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Binary outflows from young stars: interaction of co-orbital jet and wind

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ABSTRACT

Jets from young stellar objects provide insight into the workings of the beating heart at the centre of star-forming cores. In some cases, multiple pulsed outflows are detected such as the atomic and molecular jets from a proposed binary system in the T Tauri star HH 30. We investigate here the development and propagation of duelling atomic and molecular outflows stemming from the two stars in co-orbit. We perform a series of numerical experiments with the ZEUS-MP code with enhanced cooling and chemistry modules. The aim of this work is to identify signatures on scales of the order of 100 au. The jet sources are off the grid domain and so it is the propagation and interaction from ∼20 au out to 100 au simulated here. We find that the molecular flow from the orbiting source significantly disturbs the atomic jet, deflecting and twisting the jet and disrupting the jet knots. Regions of high ionization are generated as the atomic jet rams through the dense molecular outflow. Synthetic images in atomic and molecular lines are presented, which demonstrate identifying signatures. In particular, the structure within the atomic jet is lost and H α may trace the walls of the present CO cavity or where the walls have been recently. These results provide a framework for the interpretation of upcoming high-resolution observations.

Key words: hydrodynamics – stars: formation – stars: pre-main-sequence – ISM: jets and outflows.

1 INTRODUCTION

Protostellar jets offer an observational window into the births of stars. The jets protrude from the cloud that obscures the protostar and remain prominent even when the young star is visible. The characteristics of these jets may reveal something of the nature of the originating objects and the processes that govern their evolution. For example, reflection symmetry and mirror symmetry may differentiate between orbital dynamics and precessional motion such as emphasized in recent studies (Moraghan et al. 2016; Hirano & Machida 2019). In addition to the role these jets play in extracting angular momentum from infalling material in the accretion disc, the emerging outflows may also act as a feedback channel supplying turbulent energy to the surrounding molecular cloud, which will affect the star formation efficiency within the cloud (Arce et al. 2007; Dionatos & Güdel 2017).

Evidence for interacting multiple jet components introduces new challenges. The T Tauri star associated with HH 30 (Mundt & Fried 1983) is a well-known example. Each jet component can be observed in atomic or molecular tracers in the form of distinct knots within a diffuse channel. These components display behaviour that could be driven by orbiting binary stars. In fact, observations provide

evidence that HH 30 is a binary system, surrounded by a circumbinary accretion disc. Other outflows appear to be driven by multiple sources that may be gravitationally bound within a protostellar core (Anglada et al. 1991; Gueth, Schilke & McCaughrean 2001). Twin jets are also found in the L1551 IRS5 system (Fridlund & Liseau 1998) and in L 1157 (Kwon et al. 2015). In addition, C-shaped bending where two protostars are present (Lee et al. 2009) could be attributed to an orbiting jet source.

There are numerous other systems in which multicomponent outflows are identified. Some can be modelled as a disc wind and a stellar jet; there is no substantial evidence for both these sources to be centred on distinct stars or for there to be a passive binary companion. For example, the fast HH 158 jet originating from DG Tau is accompanied by a slow molecular outflow (Güdel et al. 2018). These may share a common origin commensurate with the onion-like layered radial velocity structure (Bacciotti et al. 2000).

On the other hand, T Tau S is established as a binary, Sa/Sb, with a suggested circumbinary disc and up to three distinct outflows (Duchêne et al. 2005; Kasper et al. 2016; Yang et al. 2018). This prompts the question of how these outflows may interact on the scale of their development between 10 and 100 au.

Jet simulations provide an interface between physical theory and observation (Teşileanu et al. 2014). Thus, we consider here how to simulate jets from orbiting young stellar objects. Two competing scenarios, shown in Fig. 1, are investigated in which the launch

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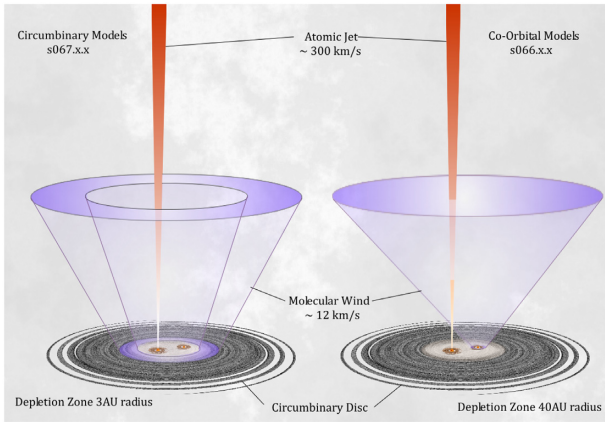
Figure 1. A schematic diagram of circumbinary (s067.x.x) versus co-orbital (s066.x.x) models, not to scale. A 'control' simulation within each model series is performed with the atomic jet only; these are simulations s067.3.3 and s066.2.4 using the terminology introduced in Section 2.

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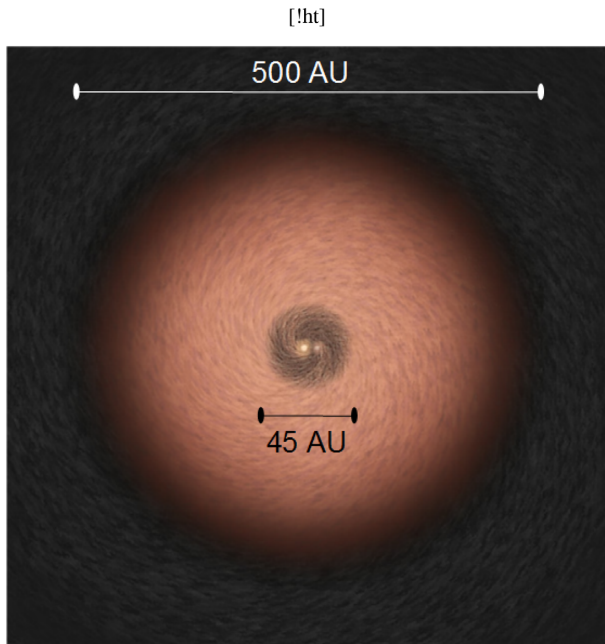
Figure 2. The underlying YSO system taken here, based on HH 30 with a circumbinary accretion disc, viewed perpendicular to the disc plane; schematic diagram only, central objects not to scale.

site of the molecular outflow differs. In the co-orbital scenario, the molecular flow is launched from the secondary binary partner, as proposed by Tambovtseva & Grinin (2008). On the other hand, in the circumbinary scenario advanced by Estalella et al. (2012), the jet launches from the inner edge of the circumbinary disc. The ALMA data for a few sources have been interpreted as favouring this second scenario provided a steady disc wind is assumed (Louvet et al. 2018; Matsushita et al. 2019).

The binary orbit and inner depletion zone of the circumbinary disc differ between the scenarios as illustrated in Fig. 2. In both the cases, a velocity-pulsed atomic jet emerges from the more massive object in the binary system. Control simulations were also carried out in which only the atomic jet was present.

The above scenarios are inspired by the HH 30 system that exhibits a well-collimated plume of hot, optically emitting atomic and partially ionized hydrogen (Gardner, Jones & Vargas 2018) and also a colder, dense, wide-angle molecular hydrogen outflow (Hartigan & Morse 2007). The HH 30 *HST* jet alone has attracted much attention, which involves variants on the above models. A two-component atomic MHD model was investigated by Teşileanu et al. (2014) employing cylindrical symmetry with a single star–disc outflow system. Staff et al. (2015) also investigated a two-component system that results after launch from a Keplerian disc.

Three-dimensional simulations are required due to the two jet sources. Such simulations were performed by Murphy et al. (2008) by taking two side-by-side nozzles through which atomic jets were injected. They found that the two jets interfere so that one ends up engulfed into the second one. In this manner, one can account for the propensity of single jet detections despite the high probability for stars to form in binaries. They were able to model L1551 features such as the bending of the secondary jet. The innovative work of Murphy et al. (2008) included attempts to implement a toroidal ambient magnetic field and also to show that slowly orbiting jets, rather than fixed nozzles, would not alter their high-speed flows that remain controlled by their dynamics.

The origins for the putative three outflows are assumed to be distributed between the three supplying discs since circumstellar discs are invariably associated with collimated outflows. The established mechanism for producing bipolar outflows involves a centrifugally driven wind flowing along a magnetic field that threads the disc (Blandford & Payne 1982). The disc can supply the material for the wind while the wind extracts most of the angular momentum from the disc. One alternative is the magnetic tower model in which a tightly coiled magnetic field rises from the spinning disc (Shibata & Uchida 1985; Lynden-Bell 1996). This tower is built on the high magnetic pressure and the steep pressure gradient along the symmetry axis (Huarte-Espinosa et al. 2012). A particular mechanism to drive fast atomic jets is based on the magnetic connection between the star and the inner edge of the disc. This is the X-wind model (Shu et al. 1988, 1994). In the context of this theory, we only consider here a uniform non-rotating supersonic flow from a moving nozzle. We assume that the nozzle scale is sufficient for centrifugal forces to be negligible.

If most stars form in binaries from accretion discs, we should be observing the effects of their combined dual outflows. However, simulations of orbiting atomic and molecular outflows have still to be performed. The fast atomic jet may be well collimated while the slow molecular flow may have a large opening angle. In this series of papers, we investigate the potential scenarios and the expected outcomes. We wish to determine if the molecules are entrained or destroyed, if the atomic jet is disrupted and ionized in the impact, how the pulses survive, and how the altered structure changes with time during an orbit.

This first study is limited to the comparison of a dual co-orbital jet-outflow scenario with a single orbiting jet within ∼120 au of the outflow source(s). The simulations were generated using the well-established Eulerian astrophysical code ZEUS-MP. In a second paper, we study and compare the precession-driven (circumbinary) scenario.

The physical and numerical assumptions made here are described in Section 2 below. We then describe the numerical experiments and the resulting physical structures, especially identifying those that differentiate the scenarios. Specific $H\alpha$ images and CO maps highlight the most distinctive signatures. In a further paper, we will

present the complete synthetic observations in terms of a range of imaging and spectroscopic diagnostics.

2 METHODS

2.1 Physics

The equations of fluid dynamics are solved numerically by the standard implementation of the ZEUS-MP astrophysics code (Hayes et al. 2006). A compressible, non-viscous, non-thermally conducting flow is assumed; in addition, there are equations for the evolution of radiation energy density and magnetic flux density (the Ideal-MHD approximation). These couple to the transport equations for momentum and energy by the inclusion of suitable source terms on the right-hand sides of these equations.

Energetic collisions lead to dissociation and ionization of molecules and atoms in addition to electronic, rotational, and vibrational transitions. Cooling occurs through radiative transitions, dissociation, and via dust grains. We use a modified implementation of ZEUS-MP that contains a simple three-species hydrogen chemistry model that tracks H_2, H, and H^+, based on the model presented in Glover & Mac Low (2007). Options are implemented in the code for other 'zoos' of chemical species, but the three-species chemistry was considered sufficient for our purposes, as the emission lines for our synthetic images (to follow) could all be computed in post-processing either directly from hydrogen species populations or by proxies based on those populations, and more complex chemistry would increase run times and storage requirements.

To model the cooling of the gas, we use cooling functions that were initially developed by Suttner et al. (1997) and Pavlovski et al. (2002) and radically improved by Glover & Mac Low (2007). Cooling and chemistry are solved simultaneously for consistency, using an implicit approach to ensure numerical stability. Compressional PdV heating is accounted for separately from radiative heating and cooling using the standard ZEUS-MP approach. The temperature exponents for the cooling terms are pre-tabulated to improve performance, and two separate temperature regimes are identified, which are handled by different routines, $T < 300\,\mathrm{K}$ and $T > 300\,\mathrm{K}$, since cooling efficiency is dominated by different components in each regime.

Though only material phases of H are dynamically traced in the simulations, the model assumes a composition including He and other elements at a concentration typical of the interstellar medium for chemistry and cooling purposes. This assumed composition permits calculation by proxy of these concentrations when determining synthetic emission properties.

The combination of orbital motion and jet magnetic field is not beyond the scope of this study but the results were very limited. First, the combination of a poloidal field and orbital motion was not handled well by the code. A toroidal field was made to work with orbiting jets, but the main effect was to modify the advancing bow shock into a conical morphology before exiting the grid, with little effect on the behaviour of the jet column. Both non-zero divergence and tiling issues were encountered. We therefore continued the study omitting the field and so all of the results presented here are derived from hydrodynamical simulations.

Nevertheless, magnetic fields must play an important role in the launch and collimation of protostellar jets with several basic mechanisms available (Blandford & Payne 1982; Shibata & Uchida 1985; Shu et al. 1994; Huarte-Espinosa et al. 2012). Purely radiation-driven or hydrodynamic models are problematic. In radiation models, photon momenta are insufficient to drive outflows even from the most luminous sources, while hydrodynamic models require the presence of a 'nozzle' formed by pressure gradients in the ambient medium (a De Laval nozzle) in order to collimate and accelerate a jet. But ambient pressures are not sufficient to provide this, nor could such a structure resist breakup from fluid dynamic instabilities. Conversely, magnetohydrodynamic launch models are able to account for both jet launching and collimation.

Finally, gravitational effects that may change the structure of the jet and ambient media are also omitted. While crucial in the initial acceleration region, we assume that the influence of gravity is negligible in the propagation region on the short dynamical time-scales considered here.

2.2 Dynamics of the HH 30 jets

The HH 30 outflow is thought to be powered two objects. Estimates put the total mass at $\sim 0.45\,\mathrm{M_\odot}$, with typical primary and secondary masses of ~ 0.31 and $\sim 0.14\,\mathrm{M_\odot}$, respectively, though these numbers vary depending on which model is chosen to explain the wiggling behaviour of the knots of bright emission in the jet, and the parameter space of each model allows a range of solutions (Anglada et al. 2007).

Based on these numbers, however, the mean orbital separation is 18 au, and the orbital period is 114 yr. The co-orbiting objects are surrounded by a circumbinary accretion disc whose optically illuminated region spans a 500 au diameter, with observations in molecular lines suggesting an extended diameter of ~ 850 au (Pety et al. 2006). Estalella et al. (2012) find that the binary components orbit within a depletion zone of ~ 40 au diameter, and that the orbital motion is the primary driving influence causing the helical wiggling appearance of the jet on parsec scales, with precession acting as a secondary influence, if present at all.

To summarize, deduced parameters for the different interpretations are accrued in Table 1. There are many missing entries where a model did not specify a value.

It has been suggested that there is variability in the jet velocity on long and short periods. The longer period variability in ejection velocity is likely to give rise to the knots of bright emission found in the HH 30 outflow on the 0.1 parsec scale (Raga et al. 1990). The short-period variability in the ejection velocity ($\sim$ months), possibly chaotic in nature, arises from variable accretion, which steepens into shock fronts that provide the main cause of heating and ionization of the jet material (Anglada et al. 2007; Hartigan & Morse 2007). In addition, the parsec scale outflow of HH 30 appears to be driven sideways in the sky plane, exhibiting a west-facing 'C'-shape curvature that may arise from systemic velocity, or impinging outflows or winds from nearby objects (Estalella et al. 2012).

A parallel strand of investigation into the HH 30 system involved observation in molecular lines (e.g. Pety et al. 2006). Mapping of HH 30 in HCO^+ and several isotopologues of the CO molecule, which emit at millimetre wavelengths, reveals a great deal, including Keplerian rotation of the accretion disc in the $^{13}CO(2\text{-}1)$ 1.35 mm line. Of particular interest for our work is the observation of a slower moving, cold, dense outflow imaged in the $^{13}CO(2\text{-}1)$ 1.35 mm line. This outflow of $\sim 12\,\mathrm{km\,s^{-1}}$ takes the form of a wide plume and is only observed emerging from the north-facing side of the accretion disc, where the atomic outflow is also most active, and is quiescent on the south face, similar to the atomic counterjet.

Later work by Tambovtseva & Grinin (2008) goes a stage further in developing a working model of the HH 30 molecular outflow, by performing simulations in which the outflow material is composed of particles that are ejected ballistically into the problem domain.

Table 1. Observed parameters of the HH30 system as deduced from several interpretations. No entry means that the value was not specified.

	Pety 2006	Anglada 2007	Hartigan 2007	De Colle 2010	Estalella 2012	
Velocity (Systemic)	7.25 ± 0.04	–	–	–	–	km s^{-1}
CB Disc Outer Radius	420 ± 25	–	–	–	∼250	au
CB Disc Inner Radius	–	–	–	–	∼40	–
CS Disc Outer Radius	–	–	–	–	≲6.00	–
CS Disc Inner Radius	–	–	–	–	∼0.07	–
Disc Axis Position Angle	32 ± 2	31.6	–	–	–	deg
Disc Inclination Angle	84 ± 3	–	–	–	–	deg
Disc Rotation Vector	North-East	–	–	–	–	–
Disc Temperature	12	–	–	–	–	K
Precession Angle (A) [2]	–	1.42 ± 0.12	–	–	–	deg
Precession Period (A) [2]	–	53 ± 15	–	–	–	yr
Half Opening Angle (A)	–	1.43 ± 0.12	2.6 ± 0.4	2.4	–	deg
Half Opening Angle (M)	30 ± 2	–	–	–	–	deg
Binary Separation [1]	–	9–18	–	–	18 ± 0.6	au
Binary Separation [2]	–	<1	–	–	–	–
Absolute Orbit (P) [1]	–	–	–	–	5.7 ± 0.9	au
Orbit Period [1]	–	53	–	–	114 ± 2	yr
Orbit Period [2]	–	<1	–	–	–	yr
Orbital Phase Angle (P)	–	–	–	–	95 ± 11	deg
Orbital Velocity (P) [1]	–	–	–	–	1.5 ± 0.2	km s^{-1}
Orbital Velocity (S) [1]	–	2-5	–	–	±	km s^{-1}
Flow Source (A) [1]	–	Secondary	–	–	Primary	–
Flow Source (A) [2]	–	Primary	–	–	–	–
Flow Velocity, Radial (A)	200 ± 0.09	100–300	–	–	–	km s^{-1}
Flow Velocity, Radial (M)	12 ± 2	–	–	–	–	km s^{-1}
Flow Velocity, Azim. (A)	–	–	–	–	–	km s^{-1}
Flow Velocity, Azim. (M)	<1.00	–	–	–	–	km s^{-1}
Flow Velocity Variability	–	–	–	–	–	–
Flow Inclination, North (A)	–	–	–	∼1	5	deg
Flow Mass (A)	2 × 10^{-8}	–	–	–	–	M$_\odot$
Flow Mass (M)	2 × 10^{-5}	–	–	–	–	M$_\odot$
Flow Mass Flux (A)	1 × 10^{-9}	–	–	1 × 10^{-8}	–	M$_\odot$ yr^{-1}
Flow Mass Flux (M)	6.3 × 10^{-8}	–	–	–	–	M$_\odot$ yr^{-1}
Flow Momentum (A)	4 × 10^{-6}	–	–	–	–	M$_\odot$ km s^{-1}^{t}
Flow Momentum (M)	2.4 × 10^{-4}	–	–	–	–	M$_\odot$ km s^{-1}^{t}
Flow Momentum Flux (A)	2.6 × 10^{-7}	–	–	–	–	M$_\odot$ km s^{-1} yr^{-1}
Flow Momentum Flux (M)	7.5 × 10^{-7}	–	–	–	–	M$_\odot$ km s^{-1} yr^{-1}
Flow Ionization (A)	–	–	0.05–0.40	–	–	–
Flow Temperature (A)	–	–	7.26 × 10^3	∼1 × 10^4	–	K
Flow No. Density (A)	–	–	1 × 10^6	–	–	–
Flow Width (A) @ 20 au	–	–	14 ± 3	15	–	au
Flow Width (A) @ 500 au	–	–	36 ± 4	30	–	au
Stellar Mass (Total) [1]	0.45 ± 0.04	0.25–2	–	–	0.45 ± 0.04	M$_\odot$
Stellar Mass (P) [1]	–	0.25–1	–	–	0.31 ± 0.04	M$_\odot$
Stellar Mass (S) [1]	–	–	–	–	0.14 ± 0.03	M$_\odot$
Stellar Mass (P) [2]	–	0.1–1	–	–	–	M$_\odot$
Stellar Mass (S) [2]	–	0.01–0.04	–	–	–	M$_\odot$
Stellar Luminosity (Total)	0.2	–	–	–	–	L$_\odot$

Note. Summarizes the findings of a number of observations and investigations into the nature of HH 30. This is not intended to be exhaustive but provides the basis for our choice of model parameters. (A) designates a parameter relating to the atomic jet; (M) to the molecular outflow. [1] and [2] are alternative scenarios: [1] is orbital and [2] is precession. Any parameter not identified as either is agnostic or else, by default, assumes the orbital scenario.

Four scenarios are investigated, based on the work of Anglada et al. (2007). The first three scenarios assume orbital motion of the molecular outflow source, with varying parametrizations. The fourth scenario assumes that the binary system is very close and that the wiggle in the atomic jet arises from tidally induced precession. In this scenario, the molecular outflow originates from the circumbinary disc, whose inner radius is much smaller (3 au) because of the tight binary orbit (0.75 au).

Tambovtseva & Grinin (2008) find that the model that assumes a very close binary orbit and a circumbinary disc source produces a closer resemblance to the observed morphology of the molecular outflow. This implies that precession of the atomic jet source is the primary origin of jet wiggling, with orbital motion a lesser influence.

In summary, analytical modelling of the atomic jet wiggle has suggested that the first scenario is the more likely, with the molecular outflow therefore being produced by an orbiting binary partner

also. However, computational modelling of the behaviour of the molecular outflow in the two cases favours the precession-driven scenario.

2.3 Implementing an orbiting source

In this paper, we present results from models of protostellar jets that are launched from a source undergoing orbital motion as part of a binary system. Numerical models of protostellar jets often assume a frame of reference in which the jet/counterjet source is at rest with respect to the surrounding medium, or else in a state of rectilinear motion. However, from *N*-body simulations of star cluster formation, and from our physical observations of the end products of star formation, these sources are often expected to occur in binary systems and will therefore be subject to acceleration. In addition, axisymmetric approximations are often imposed to reduce computational demands and to focus attention on a particular piece of internal jet physics, but in the case of jets from an orbiting source, this is not possible and a full three-dimensional model must be computed.

We assume here a binary protostellar system that possesses sufficient orbital stability to be well described by Kepler's equations. In reality, the objects are still in the process of accretion from a surrounding circumbinary disc of infalling material and so their mass will be time varying and hence their orbital characteristics will not be constant. However, the time-scales of the simulated jets are small in comparison to the accretion time-scale and so the assumption of constant mass is a reasonable one.

It is also possible that the orbits might be perturbed by anisotropy in the gravitational attraction of the surrounding accretion disc if there are local variations in the amount of matter in the disc. We assume that any accretion disc clumpiness is not biased significantly in any direction and thus any such perturbations may be ignored.

Kepler's equations describe the elliptical orbit of a member of a stable binary star system. In such a system, the binary partners co-orbit a common centre. This barycentre lies at one of the foci of each elliptical orbit. Depending on the mass ratio of the partners, the barycentre may be well outside both objects, but might also be embedded within the more massive one.

The code takes the orbital period, T, as given by $T^2 = 4\pi^2/[a^3 G(m_1 + m_2)]$, where a is the elliptical semimajor axis between masses m_1 and m_2. The relationship between eccentric anomaly, E, and mean anomaly, $M(t)$, is

$$M(t) = \frac{2\pi t}{T} = E - \epsilon \cdot \sin(E), \qquad (1)$$

where ϵ is the orbital eccentricity parameter.

The relationship between true anomaly, θ, and eccentric anomaly is

$$\cos(\theta) = \frac{\cos(E) - \epsilon}{1 - \epsilon \cdot \cos(E)}, \qquad (2)$$

with the radial distance from the barycentre as a function of the true anomaly given by $r(\theta) = a(1 - \epsilon^2)/[1 + \epsilon \cos(\theta)]$. In our implementation, the centre of the jet inlet on the $x = 0$ boundary is repositioned with each time-step, with the motion corresponding to that described by the Kepler equation. The key to this is determining the eccentric anomaly, E. Rearranging into a non-time-dependent form, we obtain

$$f(E) = 0 = E - \epsilon \cdot sin(E) - \frac{2\pi t_0}{T}, \qquad (3)$$

where t_0 is a specified, and thus constant, interval of time. This is a transcendental equation, the roots of which cannot be determined analytically. Therefore, a numerical method is required. In our implementation, the Newton–Raphson method is used, which converges to eight decimal places of accuracy after five iterations:

$$E_{n+1} = E_n - \frac{E_n - \epsilon \cdot \sin(E_n) - E_0}{1 - \epsilon \cdot \cos(E_n)}, \qquad (4)$$

in which the initial guess E_0 is taken to be the mean anomaly $\frac{2\pi t_0}{T}$. Given an accurate estimate of the eccentric anomaly, we can then calculate the true anomaly and the radial distance from the barycentre then follows. It is then simple trigonometry to calculate where the centre of the jet inlet must lie.

3 RESULTS

3.1 Orbiting atomic jet on 1000 au scale

We first simulate an extended atomic jet from an orbiting young star with the same parameters as the co-orbiting jet runs (see Table 1). This run with a single orbiting jet permits us to calibrate the code and check the dynamical behaviour. Figs 3 and 4 display all the physical outputs from this simulation at an early time of 63 yr and a late time of 189 yr, respectively. In the absence of a wide-angle molecular flow, a long problem domain is feasible; a $1280 \times 175 \times 175$ mesh was employed, running on 160 cores. The jet was $10\times$ overpressured with respect to the ambient medium and its inlet radius was 7 au.

The panels of Figs 3 and 4 show all the expected properties at an early time, with the leading bow on the domain, and at a late time when the flow has settled. A low-density low-pressure cavity is formed around the jet. Energy is transferred into the ambient medium that displays strong turbulence. The jet pulses steepen abruptly into shocks but the shocked layers expand as the shocks weaken, with the layers merging at ∼200 au. A small fraction of molecular hydrogen is seen to form in the compressed ambient medium. The initial region of high ionization is advected downstream with the advancing bow shock leaving only low-ionization regions within the jet and cavity.

3.2 Dual outflows

Although guided by HH 30, we are interested in the general problem as illustrated in Fig. 1. To quantify, a wide range of simulations were performed until a small number of relevant conditions could be taken forward for a close study.

The molecular outflow was added to the fast atomic jet through a distinct module. There was no method that could be found to sensibly implement a boundary condition where the two outflows were already interacting prior to their incursion into the domain.

Table 2 summarizes the problem domain set-up for ZEUS-MP used in the models that follow. A stack in the *x*-dimension of 20 thin slabs 8 zones in width, with each slab consisting of a 3×3 arrangement of square 115×115 zone tiles in the *y*–*z* plane, was taken. This has the advantage of fully containing the flow inlets and their orbits within the central tile. The 180 core configuration that was settled on, arising from a choice of 20 longitudinal slabs, worked well and queuing times were generally satisfactory.

Due to the wide angle of the molecular flow, simulations require a model geometry with a large lateral span. To accommodate an additional 30° molecular outflow, without truncating the problem domain, would require greater computing resources than those
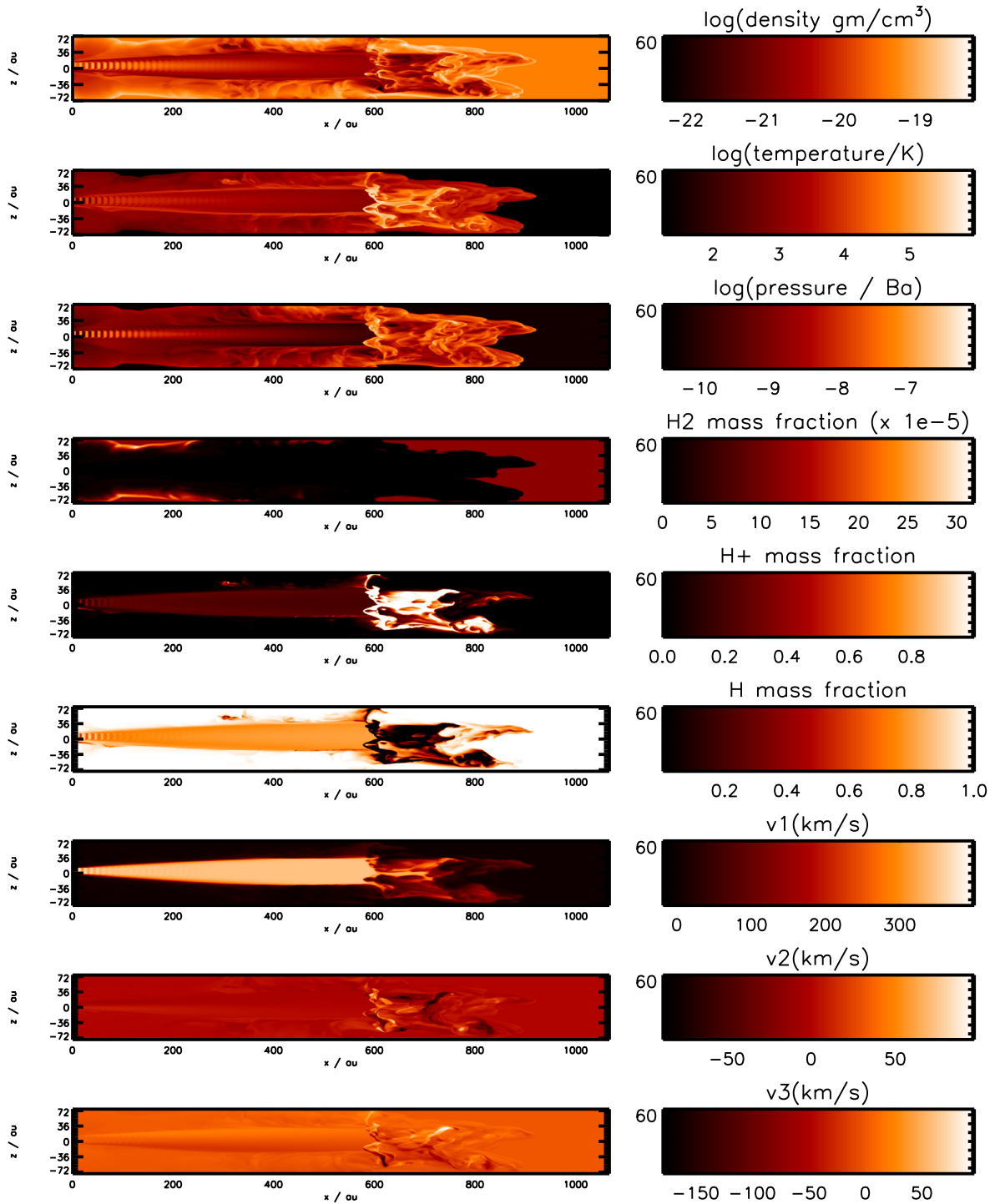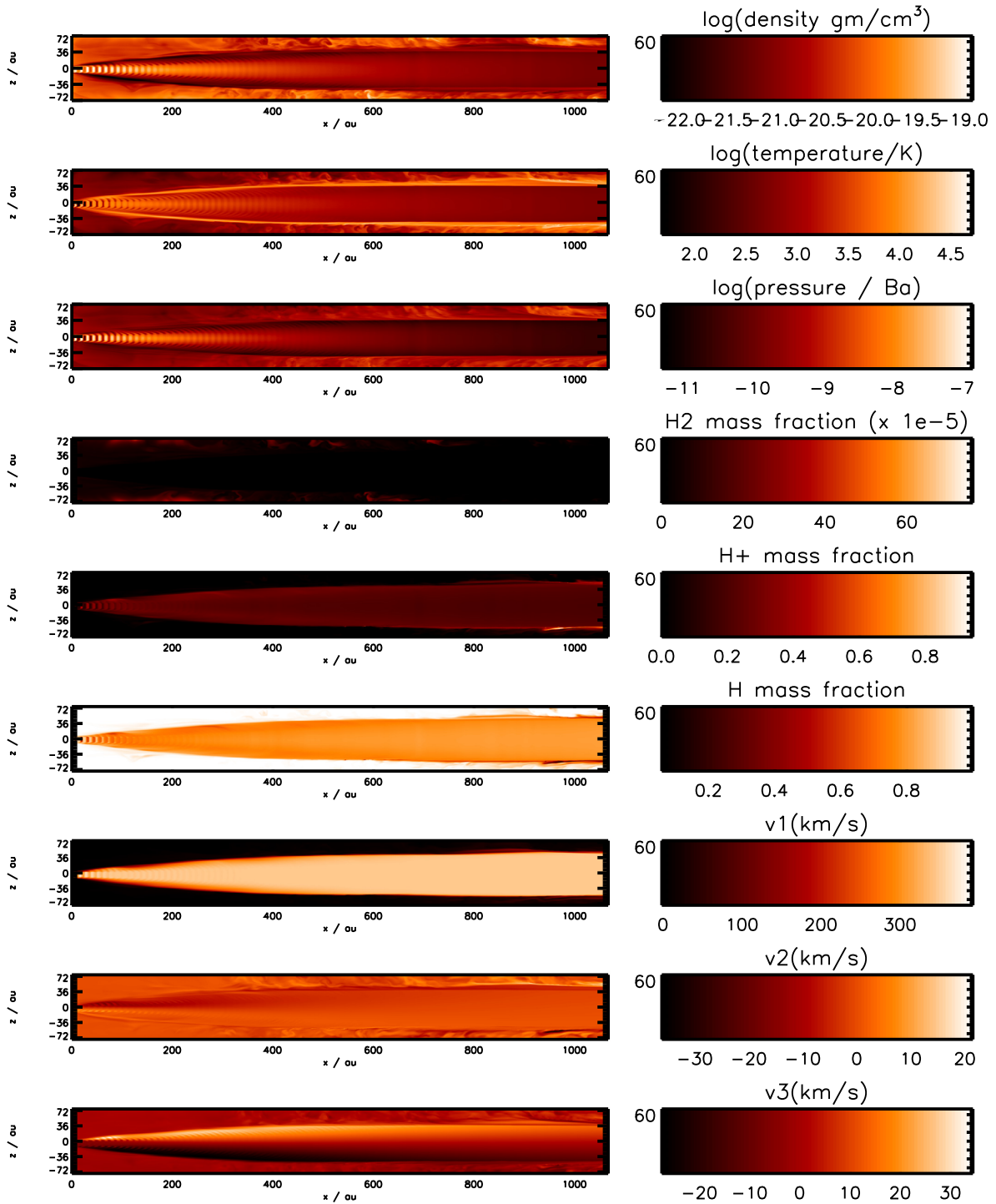
(a) Simulation Time: 63 years

Figure 3. Prototyping the HH 30 atomic jet: physical variables at the early time of 63 yr, model designated s064.2.2. This jet is from an orbiting inlet, the motion of which resembles the co-orbital series of models (s066.x.x) in the following discussion. The panels show the *x*–*z* mid-plane with the scale in au.

available. It was also desirable to increase the spatial resolution to capture the features of the flow in more detail for final results. It may well be that the most significant results relevant to future observations will appear within the first 100 au of the launch regions.

Table 3 summarizes the choices of parameters for the outflows and the initial ambient medium that was chosen to be entirely atomic with a temperature of ∼100 K. In fact, the simulations themselves were allowed to run for a sufficient length of time to 'nurture' their own problem domains, particularly in the simulation runs where

(a) Simulation Time: 189 years

Figure 4. Prototyping the HH 30 atomic jet: physical variables at the late time of 189 yr, model designated s064.2.2. This jet is from an orbiting inlet, the motion of which resembles the co-orbital series of models (s066.x.x) discussed below. The panels show the x–z mid-plane with the scale in au.

a molecular outflow component was present. The parameters in the upper section of Table 3 are configurable and could be input directly to ZEUS. The dependent parameters in the lower section are all quantities of interest that are directly calculated.

Table 4 summarizes the characteristics of the simulation runs. Co-orbital runs are designated by the prefix s066 and the circumbinary

by s067. Four major long time-scale (175 yr) runs were performed to establish fully developed flow structures. Both models were run with and without the molecular flow, designated with the suffix 2.3 and 2.4, respectively.

Six additional, shorter duration simulations were then carried out for each scenario: three with different velocity pulse time periods for

Table 2. Standard model geometry and tiling.

Coordinate	Min (cm)	Max (cm)	Span (cm)	Grid zones	Zone size (cm)	MPI Tiling	Zones /Tile
x	0	1.600E + 15	1.600E + 15	160	1.0E + 13	20	8
y	−1.725E + 15	1.725E + 15	3.450E + 15	345	1.0E + 13	3	115
z	−1.725E + 15	1.725E + 15	3.450E + 15	345	1.0E + 13	3	115

Table 3. Outflow and ambient medium parameters.

Parameter	Atomic outflow Type I	Type II	Molecular outflow Type I	Type II	Atomic medium	Units
Source Object	Primary	Primary	Secondary	CB Disc	–	–
Inner Radius	–	–	–	1.33E + 14	–	cm
Outer Radius	5.50E + 13	5.50E + 13	1.20E + 14	3.20E + 14	–	cm
Density	–	–	–	–	1.2525E−18	$g\,cm^{-3}$
Energy Density	–	–	–	–	1.3284E−08	$erg\,cm^{-3}$
Density Ratio	1	1	50	10	–	–
Pressure Ratio	10	10	7.5	1.5	–	–
Mach No.	95	95	30	30	–	–
Rotation (Solid)	1.31E−08	1.31E−08	6.53E−09	–	–	$rad\,s^{-1}$
Rotation (Kepler)	–	–	–	3.50E−09	–	$rad\,s^{-1}$
Radius (Kepler)	–	–	–	2.20E + 14	–	cm
Orbital Separation	18	0.75	–	–	–	au
Mass (Primary)	0.31	0.44	–	–	–	$M_\odot$
Mass (Secondary)	0.14	0.10	–	–	–	$M_\odot$
Precession Rate	–	3.76E−09	–	–	–	$rad\,s^{-1}$
Precession Angle	–	0.025	–	–	–	rad
Inner R (zones)	–	–	–	13.265	–	–
Outer R (zones)	5.5	5.5	12	32	–	–
Orbit Period	3.59E + 09	3.06E + 07	–	–	–	s
Precession Period	–	1.67E + 09	–	–	–	s
Adiabatic Exponent	1.666 67	1.666 67	1.428 57	1.428 57	1.666 67	–
No. Density	7.00E + 05	7.00E + 05	3.50E + 07	7.00E + 06	7.00E + 05	–
Density	1.25E−18	1.25E−18	6.26E−17	1.25E−17	1.25E−18	$g\,cm^{-3}$
Temperature	1090	1090	16.4	30	109	K
Sound Speed	3.43E + 05	3.43E + 05	3.89E + 04	4.20E + 04	1.09E + 05	$cm\,s^{-1}$
Inlet Flow Speed	3.26E + 07	3.26E + 07	1.26E + 06	1.17E + 06	–	$cm\,s^{-1}$
Inlet Area	9.50E + 27	9.50E + 27	4.52E + 28	2.66E + 29	–	cm^2
Mass Throughput	5.10E−09	5.10E−09	5.64E−08	6.15E−08	–	$M_\odot\,yr^{-1}$

Table 4. Simulation runs.

Simulations	Outflows Atomic	Molecular	Vpulse Period	Orbital ϵ	Retained dumps First	Last	T_1 (Y)	T_2 (Y)
s066.2.3	I	I	5.256E + 06	0.00	1	2599	0	175
s066.2.4	I	–	5.256E + 06	0.00	1	2599	0	175
s066.4.3	I	I	1.051E + 07	0.00	975	1300	65	87
s066.4.4	I	I	7.884E + 06	0.00	975	1300	65	87
s066.4.5	I	I	1.314E + 07	0.00	975	1300	65	87
s066.5.25	I	I	5.256E + 06	0.25	975	1300	65	87
s066.5.50	I	I	5.256E + 06	0.50	975	1300	65	87
s066.5.75	I	I	5.256E + 06	0.75	975	1300	65	87
s067.3.2	II	II	5.256E + 06	0.00	1	2599	0	175
s067.3.3	II	–	5.256E + 06	0.00	1	2599	0	175
s067.4.3	–	II	–	–	1	1199	0	80
s067.4.3	II	II	1.051E + 07	0.00	975	1300	65	87
s067.4.4	II	II	7.884E + 06	0.00	975	1300	65	87
s067.4.5	II	II	1.314E + 07	0.00	975	1300	65	87
s067.5.25	II	II	5.256E + 06	0.25	975	1300	65	87
s067.5.50	II	II	5.256E + 06	0.50	975	1300	65	87
s067.5.75	II	II	5.256E + 06	0.75	975	1300	65	87

the atomic jet and three with different values of orbital eccentricity. The data retained for these runs start from 65 yr into the outflow evolution (by which time the outflows have crossed the problem domain and exhibit fully developed flow) and track the evolution over a further 22 yr.

The velocity pulse period T_{vpulse} determines the period of the time-varying sinusoidal signal imposed on the velocity of injected material. For all the presented simulations, the Relative Amplitude parameter A_R used is set at 0.2. This generates a velocity signal:

$$V_J(t) = M \times C_J \times \frac{1 + A_R \ \cos \omega_v t}{1 + A_R}, \qquad (5)$$

where M is the Mach number of the jet, C_J is the jet sound speed, and $\omega_v = 2\pi/T_{vpulse}$. The maximum of this signal is MC_J and the minimum is 66 per cent of this value. The orbital eccentricity parameter, ϵ, is used in equation (4) to compute the Keplerian orbit of the jet inlet by the Newton–Raphson method. The HDF data dumps are produced at simulation time intervals of 2.125×10^6 s. This is chosen to be less than half of T_{vpulse}. It is a standard result in signal processing that in order to capture a sinusoidal signal the sampling frequency must be at least twice the highest frequency component of the signal. With our choice of $T_S \approx 2.5 T_{vpulse}$, 'strobe' effects are avoided in animations and the possibility exists to determine proper motions of structures within the flow.

Storage requirements were considerable. Each data dump in the listed simulations occupied 799 MB of memory space. Thus, the four 'primary' simulations required over 2 TB of storage each, while the 12 shorter simulations required 260.5 GB each.
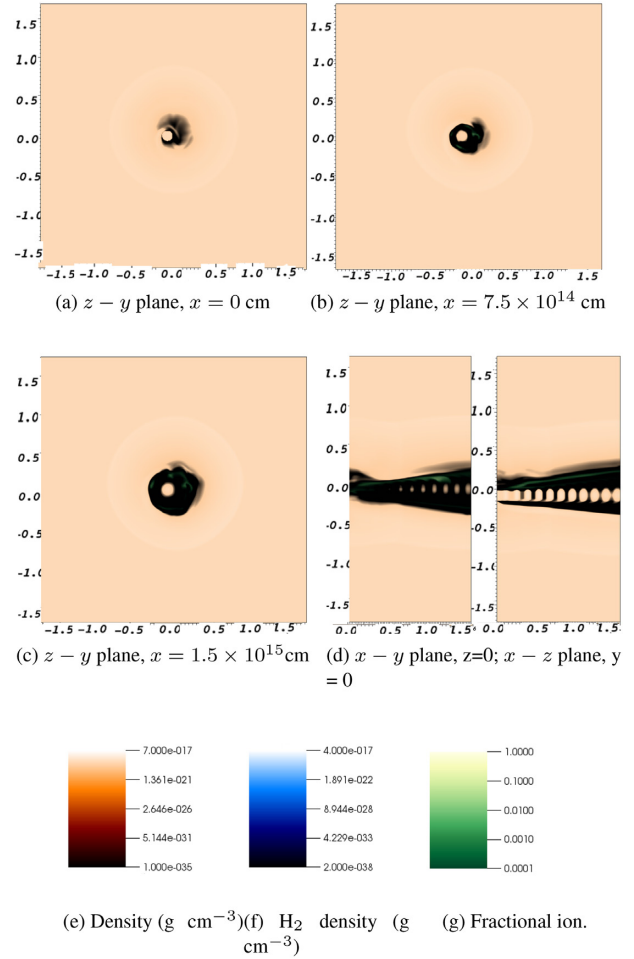
3.3 A single orbiting atomic jet at high resolution

The s066.x.x series of simulations model a scenario in which the fast-moving, atomic, optically emitting outflow of a T Tauri star is launched by the more massive binary partner in a two-star system, and the slow-moving, wide-angle molecular flow is launched from its lower mass co-orbiting partner. In this scenario, the co-radius is 18 au and the masses of the two objects are 0.31 and 0.14 $M_\odot$. Orbital eccentricity is considered in the final three simulations in this set in which values of ϵ range from 0.25 to 0.75.

The analysis of the Single Atomic Outflow Case (simulation s066.2.4) begins from Fig. 5 that shows a set of cross-sectional plots of the case where the molecular outflow is absent. The simulation time is 87.5 yr. Note that this figure shares a common colour scaling of variables and as a result some features appear faint. In spite of this, close examination reveals a spiral pattern in the density that radiates outwards from the jet column, induced in the surrounding medium by the orbital motion. Panel (a) shows the jet inlet at $x = 0$, and panels (b) and (c) show cross-sections further along the x-direction.

In subfigure (d), the jet is seen entering the domain from the left. The density distribution is sensitive to the jet location relative to the chosen mid-plane at the simulation time. An expanding cocoon of lower density material surrounds the denser jet column. We also see from panel (d) that the pulsed velocity signal has given rise to small-scale density knots within the jet column, sandwiched between regions of lower density. The knots appear to be expanding in the direction of travel as they cross the domain and exit the far x-boundary but remain well collimated in the y–z plane. The density knots are a feature of all the models though they become somewhat disrupted in the co-orbital mode where the atomic jet is in collision with the molecular outflow.

Pressure, temperature, and velocity fields of the 18 au single atomic outflow at the simulation time of 175 yr are shown in Fig. 6.


(a) $z - y$ plane, $x = 0$ cm   (b) $z - y$ plane, $x = 7.5 \times 10^{14}$ cm

(c) $z - y$ plane, $x = 1.5 \times 10^{15}$ cm   (d) $x - y$ plane, z=0; $x - z$ plane, y = 0

(e) Density (g cm^{-3})(f) H$_2$ density (g cm^{-3})   (g) Fractional ion.

Figure 5. Single atomic model under co-orbital conditions s066.2.4: 18 au binary, density cross-sections at a simulation time 87.5 yr. Axis scales are in units of 10^{15} cm. Ambient medium is atomic with trace molecular hydrogen formed during the simulation. Underlying density plot is fully opaque. For clarity, H$_2$ and ionization fraction overplots are at ramped opacity: 100 per cent opacity at maximum value, and transparent at minimum value.

Note that a scale for the partial density of H$_2$ appears on the plots since trace H$_2$ appears, arising from the molecular cooling routine that models dust grain catalytic formation of molecules. From these plots, it is evident that the low-density cocoon surrounding the jet column and also the low-density regions within the jet are populated by hot, partially ionized material.

The outer expanding shock is evident in the density distribution as a ring and in the pressure distribution as a filled circle. The shock introduces turbulence consistent with the simulations of atomic winds with strong radiative cooling (e.g. Novikov & Smith 2018). Examination of early-stage outputs in the simulation reveals that material driven outwards by the expanding outflows passes through this shock and then becomes disorganized. This can be seen occurring in the time-stepped images of the dual atomic–molecular outflow version of this model.

A common feature of all the models is lateral flow expansion. This is expected as all of the outflows (atomic and molecular) are overpressured with respect to the ambient medium (see Table 3). Based on an estimated circular radius computed from the jet cross-
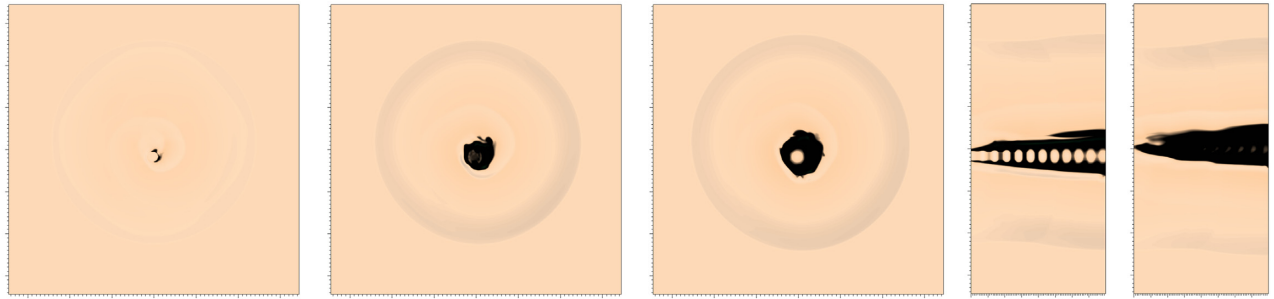
(a) density and composition at x=0 cm (b) density and composition, 50AU (c) density and composition, 100AU (d) density and composition, xz plane
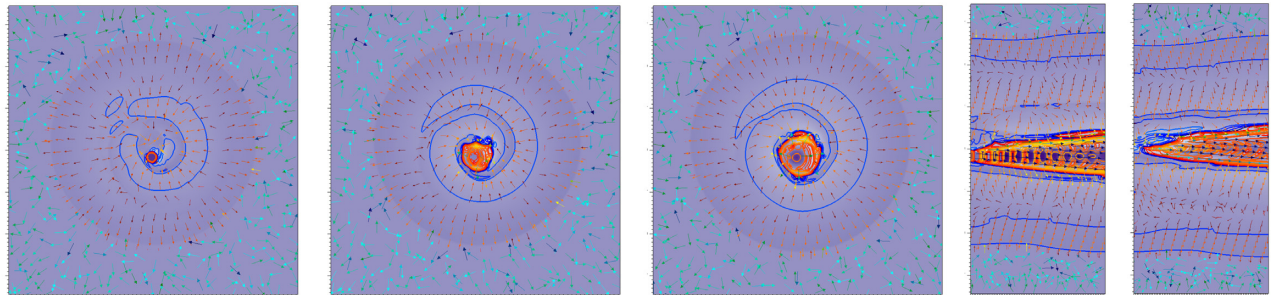
(e) p, T and v at x=0 cm (f) p, T and v at x=50AU (g) p, T and v at x=100AU (h) p, T and v in xy and xz planes
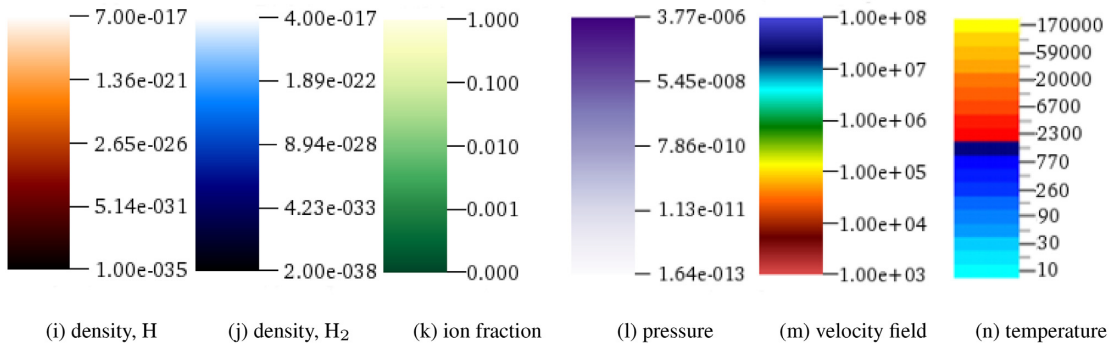
(i) density, H (j) density, H₂ (k) ion fraction (l) pressure (m) velocity field (n) temperature

Figure 6. **Cross-sections of physical parameters for the single atomic jet in the co-orbital configuration**, s066.2.4: 18 au binary, at a simulation time 175 yr. Axis scales are in units of 10^{15} cm. Ambient medium is atomic hydrogen with trace molecular hydrogen formed during the simulation. Underlying density plot is fully opaque.

sectional area (see Fig. 13), an approximate half-opening angle for the jet of the s066.2.4 simulation is 8.5°.

As the expanding material pushes outwards in its forward progress, it transfers some of its momentum in the *x*-direction to the ambient material in the boundary layer. This creates an entrained updraft. This is evident in all the simulations from the region of lower pressure surrounding the jet columns, and the surrounding velocity fields. There is no evidence of Kelvin–Helmholtz mixing of medium and jet material in the boundary layer. It should be noted that the prototyping long jet simulations did not indicate such features until after 200 au of jet propagation (see Fig. 3).
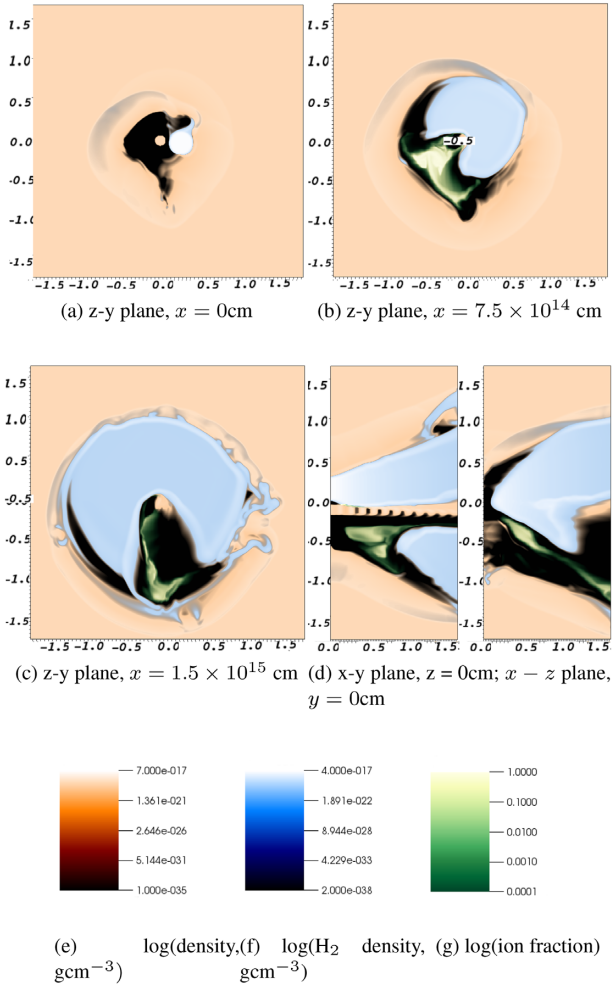
When the main outflow bow shock has left the problem domain, the return flow (backflow) is not adequately modelled. This phenomenon, a potential inaccurate feature of simulations, is discussed in Norman et al. (1982). However, the prototyping phase did not show evidence of any significant backflow. It was decided that backflow was not a feature that was likely to have a substantial influence on the dynamics of our jets in the region of interest here. This is particularly true of our main simulations that include the

dense molecular outflow, which will have a greater effect on the atomic outflow than a relatively weak backflow. Also of greater impact is the orbital motion of the inlets that will perturb the jet column to a greater extent than vortex shedding from the bow shock, which has its main impact close to the head of the jet column. Based on the prototype models, the head of the jet is 800 au+ distant during the simulation time span of our key results.

3.4 Co-orbital model s066.2.3: time evolution

The molecular outflow, emerging from the secondary binary partner, is now introduced. Fig. 7 shows a set of density cross-sectional plots at the simulation time 87.5 yr. As before, this figure gives an overview of the outflow structure, sharing a common colour scaling of variables.

As with the single-jet orbital simulation, a spiral density wave radiates outwards from the jet column, induced in the surrounding medium by the orbital motion of the jet. However, when we examine the dual-outflow system at 100 au along the *x*-axis [see panel (c) of

(a) z-y plane, $x = 0$ cm (b) z-y plane, $x = 7.5 \times 10^{14}$ cm

(c) z-y plane, $x = 1.5 \times 10^{15}$ cm (d) x-y plane, z = 0cm; $x - z$ plane, $y = 0$ cm

(e)   log(density, (f) log(H_2 density, (g) log(ion fraction) gcm^{-3})   gcm^{-3})

Figure 7. Co-orbital model s066.2.3: 18 au binary, atomic–molecular outflow, density cross-sections at a simulation time $t = 87.5$ yr. Axis scales are in units of 10^{15} cm. Ambient medium is atomic hydrogen with trace molecular hydrogen formed from cooling during the simulation. Underlying density plot is fully opaque. For clarity, H_2 and ionization fraction overplots are at ramped opacity: 100 per cent opacity at maximum value, and transparent at minimum value.

Fig. 7], we find that the expanding molecular flow is overtaking the density wave. An approximate calculation finds that the molecular flow expands at $\sim$6 km s^{-1} radially from the x-axis as it propagates across the domain, while the sound speed in the ambient medium is 1.09 km s^{-1} (see Table 3).

In the early stages of propagation, the molecular outflow is feeding material into the spiral density wave, which carries this material outwards just behind its advancing shock. This can be seen much more vividly in Fig. 8, which shows the flow at a more fully developed stage at 175 yr. By this stage, the slower moving shock has caught up with and overtaken the molecular outflow at $x = 100$ au, since the opening angle of the molecular flow remains unchanged between this and the 87.5 yr stage in Fig. 7.

In addition to a well-mixed partial density of H_2 molecules in the expanding spiral wave, there are gobbets of wholly molecular material. In panel (c) of Fig. 7, we find evidence of shearing instability pulling molecular material out from the main outflow where it interacts with the shock.

In Fig. 8, it appears that lumps of H_2 are being flung outwards, possibly by the action of the atomic and molecular outflow columns from co-orbiting sources, as they stir the surrounding medium like a giant cosmic egg whisk.

We see in Fig. 8, panels (d) and (h), that the pulsed velocity signal has given rise to small-scale density knots within the jet column, sandwiched between regions of lower density, just as we saw in the atomic-only case. However, from around $x = 20$ au onwards, the atomic jet is in direct contact with the molecular outflow along an advancing face, and the integrity of the knots is severely compromised, as a large crossing shock from the impact point sweeps through the jet while the flow carries it forwards.

There is substantial ionization occurring in this model, indicated by a green–white colour scale in the plots. The ionized material can be seen mainly in regions of very low density. However, a more detailed examination will show that these are not the regions where ionization is occurring. Panels (c) and (g) of Fig. 8 demonstrate that the atomic outflow does not pass peacefully through its wide-angle molecular companion. A large 'hole' has been blasted out of the side of the conical molecular flow. This gives a clue as to the origin of the ionized material: It is pouring out of the region where the leading face of the atomic jet is in contact with, and ploughing through, the molecular outflow.

Examination of panel (e) of Fig. 8 reveals an interesting feature. To the 'South' and 'West' of the barycentre of the two-jet system is a region of very low pressure ($\sim 10^{-13}$ Pa) and temperature (~ 10 K) that has formed. The velocity field indicates that the inflation of this cavity and drop in pressure and temperature are due to a rapid expansion of material, most of which appears to be cast off from the molecular outflow column.
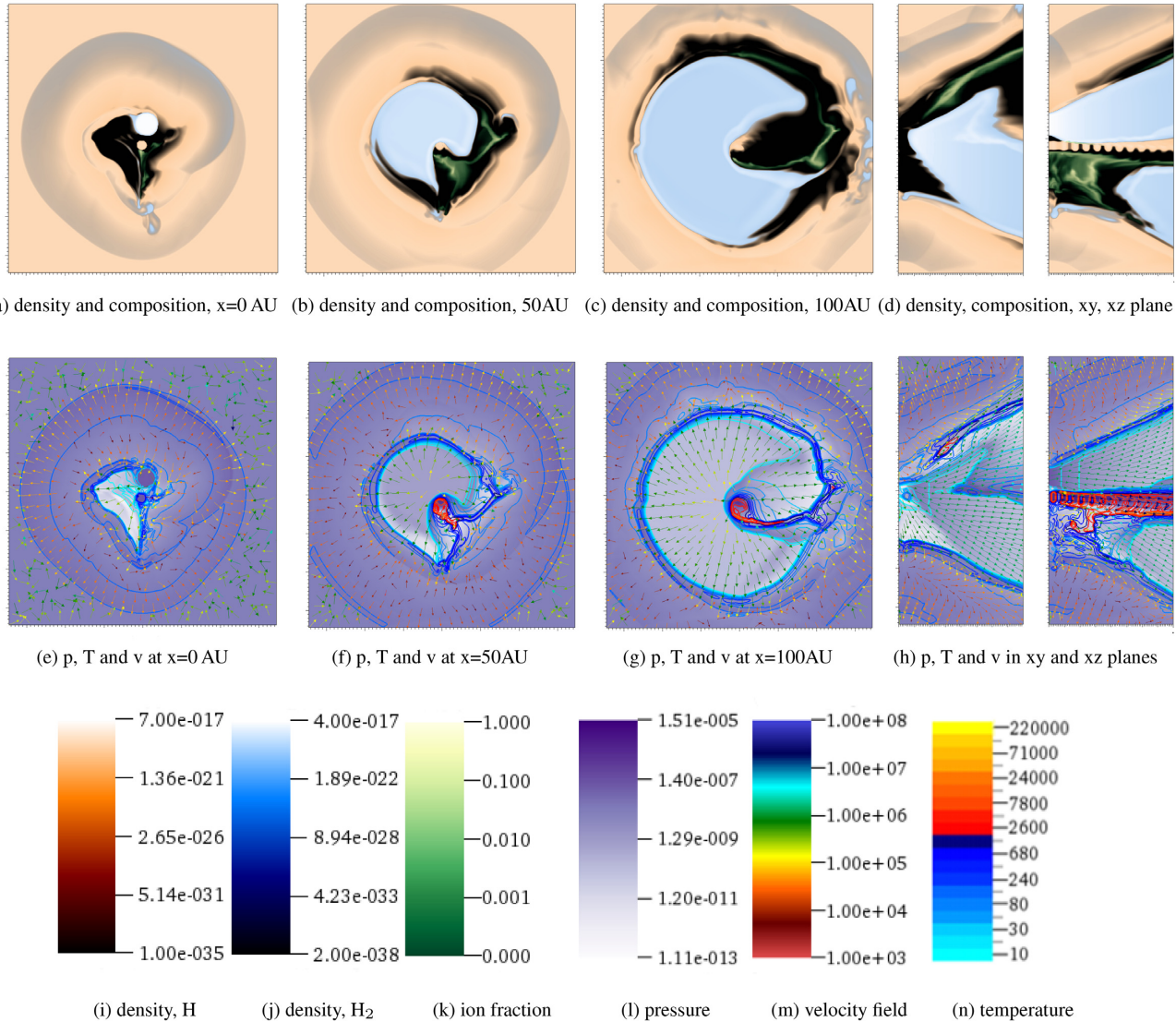
The behaviour is nicely clarified in close-up in Fig. 9 with the time sequence in Fig. 10, which depicts the two jet columns close to the inlets where these cavities form. The cavity begins as the trailing wake of the molecular outflow, and as this outflow moves off, material is drawn from the outflow by the pressure differential. The material rushes across this low-density region and meets the far 'wall' where the pressure it exerts supports the cavity against collapse. It gradually inflates radially outwards even as the continuing progress of the molecular outflow further expands the wake in the ambient medium.

This situation does not persist indefinitely. As the atomic jet swings around in its orbit, it penetrates into the cavity. The first effect of this is to disturb the flow and prevent further inflation of the cavity, though the flow separates in a portion of the cavity that carves off to form the starting point for a new cavity (see the '88 yr' subfigure). Then, some years later, a dramatic flaring event occurs as ionized material at 10^4–10^5 K erupts from the vicinity of the atomic jet column and floods out into the low-density cavity. The process then begins again.

In addition to helping to understand the formation of cavities, Fig. 10 also demonstrates how the flow is driven outwards into the ambient medium faster than the spiral wave; on passing through the forward shock, the flow becomes disorganized as previously remarked on.

3.5 Longitudinal analysis of co-orbital models

Quantities of interest pertaining to the mainly atomic jet (including its neutral, ionized, or entrained molecular hydrogen) were analysed as a function of distance along the (barycentric) x-axis of the problem domain, using IDL post-processing scripts developed for the purpose. The fast-moving material could be isolated with a filter

(a) density and composition, x=0 AU (b) density and composition, 50AU (c) density and composition, 100AU (d) density, composition, xy, xz plane

(e) p, T and v at x=0 AU (f) p, T and v at x=50AU (g) p, T and v at x=100AU (h) p, T and v in xy and xz planes

(i) density, H (j) density, H₂ (k) ion fraction (l) pressure (m) velocity field (n) temperature

Figure 8. **Cross-sections of physical parameters for the co-orbital model** s066.2.3: 18 au binary, atomic–molecular outflow at a simulation time 175 yr. Axis scales are in units of 10^{15} cm. Ambient medium is atomic hydrogen with trace molecular hydrogen formed during the simulation. Underlying density plot is fully opaque.

that selected for zones, where V_x was greater than 50 km s^{-1}. The exception was the analysis of ionized material that is simply selected material of that nature irrespective of velocity.

Generally, quantities examined were averaged over a slice of monozonal thickness for each of the 160 values of the x-coordinate, weighted by density or volume as appropriate, or in some cases the total quantity was determined. The graphs presented here include plots of the atomic–molecular co-orbital model (s066.2.3) represented by solid lines, and of the atomic-only co-orbital model (s066.2.4) by dashed lines. Given the density of data points, the single-valued nature of the functions, and lack of meaningful error that may be attached, simple connected lines have been used rather than showing separate data points.

Given the fact that quantities are being averaged over the cross-sectional area of the jet, it is reasonable to extend this principle in the x-direction also, so that each data point represents the average over a three-dimensional region of the jet. Therefore, the trend graphs smoothed in the x-direction are shown to the right of each plot.

Near-Gaussian smoothing is employed over a smoothing window equal to two jet inlet diameters. This averages out the variations related to the velocity pulsations of the atomic jet.

A dramatic change in the behaviour of the atomic jet when it encounters the molecular outflow is apparent in nearly every case. These show that the point of contact is at $x \approx 15$ au from the inlet boundary. In the smoothed graphs, the effect shifts downstream, to around 20–25 au.

The atomic jet is deflected and twisted by the presence of the molecular jet. This is evident in the lower right panel of Fig. 11, where we see, perhaps surprisingly, an inward deflection towards the x-axis, before the jet is deflected outwards, as we might naturally expect. The explanation for this is that the atomic jet entrains molecular material at the point where it first encounters the dense molecular wind; thus, sufficient x-momentum is imparted to some molecular material to push it into the >50 km s^{-1} velocity range we are selecting. This, thereby, suddenly weights the average position of the centre of mass closer to the radial origin. After this initial inward
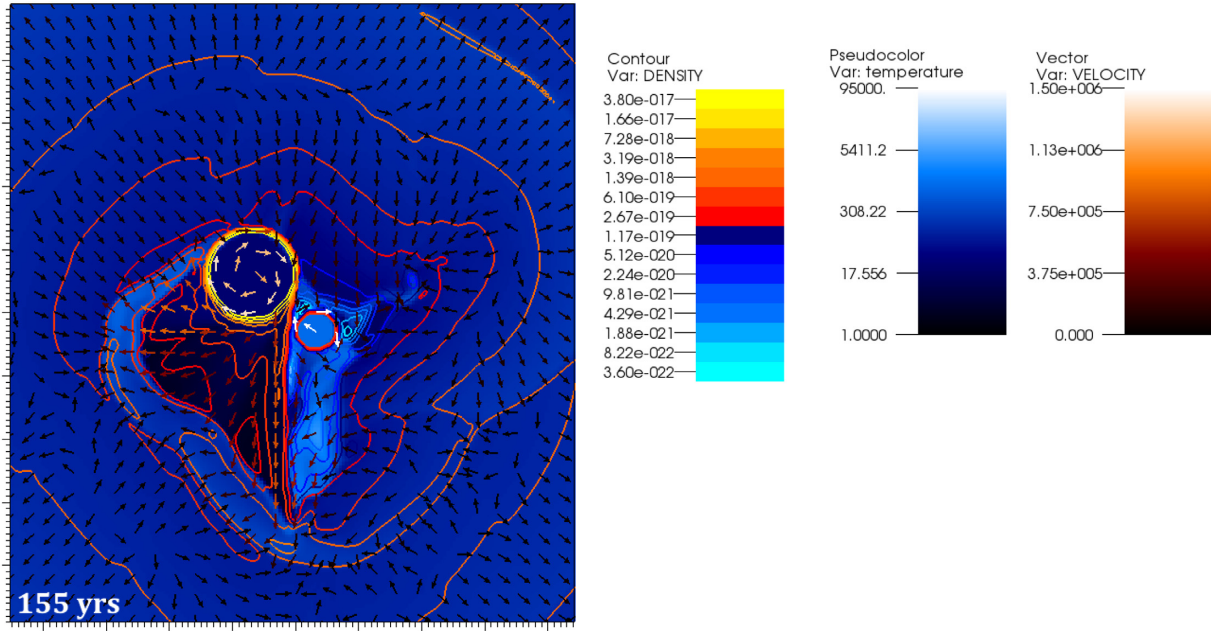
Figure 9. Atomic primary, molecular secondary, orbitally driven interaction; cross-section at 155 yr at $x = 3.2 \times 10^{13}$ cm. The background (blue) shows temperature. Velocity vectors are constant length, colour scaled. Contours show density. Model number: s066.2.3.

deflection, we see an outward deflection increasing monotonically, the radial limit of which may not be much greater than that of the undeflected atomic-only jet although a more elongated problem domain would be required to investigate this further.

However, Fig. 11 shows that the azimuth deflection angle attained by the end of the problem domain differs more conclusively from the unperturbed atomic jet (the azimuthal deflection of which is simply related to the jet's orbital motion). A difference of $9.5° \pm 0.1°$ is found. This has implications for the winding ratio or pitch of the large-scale helical outflow beyond the short domain examined here.

The forward-directed velocity of the atomic jet is analysed in Fig. 12. In the case of HH 30, this lies almost in the sky plane for the first 400 au of propagation. We see that there is a dramatic fall in the average velocity of the jet material from the point at which it encounters the molecular outflow. This is around 100 km s^{-1} negative differential against a velocity of 250–270 km s^{-1}, remaining near-constant in the velocity smoothed trend until the jet exits the domain. However, the peak velocity of material in the jet is virtually unchanged from the scenario where no molecular component is present. This implies that there is a spine to the jet that may be bent but is not impeded.

The width of the velocity distribution is displayed in Fig. 13. This shows that the perturbed jet has been squeezed, the jet cross-sectional area being only 65 per cent of the unperturbed jet.

Figs 14 and 15 show various other quantities, with self-evident differences between the perturbed and unperturbed jets. The ionization caused by the interaction is extremely large. As shown in Fig. 14, within the first 20 au of propagation, the perturbed jet produces an ionized mass 3–4 orders of magnitude greater than the unperturbed jet. In the dual-outflow co-orbital model, it is the interaction between the outflows that produces the overwhelming majority of ions. This increased ionization persists out until the exit boundary by which time the difference has fallen to one order of magnitude.

The mass per unit jet length shows a significant rise after the point of contact with the molecular outflow, approximately doubling its unperturbed value. Since the momentum (in the same figure) has increased by ∼18 per cent, we expect the velocity to reduce to ∼59 per cent of its value. Referring to Fig. 12, an approximate calculation finds the velocity to have reduced by ∼58 per cent. These calculations are approximate because the jet is not a closed system due to entrainment and mixing, and we do not expect exact conservation of momentum. None the less, we find that it is still an almost-conserved quantity over the 100 au of our problem domain.

4 SYNTHETIC IMAGING AND MAPS

Numerous atomic and molecular emission-line properties can be calculated via post-processing of the physical variables. We have developed an IDL code, MULTISNTH, to generate images, position–velocity diagrams, channel maps, and mass–velocity profiles. These are all useful as diagnostics for the specific model, and full details will be presented in a future paper.

Here, we present images of the Hα 656 nm emission superimposed on maps of the emission from the CO(2-1) rotational transition at 231 GHz. This serves to highlight the interaction at the impact zone between the outflows. Fig. 16 displays the suitably smoothed Hα images with CO contours.

The images demonstrate several defining features. First, the undisturbed atomic jet is found to generate a prominent string of Hα knots. The knots are more distinct in some snapshots but tend to blend together at other times. The jet axis clearly deviates from the domain symmetry axis, which may correspond to the underlying disc rotation axis. Secondly, the appearance of the atomic jet is disrupted by the proximity of the molecular outflow in this model. The heavy molecular outflow pushes hard on the atomic material, creating a warm bubble that is pushed out along the edge of the outflow. The molecular outflow impacts the jet at different angles, pushing the atomic material to different sides on the images. There is a considerable delay between the two emissions on the scale of
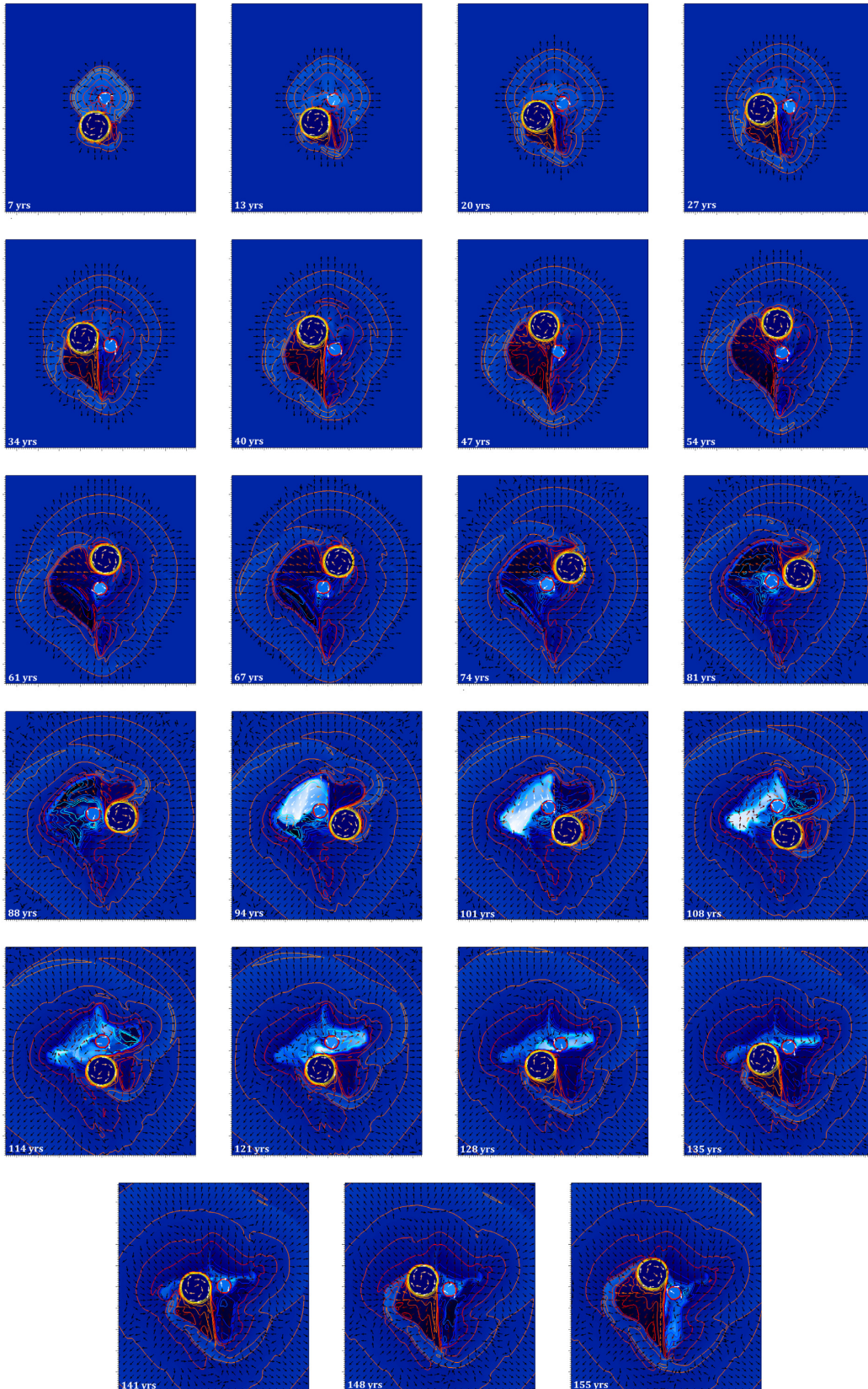
Figure 10. 18 au atomic primary, molecular secondary, orbitally driven interaction; cross sections at $x = 3.2 \times 10^{13}$ cm. The background (blue) shows temperature. Velocity vectors are constant length, colour scaled. Contours show density. Model number: s066.2.3.
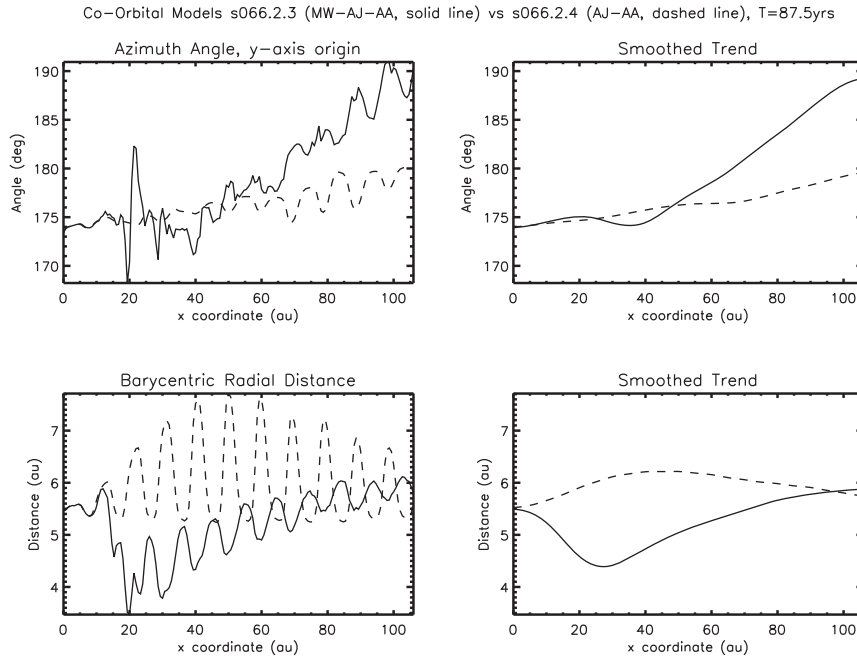
Figure 11. Atomic jet azimuth angle and radial distance, along propagation axis, at 87.5 yr for the co-orbital models. Radial distance is in the y–z plane and directed outwards with respect to the domain x-axis, which passes through the barycentre of the binary system. Azimuth angle is in the y–z plane and directed in an anticlockwise sense about the domain x-axis. Model numbers: s066.2.3 (with molecular outflow) and s066.2.4 (without).
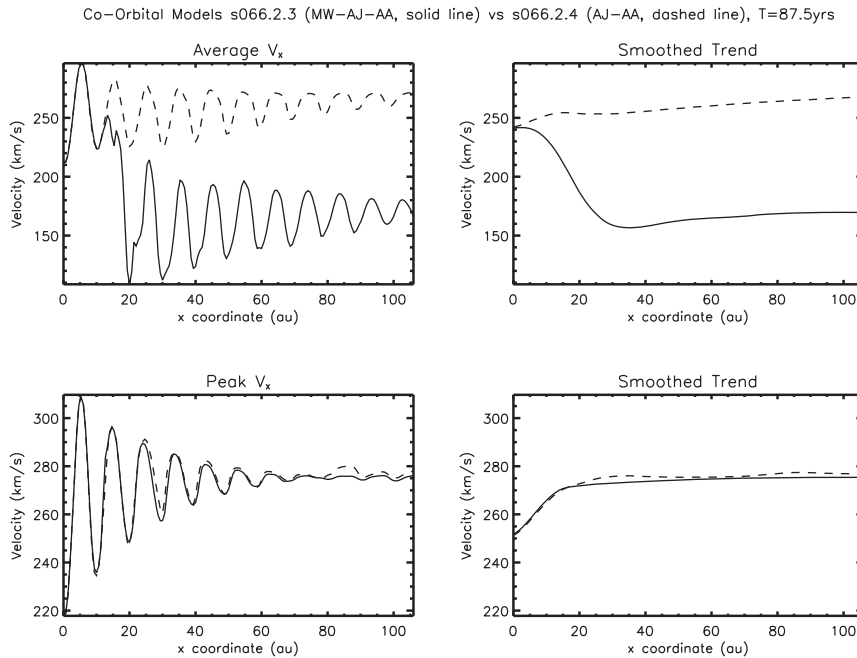
Figure 12. The average and peak velocity along the propagation axis of the atomic jet after 87.5 yr. Model numbers: s066.2.3 (with molecular outflow, solid line) and s066.2.4 (without, dashed line).

100 au. As a result, the H α and CO can be aligned or displaced. The optical jet is considerably wider with individual knots not discernible.

The CO outflow maintains its character of a limb-brightened cone. At late times, there is some indication of entrainment into the jet, which raises the possibility that molecular bullets could be catapulted out. However, it is not plausible to achieve high speeds before hydrodynamic instabilities disperse the entrained clumps.

5 CONCLUSIONS

We have discussed proposed interpretations of interacting jets associated with young stellar objects. After comparing the parameters that have been suggested, we then settled on two scenarios to take forward for a detailed study. In this paper, we have restricted the results to the co-orbital binary model in which an atomic jet and a molecular outflow are injected from discrete orbiting sources sepa-

Co-Orbital Models s066.2.3 (MW–AJ–AA, solid line) vs s066.2.4 (AJ–AA, dashed line), T=87.5yrs
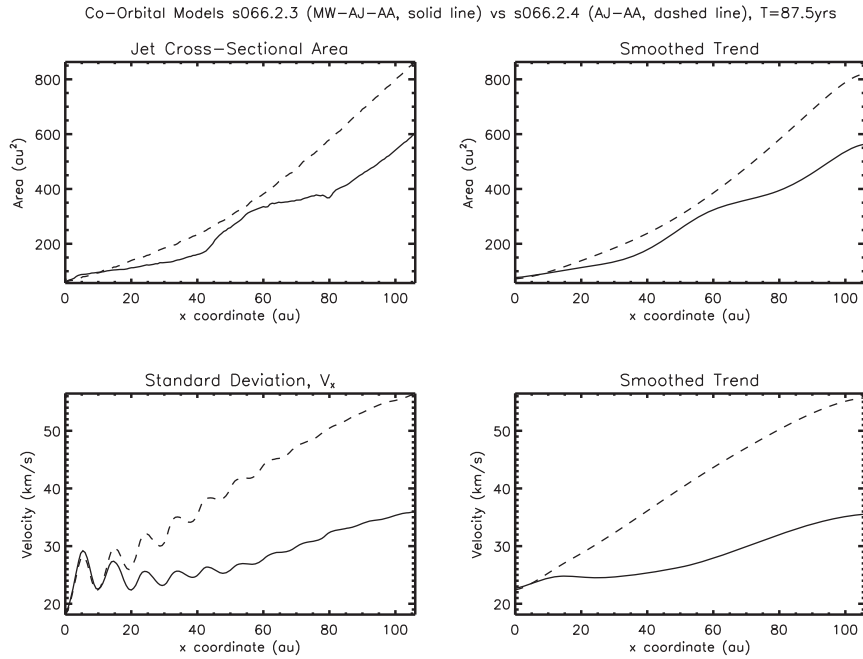
Figure 13. Co-orbital models: atomic jet *x*-sectional area and velocity standard deviation along propagation axis, $T = 87.5$ yr. Model numbers: s066.2.3 (with molecular outflow) and s066.2.4 (without).

Co-orbital Models s066.2.3 (MW–AJ–AA, solid line) vs s066.2.4 (AJ–AA, dashed line), T=87.5yrs
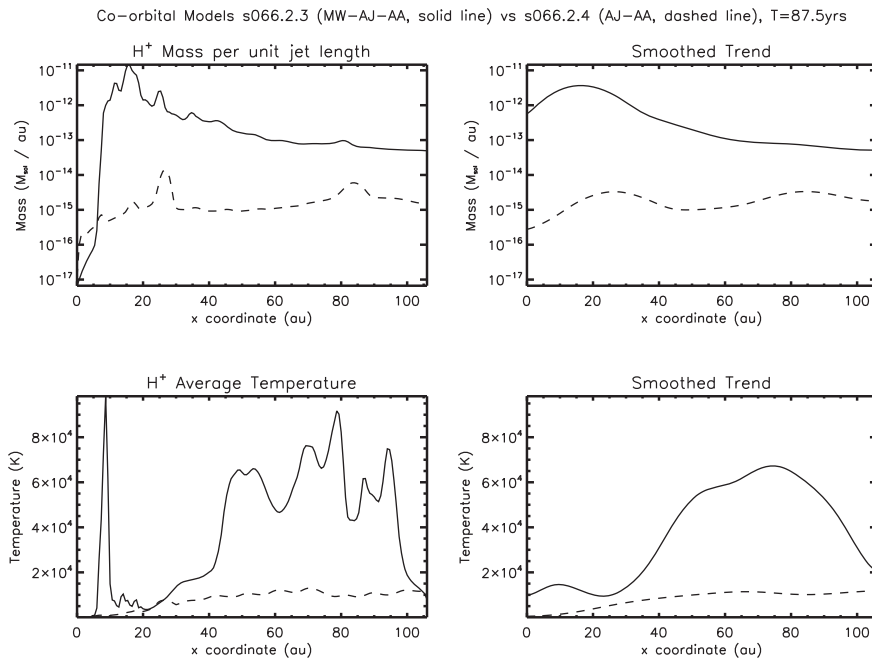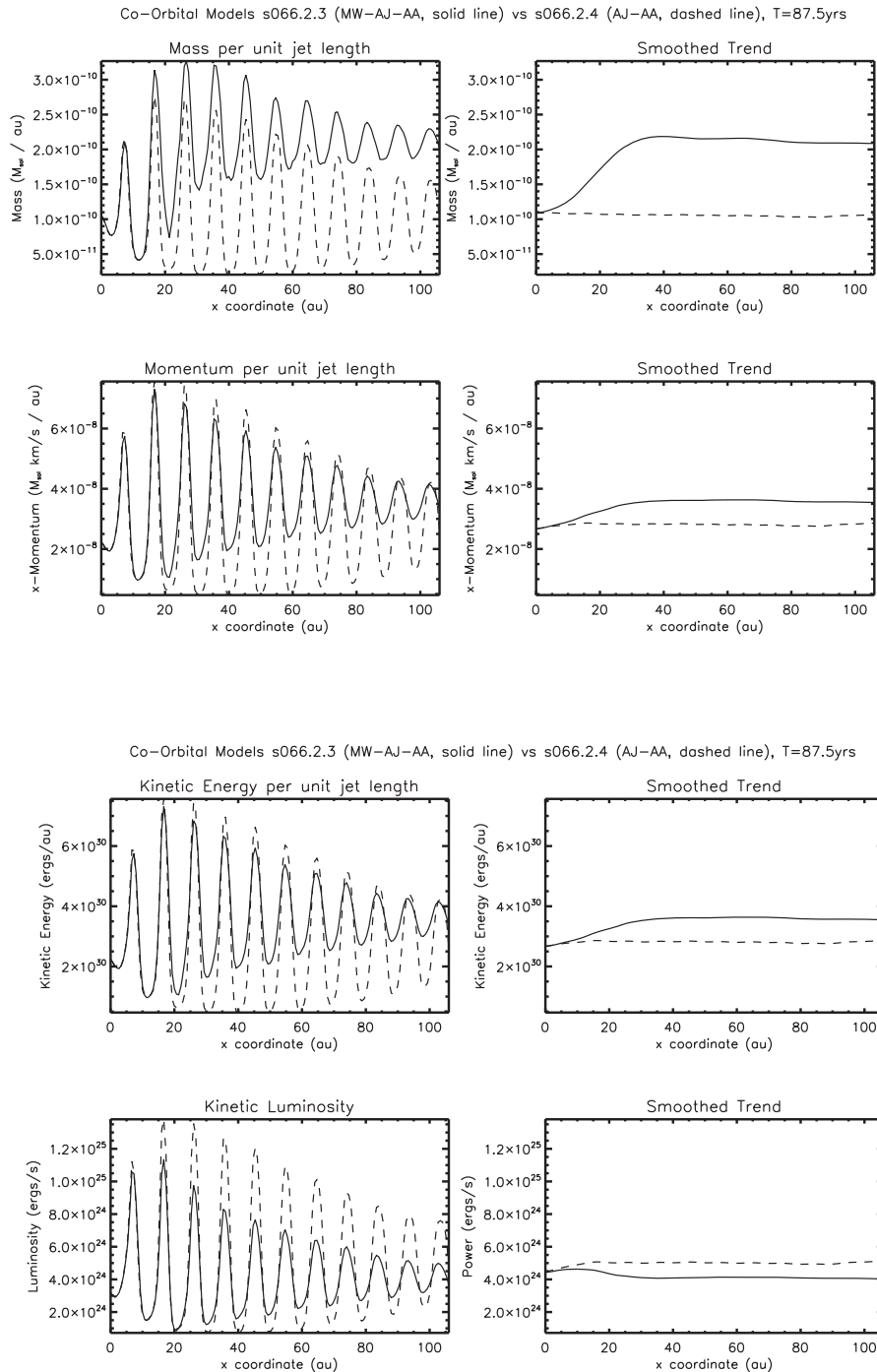
Figure 14. Ionization and ion temperature along the jet propagation axis at 87.5 yr. Model numbers: s066.2.3 (co-orbital with molecular outflow) and s066.2.4 (without).

rated by a quite short distance. The magnetic field is ignored. After some initial experiments, interesting field configurations within the jet presented computational inconsistencies. The hydrodynamic simulation begins on scales of the order of 10 au whereas the field-driven launch may be confined to within 0.1 au. Nevertheless, the field may remain crucial on the 10–100 au scale.

The hot (>1000 K) atomic jet was $10\times$ overpressured with respect to ambient and modulated with a sinusoidal velocity pulse signal. The period was 5.26×10^6 s. The jet's minimum pulsed inlet velocity was 66 per cent of its maximum velocity of 326 km s^{-1}. Further simulations were performed with different pulse characteristics. These are not included here but demonstrate that the pulse parameters do not alter the global flow dynamics. An atomic ambient medium was assumed for the T Tauri close environment; ambient medium gradients and other inhomogeneities were eschewed as the modelled outflows would be run for enough simulation time to nurture their own domains.

Co-Orbital Models s066.2.3 (MW–AJ–AA, solid line) vs s066.2.4 (AJ–AA, dashed line), T=87.5yrs

Figure 15. Atomic jet mass, momentum, kinetic energy, and kinetic power along the propagation axis, at time 87.5 yr. Model numbers: s066.2.3 (co-orbital with molecular outflow) and s066.2.4 (without).

Our main working simulations necessarily covered a shorter span of jet propagation (107 au) than the early prototyping simulations due to the need to accommodate the wide-angle molecular flow within the limits of available computing resources, and the desire for better resolution. The main atomic jet bow shock departed the problem domain well before the simulation time window in which results were calculated as we wished to examine the jet's steady-state behaviour.

The co-orbital simulations are discussed via cross-section plots of physical variables to illustrate the dynamics. In the absence of the molecular flow, the primary mode of ionization is atomic jet material processed through internal working surfaces within the pulsed jet column. When the co-orbital molecular flow is introduced, the main source of ionization is the shock boundary between the atomic and molecular outflows. The orbital dynamics of the two outflows produce some interesting structure in the surrounding medium, with low-density voids forming in the wake of the dense molecular outflow, which are then invaded and destroyed by the atomic outflow; this produces dramatic lateral flares of ionized material. An analysis of the jet's longitudinal characteristics has
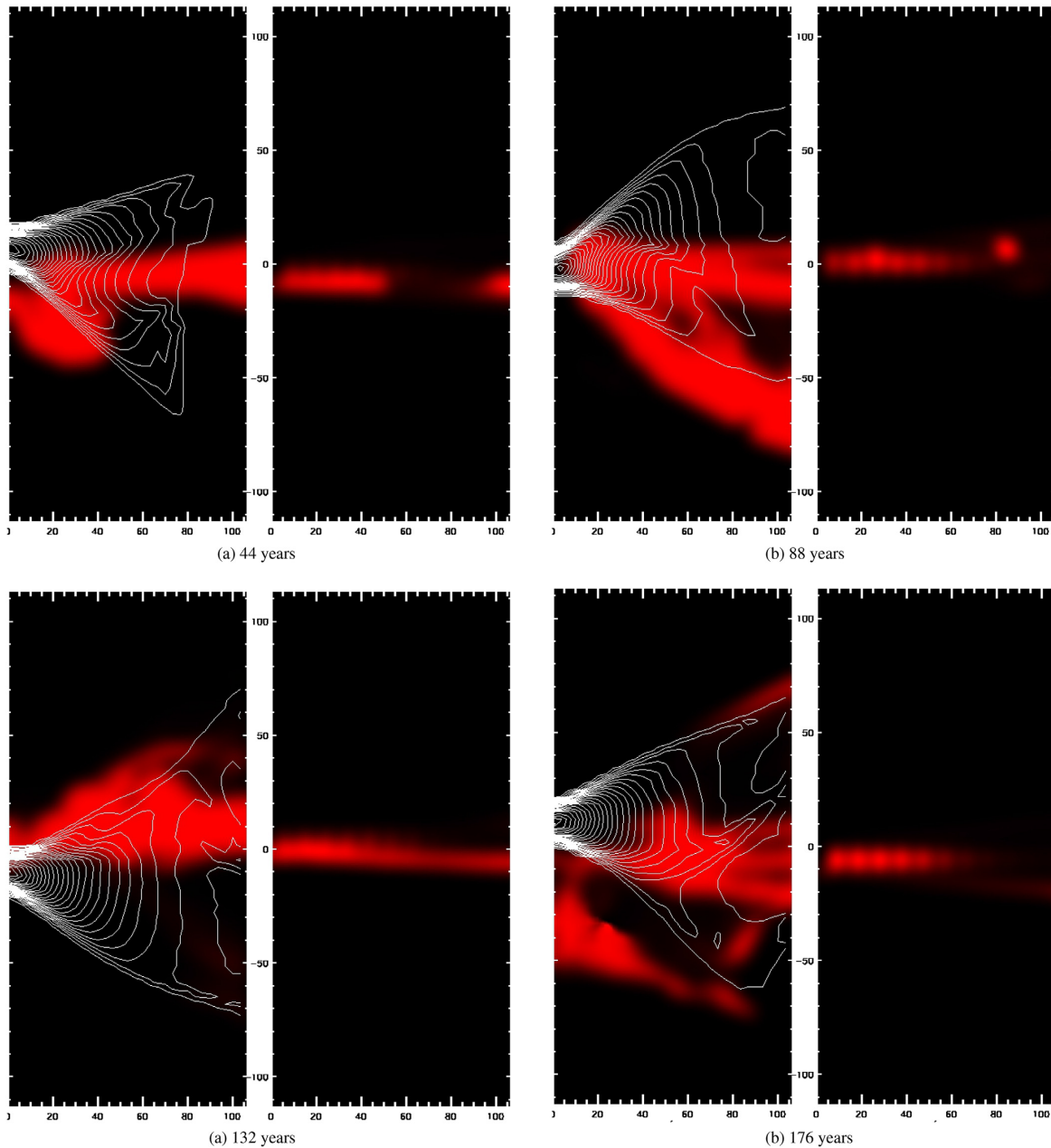
Figure 16. Synthetic Hα images and CO contours at the four indicated times from the co-orbital model with dual outflows (left-hand subpanels) and with just the atomic jet (right-hand subpanels). The optical emission is smoothed with a 14 au radius to match *HST*/WFPC2 pixel resolution of 0.1 arcsec at 140 pc. Contours indicate molecular material (unsmoothed).

shown various differences between the perturbed and unperturbed atomic jets.

Synthetic images in atomic Hα and molecular CO lines are presented, which demonstrate signatures specific to this model. In particular, the structure within the atomic jet is blurred and Hα emission is strong from sections of the walls of the CO cavity or from where the walls have been recently carved.

A surrounding circumbinary disc exists but is ignored here. This outer disc supplies the circumprimary disc, which is ultimately responsible for the fast atomic jet. In addition, it supplies the circum-secondary disc that feeds the heavier molecular outflow. We assume that the jet and outflow possess the same axial direction for the sake of these computations. However, some results can be generalized

and allow us to speculate on how misaligned outflows may interact as their sources orbit. In certain geometries, the strongly disturbed and ionized outflow may alternate on half the orbital period. In the context of the co-orbital interpretation of HH 30, a jet–counterjet asymmetry would occur out to 400 au. Additional simulations with eccentric orbits were also performed in this study (see Table 4) that show that the jet axial properties depend quite strongly on the chosen parameters.

Can the well-known HH 30 asymmetries be produced here? Asymmetric changes in brightness in the outer circumbinary disc could have a number of origins (Burrows et al. 1996; Stapelfeldt et al. 1999). One possibility raised is that of the passage of clumps in the molecular wind (Tambovtseva & Grinin 2008). These inter-

vening clumps could be sufficiently dense to provide obscuration of stellar light through dust extinction. Here, on the scales simulated, the wind itself is optically thin. Fig. 7 does show enhancements in the outflow's molecular density where it gets compressed against the atomic jet. However, there is no clear evidence for clumping and, in addition, compression in one dimension would not increase the dust column and hence the extinction.

A second known asymmetry is between the jet and counterjet. Relevant to this work on the physical parameters, we can discuss the velocity variations, the counterjet displaying more variations in the speeds of the knots (Estalella et al. 2012). Here, we remark from Fig. 12 that the passage of the atomic jet is hindered by the molecular outflow, reducing the axial flow speed considerably. Moreover, as a result of the shock interaction, mixing and dissociation, the jet luminosity falls but the momentum rises (Fig. 15).

To conclude, the wide molecular flow from the orbiting source significantly disturbs the atomic jet, deflecting and twisting it and disrupting the dense knots. Orbiting regions of high ionization are generated as the atomic jet rams through the molecular outflow. In the next work, we will place the atomic jet within the cone of the molecular outflow and study the differences.

These results provide a framework for the interpretation of upcoming subarcsecond observations. In particular, high-spatial resolution long-slit spectroscopy of forbidden lines from space will be achievable for more than just the brightest microjets. Physical quantities can be deduced along the jet axis and through the knots as done for DG Tau by Bacciotti et al. (2002) and Maurri et al. (2014) from the *Hubble Space Telescope* data. The new generation, led by the *James Webb Space Telescope* and the European Extremely Large Telescope, is capable of resolving structure on the scales of a few pixels as presented in the plots here.

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