

WR 110: A SINGLE WOLF–RAYET STAR WITH COROTATING INTERACTION REGIONS IN ITS WIND?*

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ABSTRACT

A 30 day contiguous photometric run with the *Microvariability and Oscillations of STars (MOST)* satellite on the WN5-6b star WR 110 (HD 165688) reveals a fundamental periodicity of $P = 4.08 \pm 0.55$ days along with a number of harmonics at periods P/n , with $n \approx 2, 3, 4, 5$, and 6, and a few other possible stray periodicities and/or stochastic variability on timescales longer than about a day. Spectroscopic radial velocity studies fail to reveal any plausible companion with a period in this range. Therefore, we conjecture that the observed light-curve cusps of amplitude ~ 0.01 mag that recur at a 4.08 day timescale may arise in the inner parts, or at the base, of a corotating interaction region (CIR) seen in emission as it rotates around with the star at constant angular velocity. The hard X-ray component seen in WR 110 could then be a result of a high velocity component of the CIR shock interacting with the ambient wind at several stellar radii. Given that most hot, luminous stars showing CIRs have two CIR arms, it is possible that either the fundamental period is 8.2 days or, more likely in the case of WR 110, there is indeed a second weaker CIR arm for $P = 4.08$ days, that occurs \sim two-thirds of a rotation period after the main CIR. If this interpretation is correct, WR 110 therefore joins the ranks with three other single WR stars, all WN, with confirmed CIR rotation periods (WR 1, WR 6, and WR 134), albeit with WR 110 having by far the lowest amplitude photometric modulation. This illustrates the power of being able to secure intense, continuous high-precision photometry from space-based platforms such as *MOST*. It also opens the door to revealing low-amplitude photometric variations in other WN stars, where previous attempts have failed. If all WN stars have CIRs at some level, this could be important for revealing sources of magnetism or pulsation in addition to rotation periods.

Key words: stars: individual (WR 110) – stars: rotation – stars: winds, outflows – stars: Wolf–Rayet

Online-only material: color figures

1. INTRODUCTION

While most adequately observed O stars exhibit pairs of opposing corotating interaction regions (CIRs; see Mullan 1984), manifested by propagating discrete absorption components in their resonant P Cygni absorption components (Kaper et al.

1999), little such evidence is found for CIRs in O-star descendants, the Wolf–Rayet (WR) stars. If found in WR stars, CIRs would give precious information on their rotation, which would be highly relevant in the context of rapidly rotating WR stars as prime candidates for slow gamma-ray bursts (St-Louis et al. 2009).

Recently, St-Louis et al. (2009; see also Chené & St-Louis 2010) have examined a northern sample of 25 WR stars for the presence of CIRs via large-scale variations exhibited on broad, strong emission lines. Some 20% of these stars show clear signatures of such effects, with WR 110 situated at the limit between those stars that show CIRs and those that do not. In addition, some (and possibly all, given the limited number of spectra) stars in the sample unsurprisingly show small-scale wind-clump variations. Such variations manifest themselves as stochastic propagating small-scale spectral variations on time series of virtually all WR line spectra (e.g., Lépine & Moffat 1999).

WR 110 is a moderately bright ($v = 10.30$, $b - v = 0.75$), reddened ($A_v = 3.83$ mag), southern [R.A.(2000) =

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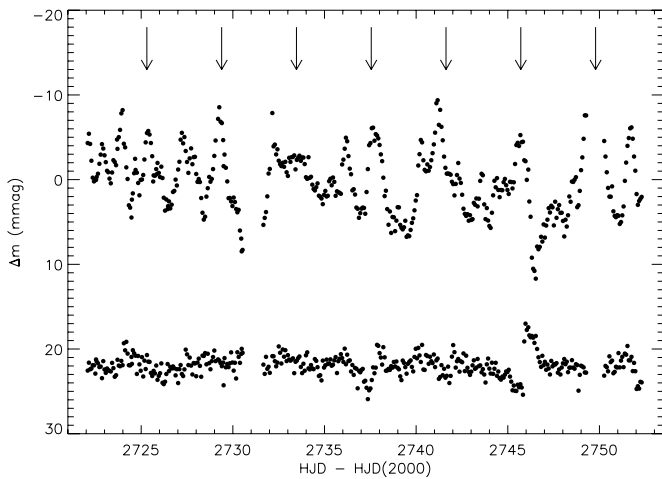


Figure 1. Upper: *MOST* light curve of WR 110 in 101 minute (*MOST* orbit) bins. The vertical lines indicate intervals of period 4.08 days starting at $t_0 = \text{HJD } 2,454,270.27$. Lower: simultaneous *MOST* light curve of HD 312657, $V = 10.8$, $Sp = A2$, reduced in the same way as WR 110, with arbitrarily displaced magnitude zero point. This star appears to be stochastically variable at the 0.005 mag peak-to-valley level and slightly larger point-to-point rms scatter of ~ 1 mmag, compared to ~ 0.5 mmag for WR 110.

18:07:56.96, decl.(2000) = $-19:23:56.8$] Galactic ($l = 10:80$, $b = +0:39$) WR star in the mid-nitrogen sequence with type WN5-6b (Smith et al. 1996). The “b” suffix indicates broad lines for its subtype and thus low H-abundance by default. Located at a distance of 1.28 kpc, it is 9 pc above the Galactic plane (van der Hucht 2001) in the normal range for most extreme population I stars. Although to date it lacks any obvious binary signature, a comprehensive radial velocity (RV) search has not yet been carried out. However, WR 110 does have relatively high X-ray flux, with a significant hard component above 3 keV compared to its lower-energy 0.5 keV emission, as seen in most WN stars (Skinner et al. 2002). The lower-energy component likely arises due to small-scale turbulent shocks in the wind (assuming thermalization of matter colliding at velocity v : $kT \approx \frac{1}{2}m_p v^2$), with velocity dispersion $\sim 10^2$ km s $^{-1}$, as seen in virtually all WR-star winds (Lépine & Moffat 1999) and other luminous hot stars (e.g., Eversberg et al. 1998; Lépine & Moffat 2008). The former hard component implies velocities of $\sim 10^3$ km s $^{-1}$, which are more difficult to explain, without invoking an additional mechanism to provide such high speeds (e.g., accretion onto a low-mass companion or colliding winds with a massive companion). However, WR 110 emits normal thermal radio emission from its wind (Skinner et al. 2002), which lends little support for its binary nature. As we show in this investigation, it is possible that CIRs are present in the wind of WR 110 and that WR 110’s hard X-rays may arise in the shocks produced by its CIRs at velocities of $\sim 10^3$ km s $^{-1}$.

2. OBSERVATIONS

The optical observations of this work consist of both high-precision, nearly continuous photometry over a month ($T_{\text{ph}} = 30$ days) from space using the *Microvariability and Oscillations of STars (MOST)* microsatellite, and ground-based optical spectroscopy from various observatories over 8 years.

2.1. Photometry

MOST (Matthews et al. 1999; Walker et al. 2003) contains a 15 cm Rumak–Maksutov telescope imaging onto a CCD detector

Table 1
Journal of Spectroscopic Observations

Telescope	Dates	No. of Spectra	Range (Å)	Å/px
Leoncito 2.2 m	2001 Apr 8–10	3	3900–5500	1.63
	2002 Apr 20–23	4	”	”
	2003 Oct 5	1	”	”
	2004 May 20–31	8	”	”
	2004 Jun 13–16	4	”	”
CFHT 3.6 m	2003 Aug 9–13	17	4300–6000	0.51
CTIO 1.5 m	2004 Jun 29–Jul 2	335	5300–6200	0.77
OMM 1.6 m	2007 Jul 13–17	5	4400–6100	0.66
Leoncito 0.6 m	2009 Mar 12–Apr 16	15	3700–5300	0.97

via a custom optical broadband (350–750 nm) filter. From its polar Sun-synchronous orbit of altitude 820 km and period 101 minutes, *MOST* has a continuous viewing zone (CVZ) about 54° wide within which it can monitor target fields for up to two months without interruption. Targets brighter than $V \sim 6$ mag are observed in Fabry imaging mode, while fainter targets (like WR 110) are observed in direct imaging mode ($3'' \times 3''$ pixels in a $\sim 1^\circ$ field), similar to standard CCD photometry with a ground-based instrument. It is not possible to obtain differential photometry, but given the orbit, thermal and design characteristics of *MOST*, experience has shown that it is a very photometrically stable platform even over long timescales (with repeatability of the mean instrumental flux from a non-variable target of the brightness of WR 110 to within about 1 mmag over a month).

Due to the failure of the tracking CCD, the science exposures on the *MOST* science CCD now take place at the same time as the guide-star exposures for satellite pointing. In the case of the WR 110 observations, the guide-star exposure time (and hence also the science target exposure time) was 0.3 s. To build up sufficient signal-to-noise ratio (S/N), the science exposures were co-added on board the satellite in “stacks” of 100 exposures. Stacks (each 30 s of total integration) were downloaded from the satellite consecutively (with no dead time between stacks), giving a sampling rate of about twice per minute.

The WR 110 field lies outside the *MOST* science CCD CVZ, so it could not be monitored (strictly speaking) continuously. This field alternated with another *MOST* primary science target (Altair) during each *MOST* satellite orbit. Typically, we monitored WR 110 for about 40%–55% of each *MOST* orbit (i.e., about 40–55 minutes of each 101 minutes) during the low-stray-light portion of the orbit, with $\sim 35,000$ individual 30 s data points between 2007 June 17 and July 16 (HJD 2,454,267–297).

Figure 1 shows the entire *MOST* light curve, binned for each *MOST* orbit. Although significant variations are seen on timescales longer than \sim a day, few variations are seen on shorter timescales and binning allows one to increase the photometric point-to-point rms precision from ~ 3 mmag before binning to ~ 0.5 mmag after binning. Interpretation of the observed variations is given below in the next section.

2.2. Spectroscopy

Nine CCD-spectroscopic runs were carried out at five telescopes over 8 years, including one parallel run with the *MOST* observations. However, this run was plagued by bad weather and instrumental problems, resulting in limited time coverage. A journal of observations is given in Table 1. Generally, the spectra cover timescales ranging from ~ 10 minutes to several years and range from near the Balmer limit to He I $\lambda 5876$,

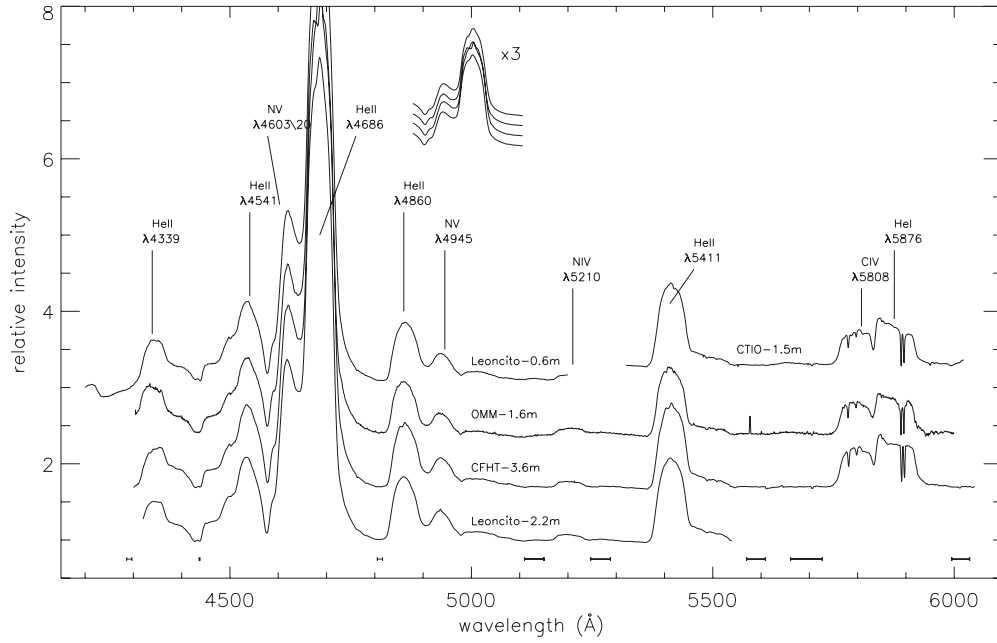


Figure 2. Mean spectra from each run as in Table 1, with some key lines identified. The vertical scale refers to the lowest spectrum, with the remaining spectra vertically displaced for clarity. The insert shows a vertical compression of the strong He II $\lambda 4686$ -line region. The horizontal bars show positions used to calculate the continuum.

although not the same for each telescope. An $S/N \sim 100$ was obtained per extracted pixel, which ranged from 0.5 to 1.6 Å, with resolution of 2–3 pixels. The bias subtraction, flat-fielding, spectrum extraction, sky subtraction, and wavelength calibration of all spectra were executed in the usual way using IRAF¹⁷ software.

Special care was taken for the normalization of the spectra. First, a mean was made for each run. Then each spectrum of a run was divided by the run mean and the ratio fitted with a low order Legendre polynomial (between the fourth and eighth orders). The original individual spectrum was divided by this fit and was therefore normalized to the run mean. When this procedure was done for each run, the run means were then put at the same level by using the same procedure as described above. This allowed us to put all individual spectra at the same level. Then, we combined all run means into a global, high-quality mean, which was then fitted in selected pseudo-continuum regions, i.e., wavelength regions where large emission lines do not dominate. These regions are shown in Figure 2. Finally, the fitted continuum function is applied to rectify each individual spectrum. The error on the continuum normalization—measured as the standard deviation of individual spectra around the continuum function—is typically of 0.5%. Mean spectra from each telescope are shown in Figure 2.

3. ANALYSIS

3.1. Light Curve

We start by examining the Fourier spectrum for the *MOST* photometry (Figure 3). Here the largest amplitude occurs at a frequency of $\nu_0 = 0.245 \pm 0.033$ cycles day⁻¹—with error equal to $1/T_{\text{ph}}$, where T_{ph} is the total duration of the light curve—($P =$

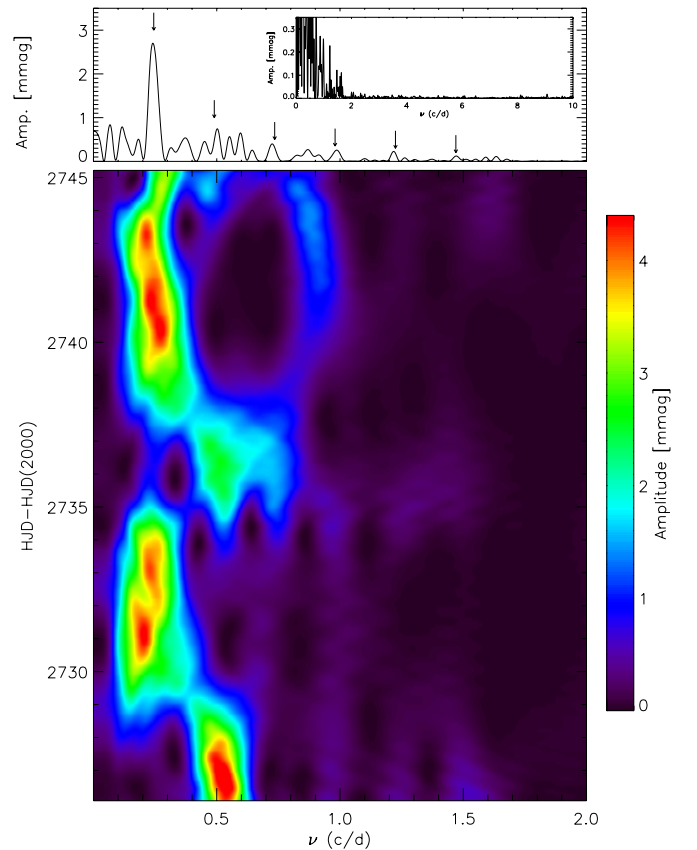


Figure 3. Upper panel: Fourier amplitude spectrum of the binned *MOST* light curve in Figure 1, with insert showing expanded scales. The highest peak refers to the adopted fundamental frequency corresponding to a period of $P = 4.08$ days. Harmonics are indicated at frequencies at multiples of two through six times fundamental (marked with downward arrows). Lower panel: time-frequency Fourier plot with 8 day running windows in time.

(A color version of this figure is available in the online journal.)

¹⁷ IRAF is distributed by the National Optical Astronomy Observatories (NOAO), which is operated by the Association of Universities for Research in Astronomy, Inc., (AURA) under cooperative agreement with the National Science Foundation (NSF).

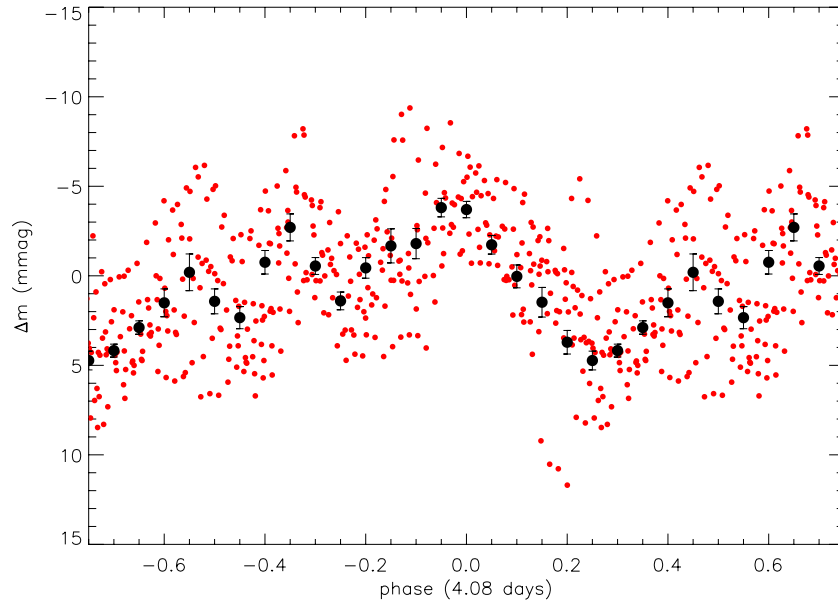


Figure 4. Phased light curve from Figure 1, using $P = 4.08$ days and $t_0 = \text{HJD } 2,454,270.27$. The large points show means in phase bins of 0.05, with 2σ error bars based on the standard error of the mean in each bin.

(A color version of this figure is available in the online journal.)

4.08 ± 0.55 days), with harmonics of this frequency at, or close to, $n \times \nu_0$, with $n \approx 2, 3, 4, 5$, and 6. A few other peaks do occur at low ($<20\%$ of the dominant peak) relative amplitudes, notably at periods longer than \sim a day, but with no obvious order or interrelationship. There is no peak due to sidelobes of the primary maximum, caused by the uniform, dense data sampling. No significant peaks beyond the above harmonics were detected with periods shorter than \sim a day. We assume that this implies a single dominating (non-sinusoidal) periodicity with $P = 4.08$ days plus mainly stochastic variability on timescales longer than \sim a day (possibly related to wind-clump activity, which survives for at least ~ 8 hr in typical WR winds; Lépine & Moffat 1999). Examination of the light curve in Figure 1 shows relatively sharp cusps occurring on this timescale (i.e., $P = 4.08$ days, starting at $t_0 = \text{HJD } 2,454,270.27$) in all seven cases but one (the third case), although even then there is a rise in brightness, though not cusp-like. The cusps have typical widths at the base of \sim a day. Such cusps are clearly non-sinusoidal, requiring a series of successive harmonics to reproduce. Figure 3 also shows a time–frequency plot with straddled 8 day (double the best period) sampling. This plot reveals that the 4.08 day period stands out at all times except near the middle and at both ends (where edge effects come into play) of the photometric time series. The former is compatible with the apparently suppressed third cusp, as noted above.

Figure 4 shows the phase-binned light curve of the orbital-mean *MOST* data from Figure 1, based on the period of 4.08 days and zero point t_0 (nearly coinciding with the first observed cusp in Figure 1). This light curve is quite noisy compared to the appearance of the non-phase-binned light curve in Figure 1. Further binning to 0.05 in phase reveals a dominating peak at phase zero, with a second lower peak at phase ~ 0.65 . Both peaks can be traced back in Figure 1 often to sharper, clearer cusps at similar repeated intervals. These cusps become somewhat smeared, although remain still visible in Figure 4, probably due to dispersion in shape and timing around some average, combined with stochastic wind variations.

3.2. Spectral Line Variations

3.2.1. Short Term

Only one spectroscopic run (CTIO 2004) is of sufficient quality and length to allow a search for variability on timescales less than a day. Figure 5 shows gray-scale time plots of the four successive CTIO nights for differences from the mean of each spectrum, for the principal WR spectral emission lines of He II $\lambda 5411$ and C IV $\lambda\lambda 5802/12 + \text{He I } \lambda 5876$. From these plots, one sees evidence for the usual stochastic, emission subpeaks, propagating outward from line center, as seen in all previously well-studied WR spectra (Lépine & Moffat 1999). Such perturbations are believed to arise in turbulent wind clumps, with no apparent periodicity on timescales of a night (~ 0.3 days) or shorter.

The simplest, most effective way to examine spectral line variations is via time dependence of line moments from order zero through four. The zeroth moment gives the line strength or equivalent width (EW) when normalized to the continuum. The first moment yields the global line position or RV. The second moment can be used to calculate the line width à la Gaussian. Line asymmetries can best be estimated by (normalized) third moments, known as skewness. Finally, normalized fourth moments yield information on how the profile differs in global shape from a Gaussian.

Figure 6 shows a time plot of normalized EW, RV, skewness, and kurtosis of the CTIO observations for He II $\lambda 5411$. The EWs were calculated by integrating the function $(1 - F_\lambda)_2$, where F_λ is the rectified line flux between $\Delta\lambda = 5348\text{--}5460$ Å. In order to highlight the variability, we prefer to divide the EWs by the mean value of all EWs. The RVs were obtained by cross-correlating individual spectra in the wavelength interval $\Delta\lambda = 5140\text{--}5980$ Å with the mean spectra. (Line bisectors were also explored to obtain RVs, but they revealed larger scatter yet similar trends as the cross-correlation technique.) The other moments of the He II $\lambda 5411$ line are calculated in the wavelength interval $\Delta\lambda = 5348\text{--}5460$ Å. The n th central moment is defined

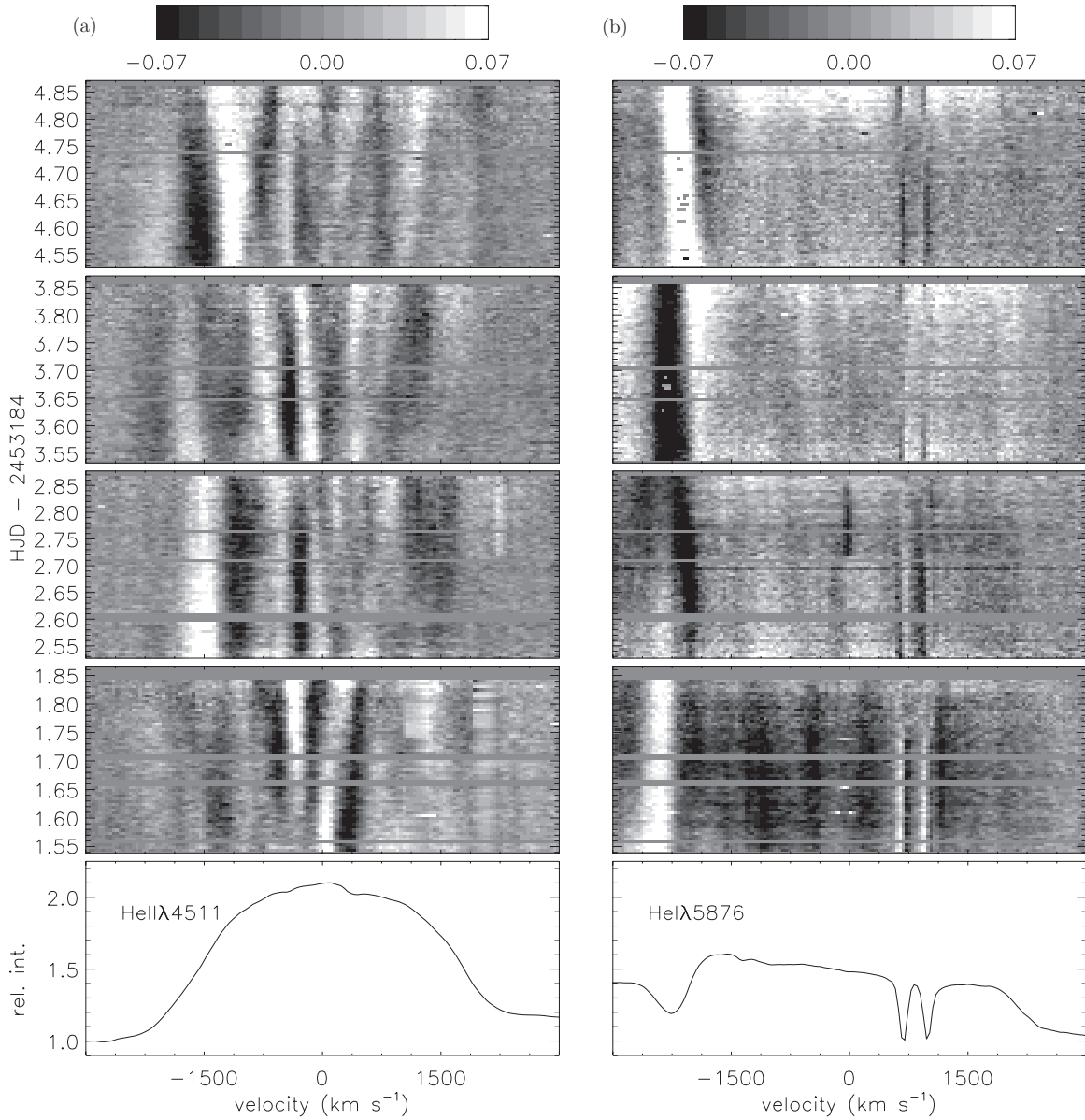


Figure 5. Upper four panels: gray-scale plots of the difference from the mean of the CTIO spectra over four successive nights (shown in the bottom panel). (a) He II $\lambda 5411$, (b) C IV $\lambda 5802/12$ + He I $\lambda 5876$.

as follows:

$$\mu_n = \sum_j (\lambda_j - \bar{\lambda})^n I_j / \sum_j I_j, \quad (1)$$

where

$$\bar{\lambda} = \sum_j \lambda_j I_j / \sum_j I_j, \quad (2)$$

with I_j being the intensity of the line and λ_j the wavelength. The skewness is $\mu_3/\mu_2^{3/2}$ and the kurtosis is μ_4/μ_2^2 . Variability is seen in all of these quantities, although with no significant periodicities. Any longer-term periodic variations of these data will be examined together with the remaining spectra in the next subsection.

3.2.2. Long Term

Figure 7 shows time plots of EW, RV, skewness, and kurtosis (measured as above) for the strongest, least blended lines: (1)

over the full 6 years for He II $\lambda 4686$ and (2) $T_{\text{sp}} = 8$ years for He II $\lambda 5411$. In the case of the latter from CTIO, we have taken nightly averages, so that the numerous original CTIO data will not dominate the period analysis. The moderately strong lines of C IV $\lambda 5802/12$ and He I $\lambda 5876$ are blended and seriously affected by telluric lines, so they were not analyzed.

While these four parameters all show variations for both lines, those of largest amplitude are in RV and skewness for He II $\lambda 4686$. These two quantities appear to vary in an anti-correlated way (correlation amplitude, $r = -0.89$), as expected if the line varies in shape; e.g., an excess on the red side of the line will cause a small positive RV shift and a negative change in skewness. A similar trend ($r = -0.86$) is also measured for these two parameters in the He II $\lambda 5411$ line, despite the apparently lower amplitude of the variations. Presumably, these are variations in profile shape and do not reflect true RV motions of the star. Furthermore, a search using phase-dispersion modulation (Stellingwerf 1978) and Fourier methods (Lenz & Breger 2005) was carried out for periodicities in the

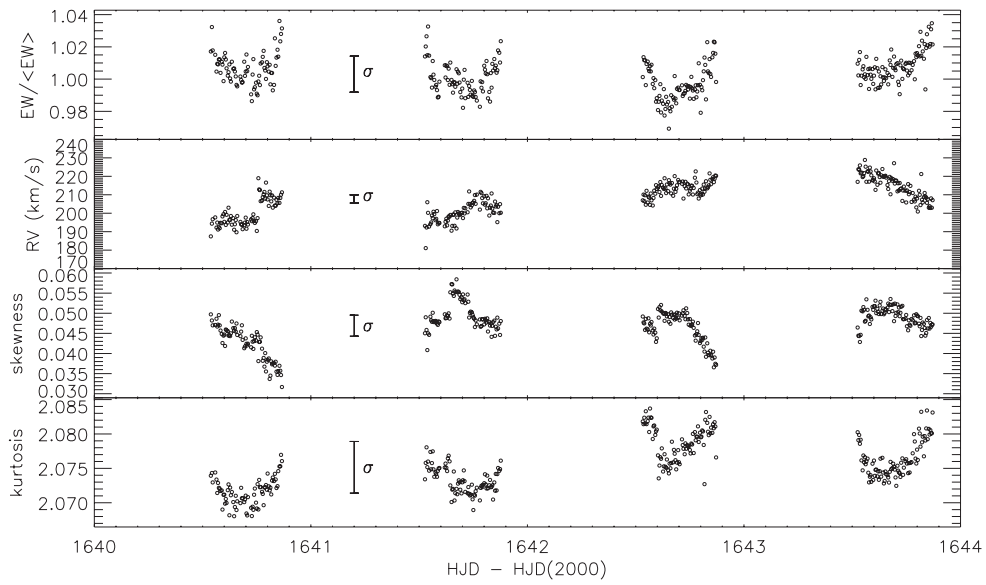


Figure 6. Normalized equivalent width (EW), line radial velocity (RV), skewness, and kurtosis vs. time for the CTIO data of He II $\lambda 5411$.

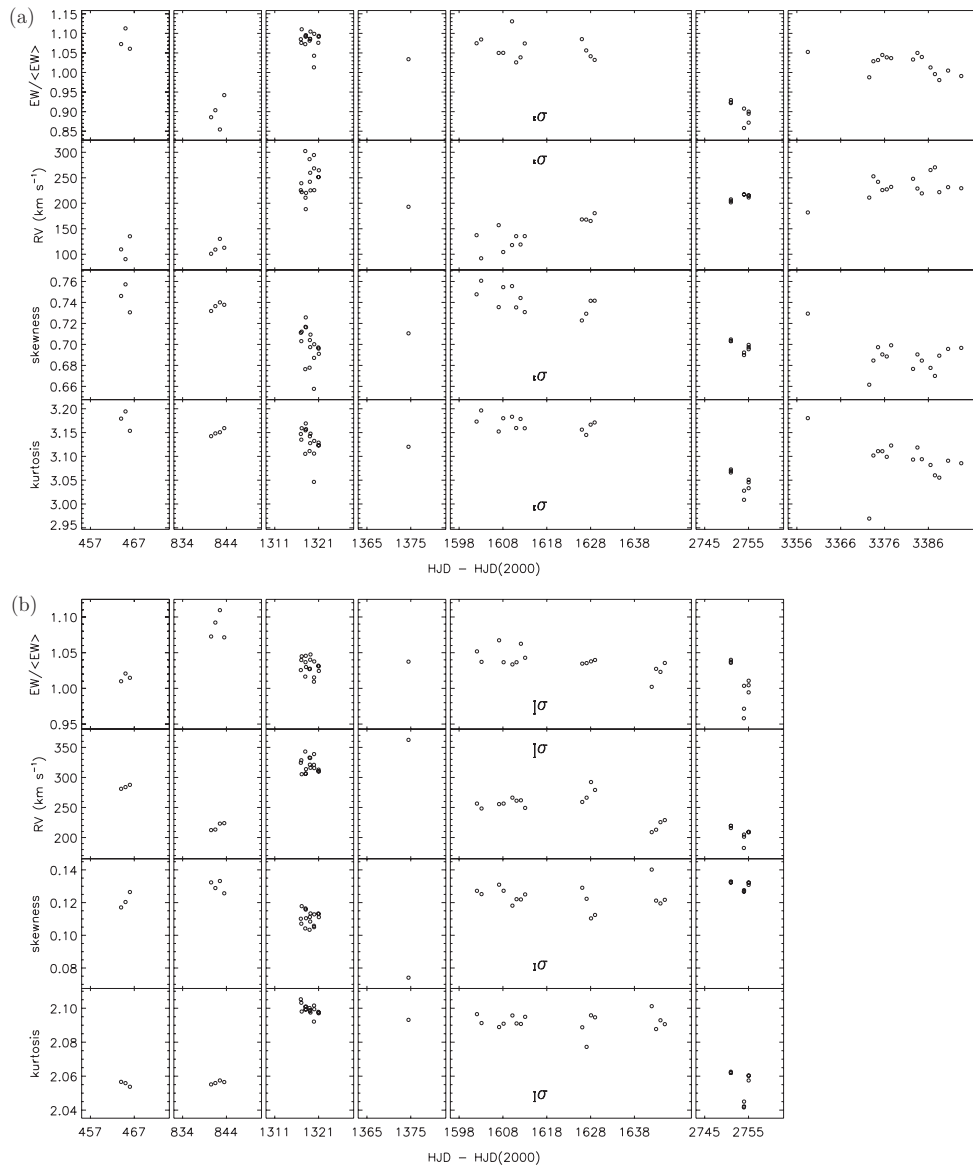


Figure 7. Same as in Figure 6 for all the spectroscopic runs for (a) He II $\lambda 4686$ and (b) He II $\lambda 5411$, with EW relative to the global mean (excluding the OMM data, due to detector zero-point intensity problems). The CTIO data for He II $\lambda 5411$ are shown in nightly means.

RVs in the range expected from the *MOST* variations. This was done in the frequency range $\nu = 0.2037\text{--}0.2885$ cycles day⁻¹, covering the whole photometric Fourier peak centered at $P = 4.08$ days, and with frequency step size 2×10^{-5} cycles day⁻¹, i.e., $\sim (10 \times 2T_{\text{sp}})^{-1}$. The search failed to reveal any significant periodicity. Alternatively, a phase RV plot of each line phased with the best period in this interval produces a scatter diagram with force-fitted sine-wave RV amplitude of $K \sim 14, 10,$ and 28 km s⁻¹ for all RV data (nightly means for CTIO, with typical errors of ~ 10 km s⁻¹) for He II $\lambda 5411$, the 335 CTIO data alone for He II $\lambda 5411$, and all RV data for He II $\lambda 4686$, respectively. The lower the scatter, the lower the K value. Combined with the above information on anti-correlation of RV and skewness, we deduce that if WR 110 is a binary system, it is not evident from the data presented here. On the other hand, the lack of a clear indication of a 4.08 day period in any of the spectroscopic data does not necessarily mean that such a period is not present. We suspect that it will take a much larger, higher-quality spectroscopic sample to reach any firm conclusions on any periodicity in spectral behavior of WR 110, as was the case with WR 134 (Morel et al. 1999).

4. CIR MODEL

Plausible scenarios to account for the 4.08 day periodicity include binarity, pulsations, and CIRs. Binarity is unlikely, due mainly to a lack of coherent RV variations that phase with the primary photometric period. Pulsations are unlikely due to the highly non-sinusoidal form of the variations at $P = 4.08$ days. Hence, the most likely scenario to explain the low-amplitude ($\sim 1\%$) 4.08 day periodicity and cusp-like behavior in the *MOST* light curve, along with mostly unrelated spectral variations, appears to be large-scale, emitting overdensity structures rotating with the wind. (One also has the impression of cusp-like minima in the light curve; however, these could be a result of the small space between two emitting cusps occurring per rotation—see below.) This would be similar to what is seen in a few other WR stars (WR 6, WR 134, WR 1—see below). CIRs (or their source at the base of the wind) as proposed for O-star winds by Cranmer & Owocki (1996) provide an attractive scenario to explain our observations. The Cranmer & Owocki model is designed for O stars, explicitly assuming CAK (Castor et al. 1975) parameters, which may not apply directly to H-free WR stars. Nevertheless, the idea of a CIR driver at the base of the wind in WR stars is a qualitatively plausible scenario, analogous to O-star winds. We now apply this idea to WR 110 in the following way.

First, in the context of the Cranmer & Owocki model, we assume that the CIRs are continuum-emitting thermal sources, heated by the associated shock action created by a hot spot on the underlying rotating star interacting with the ambient wind. Most of the CIR emission must arise close to the star, where the ambient wind density is highest, as inspired by the simulations of Cranmer & Owocki (1996)—see a simulation in their Figure 5. Alternatively, the emission could arise in a hot spot at or close to the stellar surface that produces the CIR. To simplify our calculations, we assume a point source located at some radius $R_s = \gamma R_*$ from the center of the star, with $\gamma > 1$ and R_* being the stellar core radius. For simplicity (and following Cranmer & Owocki 1996), we assume the CIR to initiate at the stellar rotation equator.

In this scenario, the changes in the light curve are caused by an assumed CIR-associated point source attenuated by the wind and seen at different angles with the line of sight as it rotates on the near side of the WR star. In this study, we use a simple WR

wind model derived by Lamontagne et al. (1996) in the context of wind eclipses for WR + O systems. Hence, the change in magnitudes is defined as

$$\Delta m = \text{constant} - 2.5 \log_{10}(I_{\text{WR}} + I_s e^{-\tau}), \quad (3)$$

with arbitrary constant and where I_{WR} is the intensity of the WR star and I_s is the intensity of the hot spot. Here the total opacity τ between the source and the observer passing through the WR wind is

$$\tau = k \int_{z_0}^{\infty} d(z/R_s) [(r/R_s)^2 (1 - R_*/r)^\beta]^{-1}, \quad (4)$$

where $(r/R_s)^2 = (\cos i \cos 2\pi\phi)^2 + \sin^2 2\pi\phi + (z/R_s)^2$, $z_0 = -(R_s \sin i) \cos 2\pi\phi$, i is the inclination of the rotation axis (assumed to coincide with the CIR rotation axis) relative to the observer, and ϕ is the rotation phase. We allow for a WR wind with a β law (like in Lamontagne et al. 1996), in which the actual β value will be fitted. Since the CIR source is much closer to the inner WR wind than the O star is in the (non photospheric eclipsing) binary case, we expect that $\beta = 0$ will not be a good choice, as it was in the case of WR + O binaries (Lamontagne et al. 1996). Following Skinner et al. (2002), we adopt a radio-based thermal mass-loss rate for the WR star of $\dot{M} = 1.6 \times 10^{-5} M_\odot \text{ yr}^{-1}$, after dividing by a factor of three to allow for wind clumping, a wind terminal speed of $v_\infty = 2100$ km s⁻¹, and a WR core radius $R_* = 4 R_\odot$. We find an opacity constant as in Lamontagne et al. (1996) $k = \alpha \dot{M} \sigma_e / [4\pi m_p v_\infty R_s] = 0.28/\gamma$, where $\alpha = 0.5$ is the number of free electrons per He nucleus in a highly ionized wind. In our case, we are primarily interested in the phases centered at phase 0.50 (i.e., when the source is between the star and the observer) since the CIR point source will not be normally visible as it rotates on the back side of the WR star.

We calculated Equation (3) numerically, centered at phase 0.5 for half a rotation period, as a function of the parameters β , I_s/I_{WR} , i , and γ . We consider two cases: (1) an optically thin point source, whose emission is isotropic in direction and (2) an optically thick source whose emission is maximum when seen perpendicularly to the stellar surface, dropping to zero when seen parallel to the surface (i.e., like pancakes on or near the stellar surface). In case (1) light-curve excesses originate from the hot spot and in case (2), we have to multiply the net spot intensity (I_s) from Equation (3) by the projection factor $(\sin i) \cos[2\pi(\phi - 0.5)]$.

Taking the second, the fourth and the fifth cusps in Figure 1 as the cleanest, most representative form of the hot spot light curve, we fit the above parameters using the IDL routine *mpfit.pro* (Markwardt 2009). The fit does not converge when the inclination i is set as a free parameter. Therefore, we fix it to different angles in the range $i = 90^\circ\text{--}40^\circ$ with increments of 10° before fitting all the other parameters. A summary of fitted parameters using different inclination angles is given in Table 2. Note that we do not include our results for an inclination i lower than 40° , since the fit does not converge when using such small angles. The values of I_s/I_{WR} tend to be slightly smaller in case (2) than in case (1), due to the projection factor in case (2) that requires less extended wind opacity to bring about the observed cuspy shape. Indeed, it is the varying opacity through the WR wind as the WR star rotates that leads to the cuspy shape; the closer the source is to the stellar surface (i.e., γ closer to unity), the smaller the β value (and less extended the wind) necessary to give the cuspy shape. The best solutions for the thin and the

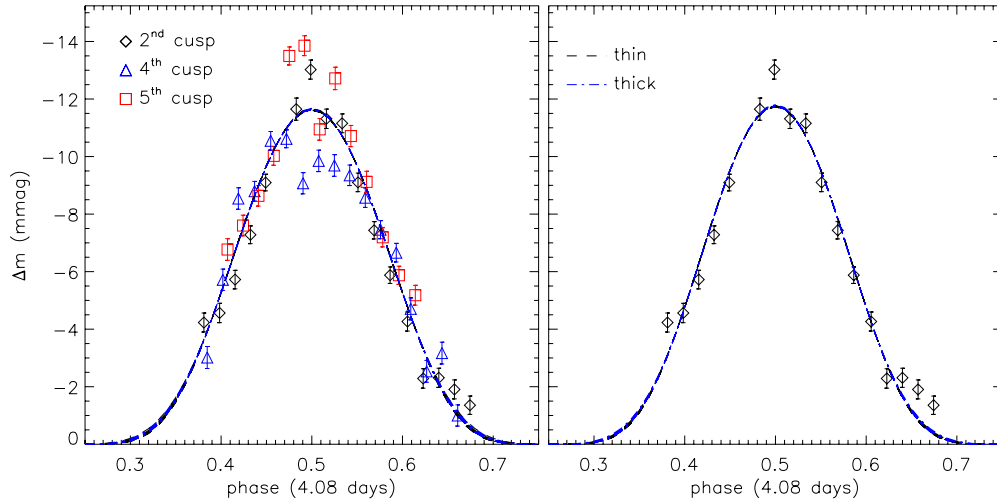


Figure 8. Left: the data points refer to the second (black diamonds), the fourth (blue triangles), and the fifth (red squares) cusps indicated by vertical lines in Figure 1. We assume that these cusps begin and end when the light curve reaches a minimum, i.e., when there is no enhancement due to the passage of a CIR. To avoid contamination from previous and subsequent features, three cusps are selected in the time intervals HJD 2,454,273.725–275.000, HJD 2,454,282.240–283.450, and HJD 2,454,285.770–286.720, respectively. To facilitate the fit, the data points are shifted vertically by an arbitrary constant, so the minimum is fixed at $\Delta m = 0$. The curves indicate the best fits at different inclination angles (see Table 2) for an optically thin point source, using black dashed lines, and an optically thick source, using blue dash-dotted lines (see the text). Note that these curves are superposed and hard to distinguish. Right: same as the left panel, but only using the second cusp.

(A color version of this figure is available in the online journal.)

Table 2
CIR Model Fit Parameters

Optically	No. of Cusps	Inc. ($^{\circ}$)	β	γ	I_s/I_{WR}	χ^2_{red}
Thin	3	90	13.36	2.825	1.351	1.489
		80	13.39	2.832	1.393	1.484
		70	13.49	2.853	1.536	1.471
		60	13.67	2.889	1.842	1.447
		50	13.98	2.947	2.497	1.412
Thick	3	40	14.48	3.032	4.145	1.363
		90	11.46	2.652	0.405	1.442
		80	11.47	2.656	0.421	1.440
		70	11.59	2.681	0.476	1.430
		60	11.76	2.716	0.594	1.412
Thin	1	50	12.08	2.778	0.847	1.386
		40	12.56	2.863	1.488	1.350
		90	14.19	2.895	2.595	1.198
		80	14.20	2.899	2.685	1.194
		70	14.30	2.920	2.993	1.180
Thick	1	60	14.51	2.960	3.666	1.155
		50	14.81	3.017	5.151	1.119
		40	15.32	3.102	9.095	1.068
		90	12.48	2.744	0.779	1.154
		80	12.51	2.750	0.813	1.150
		70	12.62	2.773	0.929	1.140
		60	12.79	2.808	1.187	1.121
		50	13.12	2.870	1.758	1.094
		40	13.62	2.956	3.294	1.055

thick cases at different inclination angles are plotted in the left panel of Figure 8. We also repeated the fits for the thin and thick cases at the different inclination angles using only the second cusp, in order to verify how our results would vary depending on this choice. The best solutions are plotted in the right panel of Figure 8 and also summarized in Table 2. The difference between the results using one or three cusps may come from an epoch dependency of the variations, smaller cusps present in only some of the considered cusps, and/or a change in the shape and amplitudes between the different cusps due to superposed stochastic variability.

All the fits give a χ^2_{red} close to unity. It is hence, at first, not possible to exclude any result only based on the quality of the fit. As seen in Figure 8, changing the inclination does not change the shape of the curve as much as the amplitude and requires higher values of β , γ , and I_s/I_{WR} (this may explain why the fit does not converge when the inclination i is set as a free parameter). However, in many cases $I_s/I_{WR} > 1$, which means that the intensity of the spot is greater than the intensity of the whole star. Since such a solution is not likely, we may exclude the optically thin case and set $i \gtrsim 70^{\circ}$. The remaining solutions imply $\beta \approx 11.5$ – 12.5 . These value would give $\beta R_* \approx 50 R_{\odot}$, which is comparable to the value obtained by Lépine & Moffat (1999) for stars of similar spectral type.

Although the overall fit is quite good, the model does not reproduce well the peaked central cusps. Tweaking the model to do this may require assumptions that go beyond the scope of this study. For example, it might be possible that the expected density-deficit channel parallel to the CIR (Cranmer & Owocki 1996) could affect the light-curve peaks.

5. DISCUSSION AND CONCLUSIONS

Assuming that CIRs do in fact account for the variability observed in the *MOST* light curve of WR 110, we can now provide a plausible explanation for the hard X-rays coming from WR 110. According to the models of CIR (Cranmer & Owocki 1996), shock velocities in the radial direction can reach over 10^3 km s $^{-1}$ at several stellar radii (beyond where the optical hot spot is seen), although highly variable from one model to another. This is several times larger than the velocity dispersion ($\sim 10^2$ km s $^{-1}$; Lépine & Moffat 1999) in clumps that is likely associated with the low-energy X-rays ($kT \approx 0.5$ keV). Scaling up proportional to the square of the velocity, one can easily account for the observed hard X-rays using CIRs without accretion onto a compact or low-mass binary companion. Note that the wind terminal velocity of WR 110, although high (2100 km s $^{-1}$ from Skinner et al. 2002, or 2300 km s $^{-1}$ from Hamann et al. 2006), is not the source per se of the X-rays; it is the velocity dispersion or the radial CIR velocity component

which counts for producing the low-energy versus high-energy X-ray shocks, respectively.

Our modeling of the photometric light curve leads us to conclude that WR 110 joins the ranks of three other WR stars, all WN, that exhibit CIRs with known rotation periods of the underlying star: WR 6, WN4, $P = 3.76$ days; WR 134, WN6, $P = 2.3$ days; WR 1, WN4, $P = 16.9$ days (Morel et al. 1997, 1999; Chené & St-Louis 2010, respectively). While the periodicity in WR 6 was easy to find because of the high photometric amplitude (~ 0.1 mag), WR 134 was more difficult with ~ 0.03 mag plus significant noise from other sources, and despite its large amplitude ~ 0.1 mag, WR 1 has a long period and complex light curve. WR 110 has an even lower amplitude, with significant noise, making the availability of *MOST* for intense, long-term monitoring indispensable for the discovery of its CIRs. In all four cases, the cuspy nature of the often multiple light-curve bumps is evident and may thus be due to the same explanation: continuum-emitting CIRs that emit mainly just outside the inner opaque WR wind, where the intensity falls off rapidly at larger radii. In most cases, there are two cusps, implying two CIRs. This is the usual number found in O-star winds (Kaper et al. 1999).

None of these rotation rates appear to be rapid enough to be directly associated with gamma-ray bursts, assuming rotational coupling between the surface and the interior (St-Louis et al. 2009).

Despite its very small amplitude and complex pattern, WR 110 has possibly revealed its CIR rotation properties due to the power of the *MOST* satellite, i.e., its ability to observe continuously for a month at high precision. This may open up the door to revealing other subtle cases, giving us a better idea of the properties of WR stars and their strong winds. It may even have given us a lead to probing the mysterious origin of the CIR phenomenon rooted at the stellar surface, e.g., as a phenomenon related to magnetism or pulsations.

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