

IL NUOVO CIMENTO
DOI 10.1393/ncc/i2005-10144-y

VOL. 28 C, N. 4-5

Luglio-Ottobre 2005

Novel X-ray telescopes for wide-field X-ray monitoring^(*)

R. HUDEC⁽¹⁾, A. INNEMAN⁽³⁾, L. PINA⁽²⁾ and L. ŠVÉDA⁽²⁾

⁽¹⁾ *Astronomical Institute, Academy of Sciences of the Czech Republic
CZ-251 65 Ondřejov, Czech Republic*

⁽²⁾ *Faculty of Nuclear Science, Czech Technical University - Prague, Czech Republic*

⁽³⁾ *Center for Advanced X-ray Technologies, Reflex sro - Prague, Czech Republic*

(ricevuto il 23 Maggio 2005; pubblicato online il 13 Ottobre 2005)

Summary. — We report on fully innovative very wide-field of view X-ray telescopes with high sensitivity as well as large field of view. The prototypes are very promising, allowing the proposals for space projects with very wide-field Lobster-eye X-ray optics to be considered. The novel telescopes will monitor the sky with unprecedented sensitivity and angular resolution of order of 1 arcmin. They are expected to contribute essentially to study and to understand various astrophysical objects such as AGN, SNe, Gamma-ray bursts (GRBs), X-ray flashes (XRFs), galactic binary sources, stars, CVs, X-ray novae, various transient sources, etc. The Lobster optics based X-ray All Sky Monitor is capable to detect around 20 GRBs and 8 XRFs yearly and this will surely significantly contribute to the related science.

PACS 95.55.Ka – X- and γ -ray telescopes and instrumentation.

PACS 01.30.Cc – Conference proceedings.

Introduction

The detection of Gamma Ray Bursts (GRBs) in X-rays confirms the feasibility of monitoring, detecting and study of these phenomena by their X-ray emission (either prompt or afterglow, *e.g.*, [1, 2]). For classical GRBs, the X-ray afterglows are detected in $\sim 90\%$ of the cases [3]. Moreover, there are X-ray-rich GRBs, (hypothetical) orphan GRBs (detectable in X-rays but not in gamma-rays due to different beaming angle) and XRFs which can be detected and studied in X-rays. However, since these events cannot be predicted, and are relatively rare, very wide-field instruments are required. They must achieve high sensitivities and provide precise localizations in order to effectively study the objects. Wide-field X-ray telescopes with imaging optics are expected to represent an important tool in future space astronomy projects in general, especially those for deep monitoring and surveys in X-rays over a wide energy range. The Lobster-Eye wide-field

^(*) Paper presented at the “4th Workshop on Gamma-Ray Burst in the Afterglow Era”, Rome, October 18-22, 2004.

TABLE I. – *LE Schmidt modules developed so far by our collaboration. Here plates have dimensions of $d \times l \times t$ and are arranged with spacing a . The modules have focal length f and field of view FOV.*

Module	size d [mm]	thickness t [mm]	distance a [mm]	length l [mm]	foc. length f [mm]	resolution r [arcmin]	FOV [°]	energy [keV]
macro	300	0.75	10.80	300	6000	7	16	3
middle	80	0.3	2	80	400	20	12	2
mini 1	24	0.1	0.3	30	900	2	5	5
mini 2	24	0.1	0.3	30	250	6	5	5
micro	3	0.03	0.07	14	80	4	3	10

X-ray optics has been suggested in 70ies by Schmidt [4], orthogonal stacks of reflectors) and by Angel ([5], array of square cells). Up to 180 deg FOV may be achieved. This novel X-ray optics offers an excellent opportunity to achieve very wide fields of view (FOV, 1000 square degrees and more) while the widely used classical Wolter grazing incidence mirrors are limited to roughly 1 deg FOV [6, 7].

Lobster-Eye X-ray telescopes

Two basic types of Lobster-Eye Wide-Field X-ray telescopes have been proposed. The telescopes in Schmidt arrangements are based on perpendicular arrays of double-sided X-ray reflecting flats. In the first prototypes developed and tested, double-sided reflecting flats produced by epoxy sandwich technology as well as gold coated glass foils have been used [8] (table I). More recently, micro Schmidt lobster eye arrays with foils thickness as low as 30 microns have been developed and tested in order to confirm the capability of these systems to achieve fine angular resolutions of order of a few arcmin. The thin



Fig. 1. – The large Lobster-Eye telescope prototype (300×300 mm plates, Schmidt arrangement) represents the largest Lobster telescope test module constructed so far.

foils are separated by 70 microns gaps in these prototypes. On the other hand, large lobster eye systems with Schmidt geometry have been designed and constructed, achieving dimensions up to $300 \times 300 \times 600$ mm (fig. 1). Their optical tests have confirmed the expected performance according to calculations (computer ray-tracing). The calculations and the measurement results indicate that the lobster eye telescope based on multi array of modules with thin and closely spaced glass foils (analogous to those already assembled and tested) can meet the requirements, *e.g.*, of the ESA ISS Lobster mission (including the angular resolution and with better transmission) and can hence represent an alternative to the recently suggested MCP technique [9].

For the alternative Angel Lobster lenses, numerous square cells of very small size (about 1×1 mm or less at lengths of order of tens of mm, *i.e.* with the length/size ratio of 30 and more) are to be produced. This demand can be also solved by modified innovative replication technology. Test modules with LE Angel cells have been successfully produced. The linear test module has 47 cells 2.5×2.5 mm, 120 mm long (*i.e.* length/size ratio of almost 50), surface microroughness 0.8 nm, $f = 1300$ mm. Another test module is represented by a L-shaped array of $2 \times 18 = 36$ cells of analogous dimension. The surface microroughness of the replicated reflecting surfaces is better than 1 nm.

From the technological point of view, the fact that the modular concepts of Schmidt LE modules, of the large segmented Wolter telescopes (such as XEUS), and of large segmented K-B telescopes are similar is important: all are based on either planparallel or curved flats and foils. This means that the development of high quality X-ray reflecting foils and flats with high mechanical stiffness and low volume density is extremely important for most of the future X-ray astronomy large aperture projects. The segmented K-B telescopes have the advantage of being highly modular on several levels. All segments are rectangular boxes with the same outer dimensions [10].

Science

Deep (limiting flux of 10^{-12} erg $\text{cm}^{-2}\text{s}^{-1}$ can be easily achieved for daily scanning observation) X-ray sky monitoring with large FOVs (*e.g.*, FOV of 6×180 deg can be easily assembled on the space station ISS) is expected to contribute significantly to various fields of modern astrophysics. A few most important examples are listed below. 1) Gamma Ray Bursts (GRBs). Detection rates of nearly 20 GRBs/year can be obtained for the prompt X-ray emission of GRBs, taking into account the expected GRB rate 300/year. 2) X-ray flashes. Detection rates of nearly 8 X-ray flashes/year are expected, assuming XRF rate of 100/year. 3) X-ray binaries. Because of their high variability in X-rays they will be one of major targets in LE observations. LE will be able to observe their short-time outbursts by long-term extended monitoring. Almost all galactic XRB are expected to be within the detection limits. 4) Stars. Because of the low X-ray luminosity of ordinary stars, only nearby stars are expected to be observable. We estimate the lower limit of ordinary stars observable by the LE telescope as 600. The sampling rate of LE observations will be sufficient enough to observe sudden X-ray flux increases during flares while still having the capability of monitoring the variability on time scales of years. 5) Supernovae. The LE telescope should be able to detect the theoretically predicted thermal flash lasting for ~ 1000 s for the first time. Together with the optical SNe detection rate and estimates of the LE FOV we estimate the total number of SNe thermal flashes observed by the LE experiment to ~ 10 /year. 6) AGNs. Active Galactic Nuclei will surely be one of the key targets of the LE experiment. LE will be able to monitor the behavior of the

large (~ 1000) sample of AGNs providing long-term observational data with good time sampling (hours). 7) X-ray transients. The LE experiment will be ideal to observe X-ray transients of various nature due to its ability to observe the whole sky several times a day for a long time with a limiting flux of about $10^{-12}\text{erg cm}^{-2}\text{s}^{-1}$. More and fainter X-ray transients are expected to be detected by the LE sky monitor enabling the detailed study of these phenomena.

Conclusions

Results of analyses and simulations of Lobster-eye X-ray telescopes have indicated that they will be able to monitor the X-ray sky at an unprecedented level of sensitivity, an order of magnitude better than any previous X-ray all-sky monitor. Limits as faint as $10^{-12}\text{ erg cm}^{-2}\text{s}^{-1}$ for daily scanning observation as well as the angular resolution < 4 arcmin in soft-X-ray range are expected to be achieved allowing monitoring of all classes of X-ray sources, not only X-ray binaries, but also fainter classes such as AGNs, coronal sources, cataclysmic variables, as well as fast X-ray transients including gamma-ray bursts and the nearby type II supernovae.

The Lobster optics based All Sky Monitor is capable to detect around 20 GRBs and 8 XRFs yearly and this will surely significantly contribute to the related science. More details on the advantages of LE X-ray telescopes in scientific analyses of SNe are given in [11].

The various prototypes of both Schmidt as well as Angel arrangements have been produced and tested successfully, demonstrating the possibility to construct these lenses by innovative but feasible technologies. Both very small Schmidt lenses (3×3 mm) as well as large lenses (300×300 mm) have been developed, constructed, and tested. This makes the proposals for space projects with very wide-field lobster-eye optics possible.

* * *

We acknowledge the support provided by the Grant Agency of the Czech Republic, grant 102/99/1546, and by the Ministry of Industry and Trade of the Czech Republic, Center of Advanced X-ray Technologies, FB-C3/29/00 and FD-K3/052. The scientific part of the study is related to the Grant Agency of the Academy of Sciences of the Czech Republic, grant A3003206.

REFERENCES

- [1] AMATI L. *et al.*, *AA*, **426** (2004) 415.
- [2] FRONTERA F. *et al.*, *ApJ*, **616** (2004) 1078.
- [3] DE PASQUALE M. *et al.*, *ApJ*, **592** (2003) 1018.
- [4] SCHMIDT W. K. H., *NucIM*, **127** (1975) 285.
- [5] ANGEL J. R. P., *Astroph. J.*, **364** (1979) 233.
- [6] PRIEDHORSKY W. C. *et al.*, *MNRAS*, **279** (1996) 733.
- [7] INNEMAN A. *et al.*, *Proc. SPIE*, **4138** (2000) 94.
- [8] INNEMAN A. *et al.*, *Proc. SPIE*, **3766** (1999) 72.
- [9] FRASER G. W. *et al.*, *Proc. SPIE*, **4497** (2002) 115.
- [10] GORENSTEIN P., *Proc. SPIE*, **3444** (1998) 382.
- [11] ŠVEDA L. *et al.*, in *Cosmic Explosions, Springer Proceedings in Physics*, edited by MARCAIDE J. M. and WEILER K. W., Vol. **99** (Springer, Heidelberg) 2005, p. 197.