

Measurements and Comparisons of Diffuse Interstellar Bands around the MBM12 Interstellar Cloud

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MBM12 Interstellar Cloud

- Studied 13 stars around
- Made of dust and gas
- $\sim 100 \times 50$ pc

(Figure provided by Adolf Witt)

1230-912
(not in photo)
0.13

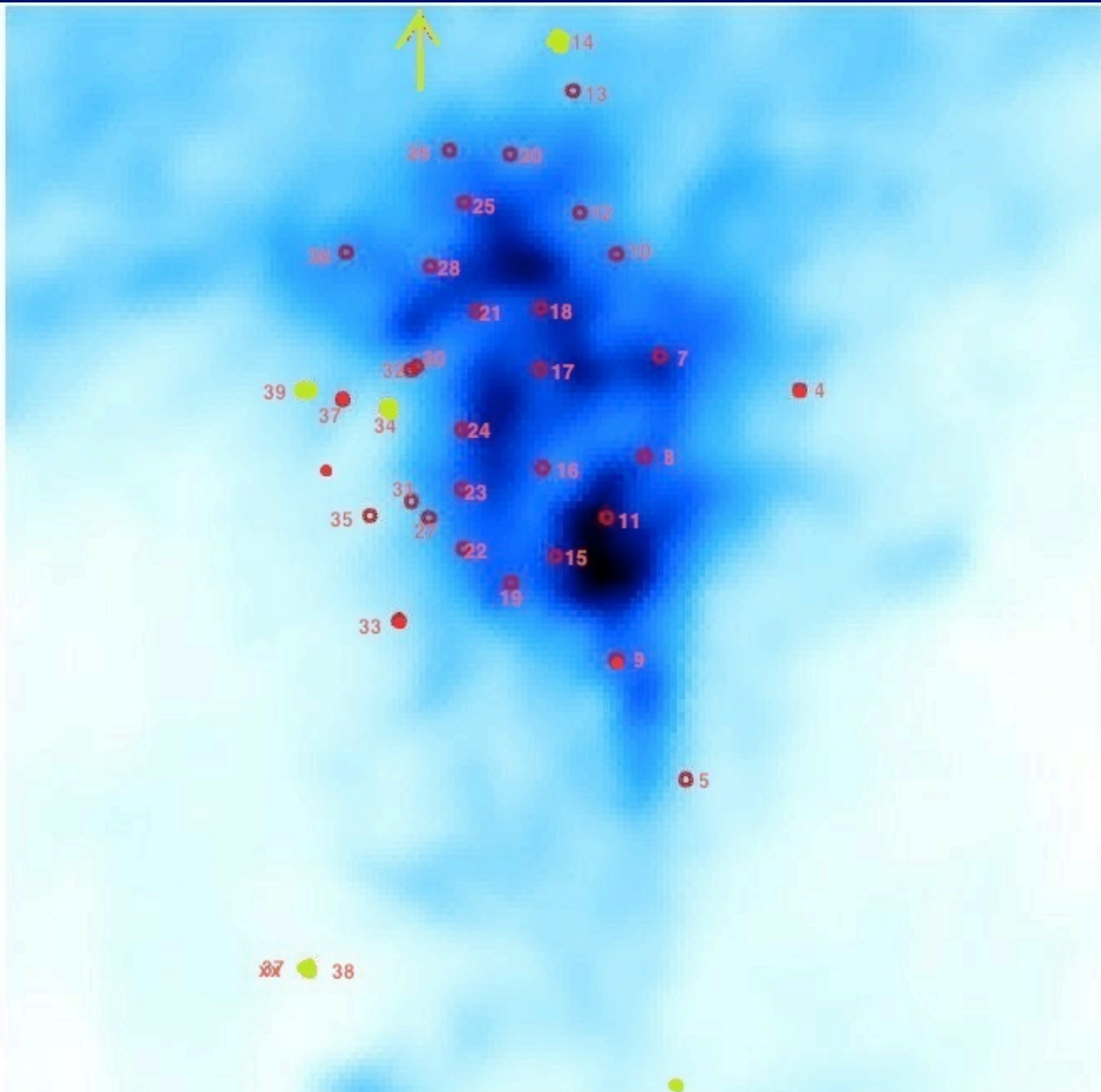
0.08 34. 1230-853
0.93 37. 1230-801
0.23 39. 1230-1010

0.39 1227-449

0.25 33. 1227-494

38. 1227-1178
0.08

0.01 1227-916



14. 1230-302 0.03

4. 1230-966 0.41

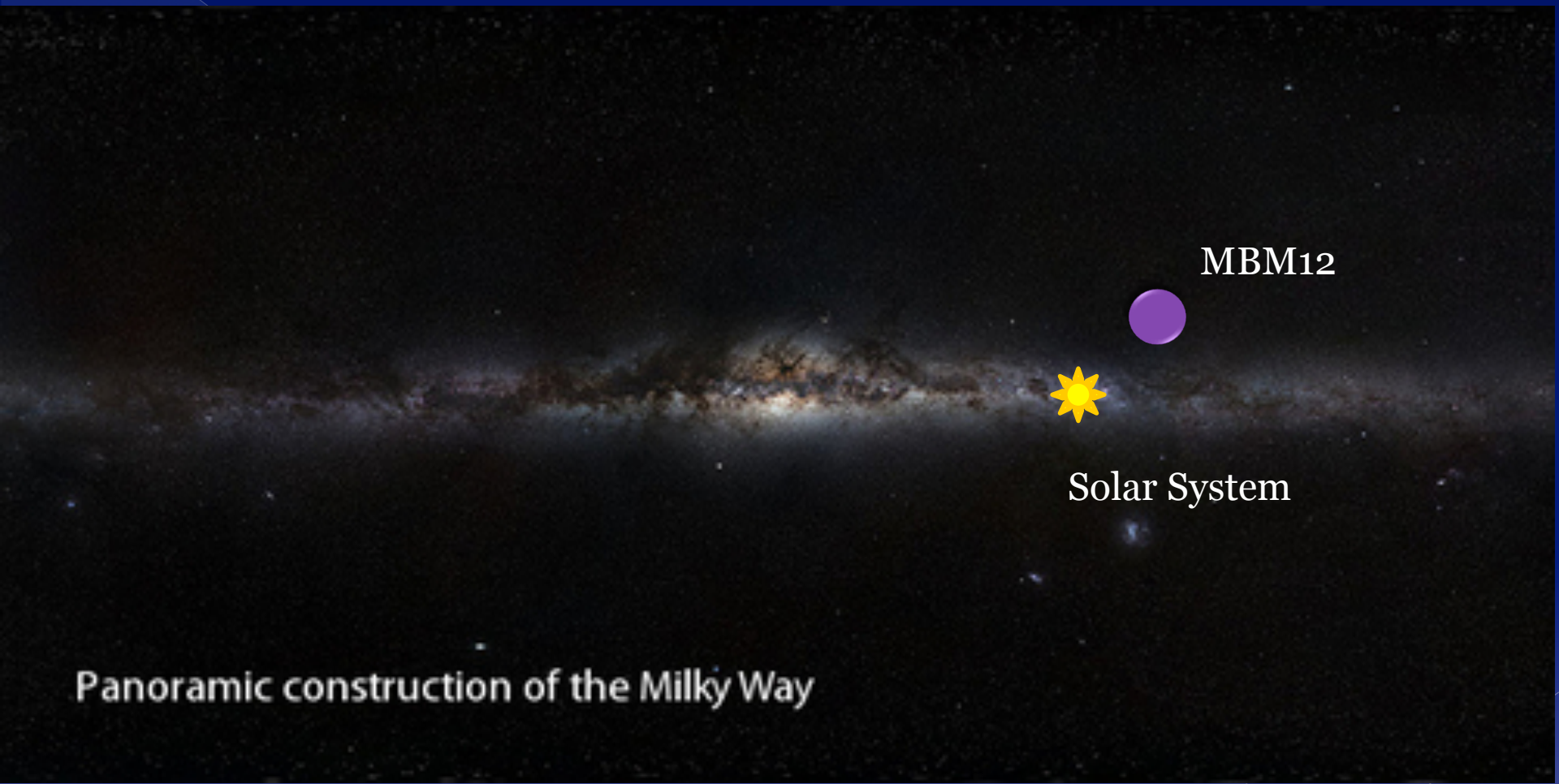
9. 1227-883 0.11

1226-1803
0.11

1226-1828
0.15

● foreground
● background
E(B-V)

(Figure provided by Adolf Witt)



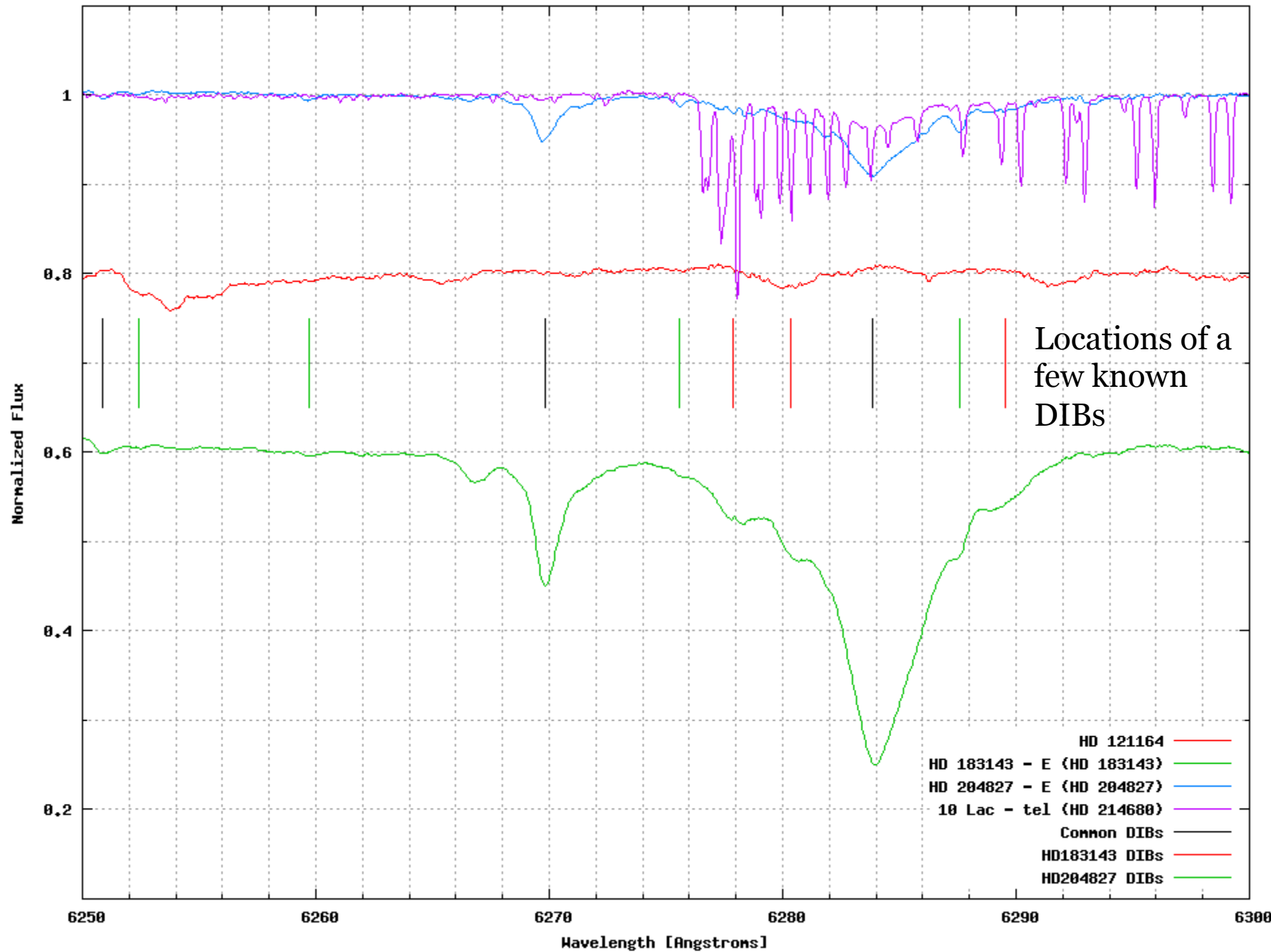
Panoramic construction of the Milky Way

Diffuse Interstellar Bands

- ⊙ Absorption in a star's spectrum
- ⊙ Caused by materials between the star and the viewing point
- ⊙ Not fully understood
 - > Found in most diffuse gas
- ⊙ Do not correlate with each other
- ⊙ Profiles have the same shape in different stars, despite having different strength
 - > Very broad compared to atomic lines
- ⊙ Slightly more than 600 found since discovery in 1922

Eyeball Estimates

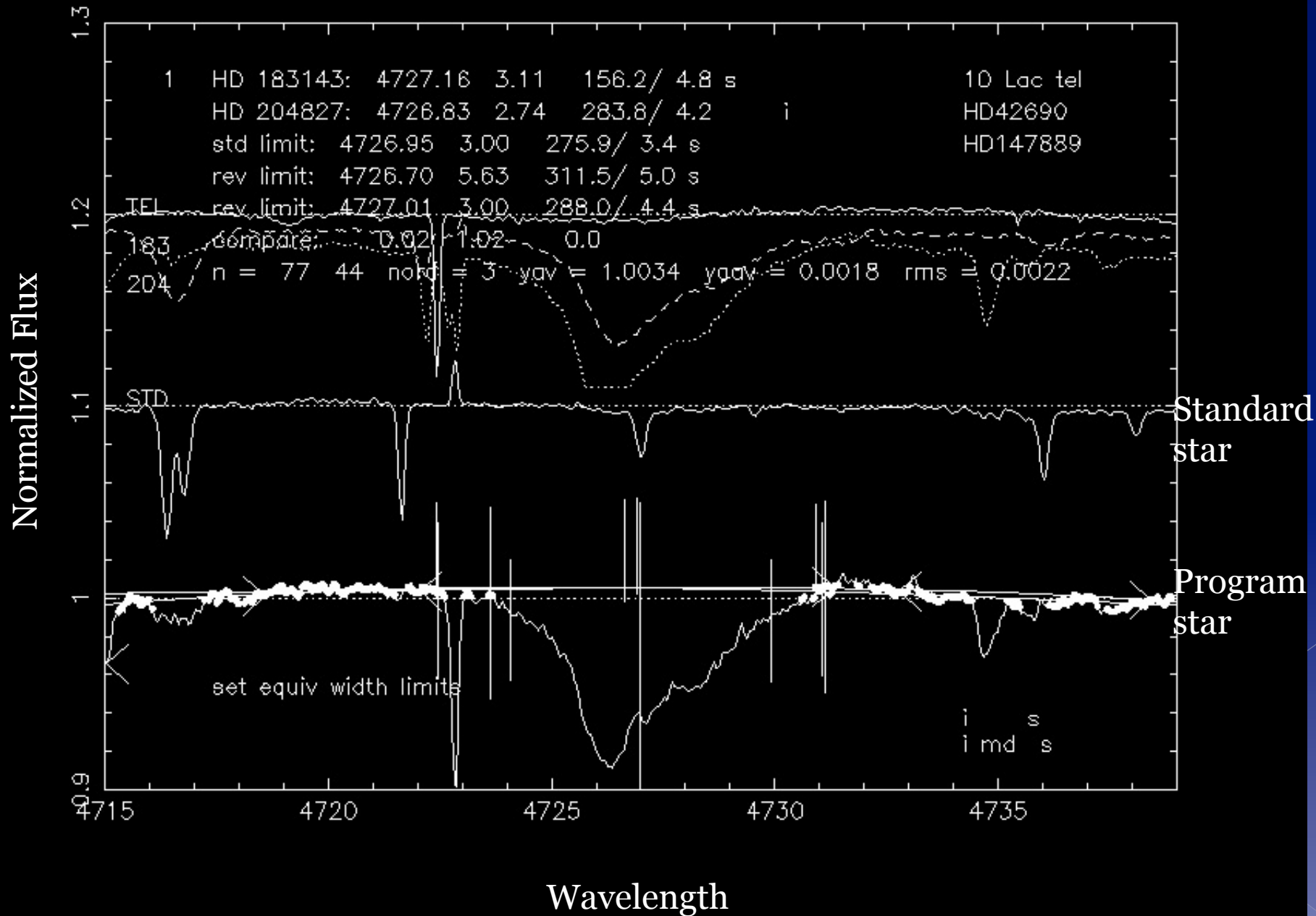
- ⊙ Plot spectrum of standard star side by side with program star
- ⊙ Telluric lines are eliminated by website
- ⊙ Velocity shift had to be accounted for so that the spectra would line up
- ⊙ Look for DIBs
 - > No absorption in standard star, DIB absorption in program star
 - > Standard star has similar spectral type and less reddening
 - > Create a list with all the suspected DIBs for the star



arcecam

- ⦿ Presents the inputted standard star and program star, and reference telluric star, two reference DIB stars
- ⦿ Measure continuum
- ⦿ Can flag measurements for stellar lines and contamination by other DIBs, etc

Example of arcexam presentation at 4725 angstroms



Foreground by strength (mÅ)

Star	1227-1178	1230-1010	1226-1828	1230-912	1230-853	1227-916	1230-302
Spectral Type	A0V (fore/off)	A9V (fore/on)	A7V (fore/off)	A3III (fore/on)	A7V (fore/on)	A9V (fore/off)	A8III (fore/on)
6284		169.1 (10.5)	46.4 (3.6)	183.3 (14.1)	81.7 (7.4)		
5780		241.6 (4.4)	57.3 (3.8)	93.9 (2.4)	69.0 (2.9)		54.5 (2.7)
6613		94.2 (2.8)	13.7 (1.8)	41.1 (0.6)	25.9 (1.4)		
6203	6.8 (0.5)	25.6 (2.0)	10.3 (0.9)	23.7 (1.6)	15.6 (1.8)		<3.7
6269		35.5 (1.2)		14.2 (0.9)	12.7 (1.8)	12.2 (1.5)	
7224			15.3 (1.9)	22.5 (0.8)	27.6 (1.5)		
6196	1.6 (0.2)	23.1 (0.9)		8.9 (0.2)	4.9 (1.0)		
5797			14.0 (1.6)	25.5 (0.6)			

Background by strength (mÅ)

Star	1226-1803	1230-966	1230-801	1227-494	1227-883	1227-449
Spectral Type	A3V (back/off)	A5V (back/on)	A0V (back/on)	A8V (back/on)	A9V (back/on)	A2III (back/on)
6284	194.7 (13.1)	234.3 (14.7)		248.2 (18.6)		
5780	120.4 (5.0)	129.6 (4.9)			86.6 (5.0)	
6613	97.7 (3.0)	61.6 (1.8)		87.4 (2.6)		
6203	25.5 (3.6)	32.6 (3.2)		32.1 (4.1)		
7224	39.4 (3.8)	42.2 (1.1)	308.0 (14.1)	78.2 (5.0)		
6269	15.3 (2.2)	12.1 (1.7)		69.0 (4.9)	19.1 (1.3)	
6196	18.5 (1.3)		12.0 (1.5)	16.7 (1.4)		
5797	39.3 (2.9)			35.6 (2.5)		

1230-912 (not in photo) 183.3
0.13

6284 Å

● foreground
● background
E(B-V)

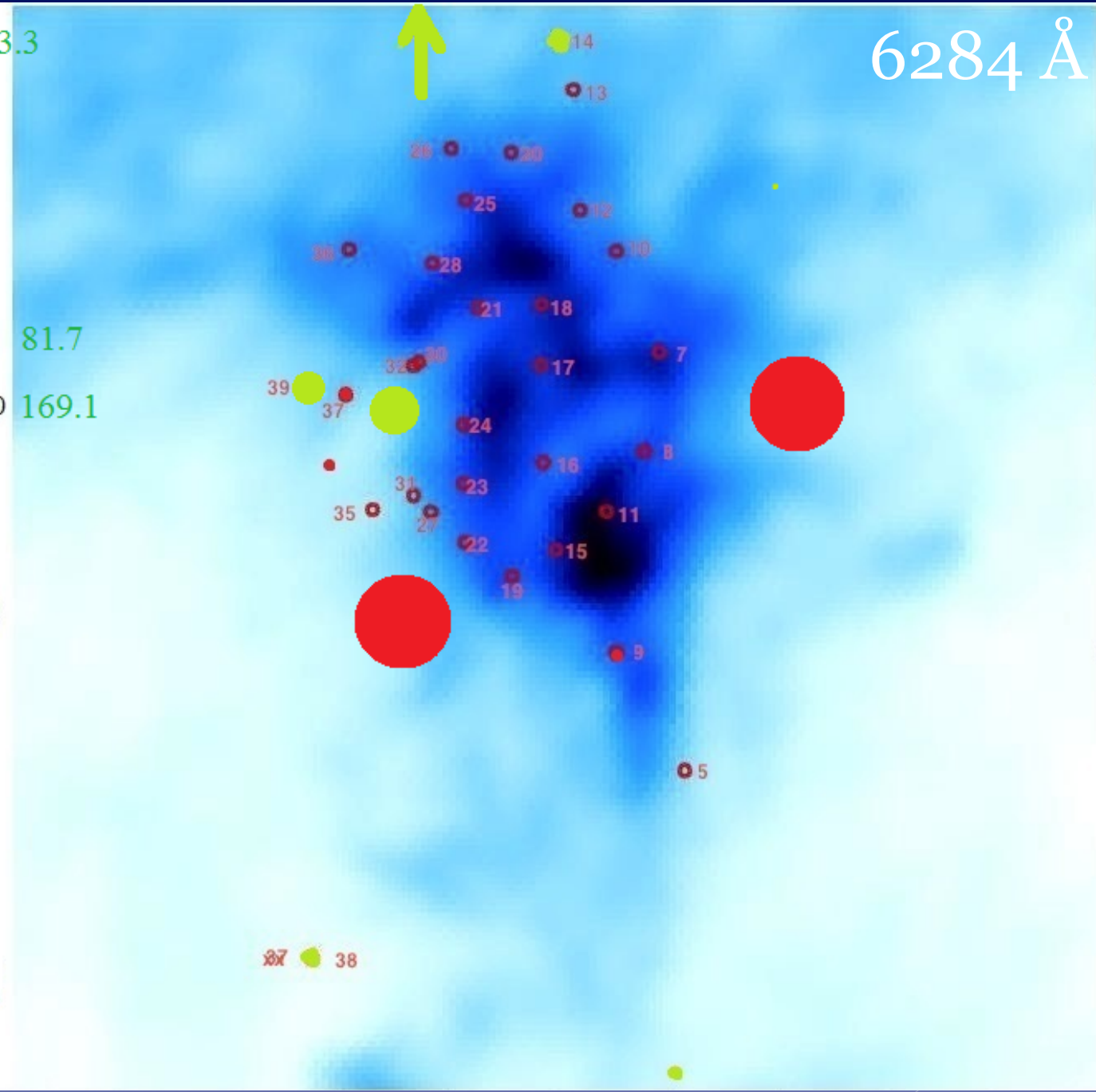
0.08 34. 1230-853 81.7
0.23 39. 1230-1010 169.1

4. 1230-966 0.41
234.3

0.25 33. 1227-494 248.2

● 1226-1803 0.11 194.7

● 1226-1828 0.15 46.4



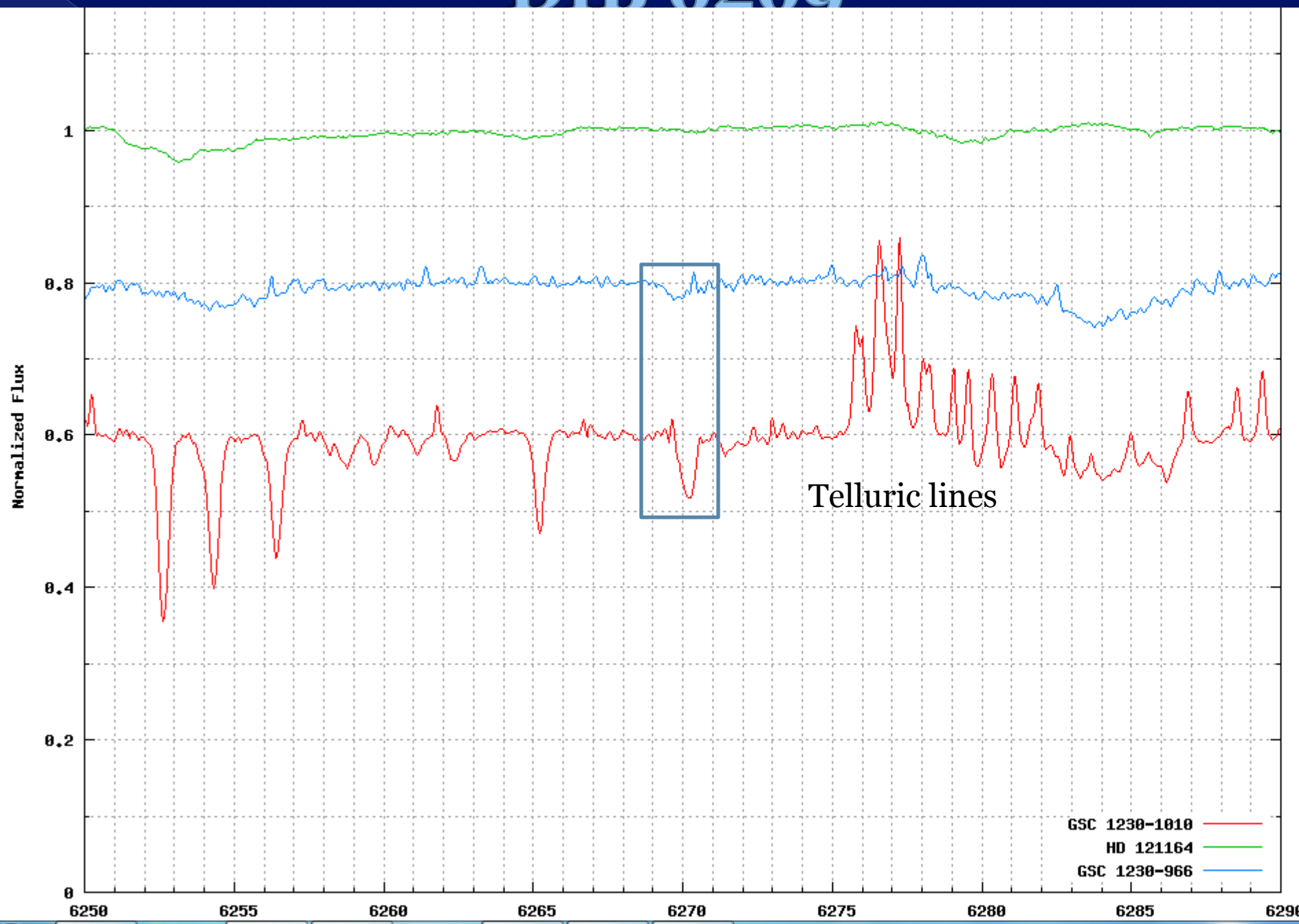
● 37 ● 38

Conclusions

- ◎ The strength of the DIBs doubled in the background stars
- ◎ 1230-1010, has DIB strength of a background star, but the low reddening of a foreground star
- ◎ Stars in center of cloud are too faint, so more observations with a bigger telescope are needed
 - > Project proves that this is possible and worth looking into
- ◎ DIB at 7224 \AA is prominent in the cloud
 - > 1230-853 (foreground) is 27.6
 - > 1230-801 (background) is 308.0



DIB 6269



Remaining Questions

- ⦿ What is the distance of 1230-1010?
- ⦿ What elements does the cloud consist of?
- ⦿ Stars in the heart of the cloud?
 - > Their DIB strengths
- ⦿ What causes DIBs?
 - > They do not correlate strongly to dust or atoms

Acknowledgements

- ◎ SIR Program and staff
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- ◎ DIB Database

Questions?

Selected Bibliography

- Friedman, S., Hobbs, L., McCall, B., Oka, T., Rachford, B., Sherman, R., . . . York, D. (Comps.). (n.d.). *DIB Database* [A source where DIB data and star spectra is stored and is able to be plotted.]. University of Chicago, Chicago, IL.
- Hearty T. (2001). MBM12: The Nearest Known Star-Forming Cloud. ASP Conference Series, 244.
- Kim, M. J., Kim, S., Youn, S., Yun, M. S., Wilson, G. W., Aretxaga, I., . . . Kang, Y. W. (2012). AzTEC 1.1 mm OBSERVATIONS OF THE MBM12 MOLECULAR CLOUD. *The Astrophysical Journal*, 746(11).
- Kaźmierczak, M., Schmidt, M., Weselak, T., Galazutdinov, G., & Krełowski, J. (2013). C₂ and Diffuse Interstellar Bands. *Proc. IAU Proceedings of the International Astronomical Union*, 9(S297), 121-124.
- Sarre P. (2014). Diffuse Interstellar Band Profiles. Proc. IAU Proceedings of the International Astronomical Union.
- Thorburn, J. A., Hobbs, L. M., Mccall, B. J., Oka, T., Welty, D. E., Friedman, S. D., . . . York, D. G. (2003). Some Diffuse Interstellar Bands Related to Interstellar C₂ Molecules. *ApJ The Astrophysical Journal*, 584(1), 339-356.
- York, D. G. (2015, September 30). Types of Lines [Personal interview].