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Observation of inverse Compton emission from a long γ -ray burst.

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251 **Long-duration gamma-ray bursts (GRBs) originate from ultra-relativistic jets launched from**
252 **the collapsing cores of dying massive stars. They are characterised by an initial phase of**
253 **bright and highly variable radiation in the keV-MeV band that is likely produced within the**
254 **jet and lasts from milliseconds to minutes, known as the prompt emission^{1,2}. Subsequently,**
255 **the interaction of the jet with the external medium generates external shock waves, respon-**
256 **sible for the afterglow emission, which lasts from days to months, and occurs over a broad**
257 **energy range, from the radio to the GeV bands¹⁻⁶. The afterglow emission is generally well**
258 **explained as synchrotron radiation by electrons accelerated at the external shock⁷⁻⁹. Re-**
259 **cently, an intense, long-lasting emission between 0.2 and 1 TeV was observed from the GRB**
260 **190114C¹⁰. Here we present the results of our multi-frequency observational campaign of**
261 **GRB 190114C, and study the evolution in time of the GRB emission across 17 orders of mag-**
262 **nitude in energy, from 5×10^{-6} up to 10^{12} eV. We find that the broadband spectral energy**
263 **distribution is double-peaked, with the TeV emission constituting a distinct spectral compo-**
264 **nent that has power comparable to the synchrotron component. This component is associ-**
265 **ated with the afterglow, and is satisfactorily explained by inverse Compton upscattering of**
266 **synchrotron photons by high-energy electrons.**

267 **We find that the conditions required to account for the observed TeV component are not**
268 **atypical, supporting the possibility that inverse Compton emission is commonly produced in**
269 **GRBs.**

270 On 14 January 2019, following an alert from the Neil Gehrels Swift Observatory (hereafter
271 *Swift*) and the *Fermi* satellite, the Major Atmospheric Gamma Imaging Cherenkov (MAGIC)

272 **telescopes** observed and detected radiation up to at least 1 TeV from GRB 190114C. Before the
273 MAGIC detection, GRB emission has only been reported at much lower energies, $\lesssim 100$ GeV, first
274 by *CGRO/EGRET* in a handful of cases, and more recently by *AGILE/GRID* and *Fermi/LAT* (see
275 ¹¹ for a recent review).

276 Detection of TeV radiation opens a new window in the electromagnetic spectrum for the
277 study of GRBs¹⁰. Its announcement¹² triggered an extensive campaign of follow-up observations.
278 Owing to the relatively low redshift $z = 0.4245 \pm 0.0005$ (see Methods) of the GRB (corresponding
279 to a luminosity distance of ~ 2.3 Gpc) a comprehensive set of multi-wavelength data could be
280 collected. We present observations gathered from instruments onboard six satellites and 15 ground
281 telescopes (radio, submm and NIR/optical/UV and very high energy gamma-rays; see Methods)
282 for the first ten days after the burst. The frequency range covered by these observations spans more
283 than 17 orders of magnitude, from 1 to $\sim 2 \times 10^{17}$ GHz, the most extensive to date for a GRB. The
284 light curves of GRB 190114C at different frequencies are shown in Fig. 1.

285 The prompt emission of GRB 190114C was simultaneously observed by several space mis-
286 sions (see Methods), covering the spectral range from 8 keV to ~ 100 GeV. The prompt light curve
287 shows a complex temporal structure, with several emission peaks (Methods; Extended Data Fig. 1),
288 with total duration ~ 25 s (see dashed line in Fig. 1) and total radiated energy $E_{\gamma,\text{iso}} = (2.5 \pm 0.1)$
289 $\times 10^{53}$ ergs (isotropic equivalent, in the energy range $1 - 10^4$ keV¹³). During the time of inter-burst
290 quiescence at $t \sim [5 - 15]$ seconds and after the end of the last prompt pulse at $t \gtrsim 25$ s, the flux
291 decays smoothly, following a power law in time $F \propto t^\alpha$, with $\alpha_{10-1000\text{keV}} = -1.10 \pm 0.01$ ¹³. The

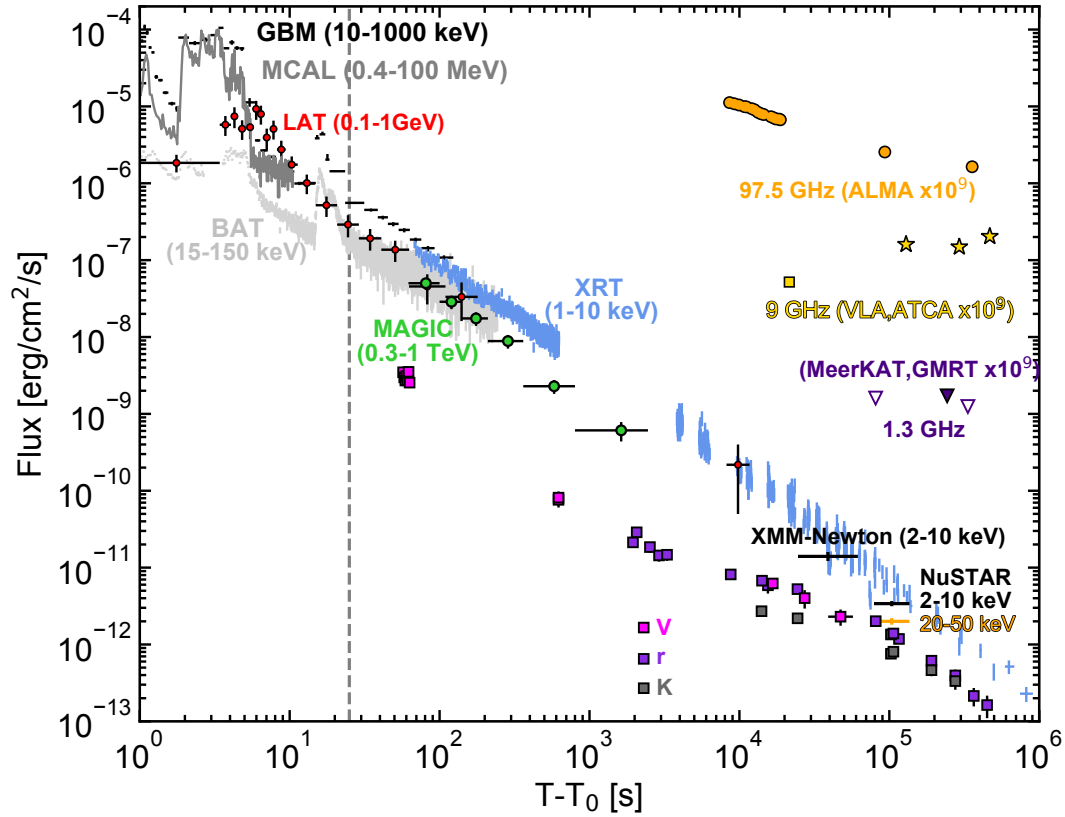


Figure 1: **Multi-wavelength light curves of GRB 190114C.** Energy flux at different wavelengths, from radio to gamma-rays, versus time since the BAT trigger time $T_0 = 20:57:03.19$ UT on 14 January 2019. The light curve for the energy range 0.3-1 TeV (green circles) is compared with light curves at lower frequencies. Those for VLA (yellow square), ATCA (yellow stars), ALMA (orange circles), GMRT (purple filled triangle), and MeerKAT (purple empty triangles) have been multiplied by 10^9 for clarity. The vertical dashed line marks approximately the end of the prompt emission phase, identified with the end of the last flaring episode. For the data points, vertical bars show the $1-\sigma$ errors on the flux, while horizontal bars represent the duration of the observation.

292 temporal and spectral characteristics of this smoothly varying component support an interpretation
293 in terms of afterglow synchrotron radiation, making this one of the few clear cases of afterglow
294 emission detected in the band $10 - 10^4$ keV during the prompt emission phase. The onset of the
295 afterglow component is then estimated to occur around $t \sim 5 - 10$ s^{13,14}, implying an initial bulk
296 Lorentz factor between 300 and 700 (Methods).

297 After about one minute from the start of the prompt emission, two additional high-energy
298 telescopes began observations: MAGIC and *Swift*/XRT. The XRT and MAGIC light curves (1-
299 10 keV, blue data points in Fig. 1, and 0.3-1 TeV, green data points, respectively) decay with time as
300 a power law, and display the following decay rates: $\alpha_X \sim -1.36 \pm 0.02$ and $\alpha_{\text{TeV}} \sim -1.51 \pm 0.04$.
301 The 0.3-1 TeV light curve shown in Fig. 1 was obtained after correcting for attenuation by the
302 extragalactic background light (EBL)¹⁰. The TeV-band emission is observable until ~ 40 minutes,
303 which is much longer than the nominal duration of the prompt emission phase. The NIR-optical
304 light curves (square symbols) show a more complex behaviour. Initially, a fast decay is seen,
305 where the emission is most likely dominated by the reverse shock component¹⁵. This is followed
306 by a shallower decay, and subsequently a faster decay at $t \gtrsim 10^5$ s. The latter behaviour might
307 indicate that the characteristic synchrotron frequency ν_m is crossing the optical band (Extended
308 Data Fig. 6), which is not atypical, but usually occurs at earlier times. The relatively late time at
309 which the break appears in GRB 190114C would then imply a very large value of ν_m , placing it in
310 the X-ray band at $\sim 10^2$ s. The millimeter light curves (orange symbols) also show an initial fast
311 decay where the emission is dominated by the reverse shock, followed by emission at late times
312 with nearly constant flux (Extended Data Fig. 3).

313 The spectral energy distributions (SEDs) of the radiation detected by MAGIC are shown in
314 Fig. 2, where the whole duration of the emission detected by MAGIC is divided into five time in-
315 tervals. For the first two time intervals, observations in the GeV and X-ray bands are also available.
316 During the first time interval (68-110 s, blue data points and blue confidence regions), *Swift*/XRT-
317 BAT and *Fermi*/GBM data show that the afterglow synchrotron component is peaking in the X-ray
318 band. At higher energies, up to $\lesssim 1$ GeV, the SED is a decreasing function of energy, as supported
319 by the *Fermi*/LAT flux between 0.1 and 0.4 GeV (see Methods). On the other hand, at even higher
320 energies, the MAGIC flux above 0.2 TeV implies a spectral hardening. This evidence is indepen-
321 dent of the EBL model adopted to correct for the attenuation (Methods). This demonstrates that
322 the newly discovered TeV radiation is not a simple extension of the known afterglow synchrotron
323 emission, but rather a separate spectral component that has never been clearly seen before.

324 The extended duration and the smooth, power-law temporal decay of the radiation detected
325 by MAGIC (see green data points in Fig. 1) suggest an intimate connection between the TeV
326 emission and the broadband afterglow emission. The most natural candidate is synchrotron self-
327 Compton (SSC) radiation in the external forward shock: the same population of relativistic elec-
328 trons responsible for the afterglow synchrotron emission Compton upscatters the synchrotron pho-
329 tons, leading to a second spectral component that peaks at higher energies. TeV afterglow emission
330 can also be produced by hadronic processes such as synchrotron radiation by protons accelerated
331 to ultra-high energies in the forward shock¹⁶⁻¹⁸. However, due to their typically low efficiency
332 of radiation⁶, reproducing the luminous TeV emission as observed here by such processes would
333 imply unrealistically large power in accelerated protons¹⁰. TeV photons can also be produced via

334 the SSC mechanism in internal shock synchrotron models of the prompt emission. However, nu-
335 merical modeling (Methods) shows that prompt SSC radiation can account at most for a limited
336 fraction ($\lesssim 20\%$) of the observed TeV flux, and only at early times ($t \lesssim 100$ s). Henceforth, we
337 focus on the SSC process in the afterglow.

338 SSC emission has been predicted for GRB afterglows^{9,11,17,19–26}. However, its quantitative
339 significance was uncertain, as the SSC luminosity and spectral properties depend strongly on the
340 poorly constrained physical conditions in the emission region (e.g., the magnetic field strength).
341 The detection of the TeV component in GRB 190114C and the availability of **multi-band** obser-
342 vations offer the opportunity to investigate the relevant physics at a deeper level. SSC radiation
343 might have been already detected in very bright GRBs, such as GRB 130427A. Photons with ener-
344 gies 10–100 GeV, as those detected in GRB 130427A are challenging to explain by the synchrotron
345 processes, suggesting a different origin^{27–29}.

346 We model the full data set (from radio band to TeV energies, for the first week after the
347 explosion) as synchrotron plus SSC radiation, within the framework of the theory of afterglow
348 emission from external forward shocks. The detailed modeling of the broadband emission and
349 its evolution with time is presented in Section Methods. We discuss here the implications for the
350 emission at $t < 2400$ s and energies above > 1 keV.

351 The soft spectra in the 0.2–1 TeV energy range (photon index $\Gamma_{\text{TeV}} < -2$; see Extended Data
352 Table 1) constrain the peak of the SSC component to be below this energy range. The relatively
353 small ratio between the spectral peak energies of the SSC ($E_p^{\text{SSC}} \lesssim 200$ GeV) and synchrotron

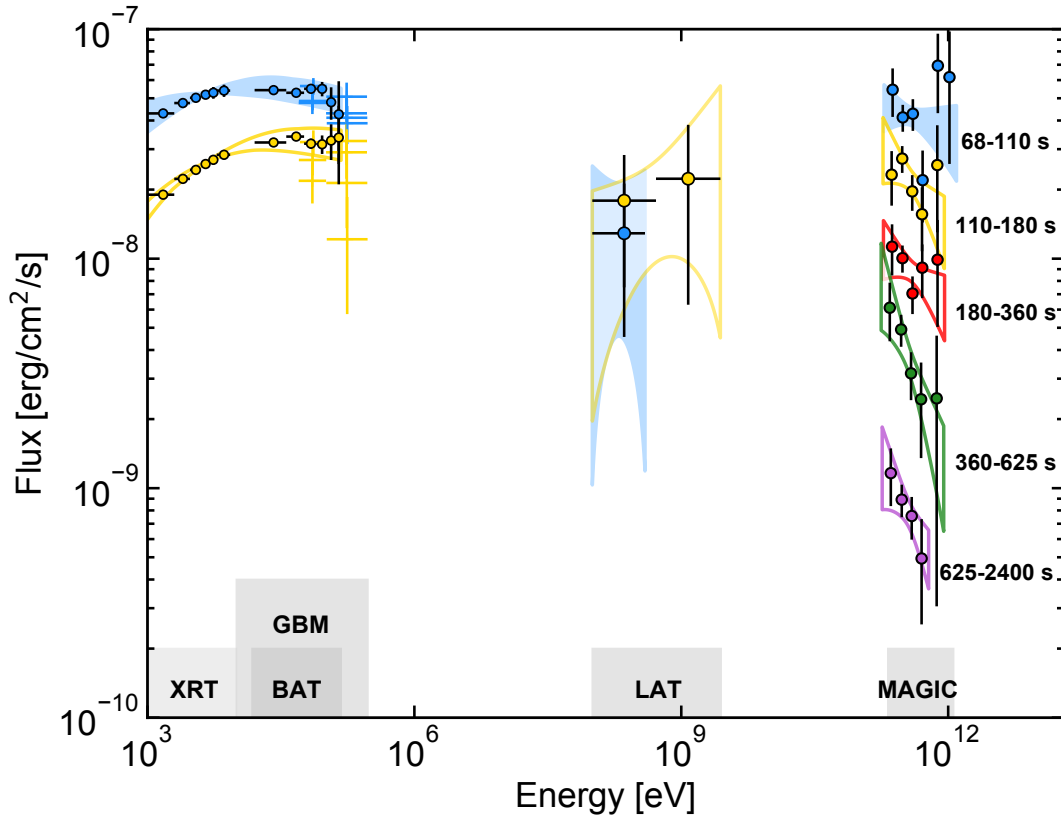


Figure 2: **Multi-band spectra in the time interval 68-2400 s.** Five time intervals are considered: 68-110 s (blue), 110-180 s (yellow), 180-360 s (red), 360-625 s (green), 625-2400 s (purple). MAGIC data points have been corrected for attenuation caused by the **extragalactic background light**. Data from other instruments are shown for the first two time-intervals: *Swift*/XRT, *Swift*/BAT, *Fermi*/GBM, and *Fermi*/LAT. For each time interval, LAT contour regions are shown limiting the energy range to the range where photons are detected. MAGIC and LAT contour regions are drawn from the $1\text{-}\sigma$ error of their best-fit power law functions. For *Swift* data, the regions show the 90% confidence contours for the joint fit XRT-BAT obtained fitting to the data a smoothly broken power law. Filled regions are used for the first time interval (68-110 s, blue color).

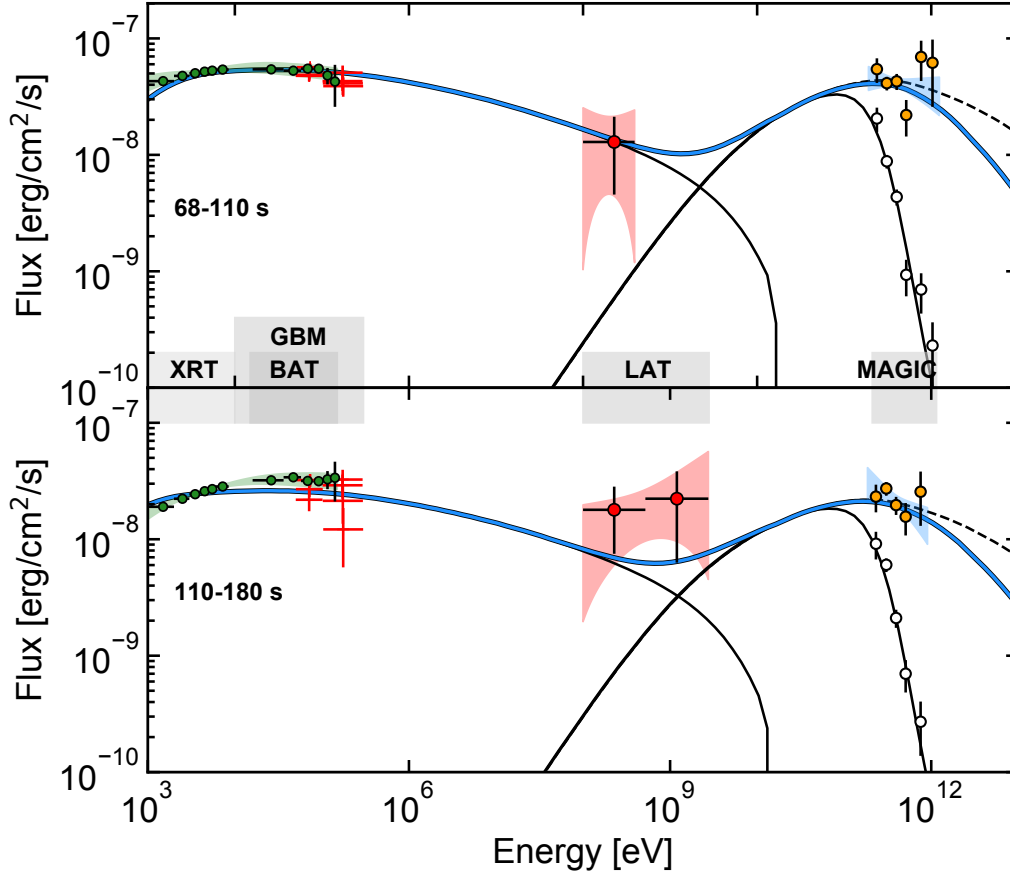


Figure 3: **Modeling of the broadband spectra in the time intervals 68-110 s and 110-180 s.**

Thick blue curve: modeling of the **multi-band** data in the synchrotron and SSC afterglow scenario.

Thin solid lines: synchrotron and SSC (observed spectrum) components; dashed lines: SSC if

internal γ - γ opacity is neglected. The adopted parameters are: $s = 0$, $\epsilon_e = 0.07$, $\epsilon_B = 8 \times 10^{-5}$,

$p = 2.6$, $n_0 = 0.5$, and $E_k = 8 \times 10^{53}$ erg, see the Text. Empty circles show the observed MAGIC

spectrum, i.e. not corrected by attenuation caused by the **extragalactic background light**. Contour

regions and data points as in Fig. 2.

354 ($E_p^{\text{syn}} \sim 10 \text{ keV}$) components implies a relatively low value for the Lorentz factor of the electrons
 355 ($\gamma \sim 2 \times 10^3$). This value is hard to reconcile with the observation of the synchrotron peak at $\gtrsim \text{keV}$
 356 energies. In order to explain the soft spectrum detected by MAGIC, it is necessary to invoke **the**
 357 **Klein-Nishina** (KN) regime scattering for the electrons radiating at the spectral peak as well
 358 as internal γ - γ absorption³⁰. While both effects tend to become less important with time, the
 359 spectral index in the 0.2-1 TeV band remains constant in time (or possibly evolves to softer values;
 360 Extended Data Table 1). This implies that the SSC peak energy is moving to lower energies and
 361 crossing the MAGIC energy band. The energy at which attenuation by internal pair production
 362 becomes important indicates that the bulk Lorentz factor is ~ 140 -160 at 100 s.

363 An example of the theoretical modeling in this scenario is shown in Fig. 3 (blue solid curve,
 364 see Methods for details). The dashed line shows the SSC spectrum when internal absorption is
 365 neglected. The thin solid line shows the model spectrum including EBL attenuation, in comparison
 366 to **the MAGIC observations** (empty circles).

367 We find that acceptable models of the broadband SED can be obtained if the conditions at
 368 the source are the following. The initial kinetic energy of the blastwave is $E_k \gtrsim 3 \times 10^{53} \text{ erg}$
 369 (isotropic-equivalent). The electrons swept up from the external medium are efficiently injected
 370 into the acceleration process, and carry a fraction $\epsilon_e \sim 0.05 - 0.15$ of the energy dissipated at
 371 the shock. The acceleration mechanism produces an electron population characterized by a non-
 372 thermal energy distribution, described by a power law with index $p \sim 2.4 - 2.6$, injection Lorentz
 373 factor $\gamma_m = (0.8 - 2) \times 10^4$ and maximum Lorentz factor $\gamma_{\text{max}} \sim 10^8$ (at $\sim 100 \text{ s}$). The magnetic

374 field behind the shock conveys a fraction $\epsilon_B \sim (0.05 - 1) \times 10^{-3}$ of the dissipated energy. At
375 $t \sim 100$ s, corresponding to $R \sim (8 - 20) \times 10^{16}$ cm, the density of the external medium is
376 $n \sim 0.5 - 5 \text{ cm}^{-3}$, and the magnetic field strength is $B \sim 0.5 - 5$ Gauss. The latter implies that
377 the magnetic field was efficiently amplified from values of a few μ Gauss that are typical of the
378 unshocked ambient medium, due to plasma instabilities or other mechanisms⁶. Not surprisingly,
379 we find that $\epsilon_e \gg \epsilon_B$, that is a necessary condition for the efficient production of SSC radiation^{17,19}.

380 The blastwave energy inferred from the modeling is comparable to the amount of energy
381 released in the form of radiation during the prompt phase. The prompt emission mechanism must
382 then have dissipated and radiated no more than half of the initial jet energy, leaving the other half
383 available for the afterglow phase. The modeling of the **multi-band** data also allows us to infer how
384 the total energy is shared between the synchrotron and the SSC components. The resultant power
385 in the two components are comparable. We estimate that the energy in the synchrotron and SSC
386 component are $\sim 1.5 \times 10^{52}$ erg and $\sim 6.0 \times 10^{51}$ erg respectively in the time interval 68-110 s, and
387 $\sim 1.3 \times 10^{52}$ erg and $\sim 5.4 \times 10^{51}$ erg respectively in the time interval 110-180 s. Thus, previous
388 studies of GRBs may have been missing a significant fraction of the energy emitted during the
389 afterglow phase that is essential to its understanding.

390 Finally, we note that the values of the afterglow parameters inferred from the modeling fall
391 within the range of values typically inferred from broadband (radio-to-GeV) studies of GRB af-
392 terglow emission. This points to the possibility that SSC emission in GRBs may be a relatively
393 common process that does not require special conditions to be produced with power similar to

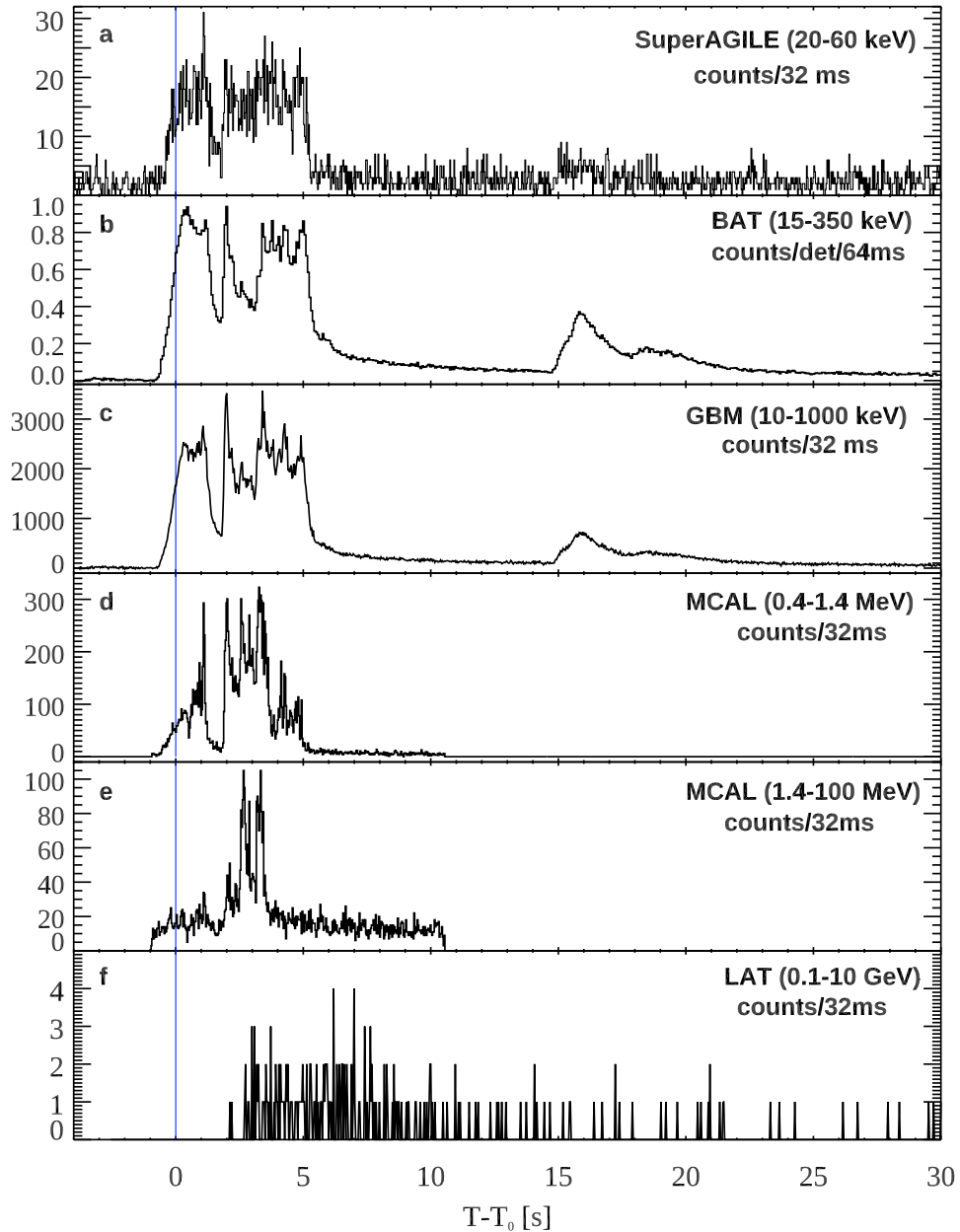
394 synchrotron radiation.

395 The SSC component may then be detectable at TeV energies in other relatively energetic
396 GRBs, as long as the redshift is low enough to avoid severe attenuation by the EBL. This also
397 provides support to earlier indications for SSC emission at GeV energies ²⁷⁻²⁹.

398 **Methods**

399 **Prompt emission observations** On 14 January 2019, the prompt emission from GRB 190114C
400 triggered several space instruments, including *Fermi*/GBM³¹, *Fermi*/LAT³², *Swift*/BAT³³, Super-
401 *AGILE*³⁴, *AGILE*/MCAL³⁴, *KONUS*/Wind³⁵, *INTEGRAL*/SPI-ACS³⁶, and *Insight*/HXMT³⁷. The
402 prompt emission light curves from *AGILE*, *Fermi*, and *Swift* are shown in Fig. 1 and in Ex-
403 tended Data Fig. 1, where the trigger time T_0 (here and elsewhere) refers to the BAT trigger time
404 (20:57:03.19 UT). The prompt emission lasts approximately for 25 s, where the last flaring emis-
405 sion episode ends. Nominally, the T_{90} , i.e. the time interval during which a fraction between
406 5% and 95% of the total emission is observed, is much longer (> 100 s, depending on the instru-
407 ment¹³), but is clearly contaminated by the afterglow component (Fig. 1) and does not provide a
408 good measure of the actual duration of the prompt emission. A more detailed study of the prompt
409 emission phase is reported in¹³.

410 **AGILE** (The Astrorivelatore Gamma ad Immagini LEggero³⁸) could observe GRB 190114C until
411 T_0+330 s, before it became occulted by the Earth. GRB 190114C triggered the Mini-CALorimeter
412 (MCAL) from $T_0-0.95$ s to $T_0+10.95$ s. The MCAL light flux curve in Fig. 1 has been pro-
413 duced using two different spectral models. From $T_0-0.95$ s to $T_0+1.8$ s, the spectrum is fit-
414 ted by a power law with photon index $\Gamma_{\text{ph}} = -1.97^{+0.47}_{-0.70}$ ($dN/dE \propto E^{\Gamma_{\text{ph}}}$). From $T_0+1.8$ s to
415 $T_0 + 5.5$ s the best fit model is a broken power law with $\Gamma_{\text{ph},1} = -1.87^{+0.54}_{-0.19}$, $\Gamma_{\text{ph},2} = -2.63^{+0.07}_{-0.07}$,
416 and break energy $E_b = 756^{+137}_{-159}$ keV. The total fluence in the 0.4–100 MeV energy range is $F =$
417 1.75×10^{-4} erg cm⁻². The Super-AGILE detector also detected the burst, but the large off-axis
418 angle prevented any X-ray imaging of the burst, as well as spectral analysis. Panels **a**, **d**, and **e**



Extended Data Figure 1: **Prompt emission light curves for different detectors.** The different panels show light curves for: **a**, SuperAGILE (20-60 keV); **b**, *Swift*/BAT (15-150 keV); **c**, *Fermi*/GBM (10-1000 keV); **d**, *AGILE*/MCAL (0.4-1.4 MeV); **e**, *AGILE*/MCAL (1.4-100 MeV); **f**, *Fermi*/LAT (0.1-10 GeV). The light curve of *AGILE*/MCAL is split into two bands to show the energy dependence of the first peak. Error bars show the $1-\sigma$ statistical errors.

419 in Extended Data Fig. 1 show the GRB 190114C light curves acquired by the Super-AGILE de-
 420 tector (20 – 60 keV) and by the MCAL detector in the low- (0.4 – 1.4 MeV) and high-energy
 421 (1.4 – 100 MeV) bands.

422 **Fermi/GBM** At the time of the MAGIC observations there are indications that some of the de-
 423 tectors are partially blocked by structure on the *Fermi* Spacecraft that is not modeled in the GBM
 424 detectors’ response. This affects the low-energy part of the spectrum³⁹. For this reason, out of cau-
 425 tion we elected to exclude the energy channels below 50 keV. The spectra detected by the *Fermi*-
 426 Gamma-ray Burst Monitor (GBM)⁴⁰ during the T_0+68 s to T_0+110 s and T_0+110 s to T_0+180 s
 427 intervals are best described by a power law model with photon index $\Gamma_{\text{ph}} = -2.10 \pm 0.08$ and
 428 $\Gamma_{\text{ph}} = -2.05 \pm 0.10$ respectively (Fig. 2 and Fig. 3). The 10-1000 keV light curve in Extended
 429 Data Fig. 1 (panel c) was constructed by summing photon counts for the bright NaI detectors.

430 **Swift/BAT** The 15 – 350 keV mask-weighted light curve of the Burst Alert Telescope (BAT⁴¹)
 431 shows a multi-peaked structure that starts at T_0-7 s (Extended Data Fig. 1, panel b). The 68 – 110 s
 432 and 110 – 180 s spectra shown in Figs. 2 and 3 were derived from joint XRT-BAT fit. The best-
 433 fitting parameters for the whole interval (68 – 180 s) are: column density $N_{\text{H}} = (7.53^{+0.74}_{-1.74}) \times$
 434 10^{22} cm^{-2} at $z = 0.42$, in addition to the galactic value of $7.5 \times 10^{19} \text{ cm}^{-2}$, low-energy photon index
 435 $\Gamma_{\text{ph},1} = -1.21^{+0.40}_{-1.26}$, high-energy spectral index $\Gamma_{\text{ph},2} = -2.19^{+0.39}_{-0.19}$, peak energy $E_{\text{pk}} > 14.5$ keV.
 436 Errors are given at 90% confidence level.

437 **Fermi/LAT** The *Fermi* Large Area Telescope (LAT)⁴² detected a gamma-ray counterpart since
 438 the prompt phase⁴³. The burst left the LAT field of view (FoV) at T_0+150 s and remained outside

439 the LAT FoV until T_0+8600 s. The count light curve in the energy range 0.1-10 GeV is shown
 440 in Extended Data Fig. 1 (panel f). The LAT spectra in the time bins 68–110 s and 110–180 s
 441 (Figs. 2 and 3) are described by a power law with pivot energies of, respectively, 200 MeV and
 442 500 MeV, photon indices $\Gamma_{\text{ph}}(68 - 110) = -2.02 \pm 0.95$ and $\Gamma_{\text{ph}}(110 - 180) = -1.69 \pm 0.42$,
 443 and corresponding normalisations of $N_{0,68-110} = (2.02 \pm 1.31) \times 10^{-7}$ ph MeV $^{-1}$ cm $^{-2}$ s and
 444 $N_{0,110-180} = (4.48 \pm 2.10) \times 10^{-8}$ ph MeV $^{-1}$ cm $^{-2}$ s. In each time-interval, the analysis has been
 445 performed limited to the energy range where photons have been detected. The LAT light curve
 446 integrated in the energy range 0.1-1 GeV is shown in Fig. 1.

447 **MAGIC** We used the Major Atmospheric Gamma Imaging Cherenkov (MAGIC) standard soft-
 448 ware⁴⁴ and followed the steps optimised for the data taking under moderate moon illumination⁴⁵
 449 to analyse the data. The spectral fitting is performed by a forward-folding method assuming a sim-
 450 ple power law for the intrinsic spectrum and taking into account the **extragalactic background**
 451 **light** (EBL) effect using the model of Domínguez et al.⁴⁶. Extended Data Table 1 shows the fitting
 452 results for various time bins (the pivot energy is chosen to minimise the correlation between nor-
 453 malisation and photon index parameters). The data points shown in both Fig. 2 and 3 are obtained
 454 from the observed excess rates in estimated energy whose fluxes are evaluated in true energy using
 455 effective time and a spill-over corrected effective area obtained as a resultant of the best fit.

456 The time resolved analysis hints to a possible spectral evolution to softer values. Although
 457 we can not exclude that the photon indices are compatible with a constant value of ~ -2.5 up to
 458 2400 s. The signal and background in the considered time bins are both in the low-count Poisson
 459 regime. Therefore, the correct treatment of the MAGIC data provided here includes along with the

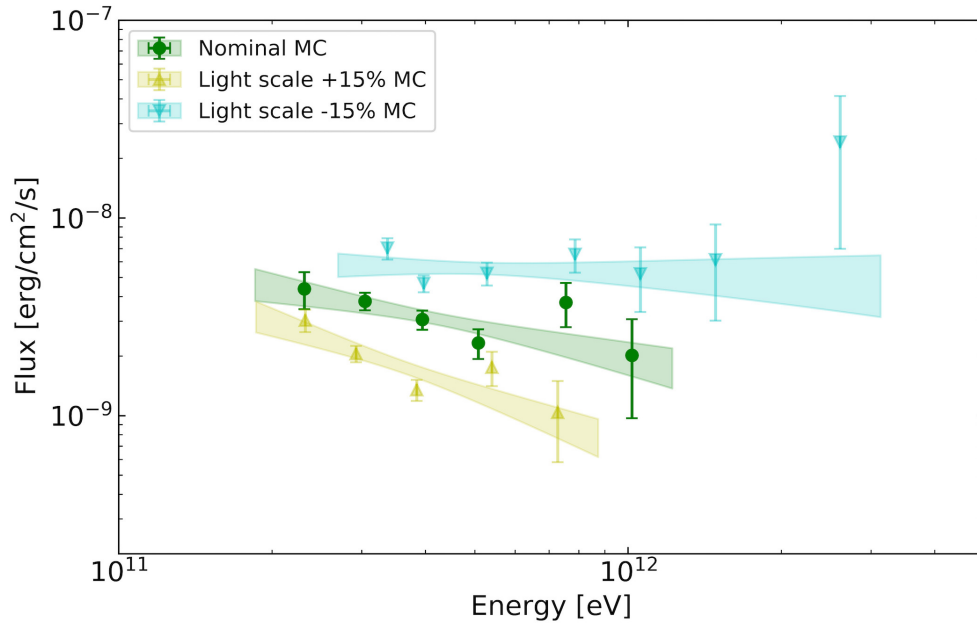
Time bin [seconds after T_0]	Normalisation [$\text{TeV}^{-1} \text{cm}^{-2} \text{s}^{-1}$]	Photon index	Pivot energy [GeV]
62 - 90	$1.95^{+0.21}_{-0.20} \cdot 10^{-7}$	$-2.17^{+0.34}_{-0.36}$	395.5
68 - 180	$1.10^{+0.09}_{-0.08} \cdot 10^{-7}$	$-2.27^{+0.24}_{-0.25}$	404.7
180 - 625	$2.26^{+0.21}_{-0.20} \cdot 10^{-8}$	$-2.56^{+0.27}_{-0.29}$	395.5
68 - 110	$1.74^{+0.16}_{-0.15} \cdot 10^{-7}$	$-2.16^{+0.29}_{-0.31}$	386.5
110 - 180	$8.59^{+0.95}_{-0.91} \cdot 10^{-8}$	$-2.51^{+0.37}_{-0.41}$	395.5
180 - 360	$3.50^{+0.38}_{-0.36} \cdot 10^{-8}$	$-2.36^{+0.34}_{-0.37}$	395.5
360 - 625	$1.65^{+0.23}_{-0.23} \cdot 10^{-8}$	$-3.16^{+0.48}_{-0.54}$	369.1
625 - 2400	$3.52^{+0.47}_{-0.47} \cdot 10^{-9}$	$-2.80^{+0.48}_{-0.54}$	369.1
62 - 2400 (Nominal MC)	$1.07^{+0.08}_{-0.07} \cdot 10^{-8}$	$-2.51^{+0.20}_{-0.21}$	423.8
62 - 2400 (Light scale +15% MC)	$7.95^{+0.58}_{-0.56} \cdot 10^{-9}$	$-2.91^{+0.23}_{-0.25}$	369.1
62 - 2400 (Light scale -15% MC)	$1.34^{+0.09}_{-0.09} \cdot 10^{-8}$	$-2.07^{+0.18}_{-0.19}$	509.5

Extended Data Table 1: **MAGIC spectral fit parameters for GRB 190114C**. For each time bin, columns represent a) start time and end time of the bin; b) normalisation of the EBL-corrected differential flux at the pivot energy with statistical errors; c) photon indices with statistical errors; d) pivot energy of the fit (fixed).

460 use of the Poisson statistic also the systematic errors. To estimate the main source of systematic
461 error caused by our imperfect knowledge of the absolute instrument calibration and the total at-
462 mospheric transmission we vary the light-scale in our Monte Carlo (MC) simulation as suggested
463 in previous studies⁴⁴. The result is reported in the last two lines of Extended Data Table 1 and in
464 Extended Data Fig. 2.

465 The systematic effects deriving from the choice of one particular EBL model were also stud-
466 ied. The analysis performed to obtain the time integrated spectrum was repeated employing other
467 three models⁴⁷⁻⁴⁹. The contribution to the systematic error on the photon index caused by the un-
468 certainty on the EBL model is $\sigma_\alpha = {}^{+0.10}_{-0.13}$ which is smaller than the statistical error only (1 standard
469 deviation) as already seen in a previous work¹⁰. On the other hand, the contribution to the system-
470 atic error on the normalisation, due to choice of the EBL model, is only partially at the same level
471 of the statistical error (1 standard deviation) $\sigma_N = {}^{+0.30}_{-0.08} \times 10^{-8}$. The chosen EBL model returns a
472 lower normalisation with respect to two of the other models and very close to the third one⁴⁷.

473 The MAGIC energy flux light curve that is presented in Fig. 1 was obtained by integrating
474 the best fit spectral model of each time bin from 0.3 to 1 TeV, in the same manner as a previous
475 publication¹⁰. The value of the fitted time constant reported here differs less than two standard
476 deviation from the one previously reported¹⁰. The difference is due to the poor constraints on the
477 spectral fit parameters of the last time bin, which influences the light curve fit.



Extended Data Figure 2: **MAGIC time integrated spectral energy distributions in the time interval 62-2400 s after T_0 .** The green (yellow, blue) points and band show the result with the nominal (+15%, -15%) light scale MC, defining the limits of the systematic uncertainties. **The contour regions are drawn from the $1-\sigma$ error of their best-fit power law functions. The vertical bars of the data points show the $1-\sigma$ errors on the flux.**

478 **X-ray afterglow observations**

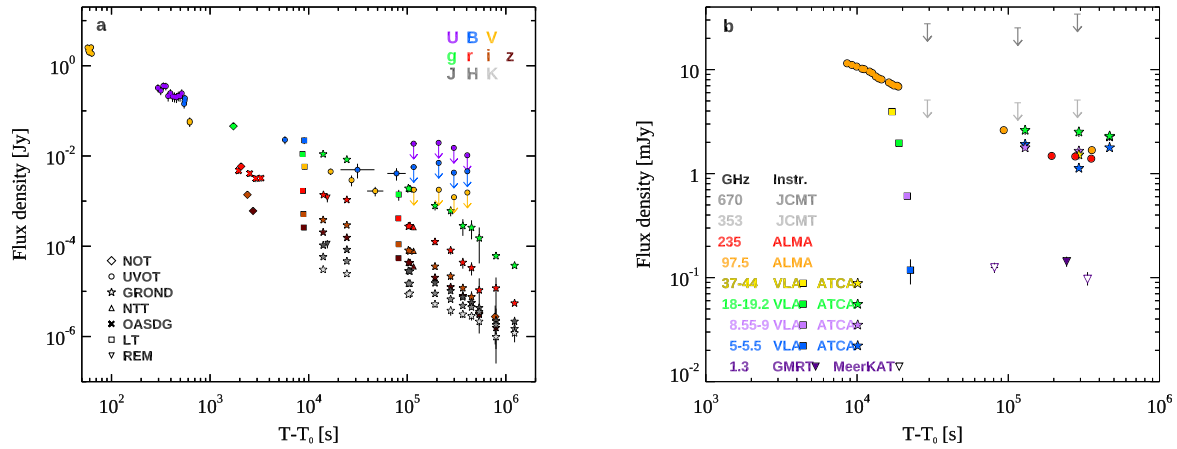
479 **Swift/XRT** The *Swift* X-Ray Telescope (XRT) started observing 68 s after T_0 . The source light
480 curve⁵⁰ was taken from the *Swift*/XRT light curve repository⁵¹ and converted into 1-10 keV flux
481 (Fig. 1) through dedicated spectral fits. The combined spectral fit XRT+BAT in Figs. 2 and 3 has
482 been described above.

483 **XMM-Newton and NuSTAR** The *XMM-Newton* X-ray Observatory and the Nuclear Spectro-
484 scopic Telescope Array (*NuSTAR*) started observing GRB 190114C under DDT ToOs 7.5 hours
485 and 22.5 hrs (respectively) after the burst. The *XMM-Newton* and NuSTAR absorption-corrected
486 fluxes (see Fig. 1) were derived by fitting the spectrum with XSPEC adopting the same power law
487 model, with absorption in our Galaxy and at the redshift of the burst.

488 **NIR, Optical and UV afterglow observations** Light curves from the different instruments pre-
489 sented in this section are shown in Extended Data Fig. 3.

490 **GROND** The Gamma-ray Burst Optical/Near-infrared Detector (GROND⁵²) started observations
491 3.8 hours after the GRB trigger, and the follow-up continued until January 29, 2019. Image re-
492 duction and photometry were carried out with standard IRAF tasks⁵³, as described in^{54,55}. JHK_s
493 photometry was converted to AB magnitudes to have a common flux system. Final photometry is
494 given in Extended Data Table 2.

495 **GTC** The BOOTES-2 ultra-wide field camera⁵⁶, took an image at the GRB 190114C location,
496 starting at 20:57:18 UT (30 s exposure time) (see Extended Data Fig. 4). The Gran Canarias



Extended Data Figure 3: **Afterglow light curves of GRB 190114C.** Flux density at different frequencies, as a function of the time since the initial burst T_0 . Panel **a**: observations in the NIR/Optical/UV bands. The flux has been corrected for extinction in the host and in our Galaxy. The contribution of the host galaxy and its companion has been subtracted. Fluxes have been rescaled (except for the r filter). Panel **b**: Radio and sub-mm observations from 1.3 GHz to 670 GHz.

T_{GROND}	AB magnitude						
	(s)	g'	r'	i'	z'	J	H
14029.94 ± 335.28	19.21 ± 0.03	18.46 ± 0.03	17.78 ± 0.03	17.33 ± 0.03	16.78 ± 0.05	16.30 ± 0.05	16.03 ± 0.07
24402.00 ± 345.66	19.50 ± 0.04	18.72 ± 0.03	18.05 ± 0.03	17.61 ± 0.03	17.02 ± 0.05	16.53 ± 0.05	16.26 ± 0.08
102697.17 ± 524.01	20.83 ± 0.06	20.00 ± 0.04	19.30 ± 0.04	18.87 ± 0.03	18.15 ± 0.05	17.75 ± 0.06	17.40 ± 0.09
106405.63 ± 519.87	20.86 ± 0.05	19.98 ± 0.03	19.34 ± 0.03	18.88 ± 0.03	18.17 ± 0.06	17.75 ± 0.06	17.34 ± 0.09
191466.77 ± 751.37	21.43 ± 0.07	20.61 ± 0.03	19.97 ± 0.03	19.52 ± 0.03	18.77 ± 0.06	18.28 ± 0.06	17.92 ± 0.14
275594.19 ± 747.59	21.57 ± 0.07	20.88 ± 0.04	20.31 ± 0.04	19.87 ± 0.04	19.14 ± 0.07	18.57 ± 0.06	18.26 ± 0.21
366390.74 ± 1105.79	21.87 ± 0.07	21.17 ± 0.04	20.62 ± 0.03	20.15 ± 0.03	19.43 ± 0.06	18.89 ± 0.06	18.46 ± 0.15
448791.55 ± 1201.33	21.90 ± 0.08	21.27 ± 0.04	20.79 ± 0.04	20.33 ± 0.03	19.66 ± 0.07	18.97 ± 0.07	18.55 ± 0.18
537481.41 ± 1132.16	22.02 ± 0.09	21.52 ± 0.05	21.00 ± 0.04	20.55 ± 0.03	19.87 ± 0.07	19.20 ± 0.07	18.83 ± 0.17
794992.63 ± 1200.69	22.14 ± 0.04	21.51 ± 0.03	21.05 ± 0.04	20.71 ± 0.05	20.31 ± 0.13	19.79 ± 0.14	19.59 ± 0.41
1226716.84 ± 1050.15	22.17 ± 0.04	21.59 ± 0.04	21.26 ± 0.04	20.97 ± 0.04	20.34 ± 0.12	19.95 ± 0.11	19.40 ± 0.34

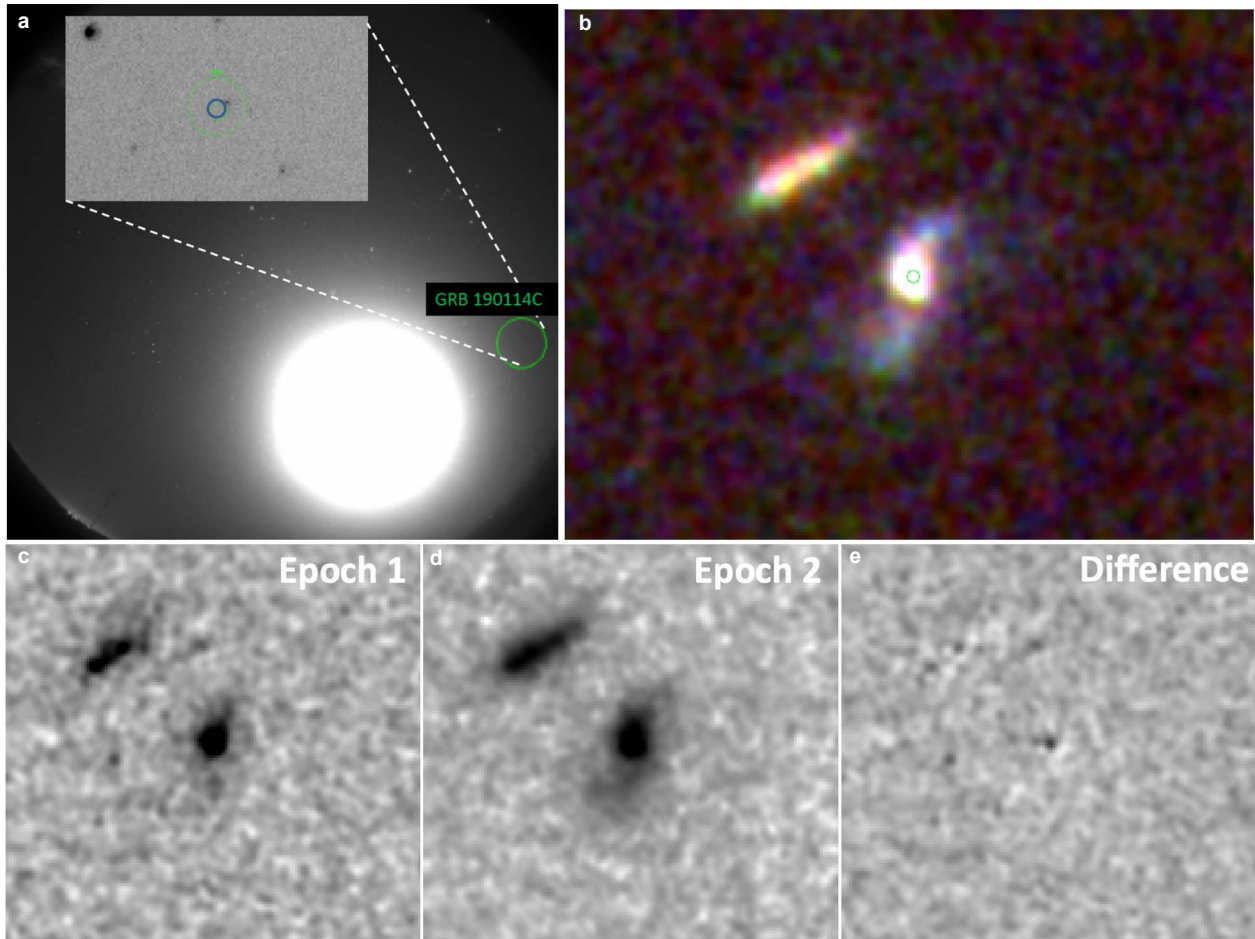
Extended Data Table 2: **GROND photometry.** T_{GROND} in seconds after the BAT trigger. The AB magnitudes are not corrected for the Galactic foreground reddening.

497 Telescope (GTC) equipped with the OSIRIS spectrograph⁵⁷ started observations 2.6 hr post-burst.
498 The gratings R1000B and R2500I were used covering the wavelength range 3,700-10,000 Å (600 s
499 exposure times for each grism). The GTC detects a highly extinguished continuum, as well as CaII
500 H and K lines in absorption, and [OII], H β , and [OIII] in emission (see Extended Data Fig. 5), all
501 roughly at the same redshift $z = 0.4245 \pm 0.0005$ ⁵⁸. Comparing the derived rest-frame equivalent
502 widths (EWs) with the work by⁵⁹, GRB 190114C clearly shows higher than average, but not
503 unprecedented, values.

504 **HST** The *Hubble Space Telescope* (HST) imaged the afterglow and host galaxy of GRB 190114C
505 on 11 February and 12 March 2019. HST observations clearly reveal that the host galaxy is spiral
506 (Extended Data Fig. 4). A direct subtraction of the epochs of F850LP observations yields a faint
507 residual close to the nucleus of the host (Extended Data Fig. 4). From the position of the residual
508 we estimate that the burst originated within 250 pc of the host galaxy nucleus.

509 **LT** The robotic 2-m Liverpool Telescope (LT⁶⁰) slewed to the afterglow location at UTC 2019-
510 01-14.974 and on the second night, from UTC 2019-01-15.814 and acquired images in B , g , V ,
511 r , i and z bands (45 s exposure each in the first night and 60 s in the second, see Extended Data
512 Table 3). Aperture photometry of the afterglow was performed using a custom IDL script with
513 a fixed aperture radius of 1.5". Photometric calibration was performed relative to stars from the
514 Pan-STARRS1 catalogue⁶¹.

515 **NTT** The ESO New Technology Telescope (NTT) observed the optical counterpart of GRB 190114C
516 under the extended Public ESO Spectroscopic Survey for Transient Objects (ePESSTO) using the



Extended Data Figure 4: **Images of the localisation region of GRB 190114C.** Panel **a**: The CASANDRA-2 at the BOOTES-2 station all-sky image. The image (30 s exposure, unfiltered) was taken at $T_0+14.8$ s. At the GRB 190114C location (circle) no prompt optical emission is detected. Panel **b**: Three-colour image of the host of GRB 190114C with the HST. The host galaxy is a spiral galaxy, and the green circle indicates the location of the transient close to its host nucleus. The image is $8''$ across, north is up and east to the left. Panels **c**, **d** and **e**: F850LP imaging of GRB 190114C taken with the HST. Two epochs are shown (images are $4''$ across), as well as the result of the difference image. A faint transient is visible close to the nucleus of the galaxy, and we identify this as the late time afterglow of the burst.

UTC	Filter	Exposure (s)	Magnitude
LT/O:O			
2019-01-14.975	<i>g</i>	45	19.08±0.06
2019-01-14.976	<i>r</i>	45	18.22±0.02
2019-01-14.977	<i>i</i>	45	17.49±0.02
2019-01-14.978	<i>z</i>	45	17.12±0.02
2019-01-14.979	<i>B</i>	45	19.55±0.15
2019-01-14.980	<i>V</i>	45	18.81±0.08
2019-01-15.814	<i>r</i>	60	19.61±0.05
2019-01-15.818	<i>z</i>	60	18.70±0.06
2019-01-15.820	<i>i</i>	60	19.04±0.04
2019-01-15.823	<i>g</i>	60	20.96±0.17
NOT/AIFOSC			
2019-01-14.89127	<i>g</i>	1 × 300	17.72±0.03
2019-01-14.89512	<i>r</i>	1 × 300	16.93±0.02
2019-01-14.89899	<i>i</i>	1 × 300	16.42 ±0.04
2019-01-14.90286	<i>z</i>	1 × 300	16.17 ±0.04
2019-01-23.8896	<i>i</i>	6 × 300	21.02±0.05

UVOT							
T_{start}	T_{stop}	Filter	Magnitude	T_{start}	T_{stop}	Filter	Magnitude
56.63	57.63	<i>V</i>	12.17±0.14	130958	142524	<i>UVM2</i>	20.37
57.63	58.63	<i>V</i>	12.34±0.14	217406	222752	<i>UVM2</i>	20.48
58.63	59.63	<i>V</i>	12.44±0.13	107573	125233	<i>U</i>	20.29
59.63	60.63	<i>V</i>	12.29±0.14	205500	210750	<i>U</i>	20.25
60.63	61.63	<i>V</i>	12.44±0.14	291188	302718	<i>U</i>	20.49
61.63	62.63	<i>V</i>	12.16±0.13	400429	412385	<i>U</i>	20.82
62.63	63.63	<i>V</i>	12.51±0.13	616	627	<i>V</i>	16.25±0.20
615.95	625.95	<i>V</i>	16.32±0.20	16295	17136	<i>V</i>	19.03±0.14
73.34	83.34	<i>white</i>	13.86	26775	27682	<i>V</i>	19.50±0.27
83.34	93.34	<i>white</i>	14.10±0.06	39149	57221	<i>V</i>	20.09±0.23
93.34	103.34	<i>white</i>	14.19±0.06	108064	125736	<i>V</i>	20.02
103.34	113.34	<i>white</i>	14.36±0.06	206689	211356	<i>V</i>	20.02
113.34	123.34	<i>white</i>	14.64±0.06	292383	303996	<i>V</i>	20.42
123.34	133.34	<i>white</i>	14.65±0.06	401305	413316	<i>V</i>	20.17
133.34	143.34	<i>white</i>	14.91±0.06	4044	51522	<i>UVW1</i>	21.17
143.34	153.34	<i>white</i>	14.99±0.06	131216	142656	<i>UVW1</i>	20.47
153.34	163.34	<i>white</i>	15.05±0.06	217984	223056	<i>UVW1</i>	20.57
163.34	173.34	<i>white</i>	15.32±0.06	592	612	<i>UVW2</i>	17.65
173.34	183.34	<i>white</i>	15.38±0.06	6056	56384	<i>UVW2</i>	21.30
183.34	193.34	<i>white</i>	15.38±0.06	130699	142346	<i>UVW2</i>	20.52
193.34	203.34	<i>white</i>	15.59±0.06	216828	222404	<i>UVW2</i>	20.55
562.0	572.0	<i>white</i>	16.96±0.10	566	586	<i>white</i>	16.90±0.07
572.0	582.0	<i>white</i>	16.90±0.10	607389	613956	<i>white</i>	22.16
535.5	555.5	<i>B</i>	17.56±0.21	624452	682416	<i>white</i>	21.99±0.18
545.5	565.5	<i>B</i>	17.25±0.18	745033	769296	<i>white</i>	21.64±0.16
285.9	305.9	<i>U</i>	17.35±0.19	818840	837216	<i>white</i>	22.50
305.9	325.9	<i>U</i>	17.50±0.20	893522	907116	<i>white</i>	22.57
325.9	345.9	<i>U</i>	17.24±0.18	991065	1004196	<i>white</i>	22.49±0.35
345.9	365.9	<i>U</i>	17.26±0.18	1077542	1094616	<i>white</i>	22.41±0.31
365.9	385.9	<i>U</i>	17.80±0.24	1140343	1170336	<i>white</i>	22.50
385.9	405.9	<i>U</i>	17.64±0.22	1220661	1274376	<i>white</i>	22.36±0.29
405.9	425.9	<i>U</i>	17.82±0.24	5851	6050	<i>white</i>	19.25±0.09
425.9	445.9	<i>U</i>	17.84±0.25	21950	22857	<i>white</i>	20.25±0.09
445.9	465.9	<i>U</i>	17.87±0.25	1353459	1359284	<i>white</i>	21.70
465.9	485.9	<i>U</i>	17.79±0.24	1502211	1548336	<i>white</i>	21.98±0.24
485.9	505.9	<i>U</i>	17.81±0.24	1692292	1703935	<i>white</i>	22.07
505.9	525.9	<i>U</i>	17.65±0.22	2132978	2146056	<i>white</i>	22.58
542	561	<i>B</i>	17.38±0.14	2299521	2317956	<i>white</i>	22.41±0.31
5646	5845	<i>B</i>	19.54±0.19	63686	80942	<i>white</i>	21.07±0.24
21038	46521	<i>B</i>	21.14±0.35	107900	125591	<i>white</i>	21.40±0.28
62774	96486	<i>B</i>	21.33±0.29	206292	211137	<i>white</i>	21.52
107737	125412	<i>B</i>	21.00	291984	303556	<i>white</i>	21.48±0.23
205896	210944	<i>B</i>	20.78	401012	413029	<i>white</i>	21.84
291586	303137	<i>B</i>	21.29	491973	505356	<i>white</i>	22.21±0.24
400721	412707	<i>B</i>	21.22	74	224	<i>white</i>	14.90±0.02
3839	50615	<i>UVM2</i>	20.88±0.28				

Extended Data Table 3: Liverpool Telescope, Nordic Optical Telescope, and UVOT observations. Magnitudes are SDSS AB-"like" for ugriz, Vega-"like" for all the other filters and are not corrected for Galactic extinction. For UVOT data, magnitudes without uncertainties are upper limits.

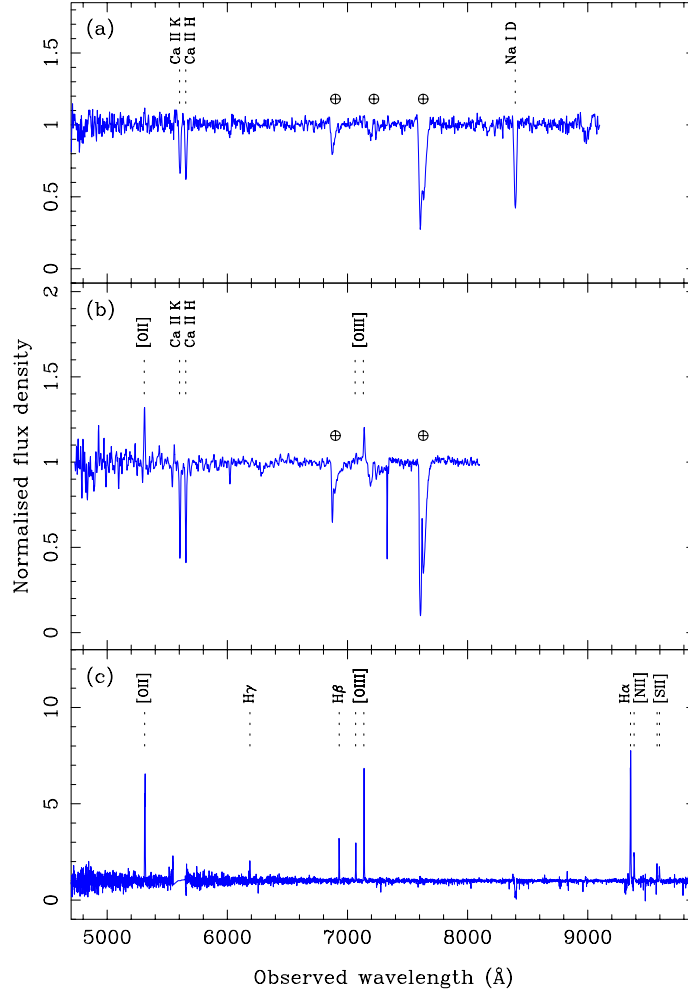
517 NTT/EFOSC2 instrument in imaging mode⁶². Observations started at 04:36:53 UT on 2019 Jan-
518 uary 16 with the g , r , i , z Gunn filters. Image reduction was carried out by following the standard
519 procedures⁶³.

520 **OASDG** The 0.5 m remote telescope of the Osservatorio Astronomico “S. Di Giacomo” (OASDG),
521 located in Agerola (Italy) started observations in the optical Rc -band 0.54 hours after the burst. The
522 afterglow of GRB 190114C was clearly detected in all the images.

523 **NOT** The Nordic Optical Telescope (NOT) observed the optical afterglow of GRB 190114C with
524 the Alhambra Faint Object Spectrograph and Camera (AlFOSC) instrument. Imaging was obtained
525 in the $griz$ filters with 300 s exposures, starting at Jan 14 21:20:56 UT, 24 minutes after the BAT
526 trigger. The normalised spectrum (Extended Data Fig. 5) reveals strong host interstellar absorption
527 lines due to Ca H & K and Na ID, which provided a redshift of $z = 0.425$.

528 **REM** The Rapid Eye Mount telescope (REM) performed optical and NIR observations with
529 the REM 60 cm robotic telescope equipped with the ROS2 optical imager and the REMIR NIR
530 camera⁶⁴. Observations were performed starting about 3.8 hours after the burst in the r , and J
531 bands and lasted about one hour.

532 **Swift/UVOT** The *Swift* UltraViolet and Optical Telescope (UVOT⁶⁵) began observations at T_0+54
533 seconds in the UVOT v band. The first observation after settling started 74 s after the trigger for
534 150 s in the UVOT *white* band⁶⁶. A 50 s exposure with the UV grism was taken thereafter, followed
535 by multiple exposures rotating through all seven broad and intermediate-band filters until switching
536 to only UVOT’s clear white filter on 2019-01-20. Standard photometric calibration and methods



Extended Data Figure 5: **Optical/NIR spectra of GRB 190114C.** Panel **a**: The NOT/AIFOSC spectrum obtained at a mid-time 1 hr post-burst. The continuum is afterglow dominated at this time, and shows strong absorption features of Ca II and Na I (in addition to telluric absorption). Panel **b**: the normalised GTC (+OSIRIS) spectrum on Jan 14, 23:32:03 UT, with the R1000B and R2500I grisms. The emission lines of the underlying host galaxy are noticeable, besides the Ca II absorption lines in the afterglow spectrum. Panel **c**: The visible light region of the VLT/X-shooter spectrum obtained approximately 3.2 d post-burst, showing strong emission lines from the star-forming host galaxy.

537 were used for deriving the aperture photometry^{67,68}. The grism zeroth order the data were reduced
538 manually⁶⁹ to derive the b -magnitude and error.

539 **VLT** The STARGATE collaboration used the Very Large Telescope (VLT) and observed GRB 190114C
540 using the X-shooter spectrograph. Detailed analysis will be presented in forthcoming papers. A
541 portion of the second spectrum is shown in Extended Data Fig. 5, illustrating the strong emission
542 lines characteristic of a strongly star-forming galaxy, whose light is largely dominating over the
543 afterglow at this epoch.

544 **Magnitudes of the underlying galaxies** The *HST* images show a spiral or tidally disrupted galaxy
545 whose bulge is coincident with the coordinates of GRB 190114C. A second galaxy is detected at an
546 angular distance of $1.3''$, towards the North East. The SED analysis was performed with LePhare
547 ^{70,71} using an iterative method that combined both the resolved photometry of the two galaxies
548 found in the *HST* and *VLT/HAWK-I* data and the blended photometry from *GALEX* and *WISE*,
549 where the spatial resolution was much lower. Further details will be given in a paper in preparation
550 (de Ugate Postigo et al.). The estimated photometry, for each object and their combination, is given
551 in Extended Data Table 4.

552 **Optical Extinction** The optical extinction toward the line of sight of a GRB is derived assuming
553 a power law as intrinsic spectral shape⁷². Once the Galactic extinction ($E_{B-V} = 0.01^{73}$) is taken
554 into account and the fairly bright host galaxy contribution is properly subtracted, a good fit to the
555 data is obtained with the LMC recipe and $A_V = 1.83 \pm 0.15$. The spectral index β ($F_\nu \propto \nu^{\beta_0}$)
556 evolves from hard to soft across the temporal break in the optical light-curve at about 0.5 days,

Filter	Host	Companion	Combined
Sloan u	23.54	25.74	23.40
Sloan g	22.51	23.81	22.21
Sloan r	22.13	22.81	21.66
Sloan i	21.70	22.27	21.19
Sloan z	21.51	21.74	20.87
2MASS J	20.98	21.08	20.28
2MASS H	20.68	20.82	20.00
2MASS K_s	20.45	20.61	19.77

Extended Data Table 4: **Observations of the host galaxy.** For each filter, the estimated magnitudes are given for the host galaxy of GRB 190114C, the companion and the combination of the two objects.

557 moving from $\beta_{o,1} = -0.10 \pm 0.12$ to $\beta_{o,2} = -0.48 \pm 0.15$.

558 **Radio and Sub-mm afterglow observations** The light curves from the different instruments are
559 shown in Extended Data Fig. 3.

560 **ALMA** The Atacama Large Millimetre/Submillimetre Array (ALMA) observations are reported
561 in Band 3 (central observed frequency of 97.500 GHz) and Band 6 (235.0487 GHz), between 2019
562 January 15 and 2019 January 19. Data were calibrated within CASA (Common Astronomy Soft-
563 ware Applications, version 5.4.0⁷⁴) using the pipeline calibration. Photometric measurements were
564 also performed within CASA. ALMA early observations at 97.5 GHz are taken from ¹⁵.

565 **ATCA** The Australia Telescope Compact Array (ATCA) observations were made with the ATCA
566 4 cm receivers (band centres 5.5 and 9 GHz), 15 mm receivers (band centres 17 and 19 GHz), and
567 7 mm receivers (band centres 43 and 45 GHz). ATCA data were obtained using the CABB con-
568 tinuum mode ⁷⁵ and reduced with the software packages MIRIAD ⁷⁶ and CASA ⁷⁴ using standard
569 techniques. The quoted errors are 1σ , which include the RMS and Gaussian 1σ errors.

570 **GMRT** The upgraded Giant Metre-wave Radio Telescope ⁷⁷ (UGMRT) observed on 17th January
571 2019 13.44 UT (2.8 days after the burst) in band 5 (1000-1450 MHz) with 2048 channels spread
572 over 400 MHz. GMRT detected a weak source with a flux density of $73 \pm 17 \mu\text{Jy}$ at the GRB
573 position ⁷⁸. The flux should be considered as an upper limit, as the contribution from the host⁷⁹ has
574 not been subtracted.

ATCA			
Start Date and Time	End Date and Time	Frequency GHz	Flux mJy
1/16/2019 6:47:00	1/16/2019 10:53:00	5.5	1.92±0.06
		9	1.78±0.06
		18	2.62±0.26
1/18/2019 1:45:00	1/18/2019 11:18:00	5.5	1.13±0.04
		9	1.65±0.05
		18	2.52±0.27
		44	1.52±0.15
1/20/2019 3:38	1/20/2019 10:25:00	5.5	1.78±0.06
		9	2.26±0.07
		18	2.30±0.23

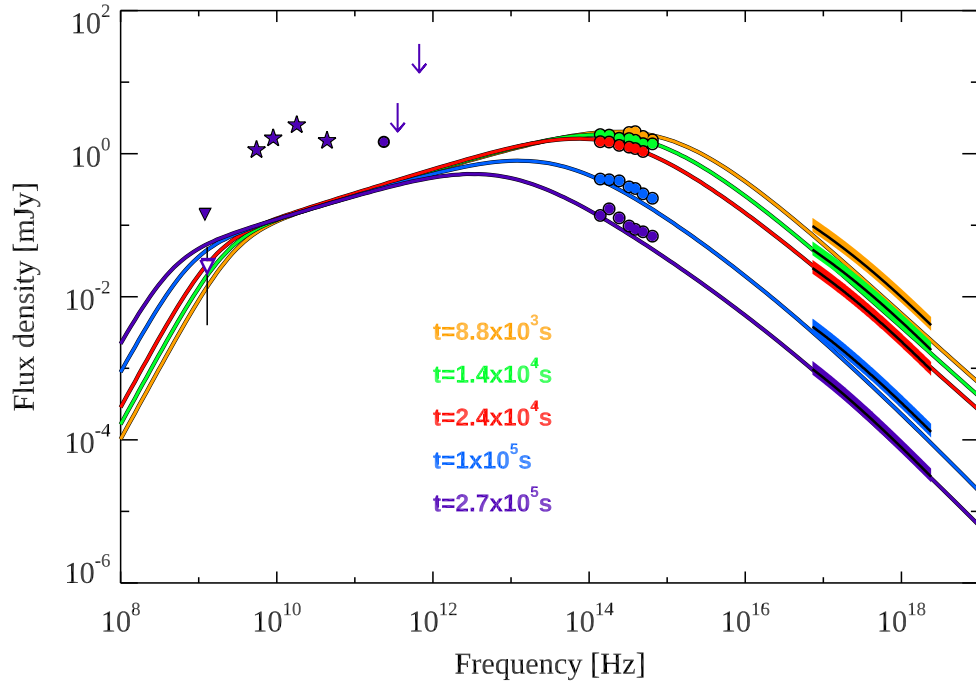
JCMT SCUBA-2						
UT Date	Time since trigger (days)	Time on source (hours)	Typical 225 GHz CSO Opacity	Typical elevation (degrees)	850 μ m RMS density (mJy/beam)	450 μ m RMS density (mJy/beam)
2019-01-15	0.338	1.03	0.026	39	1.7	9.2
2019-01-16	1.338	1.03	0.024	39	1.6	8.4
2019-01-18	3.318	0.95	0.031	37	1.7	11.4

Extended Data Table 5: **Observations of GRB 190114C by ATCA and JCMT SCUBA-2.** For ATCA data, start and end date and times (UTC) of the observations, frequency, and flux (1σ error) are reported. For JCMT SCUBA-2 data, the CSO 225 GHz opacity measures the zenith atmospheric attenuation.

575 **MeerKAT** The new MeerKAT radio observatory^{80,81} observed on 15 and 18 January 2019, with
576 DDT requested by the ThunderKAT Large Survey Project⁸². Both epochs used 63 antennas and
577 were done at L-band spanning 856 MHz and centered at 1284 MHz. MeerKAT flux estimation
578 was done by finding and fitting the source with the software PyBDSF v.1.8.15⁸³. Adding the RMS
579 noise in quadrature to the flux uncertainty leads to final flux measurements of $125 \pm 14 \mu\text{Jy}/\text{beam}$
580 on 15 January and $97 \pm 16 \mu\text{Jy}/\text{beam}$ on 18 January. The contribution from the host galaxy⁷⁹ has
581 not been subtracted. Therefore, these measurements provide a maximum flux of the GRB.

582 **JCMT SCUBA-2 Sub-millimeter** Sub-millimeter observations were performed simultaneously
583 at $850 \mu\text{m}$ and $450 \mu\text{m}$ on three nights using the SCUBA-2 continuum camera⁸⁴. GRB 190114C
584 was not detected on any of the individual nights. Combining all the SCUBA-2 continuum camera⁸⁴
585 observations, the RMS background noise is $0.95 \text{ mJy}/\text{beam}$ at $850 \mu\text{m}$ and $5.4 \text{ mJy}/\text{beam}$ at $450 \mu\text{m}$
586 at 1.67 days after the burst trigger.

587 **Prompt emission model for the early time MAGIC emission** In the standard picture the prompt
588 sub-MeV spectrum is explained as a synchrotron radiation from relativistic accelerated electrons
589 in the energy dissipation region. The associated inverse Compton component is sensitive to the
590 details of the dynamics: e.g. in the internal shock model if the peak energy is initially very high
591 and the IC component is suppressed due to Klein-Nishina (KN) effects, the peak of the IC com-
592 ponent may be delayed and become bright only at late times when scatterings occur in Thomson
593 regime. Simulations showed that magnetic fields required to produce the GeV/TeV component are
594 rather low⁸⁵, $\epsilon_B \sim 10^{-3}$. In this framework the contribution of the IC component to the observed
595 flux at early times (62-90 s, see Extended Data Table 1) does not exceed $\sim 20\%$. Alternatively, if



Extended Data Figure 6: **Radio to X-rays SED at different epochs.** The synchrotron frequency ν_m crosses the optical band, moving from higher to lower frequencies. The break between 10^8 and 10^{10} Hz is caused by the self-absorption synchrotron frequency ν_{sa} . Optical (X-ray) data have been corrected for extinction (absorption).

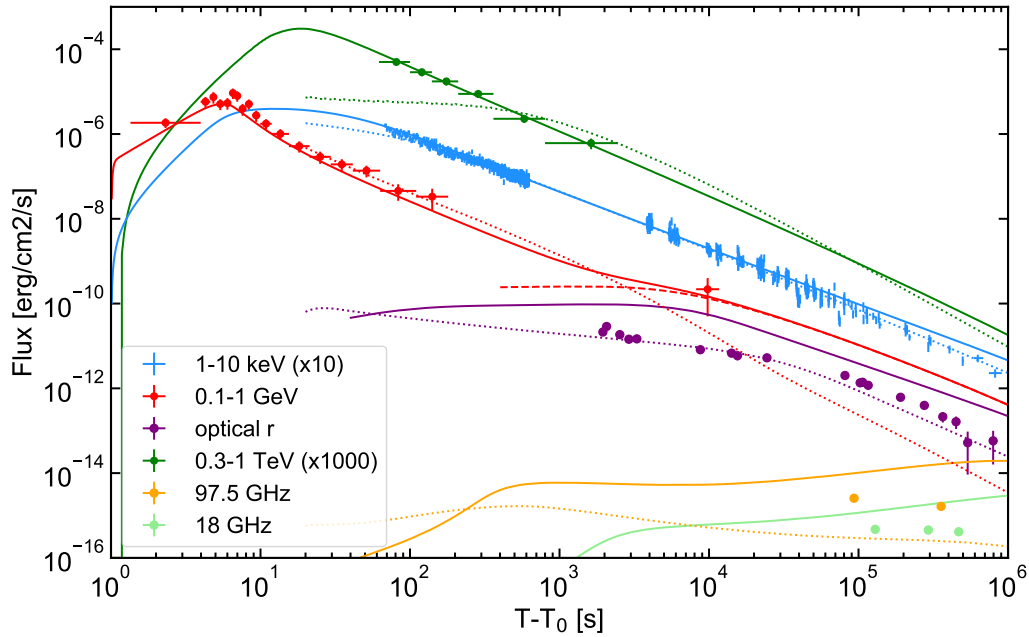
596 the prompt emission originates in reprocessed photospheric emission, the early TeV flux may arise
 597 from IC scatterings of thermal photons by freshly heated electrons below the photosphere at low
 598 optical depths. Another possibility for the generation of TeV photons might be the inverse Comp-
 599 ton scattering of prompt MeV photons by electrons in the external forward shock region where
 600 electrons are heated to an average Lorentz factor of order 10^4 at early times.

601 **Afterglow model** Synchrotron and SSC radiation from electrons accelerated at the forward shock
 602 has been modelled within the external shock scenario ^{7,8,19,24,86}. The results of the modeling are
 603 overlaid to the data in Fig. 3, and Extended Data Figs. 6 and 7.

604 We consider two types of power law radial profiles $n(R) = n_0 R^{-s}$ for the external environ-
 605 ment: $s = 0$ (homogeneous medium) and $s = 2$ (wind-like medium, typical of an environment
 606 shaped by the stellar wind of the progenitor). In the last case, we define $n_0 = 3 \times 10^{35} A_\star \text{ cm}^{-1}$.
 607 We assume that electrons swept up by the shock are accelerated into a power law distribution de-
 608 scribed by spectral index p : $dN/d\gamma \propto \gamma^{-p}$, where γ is the electron Lorentz factor. We call ν_m the
 609 characteristic synchrotron frequency of electrons with Lorentz factor γ_m , ν_c the cooling frequency,
 610 and ν_{sa} the synchrotron self-absorption frequency.

611 The early time optical emission (up to ~ 1000 s) and radio emission (up to $\sim 10^5$ s) are
 612 most likely dominated by reverse shock radiation ¹⁵. Detailed modeling of this component is not
 613 discussed in this work, where we focus on forward shock radiation.

614 The XRT flux (Fig. 1, blue data points) decays as $F_X \propto t^{\alpha_X}$ with $\alpha_X = -1.36 \pm 0.02$. If
 615 $\nu_X > \max(\nu_m, \nu_c)$, the X-ray light curve is predicted to decay as $t^{(2-3p)/4}$, that implies $p \sim 2.5$.



Extended Data Figure 7: **Modeling of the broadband light curves.** Modeling of forward shock emission is compared to observations at different frequencies (see legend). The model shown with solid and dashed lines is optimised to describe the high-energy radiation (TeV, GeV and X-ray). It has been obtained with the following parameters: $s = 0$, $\epsilon_e = 0.07$, $\epsilon_B = 8 \times 10^{-5}$, $p = 2.6$, $n_0 = 0.5$, and $E_k = 8 \times 10^{53}$ erg. Solid lines show the total flux (synchrotron and SSC), while the dashed line refers to the SSC contribution only. Dotted curves are derived to test a better modeling of observations at lower frequencies, but fail to explain the behaviour of the TeV light curve. These are obtained with the following model parameters: $s = 2$, $\epsilon_e = 0.6$, $\epsilon_B = 10^{-4}$, $p = 2.4$, $A_\star = 0.1$, and $E_k = 4 \times 10^{53}$ erg.

616 Another possibility is to assume $\nu_m < \nu_X < \nu_c$, which implies $p = 2.1 - 2.2$ for $s = 2$ and
617 $p \sim 2.8$ for $s = 0$. A broken power law fit provides a better fit (5.3×10^{-5} probability of chance
618 improvement), with a break occurring around 4×10^4 s and decay indices $\alpha_{X,1} \sim -1.32 \pm 0.03$
619 and $\alpha_{X,2} \sim -1.55 \pm 0.04$. This behaviour can be explained by the passage of ν_c in the XRT band
620 and assuming again $p = 2.4 - 2.5$ for $s = 2$ and $p \sim 2.8$ for $s = 0$.

621 The optical light curve starts displaying a shallow decay in time (with temporal index poorly
622 constrained, between -0.5 and -0.25) starting from $\sim 2 \times 10^3$ s, followed by a steepening around
623 8×10^4 s, when the temporal decay becomes similar to the decay in X-ray band, suggesting that after
624 this time the X-ray and optical band lie in the same part of the synchrotron spectrum. If the break
625 is interpreted as the synchrotron characteristic frequency ν_m crossing the optical band, after the
626 break the observed temporal decay requires a steep value of $p \sim 3$ for $s = 0$ and a value between
627 $p = 2.4$ and $p = 2.5$ for $s = 2$. Independently of the density profile of the external medium and on
628 the cooling regime of the electrons, $\nu_m \propto t^{-3/2}$, placing it the soft X-ray band at 10^2 s. The SED at
629 ~ 100 s is indeed characterised by a peak in between 5-30 keV (Fig. 3). Information on the location
630 of the self-absorption frequency are provided by observations at 1 GHz, showing that $\nu_{sa} \sim 1$ GHz
631 at 10^5 s (Extended Data Fig. 6).

632 Summarizing, in a wind-like scenario X-ray and optical emission and their evolution in time
633 can be explained if $p = 2.4 - 2.5$, the emission is initially in fast cooling regime and transitions to
634 a slow cooling regime around 3×10^3 s. The optical spectral index at late times is predicted to be
635 $(1 - p)/2 \sim -0.72$, in agreement with observations. ν_m crosses the optical band at $t \sim 8 \times 10^4$ s,
636 explaining the steepening of the optical light curve and the flattening of the optical spectrum.

637 The X-ray band initially lies above (or close to) ν_m , and the break frequency ν_c starts crossing
 638 the X-ray band around $2 - 4 \times 10^4$ s, producing the steepening in the decay rate (the cooling
 639 frequency increases with time for $s = 2$). In this case, before the temporal break, the decay rate
 640 is related to the spectral index of the electron energy distribution by $\alpha_{X,1} = (2 - 3p)/4 \sim -1.3$,
 641 for $p \sim 2.4 - 2.5$. Well after the break, this value of p predicts a decay rate $\alpha_{X,1} = (1 - 3p)/4 =$
 642 $-1.55 - 1.62$. Overall, this interpretation is also consistent with the fact that the late time ($t > 10^5$ s)
 643 X-ray and optical light curves display similar temporal decays (Fig.1), as they lie in the same part
 644 of the synchrotron spectrum ($\nu_m < \nu_{\text{opt}} < \nu_X < \nu_c$). A similar picture can be invoked to explain
 645 the emission also assuming a homogeneous density medium, but a steeper value of p is required.
 646 In this case, however, no break is predicted in the X-ray light curve.

647 We now add to the picture the information brought by the TeV detection. The modeling is
 648 built with reference to the MAGIC flux and spectral indices derived considering statistical errors
 649 only (see Extended Data Table 1 and green data points in Extended Data Fig. 2). The light curve
 650 decays in time as $t^{-1.51}$ and the photon index is consistent within $\sim 1\sigma$ with $\Gamma_{\text{ph,TeV}} \sim -2.5$ for
 651 the entire duration of the emission, although there is evidence for an evolution from harder (~ -2)
 652 to softer (~ -2.8) values. In the first broadband SED (Fig. 3, 68-110 s), LAT observations provide
 653 strong evidence for the presence of two separated spectral peaks.

654 Assuming Thomson scattering, the SSC peak is given by:

$$\nu_{\text{peak}}^{\text{SSC}} \simeq 2 \gamma_e^2 \nu_{\text{peak}}^{\text{syn}} \quad (1)$$

655 while in KN regime, the SSC peak should be located at:

$$h\nu_{\text{peak}}^{\text{SSC}} \simeq 2 \gamma_e \Gamma m_e c^2 / (1 + z) \quad (2)$$

656 where $\gamma_e = \min(\gamma_c, \gamma_m)$. The synchrotron spectral peak is located at $E_{\text{peak}}^{\text{syn}} \sim 10 \text{ keV}$, while
 657 the peak of the SSC component must be below $E_{\text{peak}}^{\text{SSC}} \lesssim 100 \text{ GeV}$ to explain the MAGIC photon
 658 index. Both the KN and Thomson scattering regimes imply $\gamma_e \lesssim 10^3$. This small value faces
 659 two problems: i) if the bulk Lorentz factor Γ is larger than 150 (that is a necessary condition to
 660 avoid strong γ - γ opacity, see below), a small γ_m translates into a small efficiency of the electron
 661 acceleration, with $\epsilon_e < 0.05$, ii) the synchrotron peak energy can be located at $E_{\text{peak}}^{\text{syn}} \sim 10 \text{ keV}$
 662 only for $B\Gamma \gtrsim 10^5 \text{ G}$. Large B and small ϵ_e would make difficult to explain the presence of a
 663 strong SSC emission. These calculations show that γ - γ opacity likely plays a role in shaping and
 664 softening the observed spectra of the SSC spectrum^{30,87}.

665 For a gamma-ray photon with energy E_γ , the $\tau_{\gamma\gamma}$ opacity is:

$$\tau_{\gamma\gamma}(E_\gamma) = \sigma_{\gamma\gamma}(R/\Gamma) n_t(E_\gamma), \quad (3)$$

666 where $n_t = L_t / (4\pi R^2 c \Gamma E_t)$ is the density of target photons in the comoving frame, L_t is the
 667 luminosity and $E_t = (m_e c^2)^2 \Gamma^2 / E_\gamma / (1+z)^2$ is the energy of target photons in the observer frame.
 668 Target photons for photons with energy $E_\gamma = 0.2 - 1 \text{ TeV}$ and for $\Gamma \sim 120 - 150$ have energies
 669 in the range $4 - 30 \text{ keV}$. When $\gamma - \gamma$ absorption is relevant, the emission from pairs can give a
 670 non-negligible contribution to the radiative output.

671 To properly model all the physical processes that are shaping the broadband radiation, we
 672 use a numerical code that solves the evolution of the electron distributions and derives the radiative

673 output taking into account the following processes: synchrotron and SSC losses, adiabatic losses,
674 $\gamma - \gamma$ absorption, emission from pairs, and synchrotron self-absorption^{88–91}. We find that for
675 the parameters assumed in the proposed modeling (see below), the contribution from pairs to the
676 emission is negligible.

677 The MAGIC photon index (Extended Data Table 1) and its evolution with time constrain the
678 SSC peak energy to be at $\lesssim 1$ TeV at the beginning of observations (Extended Data Table 1). In
679 general the internal opacity decreases with time and KN effects become less relevant. A possible
680 softening of the spectrum with time, as the one suggested by the observations, requires that the
681 spectral peak decreases with time and moves below the MAGIC energy range. In the slow cooling
682 regime, the SSC peak evolves to higher frequencies for a wind-like medium and decreases very
683 slowly ($\nu_{\text{peak}}^{\text{SSC}} \propto t^{-1/4}$) for a constant-density medium (both in KN and Thomson regimes). In fast
684 cooling regime the evolution is faster ($\nu_{\text{peak}}^{\text{SSC}} \propto t^{-1/2} - t^{-9/4}$ depending on medium and regime).

685 We model the **multi-band** observations considering both $s = 0$ and $s = 2$. The results are
686 shown in Fig. 3, Extended Data Figs. 6 and 7 where model curves are overlaid to observations.
687 The model curves shown in these figures have been derived using the following parameters. The
688 model in Fig. 3 and in 7 (solid and dashed curves) we have used $s = 0$, $\epsilon_e = 0.07$, $\epsilon_B = 8 \times 10^{-5}$,
689 $p = 2.6$, $n_0 = 0.5$, and $E_k = 8 \times 10^{53}$ erg. For the dotted curves in Extended Data Fig. 7 and the
690 SEDs in Extended Data Figs. 6 we have used $s = 2$, $\epsilon_e = 0.6$, $\epsilon_B = 10^{-4}$, $p = 2.4$, $A_\star = 0.1$, and
691 $E_k = 4 \times 10^{53}$ erg.

692 Using the constraints on the afterglow onset time ($t_{\text{peak}}^{\text{aft}} \sim 5 - 10$ s from the smooth compo-

693 nent detected during the prompt emission) the initial bulk Lorentz factor is constrained to assume
694 values $\Gamma_0 \sim 300$ and $\Gamma_0 \sim 700$ for $s = 2$ and $s = 0$, respectively.

695 Consistently with **the** qualitative description above, we find that late time optical observa-
696 tions can indeed be explained with ν_m crossing the band (see the SED modeling in Extended Data
697 Fig. 6 and dotted curves in Extended Data Fig. 7). However a large ν_m is required in this case
698 and consequently also the peak of the SSC component would be large and lie above the MAGIC
699 energy range. The resulting MAGIC light curve (green dotted curve in Extended Data Fig. 7) does
700 not agree with observations. Relaxing the requirement on ν_m , the TeV spectra (Fig. 3) and light
701 curve (green solid curve in Extended Data Fig. 7) can be explained. As noted before, a wind-like
702 medium can explain the steepening of the X-ray light curve at 8×10^4 s, while in a homogeneous
703 medium no steepening is expected (blue dotted and solid lines in Extended Data Fig. 7). We find
704 that the GeV flux detected by LAT at late time ($t \sim 10^4$ s) is dominated by the SSC component
705 (dashed line in Extended Data Fig. 7).

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