



DIGITAL ACCESS TO SCHOLARSHIP AT HARVARD

The Dust Emissivity Spectral Index in the Starless Core TMC-1C

The Harvard community has made this article openly available.
[Please share](#) how this access benefits you. Your story matters.

Citation	Schnee, Scott, Melissa Enoch, Alberto Noriega-Crespo, Jack Sayers, Susan Terebey, Paola Caselli, Jonathan Foster, et al. 2009. The dust emissivity spectral index in the starless core TMC-1C. <i>The Astrophysical Journal</i> 708(1): 127-136.
Published Version	doi:10.1088/0004-637X/708/1/127
Accessed	February 18, 2015 5:11:59 PM EST
Citable Link	http://nrs.harvard.edu/urn-3:HUL.InstRepos:4268293
Terms of Use	This article was downloaded from Harvard University's DASH repository, and is made available under the terms and conditions applicable to Open Access Policy Articles, as set forth at http://nrs.harvard.edu/urn-3:HUL.InstRepos:dash.current.terms-of-use#OAP

(Article begins on next page)

THE DUST EMISSIVITY SPECTRAL INDEX IN THE STARLESS CORE TMC-1C

SCOTT SCHNEE^{1,8}, MELISSA ENOCH², ALBERTO NORIEGA-CRESPO³, JACK SAYERS⁴, SUSAN TEREBEY⁵, PAOLA CASELLI⁶,
JONATHAN FOSTER⁷, ALYSSA GOODMAN⁷, JENS KAUFFMANN⁷, DEBORAH PADGETT³, LUISA REBULL³, ANNEILA SARGENT¹,
RAHUL SHETTY⁷

Draft version November 4, 2009

ABSTRACT

In this paper we present a dust emission map of the starless core TMC-1C taken at 2100 μm . Along with maps at 160, 450, 850 and 1200 μm , we study the dust emissivity spectral index from the (sub)millimeter spectral energy distribution, and find that it is close to the typically assumed value of $\beta = 2$. We also map the dust temperature and column density in TMC-1C, and find that at the position of the dust peak ($A_V \sim 50$), the line-of-sight-averaged temperature is ~ 7 K. Employing simple Monte Carlo modeling, we show that the data are consistent with a constant value for the emissivity spectral index over the whole map of TMC-1C.

Subject headings: stars: formation; ISM: dust, extinction

1. INTRODUCTION

Starless cores are often identified through maps of their dust emission at (sub)millimeter wavelengths, and recent surveys of nearby molecular clouds have increased the number of such objects considerably (e.g. Hatchell et al. 2007; Enoch et al. 2008; Kauffmann et al. 2008; Simpson et al. 2008). Analysis of the properties of starless cores is often restricted by sparse sampling of their spectral energy distribution (SED), uncertain dust emissivities and assumptions about their geometries. As a result, the measured (sub)millimeter fluxes are often used in conjunction with assumed values for the dust temperature and/or emissivity to determine core masses, though the systematic errors can be substantial (a factor of a few). With sufficient spectral coverage and knowledge of source geometry, it should be possible to map the column density, temperature and dust emissivity spectral index in a starless core, though this has yet to be accomplished.

One of the best-studied starless cores is TMC-1C, located in the Taurus molecular cloud at an approximate distance of 140 pc (Torres et al. 2009). Previous studies of the dust emission from TMC-1C at 450, 850 and 1200 μm have determined that it is cold ($T_d \sim 6$ K) and dense ($n \sim 10^6 \text{ cm}^{-3}$) at its center, and becomes progressively less dense and warmer at larger radii (Schnee et al. 2007a; Schnee & Goodman 2005). The

dust-derived mass of TMC-1C is a factor of a few larger than its virial mass, and self-absorbed N_2H^+ spectra show evidence for sub-sonic infall (Schnee et al. 2007b; Schnee & Goodman 2005). The chemistry of TMC-1C is characterized by depletion of CO and its isotopologues, as well as depletion of CS and N_2H^+ (Schnee et al. 2007b).

Previously reported mass and temperature profiles for TMC-1C are estimated assuming a constant value for the emissivity spectral index ($\beta = 1.8$) (Schnee et al. 2007a). However, a wide range of values for β in other regions have been measured or calculated on theoretical grounds. For instance, in the diffuse interstellar medium (ISM), an emissivity spectral index of $\beta = 2$ fits the typical far-infrared spectrum well (Draine et al. 2007; Draine & Lee 1984), while grain growth in protostellar disks results in values of $\beta \simeq 1$ (Beckwith & Sargent 1991). Since starless cores are in an intermediate stage between the diffuse ISM and protostars, one might expect that the dust in a starless core would have an emissivity spectral index in the range $1 \leq \beta \leq 2$, although the timescale of grain growth at densities $n_H \sim 10^6 \text{ cm}^{-3}$ is several hundred times the free-fall time of a typical starless core (Chakrabarti & McKee 2005). On the other hand, the growth of icy mantles on dust grains could steepen the slope of the dust SED, and this has been cited as the reason for high values of the emissivity spectral index in parts of the Orion Ridge ($\beta \sim 2.5$) (Lis et al. 1998) and in Sgr B2(N) ($\beta \sim 3.7$) (Kuan et al. 1996). In the cold and dense environment in TMC-1C, the dust grains could form similar icy mantles, which might drive the emissivity spectral index to values $\beta > 2$.

In addition to changes in β , grain growth by either coagulation or accretion of an icy mantle will change the opacity of the dust grains, κ . For instance, evidence that dust opacities are higher in dense cores than in the diffuse ISM can be found by comparing gas temperature and dust emission profiles (Keto & Caselli 2008) or by comparing dust and gas emission profiles (Keto et al. 2004). From models of dust coagulation for conditions that are plausible for a starless core, Ossenkopf & Henning (1994) show that the dust opac-

Electronic address: scott.schnee@nrc-cnrc.gc.ca

¹ Department of Astronomy, California Institute of Technology, MC 105-24 Pasadena, CA 91125, USA

² Department of Astronomy, University of California, Berkeley, CA 94720, USA

³ *Spitzer* Science Center, MC 220-6, California Institute of Technology, Pasadena, CA 91125, USA

⁴ Jet Propulsion Laboratory, California Institute of Technology, 4800 Oak Grove Drive, Pasadena, CA 9110, USA9

⁵ Department of Physics and Astronomy PS315, 5151 State University Drive, California State University at Los Angeles, Los Angeles, CA 90032, USA

⁶ School of Physics and Astronomy, University of Leeds, Leeds LS2 9JT

⁷ Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138, USA

⁸ Current address: NRC-HIA, 5071 West Saanich Road Victoria, BC V9E 2E7, Canada

ity at 1.3 mm can be a factor of five higher than in a diffuse environment. Using 160 μm data for Taurus, Terebey et al. (2009) measure an opacity of $\kappa_{160 \mu\text{m}} = 0.23 \pm 0.046 \text{ cm}^2\text{g}^{-1}$ in the cold cloud (A_V ranging from 0.4 to 4.0). The opacity is 2.6 times greater than the diffuse ISM opacity (Weingartner & Draine 2003), and provides evidence for higher opacity values in dense regions (c.f. Ossenkopf & Henning 1994).

Laboratory studies of possible dust grain materials have measured the dust emission and absorption properties at far-infrared and longer wavelengths. For instance, Aannestad (1975) measured the absorption of silicate grains and found that $\beta \simeq 2$ for grains without an ice mantle, and $\beta \simeq 3$ for grains with an ice mantle. The steepest spectral indices calculated by Aannestad (1975) were for fused quartz and olivine with ice mantles, both of which had $\beta = 3.5$, close to the extremely steep $\beta = 3.7$ reported by Kuan et al. (1996). In a study of the millimeter-wave absorption of silicate grains in the temperature range 1.2 - 30 K, Agladze et al. (1996) found that the emissivity spectral index depends on the dust temperature and for some materials can vary between $1.5 \leq \beta \leq 2.5$, with the largest value for β at a temperature $T_d \sim 10$ K. A temperature-dependent emissivity spectral index was also found by Mennella et al. (1998) for cosmic dust analog grains in the temperature range $24 \leq T_d \leq 295$ K and a wavelength range $20 \mu\text{m} \leq \lambda \leq 2$ mm, with a steeper SED at lower temperatures. The range of emissivity spectral indices is $1.2 \leq \beta \leq 2.4$ for the various grain types considered by Mennella et al. (1998). In a study of amorphous silicates in the temperature range $10 \leq T_d \leq 300$ K and wavelength range $0.1 \leq \lambda \leq 2$ mm, Boudet et al. (2005) also report an anticorrelation between T_d and β , with values of the emissivity spectral index between $1.5 \leq \beta \leq 2.5$. Based on these laboratory measurements, one might expect that in starless cores like TMC-1C, which have temperatures close to 10 K (Schnee et al. 2009), the emissivity spectral index would be close to $\beta \simeq 2.5$.

Observations of nearby star-forming regions suggest an anticorrelation between T_d and β similar to that measured in the laboratory. For instance, observations of the Orion, ρ Ophiuchi and Taurus molecular clouds at far-infrared and (sub)millimeter wavelengths by Dupac et al. (2003) show that the emissivity spectral index may vary from $2.5 \geq \beta \geq 1.0$ as the dust temperature rises from $10 \leq T_d \leq 80$ K. Similarly, in a study of the far-infrared and (sub)millimeter SEDs of luminous infrared galaxies, Yang & Phillips (2007) report that the emissivity spectral index varies from $2.4 \geq \beta \geq 0.9$ as the dust temperature varies from $30 \leq T_d \leq 60$ K. In a study of the starless core L1498, Shirley et al. (2005) find an emissivity spectral index of $\beta = 2.44 \pm 0.62$ for an assumed dust temperature of $T_d = 10.5 \pm 0.5$ K, consistent with the hypothesis that the cold dust in a starless core has a steep spectral index.

However, variations in the dust properties along the line-of-sight and noise in the observed maps can mimic an anticorrelation between T_d and β (Shetty et al. 2009a,b). Furthermore, at high densities and low temperatures, grain coagulation will decrease β while the possible T_d - β relation will increase the emissivity spectral index. Therefore, it is not at all clear that one should expect the

emissivity spectral index to vary significantly in a starless core. Observations of dust emission and extinction in the dark cloud B68 show no evidence for variations in the dust properties (Bianchi et al. 2003), suggesting that to some extent the various pressures on β may cancel out.

In this paper we present maps of the starless core TMC-1C at 160, 450, 850, 1200 and 2100 μm . By spatially resolving the emission from TMC-1C over a wide range of wavelengths, we then map the column density, temperature and emissivity spectral index of the dust, which has typically been difficult to do with previous datasets.

2. OBSERVATIONS

The 160, 450, 850 and 1200 μm maps have been presented or described in previous papers (Terebey et al. 2009; Schnee & Goodman 2005; Schnee et al. 2007a; Di Francesco et al. 2008; Kauffmann et al. 2008), so here we briefly review the observations and data reduction, and refer the reader to the earlier papers for more details. The 2100 μm map has not been previously presented, so it is discussed in detail below.

2.1. Spitzer

The Spitzer data near TMC-1C derive from the Taurus Spitzer Legacy Survey (Padgett et al. 2009), a large 44 square degree map of the Taurus star-forming region at seven wavelengths ranging from 3.6 to 160 μm . See Padgett et al. (2009) for an overview of the survey parameters and data processing. The MIPS data utilize the fast scan mode with 3 seconds integration time per frame and 5 visits per pixel. The MIPS 160 μm data obtained in fast scan mode mapping does not have enough redundancy to fill in completely all the data gaps due to a dead readout and other effects. To mitigate these data gaps and to preserve the diffuse emission as much as possible, a 3 pixel by 3 pixel ($15''$ per pixel) median filter is applied to the image. The median filter introduces some image smoothing, increasing the effective angular resolution from $40''$ to $1'$. In some cases, filtering reduces the surface brightness by 10–15%, though the reduction in the map of TMC-1C is only 2%.

Analysis of the 160 μm data in Taurus shows there are two main cloud components: emission from the extended cold cloud (14.2 ± 0.4 K, $A_V < 3$), and emission from compact cold cores ($T_d \sim 10$ K, high A_V) (Terebey et al. 2009; Flagey et al. 2009). Since our present focus is the high extinction regions sampled by (chopped) millimeter data, here we further process the 160 μm data to remove the spatially extended cloud emission. Following the procedure in Terebey et al. (2009) we use IRAS 100 μm images to model and subtract the 14.2 K cloud emission. The resulting 160 μm image traces the high extinction ridge and looks very similar to the millimeter continuum emission. The median flux removed from original 160 μm map is 96 MJy/sr, or about 90% of the total flux. Uncertainties in the temperature of the extended cloud component results in a 10% uncertainty in the final 160 μm map. Given the 12% absolute flux uncertainty of MIPS 160 μm maps (Stansberry et al. 2007), we conservatively estimate the absolute uncertainty in the 160 μm “cold core” map to be 20%.

2.2. SCUBA

Although in previous papers on TMC-1C we used sub-millimeter maps observed and reduced by us (Schnee & Goodman 2005; Schnee et al. 2007a), in this paper we use maps of TMC-1C at 450 and 850 μm from the new SCUBA Legacy Catalogues (Di Francesco et al. 2008). We choose to use the new maps because they include all of the data taken with SCUBA (Holland et al. 1999) at all epochs and use the most current calibration and reduction methods. The noise in the maps are 132 and 21 mJy beam^{-1} at 450 and 850 μm , respectively, and have resolutions of approximately $9''$ and $14''$. The uncertainties in the absolute flux calibration (primarily due to beam-shape uncertainty and fluctuations of opacity above the telescope during observations and calibration), are 20% at 850 μm and 50% at 450 μm , as described in Di Francesco et al. (2008). We set the zero point of the Legacy maps by determining the median flux density in two regions off the core but nearby and subtracting these values (-0.32 and -0.035 Jy beam^{-1}) from the 450 and 850 μm maps.

2.3. MAMBO

We observed the 1200 μm continuum emission from TMC-1C in the autumns of 2002 and 2003 using the MAMBO-2 array (Kreysa et al. 1999) on the IRAM 30 m telescope on Pico Veleta (Spain), producing a map with $10.7''$ resolution. Data were taken with on-the-fly mapping, with the telescope sub-reflector chopping in azimuth by $60''$ - $70''$, and the resultant image was reconstructed using the Emerson et al. (1979) algorithm iteratively to properly reproduce large-scale emission. The noise in map is 3 mJy beam^{-1} , and the flux calibration uncertainty, as derived from the rms of calibrator observations across pooled observing sessions and the uncertainty in the intrinsic calibrator fluxes, is about 10%. For further details on the 1200 μm observations and data reduction, see Schnee et al. (2007a) and Kauffmann et al. (2008).

2.4. Bolocam

We observed the 2100 μm continuum emission from TMC-1C in the autumn of 2007 using Bolocam on the Caltech Submillimeter Observatory (CSO). Data were collected in raster scan mode at a scan speed of $240''\text{s}^{-1}$, with a step size between rasters of $11''$. Each raster was $20'$ in length, and each of the 105 total scans were separated by $11''$. The resultant map has a coverage of $\sim 20' \times 20'$ with $60''$ resolution. Three sets of observations contain rasters parallel to RA, and the other three contain rasters parallel to Dec. Data reduction and flux calibration, accomplished through observations of Uranus and Neptune, was done as described in Sayers et al. (2009).

Observations at approximately 2 mm are dominated by sky noise. Although the Bolocam pipeline PCA cleaning algorithm efficiently removes sky noise, it also removes some of the source flux. This decreases the peak intensity, introduces negative bowls around bright sources, and reduces the brightness on scales larger than $1'$ for extended sources. We recover this lost flux using a mapping algorithm that iteratively subtracts a source model from the real data (for more description and performance evaluation, see Enoch et al. 2006). This technique is similar

to CLEAN (Högbom 1974; Schwarz 1978), but works in the image plane, and the real data are used for the source model rather than, e.g., a Gaussian component model. The first source model is formed from the original map, including only regions with signal-to-noise greater than 1.5. That source model is subtracted from the original timestream data, which is then re-cleaned and mapped without the source flux. This process is iterated until the source model stabilizes, 10 iterations in this case. TMC-1C is a faint source, so the correction is relatively small, with the source flux density increasing by up to 30% after iterative mapping. The noise in the 2100 μm map is 8 mJy beam^{-1} , and the uncertainty in the absolute flux calibration is $\sim 20\%$.

3. ANALYSIS

Here we estimate the temperature, column density and emissivity spectral index in TMC-1C using our flux density maps at 160, 450, 850, 1200 and 2100 μm . Under the assumptions that the dust is isothermal and characterized by a single emissivity spectral index, the flux density per beam in each map can be described as a modified blackbody, given by,

$$S_\nu = \Omega B_\nu(T_d) \kappa_\nu \mu m_H N_{H_2} \quad (1)$$

where

$$B_\nu(T_d) = \frac{2h\nu^3}{c^2} \frac{1}{\exp(h\nu/kT_d) - 1} \quad (2)$$

and

$$\kappa_\nu = \kappa_{230} \left(\frac{\nu}{230 \text{ GHz}} \right)^\beta \quad (3)$$

In equation (1), S_ν is the flux density per beam, Ω is the solid angle of the beam, $B_\nu(T_d)$ is the blackbody emission from the dust at temperature T_d , $\mu = 2.8$ is the mean molecular weight of interstellar material in a molecular cloud per hydrogen molecule, m_H is the mass of the hydrogen atom, N_{H_2} is the column density of hydrogen molecules, and a gas-to-dust ratio of 100 is assumed. In equation (3), $\kappa_{230} = 0.009 \text{ cm}^2 \text{ g}^{-1}$ is the emissivity of the dust grains at a gas density of 10^6 cm^{-3} covered by a thin ice mantle at 230 GHz (Ossenkopf & Henning 1994, Column 6 of Table 1) and β is the emissivity spectral index of the dust. Although Equations (1 – 3) are often used to describe the observed emission from starless cores, such observations average over a range of dust grain properties. Furthermore, the emissivity (κ_{230}) is uncertain by roughly a factor of 2. If we had chosen a different value for κ_{230} , this would affect the derived column density, but would have no impact on the derived temperature and emissivity spectral index. We do not consider variations in κ_{230} in this paper, but instead save this analysis for a future analysis.

3.1. Color Correction

The fluxes given in the dust emission maps are derived from the total energy detected convolved with the wavelength-dependent instrumental response and an assumed SED of the emitting source. The assumed SEDs in the *Spitzer*, SCUBA, MAMBO and Bolocam maps are that of a 10,000 K blackbody, a flat-spectrum source, a flat-spectrum source and a blackbody in the Rayleigh-Jeans regime, respectively. Because we are observing

emission from dust, with an assumed SED of a modified blackbody, the quoted flux at the nominal wavelength of each map ought to be corrected. However, the relatively narrow bandpasses of the detectors results in a modest correction factor.

For each map, we use the instrumental response and our best-guess dust SED of a modified blackbody with $T_d = 10$ K and $\beta = 2$ to find the “reference” wavelength of the observations. The integral of the convolution of the instrumental response and the dust SED are equal above and below this reference wavelength.

$$\int_0^{\lambda_{\text{ref}}} S_\lambda R(\lambda) d\lambda = \int_{\lambda_{\text{ref}}}^{\infty} S_\lambda R(\lambda) d\lambda \quad (4)$$

where S_λ is the flux emitted by the dust at wavelength λ and $R(\lambda)$ is the instrumental response at wavelength λ .

The reference wavelength, used in our calculations of the dust emission properties, are close to the nominal wavelength of each map, with the biggest difference in the MAMBO map because of its wide bandpass. In all maps, the flux uncertainties are larger than the correction from using the reference wavelength in our calculations rather than the nominal wavelength. The derived reference wavelengths are given in Table 1.

3.2. Single SED

One way to estimate the dust temperature, column density and emissivity spectral index is to fit modified blackbody SEDs (given by Equations 1 – 3) to the five fluxes (at 160, 450, 850, 1200 and 2100 μm) at each position. Although the resolutions of the maps at 450, 850 and 1200 μm are fairly similar (from 7.5'' – 14''), the 60'' resolution of the 160 and 2100 μm maps is not a good match to the others. Here we handle the mismatched beam sizes by making an average SED from the five flux maps to derive the best-fit dust properties (T_d , N_{H_2} and β).

We first smooth the 450, 850 and 1200 μm maps to the resolution of the 160 and 2100 μm maps by convolution with a two-dimensional Gaussian beam. We then calculate the average flux in those pixels brighter than 3σ in the 160, 850, 1200, and 2100 μm maps, and brighter than 2σ in the 450 μm map (which is noisier than the other maps). The area with significant emission in all five maps is shown in Figure 1. The noise level in the maps is determined in regions off-source but near TMC-1C (see Figure 1). We then used a χ^2 minimization routine to fit the five measured fluxes with a modified blackbody SED, finding the values of T_d , N_{H_2} and β .

In order to determine how noise and calibration uncertainties affect the accuracy of the derived dust properties, we run a simple Monte Carlo simulation. In each trial, Gaussian random noise of the amount observed is added to each flux map, and then each map is scaled by a multiplicative factor $f = 1.0 + \delta$, where δ is randomly chosen from a normal distribution of mean zero and standard deviation equal to the absolute calibration uncertainty in each map. Each map is modified by a different scale factor f , and this process was repeated 5000 times. We find that the best-fit dust properties for the average SED over TMC-1C are $\beta = 2.3 \pm 0.4$, $T_d = 8.9 \pm 1.0$, and $N_{H_2} = 10.2 \times 10^{21} \pm 2.1 \times 10^{21}$. The observed fluxes and

best-fit SED are plotted in Figure 2 and the fluxes and uncertainties are given in Table 1.

A fundamental flaw in this analysis is that equations 1 – 3 are only accurate for isothermal dust characterized by a single emissivity spectral index, but TMC-1C is known to be warmer at its edges than at the center (Schnee et al. 2007a), and it is certainly possible that β also changes within the starless core. As shown by Shetty et al. (2009a), the incorrect assumption of isothermal dust drives an anti-correlation between the derived dust temperature and emissivity spectral index. A spurious anti-correlation between T_d and β may also arise due to noise in the emission maps (Shetty et al. 2009a,b). An anti-correlation between β and T_d is seen in the Monte Carlo simulation we used to derive the dust properties, as shown in Figure 2. In the following section we attempt to mitigate the effects of variations in the dust properties on our analysis.

3.3. Maps

The dust properties (T_d , N_{H_2} and β) in TMC-1C may vary both in the plane of the sky and along the line of sight. The line-of-sight variations in the dust properties cannot be determined without knowledge of the geometry of the core, but variations in the plane of the sky are more tractable. Here we make maps of TMC-1C, first of the temperature and column density while holding the emissivity spectral index constant, and then of all three quantities.

3.3.1. Maps with Constant β

We first regrid the 450, 850 and 1200 μm maps to the resolution of the 850 μm map, 14''. We then determine the noise in the maps in the same regions selected in Section 3.2 and calculate the dust temperature and column density in those pixels brighter than 3σ in the 850 and 1200 μm maps and brighter than 2σ in the 450 μm map, using equations 1 – 3 for an assumed value of the emissivity spectral index, as was done in Schnee et al. (2007a). Temperature and column density maps are made for each value of the emissivity spectral index in the range $1.0 \leq \beta \leq 4.0$ in steps of 0.1. To account for the effects of noise and calibration uncertainties, we use the same Monte Carlo technique described in Section 3.2 to add Gaussian random noise to each pixel and adjust each map by a multiplicative scale factor, and this process is repeated 100 times for each value of β .

To determine which value of β is most appropriate for TMC-1C, we then use the derived T_d and N_H maps to predict the fluxes at 160, 450, 850, 1200 and 2100 μm . The predicted 160 and 2100 μm fluxes are convolved with a 60'' Gaussian beam to match the resolution of the observed maps. We find that the 160 and 2100 μm maps are best reproduced by an emissivity spectral index of $\beta = 2.2 \pm 0.6$. The observed, predicted and residual flux maps for $\beta = 2.2$ are shown in Figure 3, and the derived temperature and column density maps are shown in Figure 4. The column density is expressed in units of A_V , using the conversion factor $N_{H_2} = 9.4 \times 10^{20} \text{ cm}^{-2}$ (A_V/mag) (Bohlin et al. 1978).

3.3.2. Maps with Variable β

In section 3.3.1 we assumed that β is constant across TMC-1C to make 14'' maps of the dust temperature

and column density in TMC-1C. Here we make two-dimensional maps of the dust temperature, column density *and* emissivity spectral index in TMC-1C, with the tradeoff that the resultant maps have coarser resolution (60'').

We first regrid the 450, 850 and 1200 μm maps to the 60'' resolution of the 160 and 2100 μm maps. For each pixel brighter than 3σ , we then calculate T_d , N_{H_2} and β using equations 1 – 3. To estimate the uncertainties in this method due to noise and absolute calibration errors, we used the same Monte Carlo techniques as in sections 3.2 and 3.3.1 with 1000 trials. In each trial there were about 15 pixels with sufficiently high signal to noise to calculate the dust properties, and within the maps there is a clear anti-correlation between T_d and β , as shown in Figure 4. The median emissivity spectral index and dust temperature derived from the high signal to noise points in all trials are $\beta = 2.0 \pm 0.6$ and $T_d = 9.4 \pm 1.9$, similar to that derived from the composite SED of TMC-1C in section 3.2.

4. DISCUSSION

In Section 3.3.2 we found that there is an anti-correlation between the dust temperature and emissivity spectral index in the two dimensional maps of TMC-1C. Although this phenomenon has been reported in other star-forming regions (Dupac et al. 2003) and in other galaxies (Yang & Phillips 2007), this is not a known feature of starless cores. The assumption that the temperature along each line-of-sight is constant, even though the temperature profile in TMC-1C is very likely to be colder at higher densities, can lead to the mistaken conclusion that there is an anti-correlation between T_d and β (Shetty et al. 2009a,b). Noise in the dust emission maps can also drive a spurious $T_d - \beta$ anti-correlation. Here we investigate whether the dust temperature and emissivity spectral index are truly anti-correlated, and conclude that this is not necessarily the case. We then determine what constant value of β best characterizes TMC-1C.

4.1. Testing the $T_d - \beta$ Anti-correlation

To determine the effects of temperature variations along the line of sight and noise on the dust properties derived from flux maps at 160, 450, 850, 1200 and 2100 μm , we consider a spherical model of a starless core with plausible properties. The density in the core is given by

$$n(r) = \frac{n_0}{1 + (r/r_0)^{2.5}} \quad (5)$$

where r is the radius, $n_0 = 10^6 \text{ cm}^{-3}$ and $r_0 = 0.005 \text{ pc}$. The temperature in the core is given by

$$T_d(r) = T_0 (1 + r/r_1) \quad (6)$$

where $T_0 = 6.4 \text{ K}$ and $r_1 = 0.15 \text{ pc}$. The emissivity spectral index is constant in the core at a value $\beta = 2.2$. This model core is colder and denser at the center, as expected for an externally heated starless core. We do not claim that these are the temperature and density profiles for TMC-1C, but there is a general agreement between the shape of the model SED and that of TMC-1C. The density and temperature profiles are shown in Figure 5, as is the total emergent SED.

For each line of sight passing through the model core, we determine the volume of intersection between a cylinder and each spherical shell of the core. Given the density and temperature distributions from Equations 5 and 6, the flux from each layer is given by

$$S_\nu = \frac{B_\nu(T_d)\kappa_\nu n_r V_r}{d^2} \quad (7)$$

where V_r is the volume of intersection and d is the distance to the core. The two-dimensional flux distribution is calculated by summing the flux from each layer along each line of sight. For the sake of convenience, the total flux density of the model core, summed over all radii, is scaled by a common factor such that the peak of the SED has a value of unity.

We ‘‘observe’’ the model core by deriving two-dimensional flux maps at 160, 450, 850, 1200 and 2100 μm and adding Gaussian random noise equal to 5% of the peak flux in each map. Using equations 1 – 3, we then derive two-dimensional maps of the column density, dust temperature and emissivity spectral index, assuming that these quantities are constant along the line of sight (as was done in Section 3.3.2). The derived column density map peaks at the center of the core, where the temperature is seen to decrease, as one would expect from the input density and temperature profiles (see Figure 5). The derived emissivity spectral index map covers a wide range of values, mostly between $1.0 \leq \beta \leq 3.5$, even though the input emissivity spectral index was held constant. The median derived emissivity spectral index is 2.0 ± 0.4 , which is close to the input value of $\beta = 2.2$.

Furthermore, the derived temperature and emissivity spectral index are seen to be anti-correlated. The combination of noise and temperature variations along the line of sight have been shown to drive an anti-correlation between these two quantities (Shetty et al. 2009a,b), so it is not surprising that this effect is seen here. Interestingly, the portion of the $T_d - \beta$ parameter space covered by the model starless core overlaps significantly with that of TMC-1C (see Figures 3.3.2 and 5). From this we conclude that there is no definitive evidence that the emissivity spectral index in TMC-1C varies with temperature, and that the more reliable estimates of column density and temperature will come from fitting the observed fluxes while holding the emissivity spectral index constant.

4.2. Average Emissivity Spectral Index in TMC-1C

Estimates of the emissivity spectral index in TMC-1C derived from its average SED (see Section 3.2) and from two dimensional maps using the 450, 850 and 1200 μm maps to predict the lower resolution 160 and 2100 μm maps (see Section 3.3.1), $\beta = 2.3 \pm 0.4$ and $\beta = 2.2 \pm 0.6$, are in good agreement with each other. Given that an analysis of the flux maps from an idealized core, with line-of-sight temperature variations and noise, gives the correct *average* emissivity spectral index (within errors), we suggest that the value of the emissivity spectral index in TMC-1C is well-represented by $\beta \simeq 2.2 \pm 0.5$. This is consistent with the emissivity spectral index in the diffuse ISM of $\beta = 2$ (Draine et al. 2007; Draine & Lee 1984) and with the emissivity spectral index measured in the starless core L1498 ($\beta = 2.44 \pm 0.62$, Shirley et al.

2005). However, due to the large uncertainty in β in this study, we cannot rule out the possibility that the emissivity spectral index in TMC-1C could also be intermediate between the shallower SEDs seen in circumstellar disks and the SED of the ISM.

When deriving the dust temperature and mass of prestellar cores from sparsely-sampled SEDs, it is common to assume that the emissivity spectral index of the dust is equal to 2 (e.g. Kirk et al. 2007; Simpson et al. 2008). Observations of L1498 (Shirley et al. 2005) and our observations of TMC-1C are consistent with $\beta = 2$, though the best-fit values of the emissivity spectral index in both cores are slightly steeper, which would result in lower temperatures and higher masses. However, with so few measurements of β in starless cores, there is not sufficient information with which to re-examine the dust properties in previously published studies. Upcoming surveys of the Gould Belt molecular clouds using *Herschel*¹ and *SCUBA-2*² promise to make statistically significant measurements of the emissivity spectral index in starless cores.

5. SUMMARY

In this paper we have analyzed the dust emission at 160, 450, 850, 1200 and 2100 μm from TMC-1C. Our previous studies of this starless core have mapped the dust temperature and column density. Here, using new 160 and 2100 μm maps and an improved reduction of the 450 and 850 μm maps, we are able to estimate the emissivity spectral index and make an improved analysis of the temperature and column density in TMC-1C.

We find that the spectral index of the dust emission, as calculated by a variety of related methods, is in the

range $1.7 \leq \beta \leq 2.7$. This is consistent with the value of the emissivity spectral index often assumed in studies of starless cores, and is also consistent with laboratory measurements and observations of the behavior of dust grains at low temperatures ($T_d \leq 10$ K). However, there is still a relatively large uncertainty in the emissivity spectral index.

Using the new estimate of the emissivity spectral index in TMC-1C ($\beta = 2.2$), we construct temperature and column density maps, which reach values of $T_d \simeq 7$ K and $A_V \simeq 50$ at the position of the dust peak, as shown in Figure 4. We find no conclusive evidence that β varies within the core, though line-of-sight temperature variations and noise in the flux maps leave open the possibility that the emissivity spectral index is anti-correlated with dust temperature.

We thank our referee, Yancy Shirley, for comments that improved the clarity of this paper. SS acknowledges support from the Owens Valley Radio Observatory, which is supported by the National Science Foundation through grant AST 05-40399. JS was partially supported by a NASA Postdoctoral Program Fellowship. Support was provided to ME by NASA through the *Spitzer Space Telescope* Fellowship Program. The JCMT is operated by the Joint Astronomy Centre on behalf of the Particle Physics and Astronomy Research Council of the United Kingdom, the Netherlands Organisation for Scientific Research, and the National Research Council of Canada. IRAM is supported by INSU/CNRS (France), MPG (Germany), and IGN (Spain). The CSO is supported by the NSF fund under contract AST 02-29008.

¹ <http://starformation-herschel.iap.fr/gouldbelt>

² <http://www.jach.hawaii.edu/JCMT/surveys/gb/>

REFERENCES

- Aannestad, P. A. 1975, *ApJ*, 200, 30
Agladze, N. I., Sievers, A. J., Jones, S. A., Burlitch, J. M., & Beckwith, S. V. W. 1996, *ApJ*, 462, 1026
Beckwith, S. V. W., & Sargent, A. I. 1991, *ApJ*, 381, 250
Bianchi, S., Davies, J. I., Alton, P. B., Gerin, M., & Casoli, F. 2000, *A&A*, 353, L13
Bianchi, S., Gonçalves, J., Albrecht, M., Caselli, P., Chini, R., Galli, D., & Walmsley, M. 2003, *A&A*, 399, L43
Bohlin, R. C., Savage, B. D., & Drake, J. F. 1978, *ApJ*, 224, 132
Boudet, N., Mutschke, H., Nayral, C., Jäger, C., Bernard, J.-P., Henning, T., & Meny, C. 2005, *ApJ*, 633, 272
Chakrabarti, S., & McKee, C. F. 2005, *ApJ*, 631, 792
Di Francesco, J., Johnstone, D., Kirk, H., MacKenzie, T., & Ledwosinska, E. 2008, *ApJS*, 175, 277
Draine, B. T., et al. 2007, *ApJ*, 663, 866
Draine, B. T., & Lee, H. M. 1984, *ApJ*, 285, 89
Dupac, X., et al. 2003, *A&A*, 404, L11
Emerson, D. T., Klein, U., & Haslam, C. G. T. 1979, *A&A*, 76, 92
Enoch, M. L., Evans, N. J., II, Sargent, A. I., Glenn, J., Rosolowsky, E., & Myers, P. 2008, *ApJ*, 684, 1240
Enoch, M. L., et al. 2006, *ApJ*, 638, 293
Flagey, N., et al. 2009, arXiv:0906.5025
Hatchell, J., Fuller, G. A., Richer, J. S., Harries, T. J., & Ladd, E. F. 2007, *A&A*, 468, 1009
Högbom, J. A. 1974, *A&AS*, 15, 417
Holland, W. S., et al. 1999, *MNRAS*, 303, 659
Kauffmann, J., Bertoldi, F., Bourke, T. L., Evans, N. J., II, & Lee, C. W. 2008, *A&A*, 487, 993
Keto, E., Rybicki, G. B., Bergin, E. A., & Plume, R. 2004, *ApJ*, 613, 355
Keto, E., & Caselli, P. 2008, *ApJ*, 683, 238
Kirk, J. M., Ward-Thompson, D., & André, P. 2007, *MNRAS*, 375, 843
Kreysa, E., et al. 1999, *Infrared Physics and Technology*, 40, 191
Kuan, Y.-J., Mehringer, D. M., & Snyder, L. E. 1996, *ApJ*, 459, 619
Lis, D. C., Serabyn, E., Keene, J., Dowell, C. D., Benford, D. J., Phillips, T. G., Hunter, T. R., & Wang, N. 1998, *ApJ*, 509, 299
Mennella, V., Brucato, J. R., Colangeli, L., Palumbo, P., Rotundi, A., & Bussoletti, E. 1998, *ApJ*, 496, 1058
Ossenkopf, V., & Henning, T. 1994, *A&A*, 291, 943
Padgett, D. L. et al. 2009, in preparation
Pierce-Price, D., et al. 2000, *ApJ*, 545, L121
Sayers, J., et al. 2009, *ApJ*, 690, 1597
Schnee, S., Rosolowsky, E., Foster, J., Enoch, M., & Sargent, A. 2009, *ApJ*, 691, 1754
Schnee, S., Caselli, P., Goodman, A., Arce, H. G., Ballesteros-Paredes, J., & Kuchibhotla, K. 2007, *ApJ*, 671, 1839
Schnee, S., Kauffmann, J., Goodman, A., & Bertoldi, F. 2007, *ApJ*, 657, 838
Schnee, S., & Goodman, A. 2005, *ApJ*, 624, 254
Schwarz, U. J. 1978, *A&A*, 65, 345
Shetty, R., Kauffmann, J., Schnee, S., Goodman, A. A., & Ercolano, B. 2009, *ApJ*, 696, 2234
Shetty, R., Kauffmann, J., Schnee, S., & Goodman, A. A. 2009, *ApJ*, 696, 676
Shirley, Y. L., Nordhaus, M. K., Grcevich, J. M., Evans, N. J., II, Rawlings, J. M. C., & Tatematsu, K. 2005, *ApJ*, 632, 982
Simpson, R. J., Nutter, D., & Ward-Thompson, D. 2008, *MNRAS*, 391, 205

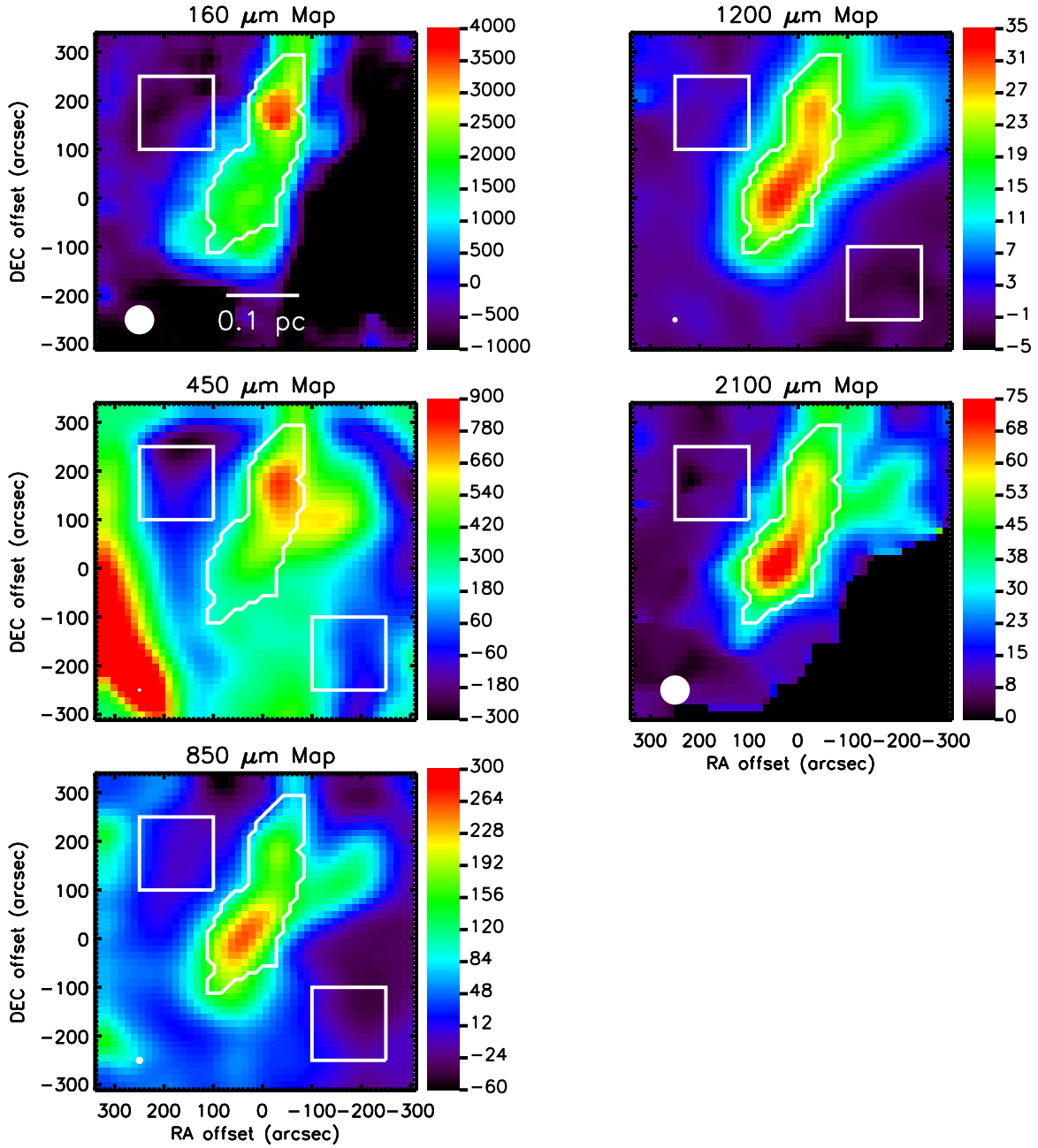


FIG. 1.— Flux maps of TMC-1C at 160, 450, 850, 1200 and 2100 μm , in units of mJy beam^{-1} , smoothed with the 2100 μm map's beam. The white contour shows the region that is above 3σ in the 160, 850, 1200 and 2100 μm maps and above 2σ in the 450 μm map, and is the region used for determining the dust properties. The white boxes show the regions over which the noise in the maps was calculated. The beam size of each map is shown in the bottom left corners of each panel.

Stansberry, J. A., et al. 2007, *PASP*, 119, 1038
 Terebey, S., et al. 2009, *ApJ*, 696, 1918
 Torres, R. M., Loinard, L., Mioduszewski, A. J., & Rodríguez,
 L. F. 2009, *ApJ*, 698, 242

Weingartner, J. C., & Draine, B. T. 2003, *ApJ*, 589, 289
 Yang, M., & Phillips, T. 2007, *ApJ*, 662, 284

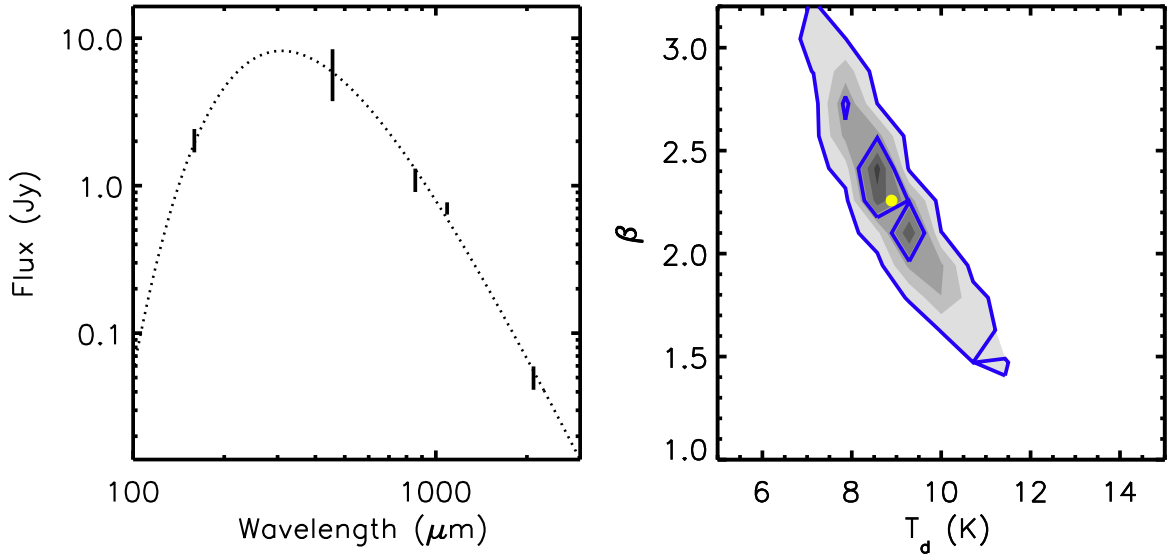


FIG. 2.— (*left*) The fluxes of TMC-1C at 160, 450, 850, 1200 and 2100 μm , in a $60''$ Gaussian beam, averaged over the high signal-to-noise portion of the maps, shown with 1σ error bars. The dotted line shows the best-fit modified blackbody spectrum, which has column density $N_{H_2} = 10.2 \times 10^{21}$, dust temperature $T_d = 8.9$ and emissivity spectral index $\beta = 2.3$. See section 3.2 for details. (*right*) Filled contours show the density of $T_d - \beta$ values derived in the Monte Carlo simulation of the average TMC-1C fluxes, with the open blue contours containing 50% and 95% of all points. The yellow circle shows the position of the median dust temperature and emissivity spectral index, used to calculate the SED in the left side of this figure. The fluxes and uncertainties are given in Table 1.

TABLE 1
TMC-1C FLUXES

Nominal Wavelength (μm)	Reference Wavelength (μm)	Average Flux (Jy/beam ¹)	Flux Uncertainty
160	159	2.0	20%
450	456	5.6	50%
850	856	1.1	20%
1200	1090	0.70	10%
2100	2103	0.050	20%

¹ in a $60''$ Gaussian beam

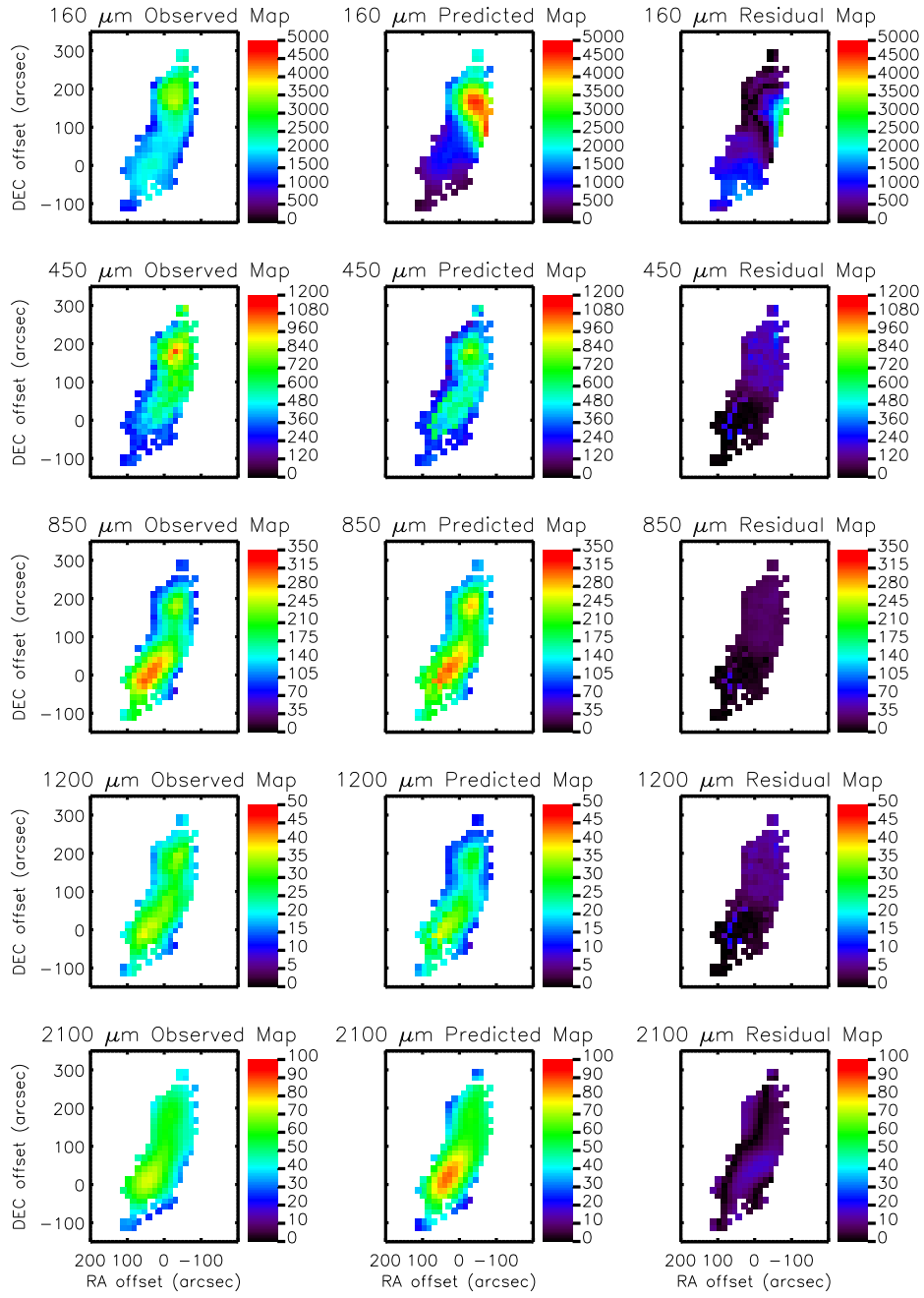


FIG. 3.— (left column) The observed maps of TMC-1C at 160, 450, 850, 1200 and 2100 μm , shown in units of mJy beam^{-1} . (center column) The predicted maps of TMC-1C, based on the derived temperature and column density maps (see Figure 4), with $\beta=2.2$. (right column) The absolute value of the residuals. See section 3.3.1 for details.

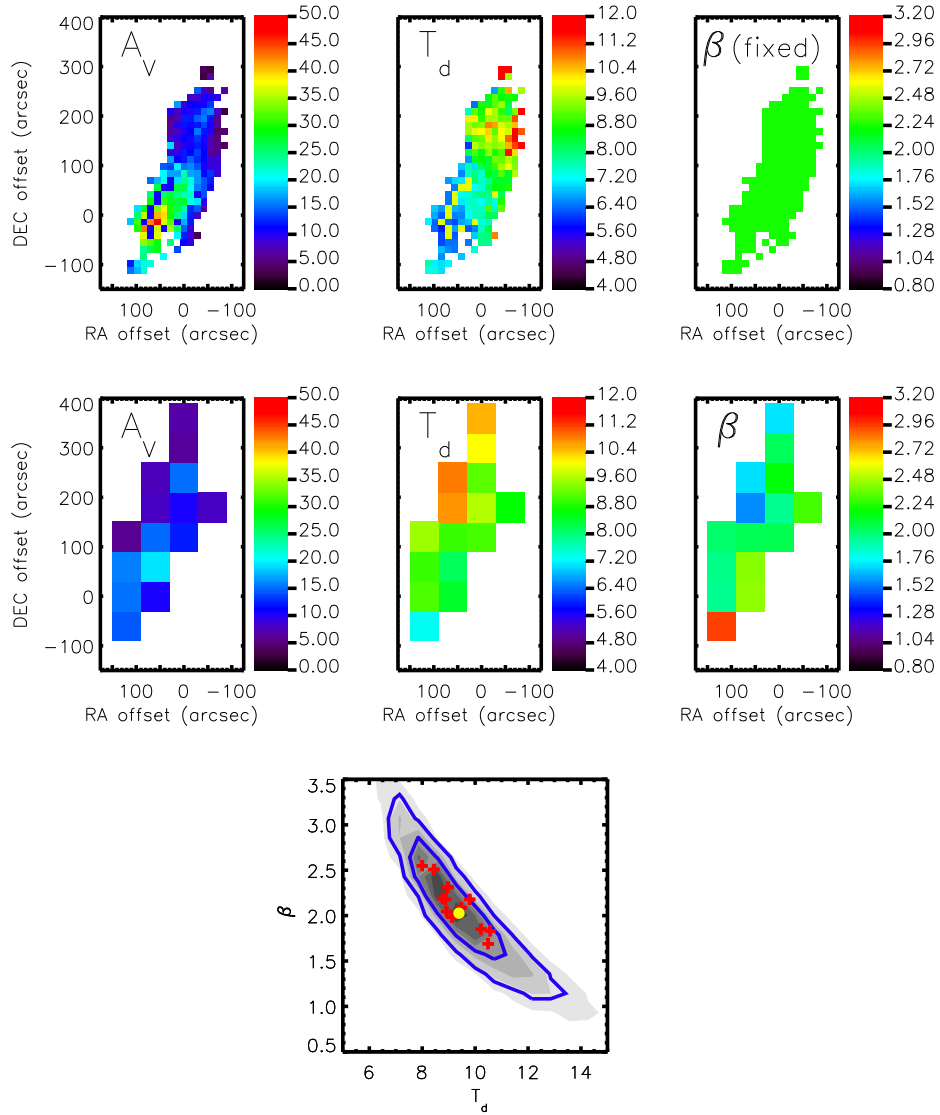


FIG. 4.— (*top row*) Column density and dust temperature maps, expressed in units of A_V and K, derived from $14''$ maps of TMC-1C at 450, 850 and $1200\ \mu\text{m}$ using $\beta=2.2$. See section 3.3.1 for details. (*center row*) Column density, dust temperature and emissivity spectral index maps of TMC-1C, derived from $60''$ maps at 160, 450, 850, 1200 and $2100\ \mu\text{m}$. See section 3.3.2 for details. (*bottom row*) Filled contours show the density of $T_d - \beta$ values, with the open blue contours containing 50% and 95% of all points. The yellow circle shows the position of the median dust temperature ($T_d = 9.4$) and emissivity spectral index ($\beta = 2.0$). The red crosses show the values of β and T_d derived from each independent position in the $60''$ flux density maps with high signal to noise.

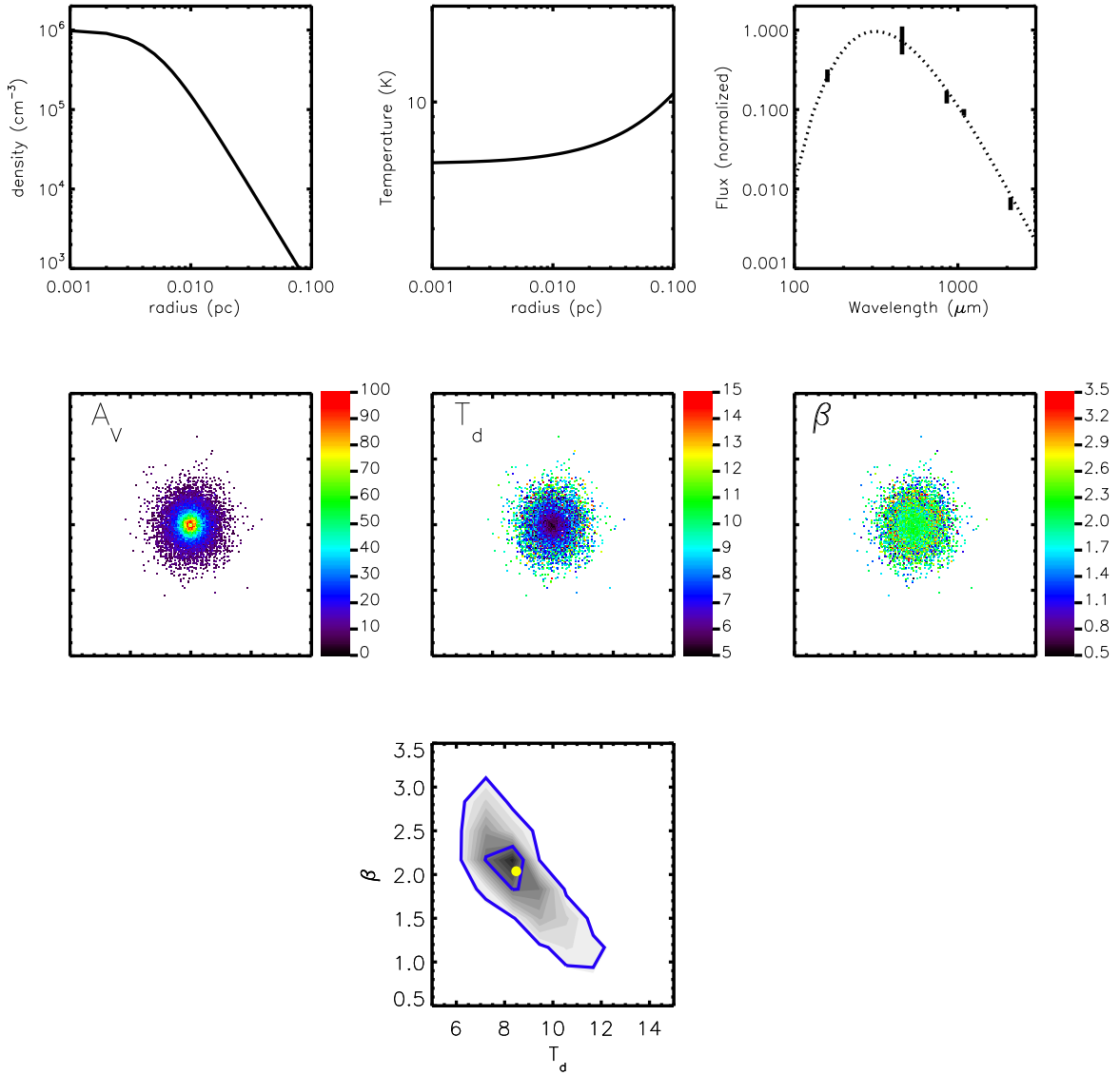


FIG. 5.— (*top left and center*) Density and temperature profiles of a spherical model of a starless core. (*top right*) The SED of the model starless core with $\beta = 2.2$, normalized such that the SED peak is at unity. The TMC-1C observations are also plotted for reference, and scaled by a common factor such that the $450 \mu\text{m}$ flux from TMC-1C is equal to that of the model core. (*center row*) Column density, temperature and emissivity spectral index maps that would be derived from the noisy maps of the model starless core. (*bottom row*) Filled contours show the density of $T_d - \beta$ values, with the open blue contours containing 50% and 95% of all points. The yellow circle shows the position of the median dust temperature and emissivity spectral index. The anti-correlation between T_d and β is similar to that observed in TMC-1C (see Figure 4).