



*Mapana Journal of Sciences*

2018, Vol. 17, No. 1, 11–23

ISSN 0975-3303 | <https://doi.org/10.12723/mjs.44.2>

# Gaia: Surveying Heavens

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## Abstract

In this paper we attempt to study an ongoing astrometry mission of the European Space Agency (ESA), named Gaia, whose aim is to make the largest and most precise three-dimensional map of our Galaxy. We present the scientific goals of Gaia and give a brief description of the spacecraft. We also present a preliminary analysis of comparing distance estimates of Be stars from the first Gaia data release, Gaia DR1, and Hipparcos mission. From our analysis, we confirm that Gaia stands out as a promising mission in terms of the distance measurements when compared to Hipparcos, particularly for distances greater than 1 kpc.

**Keywords:** Gaia, parallax, Be star, Hipparcos, Distance

## 1. Introduction

On 19 December 2013, Gaia space observatory was launched by ESA (European Space Agency) from Kourou in French Guiana. The astrometric mission aimed to make largest and most precise three-dimensional (3D) map of our Galaxy, the Milky Way [4]. It was proposed in 2000 as an ESA-only mission whose phase implementation was from 2006 by Airbus Defense and Space. The analysis and processing of data were entrusted to Data Processing and Analysis Consortium (DPAC). Before Gaia, there have been attempts to make such maps of Milky Way Galaxy. Hipparcos was the predecessor mission by ESA for mapping the Galaxy [6]. The mission focused on nearby stars which

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measured the absolute parallax and proper motion with milli-arc second accuracy. A new set of reduction process for Hipparcos data was proposed in 2007 which increased the accuracy of distance to the stars with magnitudes brighter than 8 by a factor of 4 [24]. Using the star mapper of the ESA Hipparcos satellite, another astrometric catalogue (including position, proper motion, parallax, magnitudes in two bands), Tycho-2, of 2.5 million brightest (scanned the sky to a  $V = 11.5$  magnitude) stars are generated and presented in [9]. On the other hand, Gaia intends to prepare a three dimensional catalogue that determines the proper motions and stellar parallaxes with very high accuracy.

## 2. Scientific Goals

Gaia is an ambitious mission to chart a 3D map of our Galaxy and in the process revealing the composition, formation and evolution of the Galaxy [4]. With respect to ground-based facilities, Gaia could provide more accuracy, sensitivity, dynamic range, and sky coverage. During the mapping, it will examine the kinematical, dynamical and chemical structure of stars and the evolution of our Milky Way Galaxy [16]. Many other discoveries such as planets around stars [7], moving bodies in Solar System [23], far-distant supernovae and quasars [18] and information about interstellar gas and dust, and dark matter can be obtained through this mission.

During the initial lifetime of five years, Gaia will examine the motions of about 1 billion stars and it will observe a single object approximately 70 times during its expected mission period [4]. At each epoch, photometric measurements are also made. Continuous space survey will provide data regarding the positional and kinematic distribution of stars in different regions extending from the Galactic center to the halo [18]. Combining the absolute luminosities of stars with the information of the physical properties and corresponding metallicities, it is possible to deduce the star formation histories of the stellar populations in the Galaxy [16].

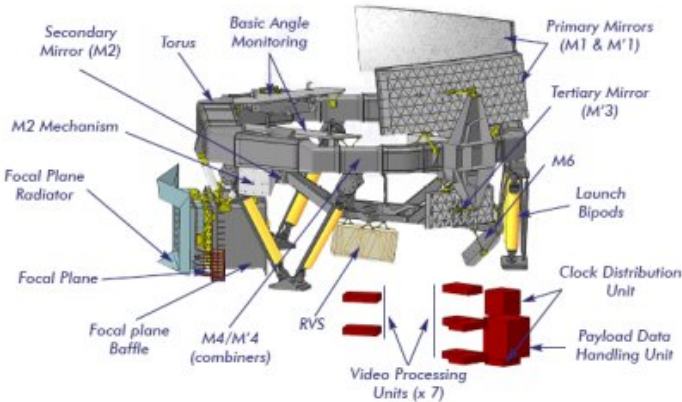
The revolutionary accuracy of Gaia in parallaxes will help in deriving highly precise color-magnitude diagrams (CMDs) so that the CMDs can be calibrated with a very minimum error rate. With 1 billion parallaxes observed, Gaia will cover most phases of evolution across the stellar mass range. The photometric and spectroscopic data can reinforce our knowledge of multiple stars, Cepheid variables, quasars, supernovae in distant galaxies, brown dwarfs and other rare objects in space [17]. Also, astrometric data will be a strong tool in exoplanets research area such as the discovery of many Jupiter sized planets in multi-year orbits around their host stars in our Galaxy [7]. The full-sky coverage of Gaia will provide the detection of movement

of faint objects in the Solar system such as asteroids and comets [23]. The spatial resolution of Gaia is sufficient to resolve and observe some Local Group galaxies.

Gaia data will be a strong tool for testing relativistic effects like light bending due to the gravitational field of Solar system on stellar objects and their motions [13]. So with an extensively accurate data on positions, parallaxes, photometric and spectroscopic measurements, the scientific goals of Gaia are so vast and immense which will probe deep into the star formation history, chemical evolution in the disk, bulge, halo and even in the outer parts of the Milky Way [4]. Exploration of inter-galactic interactions may reveal mass distributions of unresolved and invisible dark matter.

### 3. Spacecraft Overview

Gaia is placed in an orbit around the Sun, at the second Lagrange point L2 where it scans the entire celestial sphere during the course of one year [4]. Figure 1 represents the schematic diagram of the instruments loaded in the Gaia spacecraft. It is equipped with two identical telescopes that point  $106.5^\circ$  apart, around an optical bench. The single integrated focal plane assembly, besides wave-front-sensing and basic angle metrology, comprises three science functions: astrometry, photometry, and spectroscopy [4].



**Figure 1.** A pictorial representation of Gaia mission payload. The payload module is built around an optical bench which provides the structural support for the single integrated instrument that performs three functions: astrometry, photometry and spectrometry. Courtesy: [www.sci.esa.int](http://www.sci.esa.int)

Finding stars positions in 3D require further measurements which are performed by the payload module. It contains a single integrated

instrument that performs the above three major functions. The astrometric instrument measures the position, parallax and proper motion of the star [5]. The spectral energy distribution is measured by the photometric instrument while the radial velocities of stars through Doppler shift measurements are taken from the spectroscopic instrument [4]. The mechanical and the electrical service module deals with the external elements supporting the instrument and the spacecraft electronics. It helps the Gaia payload and spacecraft for pointing. It supports electrical power control and distribution, data management and communications with earth [4].

#### 4. Analysis

We carried out an analysis about the first Gaia data release, Gaia DR1 2016 (Gaia Collaboration 2016, hereafter Gaia DR1), consists of astrometric and photometric data for over 1 billion sources. The astrometric dataset contains the positions, parallaxes and mean proper motions for about 2 million of the brightest stars in common with the Hipparcos catalogue [6] and Tycho-2 catalogue [9] and the positions for an additional 1.1 billion sources. Compared to Hipparcos data, Gaia data provides much better catalogue in terms of the expected accuracies at current epochs with outstanding contributions.

We performed the analysis with a sample of classical Be stars using the first version of Gaia DR1 data release. A classical Be star is a non-supergiant B-type star whose spectrum has, or had at some time, one or more Balmer lines in emission [10, 20]. They are main sequence stars presumed to be fast rotators spinning at 200 km/s or more [22]. It is having a decretion disk fed by mass ejected from the central star, which more precisely is an outwardly diffusing gaseous Keplerian disk [21]. The formation and structure of circumstellar disk [14, 25] and the evolutionary status of classical Be stars [19, 11, 12] are some of the unresolved problems.

##### 4.1 Data Analysis

The selected Be catalogue [8], contains 1159 objects which has been cross-matched with Gaia DR1. Among this only 1007 Be stars has been matched with Gaia DR1 in which parallaxes of 736 are available. After rejecting the objects with negative parallaxes, 318 objects are obtained from Gaia DR1, and the parallaxes are also obtained from the Hipparcos catalogue. The data extracted from the two catalogues are listed in Table 1.





Table 1 – *Continued from previous page*

RA	Dec	Gaia distance	Hipparcos distance	RA	Dec	Gaia distance	Hipparcos distance
(deg)	(deg)	(pc)	(pc)	(deg)	(deg)	(pc)	(pc)
168.0242	-71.2172	226 ± 14	195 ± 27	344.0479	58.6672	398 ± 37	286 ± 176
168.1158	-76.7394	199 ± 9	318 ± 750	344.1774	62.6244	1205 ± 697	763 ± 559
168.2537	-57.0349	909 ± 231	990 ± 1020	345.2276	38.708	395 ± 75	273 ± 55
170.6325	-53.3701	649 ± 156	1042 ± 673	347.3196	49.6504	709 ± 347	649 ± 308
171.5055	-59.3524	676 ± 187	685 ± 366	349.8608	79.0036	19 ± 0	20 ± 0
171.9875	-62.805	971 ± 273	508 ± 211	350.0789	55.8075	2273 ± 1291	588 ± 315
172.779	-62.9467	680 ± 333	1538 ± 1775	350.5432	56.3479	526 ± 86	403 ± 135
172.9615	-68.0573	1250 ± 391	1538 ± 1846	357.471	62.2138	769 ± 195	1149 ± 898
173.3571	-70.1950	109 ± 4	103 ± 6	358.7671	28.6336	39 ± 0	42 ± 2

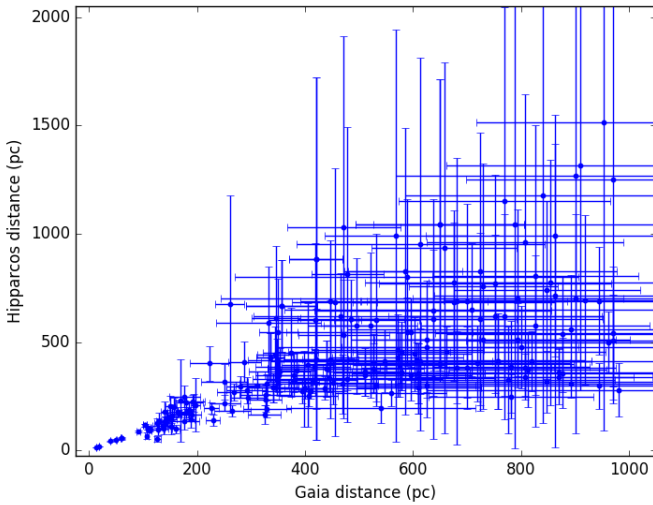
The right ascension (RA) and declination (Dec) of 318 objects with their distance estimations using Gaia and Hipparcos are listed in Table 1. We also carried out error estimations for both catalogue and are given along with the corresponding distance.

## 5. Discussion

We plotted the distance estimated from the Hipparcos data versus Gaia DR1 for a sample (Gaia distance  $\leq 1$  kpc) of 236 Be stars and is illustrated in Figure 2. Within the Gaia distance of 1 kpc, the objects with percentage error in distance more than 80% are eliminated from plotting in this figure and the rest are plotted with error bars. In the X-axis distances from Gaia DR1 data (distance limited to 1 kpc) and in Y-axis the corresponding distances calculated from Hipparcos data is taken. It is clear from the Figure 2 that, for nearby objects (in the range of 0 – 300 pc) the error is small, when distance increases for both Hipparcos and Gaia DR1, the error in distance gradually increases. But increase in error is more prominent for Hipparcos compared to Gaia. From the figure it can be concluded that for distance less than 250 pc both Gaia and Hipparcos missions produce reliable estimates of distance. Also, the distances obtained from Gaia DR1 data is more preferable, within the distance range 250 pc to 600 pc, than the Hipparcos data. For farther objects, distances calculated using both Gaia and Hipparcos contain significant error.

We also plotted histograms with distances of 318 Be stars from Gaia DR1 and Hipparcos data taken along X-axis and the number of objects along Y-axis. The clustered bars shows the objects with different percentage error corresponding to a particular distance range.

In Figure 3, the distance for 318 objects obtained from Gaia DR1 data are classified to 4 bins and are taken along the X-axis while the number of objects are taken along the Y-axis. The histogram shows clustered bars in different colours, representing the number of objects within the different ranges of percentage error (%) in distance. From the plot, it can be observed that, for distances ranging from 0 to 0.5 kpc, 97 objects with percentage error less than 20% and 7 objects with percentage error in the range 20 to 50% are present. Within this range of distance there are no objects present with percentage error

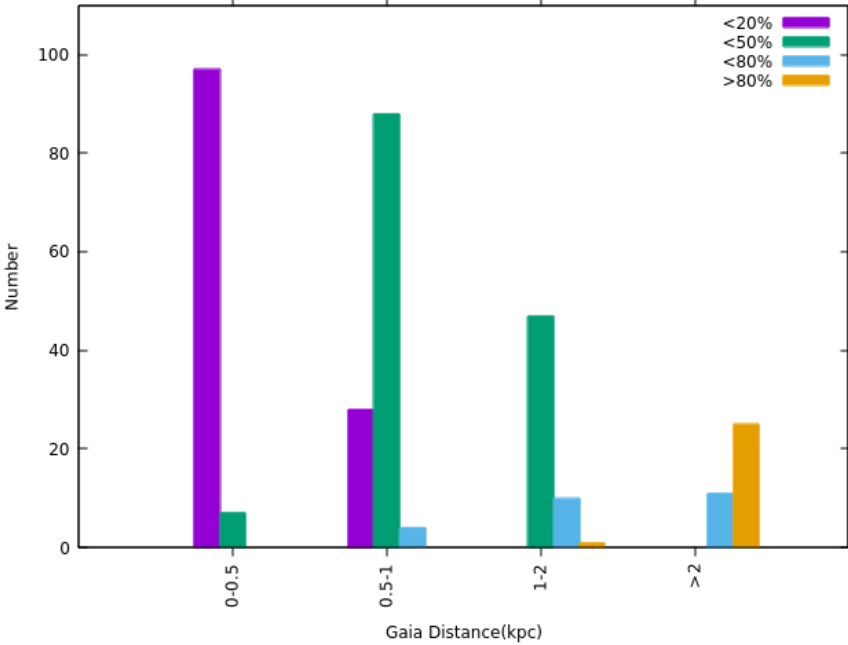


**Figure 2.** Figure represents Hipparcos distance versus Gaia DR1 distance for 318 Be stars from [8]. The X-axis shows Gaia distance within 1 kpc and Y-axis shows the corresponding Hipparcos distance. The objects are plotted with corresponding percentage error in both axes.

greater than 50%. For the distances ranging from 0.5 to 1 kpc, 28 objects fall in the range of  $< 20\%$ , 88 objects fall in the range of 20 to 50%, and just 4 objects fall in the range of 50 to 80%. Now, considering the distances ranging from 1 to 2 kpc, 47 objects are present within the range of 20 to 50% while there is no object present with a percentage error less than 20%. In this distance range, 10 objects have their percentage error in the range 50 to 80% and 1 object have percentage error  $> 80\%$ . When moving to farther distances, greater than 2 kpc, Gaia DR1 data have a significant amount of error indicating 25 objects to be present in range of percentage error greater than 80%. Also 17 objects fall in the range of 50 to 80% and only no object have its percentage error less than 50%.

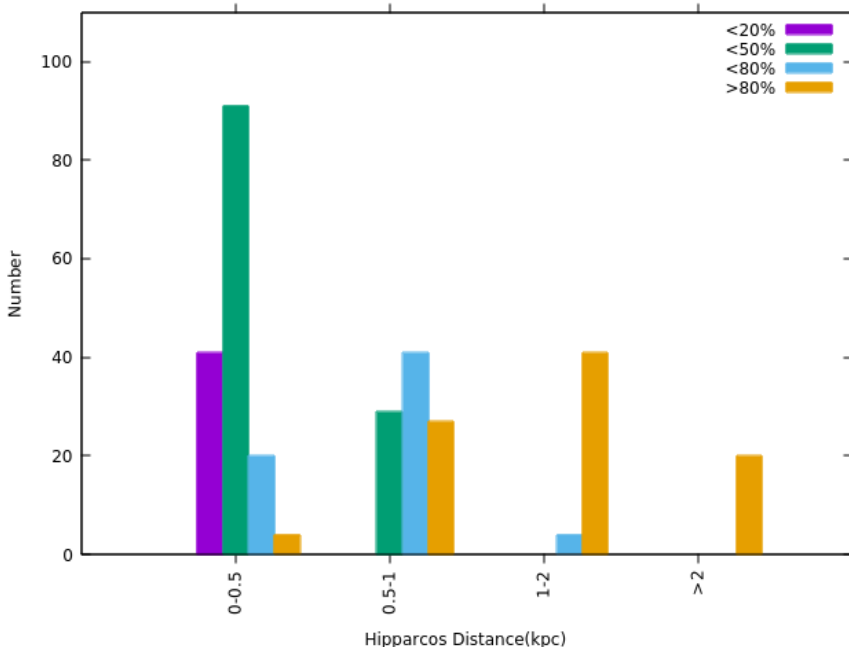
Similarly, in Figure 4, the distances obtained from Hipparcos data are arranged in 4 bins for 318 objects are taken along the X-axis and the number of objects are taken along the Y-axis. Here, in the distance range of 0 to 0.5 kpc, 41 objects are seen with percentage error  $< 20\%$  which is much less than that observed in the case of Gaia. On the other hand, there are about 91 objects present within the range of 20 to 50% which is very much greater than in the case of Gaia. Also it can be seen that 20 objects have a percentage error greater than 50% and around 4 objects have a percentage error greater than 80%. Considering the distance range of 0.5 to 1 kpc, there is no object





**Figure 3.** Histogram illustrates Gaia distances [4] of 318 Be stars. X-axis contains the distance bins (0 – 0.5, 0.5 – 1.0, 1 – 2, >2 kpc) for Gaia distances and the respective number of objects in Y-axis. The clustered bars show the objects with different percentage error corresponding to a particular distance range.

present with percentage error less than 20%. But it is observed that, 29 objects are present with percentage error between 20% and 50%, 41 objects are present with percentage error between 50% and 80%, and 27 objects are having a much greater percentage error of more than 80%. The percentage error in distance calculated from Hipparcos data for our selected objects increases drastically with increase in distance. It is clearly visible that there are no objects with percentage error less than 50% in the distance ranging from 1 to 2 kpc. In the case of objects in our selected list with distance greater than 2 kpc, all the objects have their percentage error > 80%. From the Figure 3 & Figure 4, we can conclude that the Gaia DR1 data is more preferable than the Hipparcos data for our selected objects. For nearer distances (0 to 1 kpc), Gaia data consists more objects with percentage error less than 20% than Hipparcos data i.e., in this range, distances calculated using Hipparcos data contains percentage error greater than 50% and some objects have percentage error greater than 80%. At larger distances, both Gaia and Hipparcos consist significant error. But considering the objects in our list with distances greater than 2 kpc, Hipparcos data gives only objects with percentage error greater



**Figure 4.** Histogram illustrates Hipparcos distances ESA [6] of 318 Be stars. X-axis contains the distance bins (0 – 0.5, 0.5 – 1.0, 1 – 2, >2 kpc) for Hipparcos distances and the respective number of objects in Y-axis. The clustered bars shows the objects with different percentage error corresponding to a particular distance range.

than 80% and there are no objects with less percentage error. Hence, we make an inference that, for our selected list of Be stars the Gaia data is more preferable than the corresponding Hipparcos data.

## 6. Conclusion

Gaia is the space-astrometry mission of the European Space Agency which is aimed to prepare a 3D spatial and velocity distribution of stars. The data of Gaia will have a strong impact on many other areas of astrophysical research. The Gaia satellite has been built under an ESA contract by Airbus Defense and Space and the raw data has been processed by the Gaia Data Processing and Analysis Consortium. The first intermediate release of Gaia DR1 comprises astrometry and photometry data. The advancement in technology makes Gaia DR1 more reliable than the Hipparcos data. We carried out a comparison of distances of Be stars from [8] with the available Gaia and Hipparcos data. Analysis of the data set through an illustration revealed that the obtained Gaia and Hipparcos distances show very low errors for

nearby objects ( $< 250$  pc). Two histogram representation of percentage error with distances of the selected objects from Gaia DR1 and Hipparcos data. The histogram implies that objects with higher errors ( $> 50\%$ ) are comparatively low in the data taken from Gaia ( $< 5$  objects) than Hipparcos ( $> 65$  objects) within a distance of 1 kpc. In summary, we infer that Gaia data is more preferable than the corresponding Hipparcos data. The second data release, in 2018, will publish the final Gaia catalogue which will be a census of our Galaxy of very high precision and more information about various astronomical objects.

## Acknowledgements

We thank Dr. Paul K.T for his great support and guidance through this work and we also thank Dr. Blesson Mathew who provided insight and expertise that greatly assisted the project. We specifically acknowledge the resource softwares VizieR, X-match, Python and Overleaf.

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