

IL NUOVO CIMENTO
DOI 10.1393/ncc/i2005-10046-0

VOL. 28 C, N. 3

Maggio-Giugno 2005

The updated $E_{\text{peak}}-E_{\gamma}$ correlation in GRBs^(*)

G. GHIRLANDA⁽¹⁾, G. GHISELLINI⁽¹⁾, D. LAZZATI⁽²⁾, C. FIRMANI⁽¹⁾

⁽¹⁾ *INAF - Osservatorio Astronomico Brera - via E. Bianchi 46, Merate, Italy*

⁽²⁾ *JILA, 440 UCB, University of Colorado - Boulder, CO, USA*

(ricevuto il 23 Maggio 2005; pubblicato online il 20 Settembre 2005)

Summary. — The recently discovered correlation between the rest frame GRB peak spectral energy E_{peak} and the collimation corrected energy E_{γ} in long GRBs is potentially very important, yet awaits confirmation from an independent sample. It may help to shed light on the radiation mechanism of the prompt GRB phase and on the way—and in which form—the energy is released from the central engine. We here present some additional evidence for the correlation (two new bursts) and re-derive the best-fit parameters. The tightness of the correlation is confirmed ($\sigma = 0.1$ dex). We show that this correlation allows us, for the first time, to use GRBs as cosmological probes to constrain the expansion history of the universe.

PACS 98.70.Rz – γ -ray sources.; γ -ray bursts.

PACS 01.30.Cc – Conference proceedings.

1. – Introduction

Since their discovery, GRBs turned out to be incredibly powerful sources, with detected fluences up to $> 10^{-4}$ erg/cm² in the γ -ray band, above few tens of keV. The first spectroscopic measurements of their redshifts (*e.g.*, [1]), besides confirming their cosmological nature, indicated that these events release, in the γ -ray band, a huge amount of energy, up to $E_{\text{iso}} = 10^{55}$ erg. This extraordinary energetic content became a challenge for the proposed GRB models. One implicit hypothesis in deriving the GRB energy, from the observed fluence and measured redshift, consisted in assuming that GRBs emit isotropically. However, it was suggested [2] that GRBs might be collimated into a cone of semiaperture θ_j . The jet opening angle could be directly estimated, under some simplifying assumptions on few other parameters, from the measure of the achromatic break time t_b in the afterglow light curve [3]. The presence of a jet in GRB outflows, supported by observations in most events [4], allowed to correct the isotropic equivalent energy E_{iso} for the collimation factor, therefore obtaining the collimation corrected energy $E_{\gamma} = E_{\text{iso}}(1 - \cos \theta_j)$ [5, 6]. These results suggested that GRBs might have a unique energy $\sim 10^{51}$ erg. The study of the rest frame spectral properties of a sample

(*) Paper presented at the “4th Workshop on Gamma-Ray Burst in the Afterglow Era”, Rome, October 18-22, 2004.

of *BeppoSAX* GRBs [7] (see also [8]) led to the discovery of a correlation between the isotropic equivalent energy E_{iso} and the νF_ν spectral peak energy, E_{peak} (the Amati correlation).

With the largest sample of bursts with spectroscopically measured redshifts, published spectra and well determined jet break time, we estimated the jet opening angles and derived E_γ . We discovered a very tight correlation between E_γ and the spectral peak energy E_{peak} [9] (the Ghirlanda correlation). This correlation relates the GRB prompt emission energy—properly corrected for the burst geometry—to its peak frequency. It might be the key to understand some still obscure aspects of the physics and origin of GRBs. Besides, its small scatter and good power law fit allowed to use, for the first time, GRBs as standard candles to constrain the cosmological parameters (see [10-12]).

2. – The $E_{\text{peak}}-E_\gamma$ Correlation

We collected all the GRBs with redshift measurements and published spectral parameters. In most cases the spectrum is represented by the empirical Band function. The spectrum allows to compute the burst bolometric fluence F (*e.g.* the time integrated flux) and, hence, the isotropic equivalent energy, *i.e.* $E_{\text{iso}} = 4\pi D_L^2(z)F/(1+z)$. The jet break time, typically observed between 0.1 and 10 days since the burst trigger, is due to the deceleration, by the external ISM, of the GRB relativistic fireball. When the Lorentz factor of the fireball is $\Gamma \propto 1/\theta_j$ a change in the time decay slope of the afterglow flux is observed. This characteristic time depends also on the ISM density n and on the kinetic energy which is still in the fireball after the prompt emission phase (parametrized through the efficiency η_γ). Therefore, the measure of t_b allows to estimate $\theta_j \propto t_b^{3/8} E_{\text{iso}}^{-1/8} (\eta_\gamma n)^{1/8}$. After the publication of the original work of GGL04, the redshifts and spectral parameters of four more bursts were published. We present here the updated correlations (either $E_{\text{peak}}-E_{\text{iso}}$ and $E_{\text{peak}}-E_\gamma$). A continuously updated version of the correlations and the relative tables can be found at <http://www.merate.mi.astro.it/~ghirlanda/deep/blink.htm> (with the complete reference list). We use $\Omega_M = 0.3$, $\Omega_\Lambda = h_0 = 0.7$. The GRBs added to the original sample of 23 (see also [13]) are

- GRB 021004 [14], with a rest frame peak energy $E_{\text{peak}} = 266 \pm 116$ keV and $E_{\text{iso}} = 3.27 \pm 0.39 \times 10^{52}$ erg. The jet break time is $t_b = 4.74 \pm 0.5$ days, the jet opening angle $\theta_j = 8.51^\circ \pm 1.04^\circ$ and, therefore, $E_\gamma = 3.6 \pm 1.0 \times 10^{50}$ erg.
- GRB 030323 [15], with a rest frame peak energy $E_{\text{peak}} = 272 \pm 188$ keV and $E_{\text{iso}} = 3.0 \pm 0.8 \times 10^{52}$ erg. The afterglow light curve is relatively flat *i.e.* indicating a jet break time $t_b > 4.8$ days, implying $\theta_j > 7.8^\circ$ and, therefore, $E_\gamma > 2.77 \times 10^{50}$ erg.
- GRB 040924 [16], with a rest frame peak energy of $E_{\text{peak}} = 96 \pm 20$ keV and $E_{\text{iso}} = 9.5 \pm 1.0 \times 10^{51}$ erg. The afterglow light curve is relatively flat [$F(t) \propto t^{-1}$] up to 1 day, indicating a jet break time $t_b > 1$ day, implying $\theta_j > 6.9^\circ$ and, therefore, $E_\gamma > 6.8 \times 10^{49}$ erg.
- GRB 041006 [17], with a rest frame peak energy $E_{\text{peak}} = 109 \pm 22$ keV and $E_{\text{iso}} = 4.0 \pm 0.4 \times 10^{52}$ erg. The jet break time is $t_b = 0.14 \pm 0.02$ days, implying $\theta_j = 2.9^\circ \pm 0.4^\circ$ and, therefore, $E_\gamma = 4.9 \pm 1.3 \times 10^{49}$ erg.

We added the four GRBs to the 23 GRBs of tables 1 and 2 of GGL04. In fig. 1 we report the updated correlations. With 27 GRBs (black symbols is fig. 1) the best fit power law (weighting for the errors on both coordinates) of the $E_{\text{peak}}-E_{\text{iso}}$ correlation is

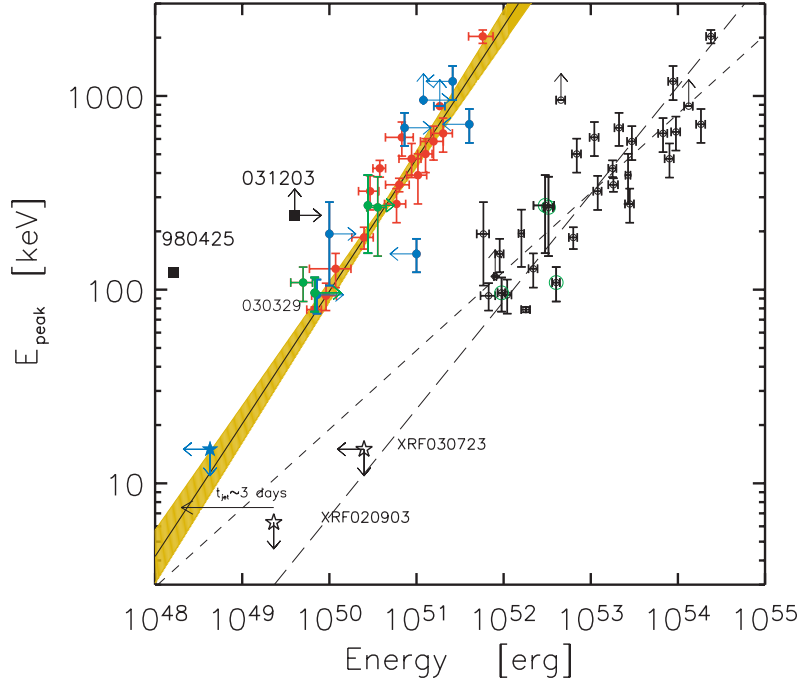


Fig. 1. – (Colour on-line) The rest frame $E_{\text{peak}}-E_{\text{iso}}/E_{\gamma}$ plane. Black open symbols represent the isotropic equivalent energy. Red filled symbols are the 15 GRBs for which a jet break was measured in their afterglow light curves (from tables 1 and 2 of GGL04). Blue symbols are upper/lower limits for E_{γ} . The four new GRBs are represented as open green circles for E_{iso} and filled green symbols for E_{γ} . Also shown are two outliers (black squares) for either the Amati and the Ghirlanda correlation. Stars are the two XRF with known redshift. The Amati correlation is also reported either fitting with the errors on both coordinates (long dashed line) and with the least square method (dashed line). The best fit Ghirlanda correlation (solid black line), giving a reduced $\chi^2 = 1.33$ and a slope ~ 0.7 , is also shown with its uncertainty region (shaded area).

$$(1) \quad E_{\text{peak}}/100 \text{ keV} = (3.21 \pm 0.11)(E_{\text{iso}}/1.1 \times 10^{53} \text{ erg})^{0.56 \pm 0.02}$$

with a reduced $\chi^2 = 5.19$ (long dashed line in fig. 1). The least square fit to the same data points (*i.e.* ignoring the errors on the two coordinates) gives a slope of 0.41 ± 0.05 (dashed line in fig. 1).

The Ghirlanda $E_{\text{peak}}-E_{\gamma}$ correlation, updated with two of the four GRBs with known θ_j (green filled symbols in fig. 1), is

$$(2) \quad E_{\text{peak}}/100 \text{ keV} = (2.5 \pm 1.0)(E_{\gamma}/3.8 \times 10^{50} \text{ erg})^{0.69 \pm 0.04}$$

with a reduced $\chi^2 = 1.33$ (solid line in fig. 1, shaded region is the uncertainty of this correlation). The least square fit gives a slightly flatter slope of 0.6. The present Gaussian fit of the distribution of the perpendicular scatter of the 27 GRB around their best fit (eq. (1)) is $\sigma = 0.22$. The scatter of the Ghirlanda correlation (eq. (2)), instead, is only $\sigma = 0.1$ (*i.e.* consistent with what found with the 15 GRBs in GGL04). We stress

that the Ghirlanda correlation is well fitted (also with the 2 new GRBs) by a single power law and that its slope, 0.69, is consistent with what found with the 15 GRBs in GGL04. Moreover, the reduced $\chi^2 = 1.33$ allows its use in cosmology (GGLF04, GGF05 and [12]). This is in net contrast with what claimed by [18] who find $\chi^2 \sim 3.71$ for the $E_{\text{peak}}-E_\gamma$ fit. Although the authors do not investigate the reason of their statistically unacceptable results, we note that their severe underestimation of the uncertainty on two relevant parameters (*i.e.* the ISM density and the jet break time) is driving their conclusions. They assume, when unknown, an ISM density with an uncertainty of 50%. This parameter is highly uncertain due to the few measured values (and also in these few cases highly debated). Until precise measurements of this parameter will not be available, it is preferable to let it vary within a relatively large range. This is indeed what GGL04 did in their original work (where $1 < n < 10$). Moreover, in [18] a few jet break times are reported with unreliably small errors (down to 1%). This means that for some GRBs we could determine a break in their light curves at—say—1.5 days with an uncertainty of 28 minutes (*e.g.*, GRB 011211), which may be a challenge for the future but which was extremely hard in the past afterglow observational campaigns. These small errors are responsible for their statistically unacceptable results.

3. – Conclusions

We have presented the updated Ghirlanda correlation with 17 GRBs with firm redshift measurements and published spectral parameters. These 2 more events perfectly fit the Ghirlanda correlation as found by GGL04. While waiting for future events, even the present small sample of GRBs have important and intriguing implications for the use of GRBs as standard candles to measure our universe.

* * *

We thank A. CELOTTI and F. TAVECCHIO for useful discussions. GG and GG thank the MIUR for the Cofin grant 2003020775_002.

REFERENCES

- [1] GALAMA T. *et al.*, *IAUC*, **6655** (1997).
- [2] SARI R., PIRAN T. and HALPERN J. P., *ApJ*, **519** (1999) L17.
- [3] RHOADS J. E., *ApJ*, **487** (1997) L1.
- [4] PANAITESCU A. and KUMAR P., *ApJ*, **554** (2001) 667.
- [5] FRAIL D. A. *et al.*, *ApJ*, **562** (2001) L55.
- [6] BLOOM J. S. *et al.*, *ApJ*, **594** (2003) 674.
- [7] AMATI L. *et al.*, *A&A*, **390** (2002) 81 - A02.
- [8] LLOYD-RONNING N. M. and RAMIREZ-RUIZ E., *ApJ*, **576** (2002) 101.
- [9] GHIRLANDA G., GHISELLINI G. and LAZZATI D., *ApJ*, **616** (2004) 331 - GGL04.
- [10] GHIRLANDA G. *et al.*, *ApJ*, **613** (2004) L13 - GGLF04.
- [11] FIRMANI C. *et al.*, *MNRAS*, **360** (2005) L1, astro-ph/0501395.
- [12] GHISELLINI G. *et al.*, these proceedings, (2005).
- [13] GHIRLANDA G., GHISELLINI G. and FIRMANI C., *MNRAS*, **361** (2005) L10, astro-ph/0502186.
- [14] MIRABAL N. *et al.*, *GCN*, (2002) 1618.
- [15] ATTEIA J. L. *et al.*, *ApJ*, **626** (2005) 292, astro-ph/0502494.
- [16] WIERSEMA K. *et al.*, *GCN*, (2004) 2800.
- [17] PRICE P. *et al.*, *GCN*, (2004) 2791.
- [18] FRIEDMAN A. and BLOOM J. S., *ApJ*, **627** (2004) 1, astro-ph/0408413.