al. 1994, *ApJS*, 94, 551
- Fomalont, E. B., Petrov, L., MacMillan, D. S., Gordon, D., & Ma, C. 2003, *AJ*, 126, 2562
- Förster Schreiber, N. M., Roussel, H., Sauvage, M., & Charmandaris, V. 2004, *A&A*, 419, 501
- Fossati, G., Maraschi, L., Celotti, A., Comastri, A., & Ghisellini, G. 1998, *MNRAS*, 299, 433
- Ghisellini, G., Celotti, A., Fossati, G., Maraschi, L., & Comastri, A. 1998, *MNRAS*, 301, 451
- Ghisellini, G., Maraschi, L., & Tavecchio, F. 2009, *MNRAS*, 396, L105
- Ghisellini, G., & Tavecchio, F. 2008, *MNRAS*, 387, 1669
- Ghisellini, G., et al. 2010, *MNRAS*, in press (arXiv:0912.0001)
- Giommi, P., Colafrancesco, S., Padovani, P., Gasparini, D., Cavazzuti, E., & Cutini, S. 2009, *A&A*, 508, 107
- Giommi, P., Menna, M. T., & Padovani, P. 1999, *MNRAS*, 310, 465
- Giommi, P., et al. 2007, *A&A*, 468, 97
- Giroletti, M., Reimer, A., Fuhrmann, L., Pavlidou, V., & Richards, J. L. 2010, arXiv:1001.5123
- Hartman, R. C., et al. 1999, *ApJS*, 123, 79
- Healey, S. E., Romani, R. W., Taylor, G. B., Sadler, E. M., Ricci, R., Murphy, T., Ulvestad, J. S., & Winn, J. N. 2007, *ApJS*, 171, 61
- Healey, S. E., et al. 2008, *ApJS*, 175, 97
- Hill, G. J., Nicklas, H. E., MacQueen, P. J., Tejada, C., Cobos Duenas, F. J., & Mitsch, W. 1998, *Proc. SPIE*, 3355, 375
- Hinshaw, G., et al. 2007, *ApJS*, 170, 288
- Hofmann, W., & Fegan, S. 2009, *ATel*, 2293, 1
- Inoue, Y., & Totani, T. 2009, *ApJ*, 702, 523
- Jones, D. H., et al. 2004, *MNRAS*, 355, 747
- Jones, D. H., et al. 2009, *MNRAS*, 399, 683
- Komatsu, E., et al. 2009, *ApJS*, 180, 330
- Kovalev, Y. Y. 2009, *ApJ*, 707, L56
- Kovalev, Y. Y., Petrov, L., Fomalont, E. B., & Gordon, D. 2007, *AJ*, 133, 1236
- Kovalev, Y. Y., et al. 2009, *ApJ*, 696, L17
- Laing, R. A., Riley, J. M., & Longair, M. S. 1983, *MNRAS*, 204, 151
- Lamb, R. C., & Macomb, D. J. 1997, *ApJ*, 488, 872
- Landt, H., Padovani, P., Perlman, E. S., & Giommi, P. 2004, *MNRAS*, 351, 83
- Lister, M. L., Homan, D. C., Kadler, M., Kellermann, K. I., Kovalev, Y. Y., Ros, E., Savolainen, T., & Zensus, J. A. 2009, *ApJ*, 696, L22
- Lott, B., Ballet, J., Chiang, J., Lonjou, V., & Funk, S. 2007, in AIP Conf. Ser. 921, The First GLAST Symposium, ed. S. Ritz, P. Michelson, & C. A. Meegan (Melville, NY: AIP), 558
- Malmquist, K. 1920, *Medd. Lunds Astron. Obs. Ser.*, 2, 22
- Maraschi, L., Ghisellini, G., & Celotti, A. 1992, *ApJ*, 397, L5
- Marchã, M. J. M., Browne, I. W. A., Impey, C. D., & Smith, P. S. 1996, *MNRAS*, 281, 425
- Mason, K. O., et al. 2000, *MNRAS*, 311, 456
- Massaro, E., Giommi, P., Leto, C., Marchegiani, P., Maselli, A., Perri, M., Piranomonte, S., & Sclavi, S. 2009, *A&A*, 495, 691
- Mattox, J. R., Hartman, R. C., & Reimer, O. 2001, *ApJS*, 135, 155
- Mattox, J. R., et al. 1996, *ApJ*, 461, 396
- Mauch, T., Murphy, T., Buttery, H. J., Curran, J., Hunstead, R. W., Piestrzynski, B., Robertson, J. G., & Sadler, E. M. 2003, *MNRAS*, 342, 1117
- Mazin, D., Lindfors, E., Berger, K., Galante, N., Prandini, E., & Saito, T. 2009, arXiv:0907.0366
- Meisner, A. M., & Romani, R. W. 2010, *ApJ*, 712, 14
- Mitton, S., Hazard, C., & Whelan, J. A. J. 1977, *MNRAS*, 179, 569
- Monet, D. G., et al. 2003, *AJ*, 125, 984
- Mücke, A., & Pohl, M. 2000, *MNRAS*, 312, 177
- Myers, S. T., et al. 2003, *MNRAS*, 341, 1
- Narumoto, T., & Totani, T. 2006, *ApJ*, 643, 81
- Nieppola, E., Tornikoski, M., & Valtaoja, E. 2006, *A&A*, 445, 441
- Ong, R. A. 2009a, *ATel*, 2301, 1
- Ong, R. A. 2009b, *ATel*, 2309, 1
- Ong, R. A., & Fortin, P. 2009, *ATel*, 2272, 1
- Padovani, P., & Giommi, P. 1995, *ApJ*, 444, 567
- Padovani, P., Perlman, E. S., Landt, H., Giommi, P., & Perri, M. 2003, *ApJ*, 588, 128
- Perley, R. A., Bridle, A. H., & Willis, A. G. 1984, *ApJS*, 54, 291
- Petrov, L., Kovalev, Y. Y., Fomalont, E. B., & Gordon, D. 2005, *AJ*, 129, 1163
- Petrov, L., Kovalev, Y. Y., Fomalont, E. B., & Gordon, D. 2006, *AJ*, 131, 1872
- Petrov, L., Kovalev, Y. Y., Fomalont, E. B., & Gordon, D. 2008, *AJ*, 136, 580
- Pittori, C., et al. 2009, *A&A*, 506, 1563
- Plotkin, R. M., et al. 2010, *AJ*, 139, 390
- Pushkarev, A. B., Kovalev, Y. Y., Lister, M. L., & Savolainen, T. 2009, *A&A*, 507, L33
- Raue, M., Behera, B., Charbonnier, A., Giebels, B., Hauser, M., Kosack, K., PUNCH, M., & Zechlin, H. 2009, *ATel*, 2350, 1
- Reyes, L. C., & Cheung, C. C. 2009, *ATel*, 2226, 1
- Savolainen, T., Homan, D. C., Hovatta, T., Kadler, M., Kovalev, Y. Y., Lister, M. L., Ros, E., & Zensus, J. A. 2010, *A&A*, 512, 24
- Sowards-Emmerd, D., Romani, R. W., & Michelson, P. F. 2003, *ApJ*, 590, 109
- Spinrad, H., Marr, J., Aguilar, L., & Djorgovski, S. 1985, *PASP*, 97, 932
- Starck, J.-L., & Pierre, M. 1998, *A&AS*, 128, 397
- Stecker, F. W., & Salamon, M. H. 1996, *ApJ*, 464, 600
- Stoeckel, J. T., Morris, S. L., Gioia, I. M., Maccacaro, T., Schild, R., Wolter, A., Fleming, T. A., & Henry, J. P. 1991, *ApJS*, 76, 813
- Sutherland, W., & Saunders, W. 1992, *MNRAS*, 259, 413
- Tavani, M., et al. 2008, *Nucl. Instrum. Methods Phys. Res. A*, 588, 52
- Thompson, D. J., et al. 1993, *ApJS*, 86, 629
- Ubertini, P., Sguera, V., Stephen, J. B., Bassani, L., Bazzano, A., & Bird, A. J. 2009, *ApJ*, 706, L7
- Unwin, S. C., Cohen, M. H., Pearson, T. J., Seielstad, G. A., Simon, R. S., Linfield, R. P., & Walker, R. C. 1983, *ApJ*, 271, 536
- Urry, C. M., & Padovani, P. 1995, *PASP*, 107, 803
- Vermeulen, R. C., Taylor, G. B., Readhead, A. C. S., & Browne, I. W. A. 1996, *AJ*, 111, 1013
- Véron-Cetty, M., & Véron, P. 1993, *A&AS*, 100, 521
- Véron-Cetty, M., & Véron, P. 2006, *A&A*, 455, 773
- Wills, K. A., Morganti, R., Tadhunter, C. N., Robinson, T. G., & Villar-Martin, M. 2004, *MNRAS*, 347, 771
- Wills, K. A., Pedlar, A., Muxlow, T. W. B., & Stevens, I. R. 1999, *MNRAS*, 305, 680
- Zensus, J. A., Cohen, M. H., & Unwin, S. C. 1995, *ApJ*, 443, 35