

Dartmouth College Dartmouth Digital Commons

Open Dartmouth: Faculty Open Access Articles

6-20-2012

Analysis of the Early-Time Optical Spectra of SN 2011fe in M101

J. T. Parrent

Dartmouth College

D. A. Howell

University of California, Santa Barbara

B. Friesen

University of Oklahoma

R. C. Thomas

Lawrence Berkeley National Laboratory

Follow this and additional works at: <https://digitalcommons.dartmouth.edu/facoa>



Part of the [Stars, Interstellar Medium and the Galaxy Commons](#)

Recommended Citation

Parrent, J. T.; Howell, D. A.; Friesen, B.; and Thomas, R. C., "Analysis of the Early-Time Optical Spectra of SN 2011fe in M101" (2012). *Open Dartmouth: Faculty Open Access Articles*. 1790.

<https://digitalcommons.dartmouth.edu/facoa/1790>

This Article is brought to you for free and open access by Dartmouth Digital Commons. It has been accepted for inclusion in Open Dartmouth: Faculty Open Access Articles by an authorized administrator of Dartmouth Digital Commons. For more information, please contact dartmouthdigitalcommons@groups.dartmouth.edu.

ANALYSIS OF THE EARLY-TIME OPTICAL SPECTRA OF SN 2011fe IN M101

J. T. PARRENT^{1,2}, D. A. HOWELL^{2,3}, B. FRIESEN⁴, R. C. THOMAS⁵, R. A. FESEN¹, D. MILISAVLJEVIC⁶, F. B. BIANCO^{3,2},
B. DILDAY^{2,3}, P. NUGENT^{5,7}, E. BARON^{4,8,9}, I. ARCAVI¹⁰, S. BEN-AMI¹⁰, D. BERSIER¹¹, L. BILDSTEN³, J. BLOOM⁷, Y. CAO¹²,
S. B. CENKO⁷, A. V. FILIPPENKO⁷, A. GAL-YAM¹⁰, M. M. KASLIWAL¹³, N. KONIDARIS¹², S. R. KULKARNI¹², N. M. LAW¹⁴,
D. LEVITAN¹², K. MAGUIRE¹⁵, P. A. MAZZALI^{16,17}, E. O. OFEK¹⁰, Y. PAN¹⁵, D. POLISHOOK¹⁸, D. POZNANSKI¹⁹, R. M. QUIMBY²⁰,
J. M. SILVERMAN⁷, A. STERNBERG¹⁶, M. SULLIVAN¹⁵, E. S. WALKER²¹, DONG XU¹⁰, C. BUTON²², AND R. PEREIRA²³

¹ 6127 Wilder Lab, Department of Physics & Astronomy, Dartmouth College, Hanover, NH 03755, USA

² Las Cumbres Observatory Global Telescope Network, Goleta, CA 93117, USA

³ Department of Physics, University of California, Santa Barbara, Santa Barbara, CA 93117, USA

⁴ Homer L. Dodge Department of Physics and Astronomy, University of Oklahoma, 440 W Brooks, Norman, OK 73019, USA

⁵ Physics Division, Lawrence Berkeley National Laboratory, 1 Cyclotron Road, Berkeley, CA 94720, USA

⁶ Harvard-Smithsonian Center for Astrophysics, 60 Garden St., Cambridge, MA 02138, USA

⁷ Department of Astronomy, University of California, Berkeley, CA 94720, USA

⁸ Hamburger Sternwarte, Gojenbergsweg 112, D-21029 Hamburg, Germany

⁹ Physics Department, University of California, Berkeley, CA 94720 USA

¹⁰ Department of Particle Physics and Astrophysics, The Weizmann Institute of Science, Rehovot 76100, Israel

¹¹ Astrophysics Research Institute, Liverpool John Moores University, Birkenhead CH41 1LD, UK

¹² Cahill Center for Astronomy and Astrophysics, California Institute of Technology, 1200 E California Blvd., Pasadena, CA 91125, USA

¹³ Observatories of the Carnegie Institution for Science, 813 Santa Barbara Street, Pasadena, CA 91101 USA

¹⁴ Dunlap Institute for Astronomy and Astrophysics, University of Toronto, 50 St. George Street, Toronto M5S 3H4, Ontario, Canada

¹⁵ Department of Physics (Astrophysics), University of Oxford, Keble Road, Oxford OX1 3RH, UK

¹⁶ Max-Planck-Institut für Astrophysik, Karl-Schwarzschild-Strasse 1, D-85748 Garching, Germany

¹⁷ INAF-Osservatorio Astronomico, vicolo dell'Osseatorio, 5, I-35122 Padova, Italy

¹⁸ Department of Earth, Atmospheric, and Planetary Sciences, Massachusetts Institute of Technology, Cambridge, MA 02139, USA

¹⁹ School of Physics and Astronomy, Tel-Aviv University, Tel-Aviv 69978, Israel

²⁰ Kavli Institute for the Physics and Mathematics of the Universe, University of Tokyo, Kashiwa, Japan 277-8583

²¹ Scuola Normale Superiore, Piazza dei Cavalieri 7, 56126 Pisa, Italy

²² Physikalisches Institut Universität Bonn (Bonn), Nussallee 12, D-53115 Bonn, Germany

²³ Université de Lyon, F-69622 France; Université de Lyon 1; CNRS/IN2P3, Institut de Physique Nucléaire de Lyon, France

Received 2012 April 18; accepted 2012 May 15; published 2012 May 31

ABSTRACT

The nearby Type Ia supernova (SN Ia) SN 2011fe in M101 ($cz = 241 \text{ km s}^{-1}$) provides a unique opportunity to study the early evolution of a “normal” SN Ia, its compositional structure, and its elusive progenitor system. We present 18 high signal-to-noise spectra of SN 2011fe during its first month beginning 1.2 days post-explosion and with an average cadence of 1.8 days. This gives a clear picture of how various line-forming species are distributed within the outer layers of the ejecta, including that of unburned material (C+O). We follow the evolution of C II absorption features until they diminish near maximum light, showing overlapping regions of burned and unburned material between ejection velocities of 10,000 and 16,000 km s^{-1} . This supports the notion that incomplete burning, in addition to progenitor scenarios, is a relevant source of spectroscopic diversity among SNe Ia. The observed evolution of the highly Doppler-shifted O I $\lambda 7774$ absorption features detected within 5 days post-explosion indicates the presence of O I with expansion velocities from 11,500 to 21,000 km s^{-1} . The fact that some O I is present above C II suggests that SN 2011fe may have had an appreciable amount of unburned oxygen within the outer layers of the ejecta.

Key words: supernovae: general – supernovae: individual (SN 2011fe)

Online-only material: color figures

1. INTRODUCTION

SN 2011fe in M101 is the closest (6.9 Mpc) Type Ia supernova (SN Ia) in 25 years. It was discovered only 11 hr after explosion, the earliest of any SN Ia (Nugent et al. 2011), and has revealed more about its progenitor than any other recent SN Ia. In the canonical picture of an SN Ia, a CO white dwarf (WD) gains mass until it approaches the Chandrasekhar mass, either by mass transfer from a non-degenerate companion, or by merging with another WD (for a review, see Howell 2011).

The merger of two WDs has been strongly suspected in the case of “super-Chandra” SNe Ia, those that are so luminous they appear to require a progenitor above the Chandrasekhar mass (Howell et al. 2006; Scalzo et al. 2010). Unlike most SN Ia, these subtypes are known for showing conspicuous and relatively long-lasting spectroscopic signatures of carbon in their outer

layers, possibly left over from the violent CO-WD merger (Silverman et al. 2011). But this scenario cannot explain all SNe Ia as some show narrow lines of sodium from circumstellar material (CSM), thought to be from a non-degenerate secondary star in a single-degenerate system (Dilday et al. 2012 and references therein). Sternberg et al. (2011) estimate that at least 20% of SNe Ia in spiral galaxies arise from this channel. Furthermore, it remains to be seen if circumstellar sodium is concurrent with carbon detections. Could abundance studies of carbon lines reveal the progenitor system?

With SN 2011fe, it was confirmed observationally for the first time that the exploding star was a WD (Nugent et al. 2011; Bloom et al. 2012). Pre-explosion *Hubble Space Telescope* images revealed that any secondary star could not be a luminous red giant (Li et al. 2011). Furthermore, the lack of an observed shock, expected when the ejecta hits any non-degenerate

Table 1
Summary of Observations

UT Date	MJD 55,000	Epoch ^a (days)	Phase ^b (days)	Telescope + Instrument	Exp. Time (s)	Range (Å)
2011 Aug 24	797.86	1.2	-16.4	LT+FRDOSpec	1800	3900-9000 ^c
2011 Aug 25	798.16	1.5	-16.1	Lick3-m+Kast	1200	3420-10290 ^c
2011 Aug 25	798.95	2.3	-15.3	TNG+DOLORES	300	3780-8070
2011 Aug 26	799.26	2.6	-15.0	UH88+SNIFS	300	3300-9700
2011 Aug 28	801.19	4.5	-13.1	P200+DBSP	120	3300-10000
2011 Aug 29	801.74	5.1	-12.5	GN+GMOS	120(×4)	3500-9700
2011 Aug 30	803.86	7.2	-10.4	WHT+ISIS	60(×2)	3500-9500
2011 Aug 31	804.23	7.5	-10.0	Keck1+LRIS	45	3050-10200
2011 Aug 31	804.94	8.3	-9.3	WHT+ISIS	60(×2)	3500-9300
2011 Sep 1	805.75	9.1	-8.5	GN+GMOS	60(×2)	3500-9700
2011 Sep 1	805.90	9.2	-8.4	WHT+ISIS	120	3500-9300
2011 Sep 3	807.13	10.4	-7.2	P200+DBSP	120	3680-9700
2011 Sep 8	812.23	15.5	-2.0	GN+GMOS	60(×2)	3500-9700
2011 Sep 9	813.23	16.5	-1.0	GN+GMOS	60(×2)	3500-9700
2011 Sep 12	816.22	19.5	+1.9	GN+GMOS	60(×2)	3500-9700
2011 Sep 13	817.27	20.6	+3.0	GN+GMOS	60(×2)	3500-9700
2011 Sep 17	821.27	24.6	+6.9	GN+GMOS	60(×2)	3500-9700
2011 Sep 25	829.27	32.6	+14.9	GN+GMOS	60(×2)	3500-9700

Notes.

^a Relative to an estimated explosion JD of $2,455,797.187 \pm 0.014$ (Nugent et al. 2011).

^b Relative to an estimated $M_{b,\max}$ JD of $2,455,814.28 \pm 0.02$ (F. Bianco et al. 2012, in preparation).

^c Spectra presented in Nugent et al. (2011).

secondary (Kasen 2010), suggests that the secondary could not be a red giant (Nugent et al. 2011), and probably not any non-degenerate star (Bloom et al. 2012). Nondetections of radio and X-ray emission place strong limits on the amount of CSM, disfavoring most single-degenerate hypotheses, but consistent with the merger of two WD stars (Horesh et al. 2012; Chomiuk et al. 2012; Margutti et al. 2012).

In addition, thorough and complete spectroscopic observations of SN 2011fe offer a chance to study the distribution of ejected material. For instance, as many as a third of SNe Ia show evidence of unburned material if spectra are obtained early enough (Parrent et al. 2011; Thomas et al. 2011a; Folatelli et al. 2012; Silverman & Filippenko 2012; Blondin et al. 2012). Carbon features may not always, if at all, be the signature of a merger (Scalzo et al. 2010), but rather signal some difference in the explosion physics of the primary WD (Baron et al. 2003; Thomas et al. 2007). Abundance studies of unburned material can be coupled with constraints on material that is synthesized in the explosion, as some explosion models predict the complete burning of the WD, leaving no unburned carbon (Kasen et al. 2009; Maeda et al. 2010b).

To investigate these questions, it is necessary to combine our limited progenitor knowledge with a map of the elemental distribution of ejected material in SNe, including unburned material. Here we present 18 high signal-to-noise (S/N) optical spectra of SN 2011fe during its first month, starting 1 day after explosion. We fit all spectra with the automated spectrum synthesis code SYNAPPS²⁴ to trace the compositional structure of the ejected material.

2. OBSERVATIONS

On 2011 August 24.167 (UT dates are used throughout this Letter), the Palomar Transient Factory (PTF) discovered SN 2011fe in M101. The SN was discovered early enough

that the time of explosion, t_{exp} , could be estimated to within ~ 20 minutes of 2011 August 23.69 (Nugent et al. 2011). The SN reached a *B*-band peak brightness on September 10.3 (F. Bianco et al. 2012, in preparation). Optical spectra were obtained on eight different telescopes (Table 1; Figure 1). Reductions were done with IRAF²⁵.

3. SPECTROSCOPIC ANALYSIS

From the onset, the spectroscopic features of SN 2011fe appear broad and comprised of P-Cygni absorption/emission profiles of intermediate-mass elements (IMEs) and iron-peak elements (IPEs) typically seen in supernova spectra (Branch & Patchett 1973; Filippenko 1997). Over time, the features narrow, depicting an evolution similar to “normal” subtype objects. Most notably the spectra exhibit: (1) high velocity features (HVF; $v > 16,000 \text{ km s}^{-1}$) of Si II, Ca II and Fe II absorption lines from the onset of the explosion, (2) unburned material in the form of a C II $\lambda 6578$ absorption feature, and (3) O I $\lambda 7774$ at velocities higher than previously detected for an SN Ia (Nugent et al. 2011). From the lack of conspicuous Ti II/Fe III absorption features sometimes observed in sub- and over-luminous SNe Ia, respectively, as well as the slowly decreasing blueshift of the Si II 6150 Å feature, our spectroscopic observations of SN 2011fe in Figure 1 are consistent with characteristics of a “low velocity gradient”/“core normal” subtype (see Blondin et al. 2012 and references therein).

3.1. Line Identifications: SYNAPPS

Given that supernova expansion velocities are $\sim 10^4 \text{ km s}^{-1}$, Doppler line blending can be an obstacle for mapping the underlying compositional structure. SYNAPPS is a useful tool for identifying line features in photospheric phase spectra through

²⁴ Derived from SYNOW; see Branch et al. (2005) and Thomas et al. (2011b).

²⁵ IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation (NSF).

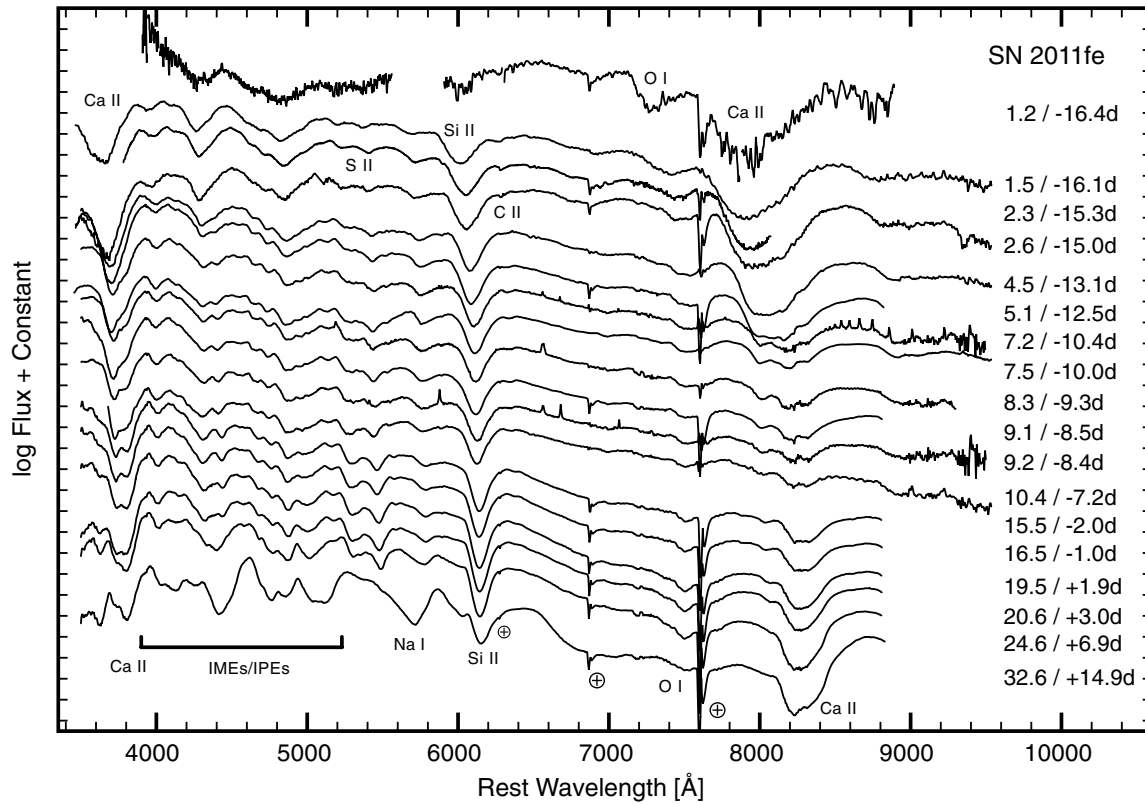


Figure 1. Optical spectra of SN 2011fe during the first month of observations, labeled with respect to an explosion date of 2011 August UT 23.69 and a *B*-band maximum light date of 2011 September UT 10.3. Most of the data span 3500–9700 Å, roughly every 2 days between the pre-maximum and post-maximum phases. Line identifications are labeled for some of the most conspicuous features. The \oplus symbol indicates the location of telluric features near 6279, 6880, and 7620 Å.

its empirical spectrum synthesis procedure which requires no a priori compositional structure as input. Under the assumption of pure-resonance line transfer and a sharp blackbody-emitting photosphere, an atomic line list is used to construct a basic spectrum from a pseudo-blackbody continuum level. The optical depth profile for a line is modeled as $\tau(v) = \tau_{\text{ref}} \times \exp(v_{\text{ref}} - v/v_e)$, where v is the expansion velocity, v_{ref} is an arbitrary normalization velocity, v_e is the e-folding velocity, and $\tau_{\text{ref}} = \tau(v_{\text{ref}})$ (Thomas et al. 2011b).

Because the sources of line formation influence nearly all regions of the spectrum from a wide range of depths within the ejecta, we model the spectrum approximately by treating the minimum reference velocity, v_{min} , for every ion as a free parameter. For a given set of ions in a fit, we set the photospheric velocity, v_{phot} , equal to the lowest measured v_{min} . This follows the notion that not all blended absorption minima coincide with the Doppler-shifted location of the photosphere over time (Bongard et al. 2008). Furthermore, we use a quadratic warping function to simplify the treatment of the continuum, calibration errors, and uncorrected reddening. Consequently, our fits are constrained by how we are able to best match absorption minima and the relative strengths of all features.

For the spectra in Figure 1, we converged a series of SYNAPPS fits. In Figure 2 we plot the spectra for the days 1.5, 10.4, 20.6 and 32.6 post-explosion, as they best document the relevant temporal changes on which we focus our attention. For each epoch, we compare our SYNAPPS fit and overplot the continuum fit for reference. We have loosely highlighted various regions of the fits to indicate which parts of the data we use to infer, and thereby constrain, the presence of particular ions. For example, the two broad absorption features centered around 4300 and 4800 Å are complex blends of three or more ions. When we

remove any one of these ions from the fit, we are unable to compensate for it with any reasonable perturbations of the remaining parameters. That is to say, while one or two ions may dominate it, the shape of a particular blended feature is still subject to the minor but observable influences of other species.

For the 1.5 day fit we set $v_{\text{phot}} = 16,000 \text{ km s}^{-1}$ and include photospheric components of C II, O I, Mg II, Si II, S II, Ca II, and Fe III. In addition, we include *higher* velocity components of O I, Si II, Ca II, and Fe II *above* 16,000 km s^{-1} in order to follow the evolution of the corresponding HVFs detected during the earliest observed epochs (Nugent et al. 2011). Overall, many of the features are accounted for in our subsequent fits, and the relative fluxes and respective minima match nicely.

Over the course of their detection, we model the Ca II HVFs as material “detached” 10,000 km s^{-1} above the photospheric line-forming regions. It is clear when the Ca II HVFs disappear for SN 2011fe at ~ 5 days after t_{expl} ; however, the same is not true for Si II. We model Si II as present in both photospheric and higher velocity regions, despite the separation of 2000 km s^{-1} near maximum light. It is sometimes clear that the 6150 Å feature requires two distinctly separate components of Si II for a best fit (most notably SN 2009ig; see Foley et al. 2012). Recently however, Smith et al. (2011) presented spectropolarimetry of SN 2011fe showing two Si II absorption line polarizations at ~ 7300 and $12,000 \text{ km s}^{-1}$ just after maximum light (see their Figure 3). During this phase our two components of Si II are at 8000 and $11,000 \text{ km s}^{-1}$. While the overlap is not exact, the finding of a two-component Si II feature is consistent with the spectropolarimetry.

The set of ions mentioned above is primarily the same used until maximum light, at which point Na I, S II, and Fe II begin to account best for most of the features. For SN 2011fe, while

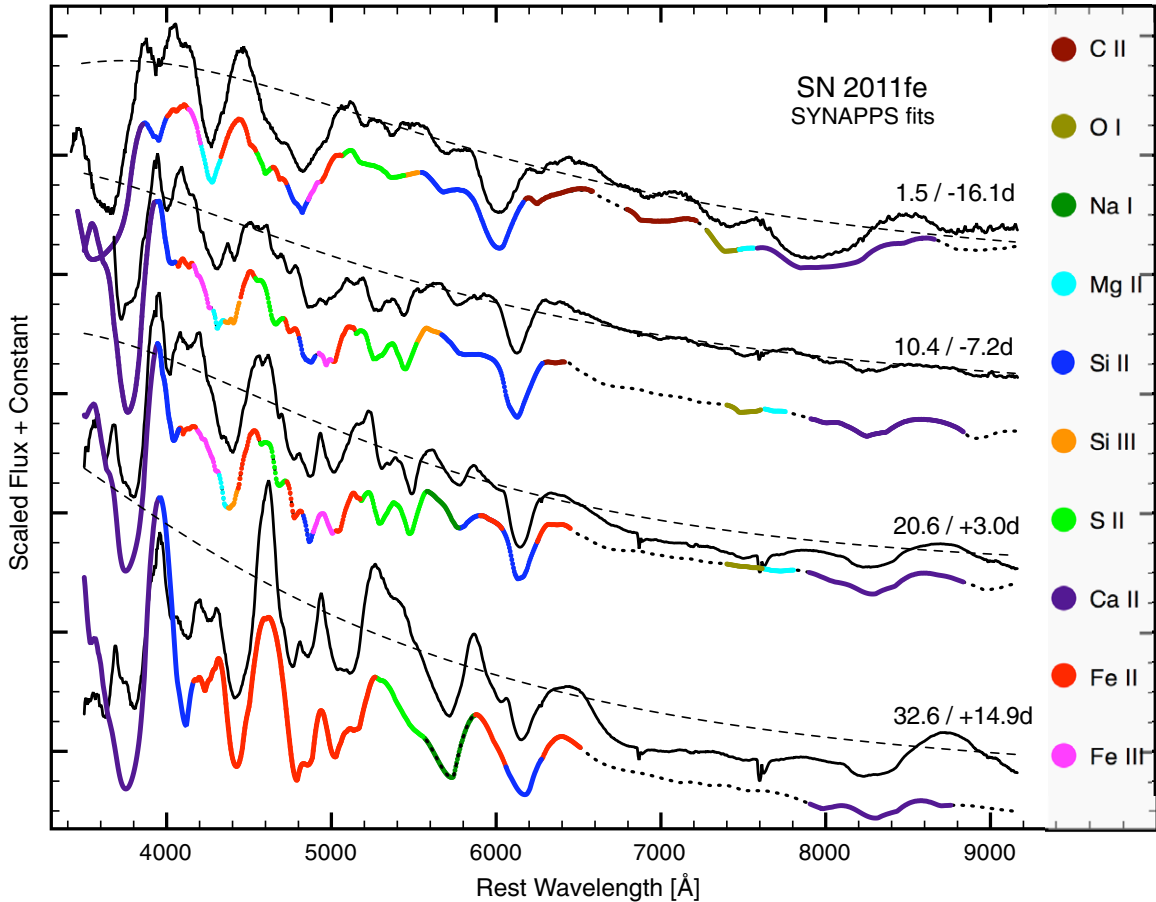


Figure 2. Synthetic spectra comparisons are shown for days 1.5, 10.4, 20.6, and 32.6 post-explosion (-16.1 , -7.2 , $+3.0$, and $+14.9$ days with respect to maximum light). The dashed line represents the fitted continuum level. We color various pieces of our fits to indicate the dominant ion we attribute to the features. The black dotted line of the fit indicates where we do not attribute features to any of the assumed ions.

(A color version of this figure is available in the online journal.)

using only SYNAPPS, we cannot definitively infer the presence of ions such as cobalt and nickel from the first month of optical observations alone, i.e., species responsible for powering SN Ia light curves (Colgate & McKee 1969). However, early-phase broadband photometry, as well as spectroscopic analyses of infrared data, could help in this regard (Marion et al. 2009; Piro 2012).

3.2. Spectroscopic Modeling: SYNAPPS

The result of any series of SYNAPPS fits is the estimation of velocities spanned by a particular set of ions. Therefore, we can use spectra in two different ways, as seen in Figure 3. From our time-series fits, in Figure 3(a) we show our $v_{\min}(t_{\text{obs}})$ estimates for each of the proposed ions in Section 3.1. A clear separation between outer and inner regions of ejected material can be seen. The material near photospheric velocities that span roughly 7000 – $16,000$ km s^{-1} includes Na I, C II, O I, Mg II, Si II, S II, Ca II, and Fe III, whereas most of the higher velocity material between $16,000$ and $30,000$ km s^{-1} is comprised of O I, Ca II, Si II, and Fe II.

Since we treat v_{\min} as a free parameter, there is some dispersion among ions near the modeled photosphere for a given epoch (black dashed line in Figure 3(a)). However, SYNAPPS fits are generally good to within 250 km s^{-1} and the scatter of our inferred v_{\min} is no less than 500 km s^{-1} . Therefore, we consider the dispersion consistent with what should be expected in an environment of multiple line-forming regions.

When we take into account the Doppler widths of our fits, we are able to place tighter constraints on the location of species by constructing a velocity-scaled map that traces how far each ion extends. In Figure 3(b) we plot a single normalized absorption profile for each ion, summed over the first month. Clearly, a photospheric and higher velocity region can be seen. In addition, Figure 3(b) yields some insight into where species are likely to be present, with respect to one another, throughout the ejecta. This is particularly important for investigating unburned material since its location with respect to burned material can be used to constrain the details of the explosion nucleosynthesis.

Figure 4 highlights regions of interest for tracing unburned material with C II $\lambda 6578$ and potentially O I $\lambda 7774$ (see also Figure 3). For our C II fits, we are able to follow the suspected absorption feature as it evolves across the emission component of the Si II $\lambda 6355$ line over time. This implies that unburned carbon is present at least between $10,000$ and $16,000$ km s^{-1} . Our O I fits in Figure 4 require both photospheric and higher velocity components of O I for the 1.2, 1.5, 2.6, and 4.5 day spectra as the O I HVF disappears quickly and within 5 days after t_{exp} . Within the time frame of our observations, C II $\lambda 6578$ is present from 1.5 to 20.6 days after t_{exp} , i.e., from the onset and until at least maximum light. The same is true for O I; however, we find that O I initially extends at least 5000 km s^{-1} above C II, within the outermost layers. This suggests the O I detected during the earliest epochs is likely to be unburned progenitor material rather than a product of carbon burning in the SN.

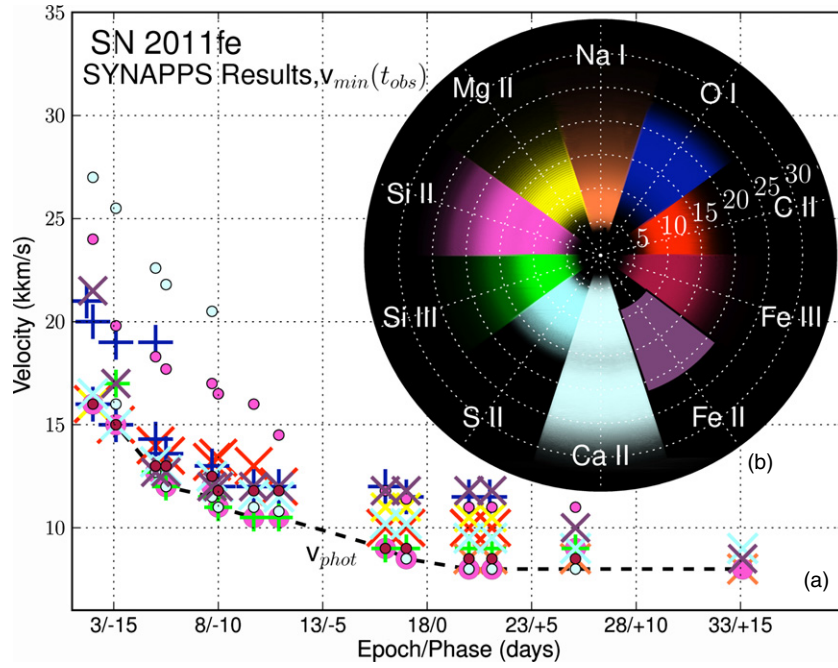


Figure 3. Symbols are for clarity within the plot and different ions are represented by different colors, as labeled in panel (b). (a) Based on our time series of synthetic SYNAPPS spectra, v_{min} estimates of each ion are plotted vs. time. The dashed line represents the location of our best modeled photosphere. (b) Here we reproduce panel (a), but include the Doppler widths of our fits. Color intensities are mapped to the normalized sum of a single absorption line component for each ion in every fit. Therefore, panel (b) only depicts where light has been scattered by a particular ion. Due to telluric features, we have only included our fitted Doppler widths for the high velocity O I component. The numbered rings (dashed-white) are in increments of 5 kkm s⁻¹. (A color version of this figure is available in the online journal.)

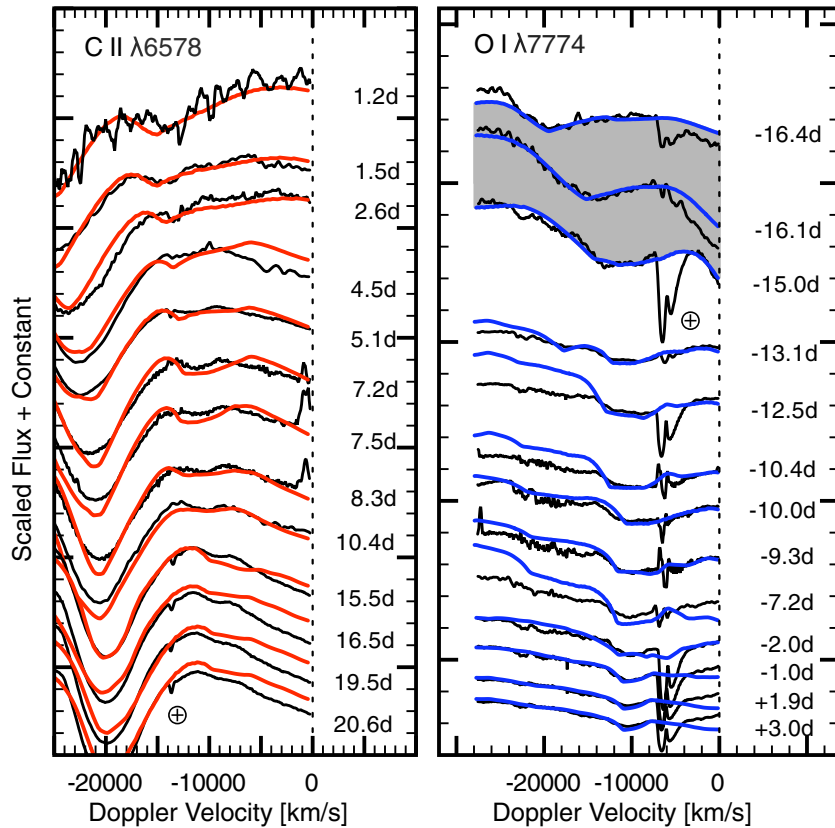


Figure 4. “Snap-shots” of unburned material, i.e., spectroscopic signatures of unburned material as best traced by C II λ6578 and possibly O I λ7774. On the left, we compare our SYNAPPS fits (red line) to the observations (black line) that show a feature near 6300 Å. At 1.2 days after t_{expl} the C II absorption minimum corresponds to a Doppler velocity of 16,000 km s⁻¹, at the upper reaches of photospheric velocities. On the right, we plot the same but instead for the O I λ7774 absorption feature. Our fits (blue line) place O I in both photospheric and higher velocity regions. The gray band highlights the early evolution of the O I feature and the ⊕ symbol indicates the location of telluric features.

(A color version of this figure is available in the online journal.)

4. DISCUSSION AND CONCLUSIONS

Thanks to the proximity of SN 2011fe, we have been able to obtain frequent high S/N spectra and use SYNAPPS to produce what may be the best-ever map of the material distribution in a “normal” SN Ia for the first month of its evolution. For SN 2011fe, it is important to note that, apart from differences in relative expansion velocities (Figure 3(a)), the spectra do not vary significantly from one epoch to another. Furthermore, as in most SNe Ia, in SN 2011fe we find O I, Mg II, Si II, S II, Ca II, and Fe II present from 16,000 km s⁻¹ or more down to about 7000 km s⁻¹ (later spectra will probe deeper). This implies similar compositional constituents over a large region of ejected material and is supported by our SYNAPPS fitting results (Figure 3(b)).

Stretching to nearly 30,000 km s⁻¹, high velocity regions of silicon and calcium are inferred from pre-maximum light spectra. A region of sodium is inferred to be present at high velocities as well, but is detected post-maximum light when the outer layers begin to cool. However, rather than large-scale asymmetries, as predicted by Maeda et al. (2010a), these high velocity features may trace only a small amount of material, as little as 10⁻⁷ M_⊙ in the case of Ca II (Hatano et al. 1999; Mazzali et al. 2005; Tanaka et al. 2008). Such a small inferred Ca mass supports the notion that strong features in SN spectra do *not* always correspond to large masses of line-forming ejecta.

We also find Si III and Fe III present at early times (before 3 days after maximum light), indicating that some parts of the ejecta achieved high temperatures, although the entire photosphere was not excessively hot. Even a day after explosion the spectra were not particularly blue, nor were they dominated by high-ionization species. Various scenarios that predict an early shock from interaction with a non-degenerate secondary star (Kasen 2010), a torus of material from a disrupted secondary (Scalzo et al. 2010), or CSM, are not supported by the data. This is consistent with a lack of radio and X-ray detections (Horesh et al. 2012; Chomiuk et al. 2012; Margutti et al. 2012), although the spectra here start a day earlier, providing additional constraints.

Regarding unburned material, other SNe Ia show distributions of unburned carbon spanning a range of velocities. This suggests that carbon in some SNe Ia arises from incomplete burning, in contrast to the predictions of some delayed detonation models and merger scenarios (Marion et al. 2009; Fryer et al. 2010; Roepke et al. 2012). In general, due to the rarity of early SN Ia observations, little is certain about the conditions under which carbon is observable in all subtypes. Noise makes weak carbon lines difficult to detect, especially when supernovae are faint at early times, and different velocities cause it to be blended with other features. Furthermore, SN Ia subtypes have different temperatures (Nugent et al. 1995), and these change with time, affecting whether carbon signatures exist as C I, C II, or C III. Thus, carbon may be fairly common in SNe Ia, but only apparent if spectra are obtained early (Parrent et al. 2011; Thomas et al. 2011a; Folatelli et al. 2012; Silverman & Filippenko 2012; Blondin et al. 2012). In addition, asymmetry may play a part in whether carbon is detected given that both explosive burning and WD mergers are inherently asymmetric processes (Livio & Pringle 2011; Pakmor et al. 2012). Spectropolarimetry and late-time observations, which are only possible for bright, nearby SNe like SN 2011fe, can shed light on these links.

For the carbon in SN 2011fe, we find C II to be strongest in the earliest spectra, suggesting it is mostly concentrated in the outer layers. The C II λ 6578 detection persists until just

before maximum light, although at later times it is very weak and would have been missed in any but the highest S/N spectra. Note, however, that its presence *throughout* regions of freshly synthesized material shows that it is indeed distributed within the ejecta and is not just concentrated within the outermost layers. Because the high velocity O I is located above the outermost reaches of the C II, it is likely unburned progenitor material (Nugent et al. 2011).

For the purposes of investigating spectroscopic diversity, it is convenient that SN 2011fe exhibits the characteristic signatures of a “normal” SN Ia. Future comparative studies of SNe Ia will benefit from these results and that of other well-observed objects. The map of elements in SN 2011fe can now be used as input for model atmosphere codes, such as PHOENIX (see Hauschildt & Baron 1999 and references therein), enabling us to better determine abundances, test theories of the explosion physics, and perhaps gain further insight into progenitor scenarios.

This work was supported by the Las Cumbres Observatory Global Telescope Network. This research used resources of the National Energy Research Scientific Computing Center, which is supported by the Office of Science of the U.S. Department of Energy under contract no. DE-AC02-05CH11231.

Observations were obtained as part of the Palomar Transient Factory project. The TNG is operated by the Fundacion Galileo Galilei—INAF (Istituto Nazionale di Astrofisica). The LT and WHT are operated by the Isaac Newton Group. The three telescopes are located in the Spanish Observatorio of the Roque de Los Muchachos of the Instituto de Astrofisica de Canarias. Some observations were obtained at the Gemini Observatory (GN-2011B-Q-215), which is operated by the Association of Universities for Research in Astronomy, Inc., under a cooperative agreement with the NSF on behalf of the Gemini partnership. Other observations were obtained with the 200 inch telescope at Palomar Observatory, the UH 2.2 m telescope, which is part of the Nearby Supernova Factory II project, and the Lick 3 m telescope, which is operated by the University of California. Some of the data presented herein were obtained at the W. M. Keck Observatory.

REFERENCES

- Baron, E., Lentz, E. J., & Hauschildt, P. H. 2003, *ApJ*, 588, L29
 Blondin, S., Matheson, T., Kirshner, R. P., et al. 2012, *AJ*, 143, 126
 Bloom, J. S., Kasen, D., Shen, K. J., et al. 2012, *ApJ*, 744, L17
 Bongard, S., Baron, E., Smadja, G., Branch, D., & Hauschildt, P. H. 2008, *ApJ*, 687, 456
 Branch, D., Baron, E., Hall, N., Melakayil, M., & Parrent, J. 2005, *PASP*, 117, 545
 Branch, D., & Patchett, B. 1973, *MNRAS*, 161, 71
 Chomiuk, L., Soderberg, A. M., Moe, M., et al. 2012, *ApJ*, 750, 164
 Colgate, S. A., & McKee, C. 1969, *ApJ*, 157, 623
 Dilday, B., et al. 2012, *Science*, submitted
 Filippenko, A. V. 1997, *ARA&A*, 35, 309
 Folatelli, G., Phillips, M. M., Morrell, N., et al. 2012, *ApJ*, 745, 74
 Foley, R. J., Challis, P. J., Filippenko, A. V., et al. 2012, *ApJ*, 744, 38
 Fryer, C. L., Ruiter, A. J., Belczynski, K., et al. 2010, *ApJ*, 725, 296
 Hatano, K., Branch, D., Fisher, A., Millard, J., & Baron, E. 1999, *ApJS*, 121, 233
 Hauschildt, P. H., & Baron, E. 1999, *J. Comput. Appl. Math.*, 109, 41
 Horesh, A., Kulkarni, S. R., Fox, D. B., et al. 2012, *ApJ*, 746, 21
 Howell, D. A. 2011, *Nature Commun.*, 2, 350
 Howell, D. A., Sullivan, M., Nugent, P. E., et al. 2006, *Nature*, 443, 308
 Kasen, D. 2010, *ApJ*, 708, 1025
 Kasen, D., Röpke, F. K., & Woosley, S. E. 2009, *Nature*, 460, 869
 Li, W., Bloom, J. S., Podsiadlowski, P., et al. 2011, *Nature*, 480, 348

- Livio, M., & Pringle, J. E. 2011, [ApJ](#), 740, L18
- Maeda, K., Benetti, S., Stritzinger, M., et al. 2010a, [Nature](#), 466, 82
- Maeda, K., Röpke, F. K., Fink, M., et al. 2010b, [ApJ](#), 712, 624
- Margutti, R., Soderberg, A. M., Chomiuk, L., et al. 2012, [ApJ](#), 751, 134
- Marion, G. H., Höflich, P., Gerardy, C. L., et al. 2009, [AJ](#), 138, 727
- Mazzali, P. A., Benetti, S., Altavilla, G., et al. 2005, [ApJ](#), 623, L37
- Nugent, P., Phillips, M., Baron, E., Branch, D., & Hauschildt, P. 1995, [ApJ](#), 455, L147
- Nugent, P. E., Sullivan, M., Cenko, S. B., et al. 2011, [Nature](#), 480, 344
- Pakmor, R., Kromer, M., Taubenberger, S., et al. 2012, [ApJ](#), 747, L10
- Parrent, J. T., Thomas, R. C., Fesen, R. A., et al. 2011, [ApJ](#), 732, 30
- Piro, A. L. 2012, arXiv:1201.5398
- Roepke, F. K., Kromer, M., Seitenzahl, I. R., et al. 2012, [ApJ](#), 750, L19
- Scalzo, R. A., Aldering, G., Antilogus, P., et al. 2010, [ApJ](#), 713, 1073
- Silverman, J. M., & Filippenko, A. V. 2012, arXiv:1202.3788
- Silverman, J. M., Ganeshalingam, M., Li, W., et al. 2011, [MNRAS](#), 410, 585
- Smith, P. S., Williams, G. G., Smith, N., et al. 2011, arXiv:1111.6626
- Sternberg, A., Gal-Yam, A., Simon, J. D., et al. 2011, [Science](#), 333, 856
- Tanaka, M., Mazzali, P. A., Benetti, S., et al. 2008, [ApJ](#), 677, 448
- Thomas, R. C., Aldering, G., Antilogus, P., et al. 2007, [ApJ](#), 654, L53
- Thomas, R. C., Aldering, G., Antilogus, P., et al. 2011a, [ApJ](#), 743, 27
- Thomas, R. C., Nugent, P. E., & Meza, J. C. 2011b, [PASP](#), 123, 237