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MMS Multi-Point Analysis of FTE Evolution: Physical Characteristics and Dynamics

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Key Points:

- The forces inside FTEs observed by MMS suggest plasma acceleration toward and along the FTE's central axis causing plasma to escape
- The roles of adiabatic expansion and reconnection in FTE growth are explored using MMS observations
- The observed sub-adiabatic decrease of plasma pressure as FTE size increases requires plasma heating mechanisms such as reconnection

Supporting Information:

- Supporting Information S1

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Abstract Previous studies have indicated that flux transfer events (FTEs) grow as they convect away from the reconnection site along the magnetopause. This increase in FTE diameter may occur via adiabatic expansion in response to decreasing external pressure away from the subsolar region or due to a continuous supply of magnetic flux and plasma to the FTEs' outer layers by magnetic reconnection. Here we investigate an ensemble of 55 FTEs at the subsolar magnetopause using Magnetospheric Multiscale (MMS) multi-point measurements. The FTEs are initially modeled as quasi-force-free flux ropes in order to infer their geometry and the spacecraft trajectory relative to their central axis. The MMS observations reveal a radially-inward net force at the outer layers of FTEs which can accelerate plasmas and fields toward the FTE's core region. Inside the FTEs, near the central axis, plasma density is found to decrease as the axial net force increases. It is interpreted that the axial net force accelerates plasmas along the axis in the region of compressing field lines. Statistical analysis of the MMS observations of the 55 FTEs indicates that plasma pressure, P_{th} , decreases with increasing FTE diameter, λ , as $P_{th,obs} \propto \lambda^{-0.24}$. Assuming that all 55 FTEs started out with similar diameters, this rate of plasma pressure decrease with increasing FTE diameter is at least an order of magnitude slower than the theoretical rate for adiabatic expansion (i.e., $P_{th,adiab.} \propto \lambda^{-3.3}$), suggesting the presence of efficient plasma heating mechanisms, such as magnetic reconnection, to facilitate FTE growth.

1. Introduction

Magnetic reconnection is the process by which topological rearrangement of field lines relieves stress in magnetized plasmas (e.g., Biskamp, 1996; Gonzalez & Parker, 2016; Yamada et al., 2010). Magnetic flux ropes are a crucial component of magnetic reconnection (Daughton et al., 2006; Drake et al., 2006). At the Earth's magnetopause, they take the form of FTEs (Eastwood et al., 2007; Lee & Fu, 1985; Russell & Elphic, 1978; Scholer, 1988). Theory suggests that FTEs can begin as small (electron- or ion-scale) flux ropes generated inside the magnetopause current layer as a result of multiple X-line reconnection (e.g., Lee & Fu, 1985; Deng et al., 2004; Daughton et al., 2006; Chen et al., 2008; Daughton et al., 2011b).

Kinetic simulations and in situ observations have suggested that, upon generation, flux ropes in three dimensions evolve in response to their environment moving tailward with the magnetosheath flow and expanding as the external pressure relaxes toward solar wind values (e.g., Berchem & Russell, 1984; Rijnbeek et al., 1984; Zhao et al., 2016; Farrugia et al., 2016; Teh, Nakamura, et al., 2017) and coalescence (e.g., Finn & Kaw, 1977; Karimabadi et al., 2011; Pritchett, 1992; Song et al., 2012; Wang et al., 2017; Zhou et al., 2014).

Global magnetohydrodynamic (MHD), kinetic, and Vlasov simulations indicate that subsolar FTEs tend to grow larger as they travel away from the reconnection site along the magnetopause (Akhavan-Tafti et al., 2018; Dorelli & Bhattacharjee, 2009; Hoilijoki et al., 2017; Hoilijoki et al., 2019; Omidi et al., 2006; Omidi & Sibeck, 2007; Raeder, 2006). The growth slows down when the FTEs approach their lowest-energy state (i.e., "force-free"; Eastwood et al., 2012). The present study seeks to distinguish between FTE "expansion" and "growth." FTE expansion is a type of FTE growth (i.e., increase in diameter) in which magnetic flux remains unchanged throughout the process. Figure 1 illustrates three potential FTE growth mechanisms: (i) adiabatic expansion, (ii) continuous reconnection, and (iii) coalescence. The concentric circles show magnetic islands which are two-dimensional projections of flux ropes. The shading of the islands determines the thermal pressure p inside them with the darkest shades indicating the region of largest p . The number of circles per area indicates magnetic field magnitude and the gradient shading inside the islands shows plasma

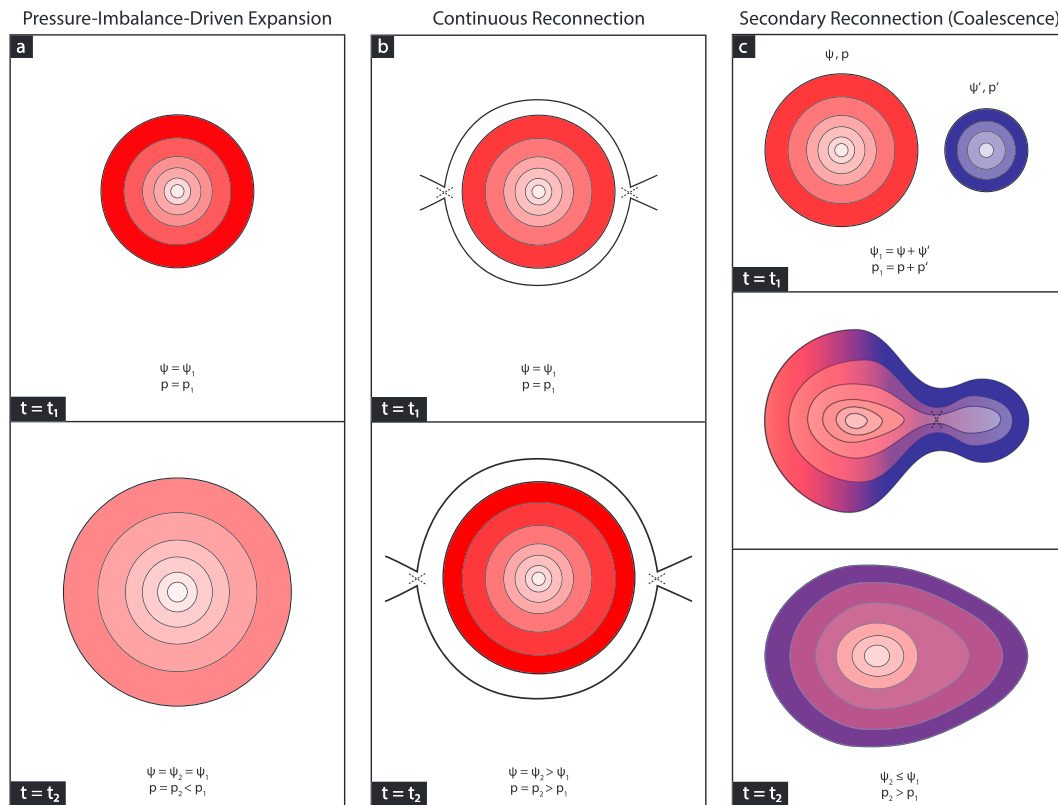


Figure 1. Three candidate FTE growth mechanisms: (a) adiabatic expansion, (b) continuous reconnection, and (c) coalescence. The concentric circles show magnetic islands which are two-dimensional projections of flux ropes. The shading of the islands determines the thermal pressure p inside them with the darkest shades indicating the region of largest p . The number of circles per area indicates magnetic field magnitude and the gradient shading inside the islands shows plasma pressure variability across the structure. The number of black circles represents the islands' magnetic flux content ψ . The X markers indicate X points at which magnetic field lines reconnect. The subscripts 1 and 2 indicate earlier and later evolution stages, respectively.

pressure variability across the structure. The number of black circles represents the islands' magnetic flux content ψ . The X markers indicate X points at which magnetic field lines reconnect. Due to reconnection, plasma is often heated resulting in an enhanced thermal pressure, that is, darker shades, and increased magnetic flux content, that is, more circles. Evolution of the islands is represented with increasing time, $t_1 > t_2$, where subscripts 1 and 2 indicate earlier and later evolution stages, respectively.

In the first scenario, in Figure 1a, FTEs grow due to pressure imbalance across the FTE boundary. As the FTE is carried away from the subsolar region, the external magnetosheath pressure will reduce resulting in the temporal expansion of the FTE. Inside these expanding FTEs, plasma pressure will drop but the magnetic flux content remains constant with time ($\psi = BA$, where ψ and A denote the magnetic flux content and the cross-sectional area of an FTE, respectively).

In the other two scenarios in Figure 1, FTEs also grow via reconnection. In one scenario, see Figure 1b, multiple X-line reconnection may continue after the initial formation of the magnetic island (Lee & Fu, 1985, 1986). Multiple X-line reconnection may or may not be a steady process (Phan et al., 2007). At $t = t_1$, the reconnected magnetic field lines form magnetic islands. But, due to the ongoing magnetic reconnection (Hoilijoki et al., 2019; Raeder, 2006), the magnetic islands grow to a size much larger than the original current sheet thickness. During this growth process, plasma is heated due to reconnection (i.e., $p_2 > p_1$). Similarly, ongoing reconnection contributes additional magnetic flux to the original magnetic island (i.e., $\psi_2 > \psi_1$).

In the other scenario, depicted in Figure 1c, the magnetic islands grow via coalescence. Chains of ion-scale FTEs (Milan et al., 2000; Slavin et al., 2012; Dong et al., 2017; Teh, Denton, et al., 2017; Wang et al., 2017; Hwang et al., 2018) are formed at the dayside magnetopause due to multiple X-line reconnection (Lee & Fu, 1985). In this scenario, the neighboring FTEs are forced to merge, reconnect, and grow into larger

structures, a process called “coalescence” (Biskamp, 1982; Finn & Kaw, 1977; Pritchett & Wu, 1979). Some in situ Magnetospheric Multiscale (MMS) studies have found evidence for FTE coalescence at the magnetopause (Alm et al., 2017; Wang et al., 2017; Zhou et al., 2017).

As shown in Figure 1c, repeated formation and convection of magnetic flux ropes during the multiple reconnection process will sometimes force neighboring FTEs to interact and merge via secondary reconnection. This interaction results in heating plasma (i.e., $p_2 > p + p'$, where p and p' represent plasma pressures of individual magnetic islands) and has been the subject of various simulation studies (e.g., Ding et al., 1992, and references therein). In two dimensions (i.e., anti-parallel reconnection), theory suggests that magnetic flux in the two magnetic islands do not add during the coalescence process, $\psi_2 < \psi + \psi'$, where ψ and ψ' represent the flux contents of the two magnetic islands. Rather, the flux in the smaller structure merges with the larger one (Fermo et al., 2011). In this manner, magnetic flux tubes from the smaller magnetic island connect on both ends to the flux tubes in the larger magnetic island, and in doing so create a larger magnetic island with a greater volume and a larger total mass of plasma.

The three-dimensional picture of coalescing FTEs is more complex (e.g., Yamada et al., 1990). Magnetic helicity, a measure of twist of magnetic field lines inside flux ropes, influences the coalescence process. In particular, the merging FTEs' shear angle, which is controlled by the axial component of their magnetic fields, will determine the rate and the extent to which FTEs coalesce. Correspondingly, the resulting FTE will contain equal, $\psi_2 = \psi < \psi_1$, in the case of anti-parallel reconnection (e.g., Tian et al., 2010), or different magnetic flux than either of the two original FTEs, $\psi - \psi' \leq \psi_2 \leq \psi + \psi'$, in the case of component or guide-field reconnection (e.g., Pritchett & Coroniti, 2004; Furno et al., 2005; Wang et al., 2016; Stanier et al., 2017).

The high temporal and spatial resolution MMS observations enable detailed, multi-point investigations of magnetic reconnection. Here we use a database of 55 previously identified FTEs (Akhavan-Tafti et al., 2018) to investigate the structure and temporal evolution of FTEs. The previous study used the constant-alpha force-free model to determine the impact parameter, which is defined as the relative spacecraft distance to the central axis of the flux rope at the closest approach, as well as the radius, core magnetic field intensity, and magnetic flux content of the FTEs. They found that the FTE size distribution followed an exponential law. This result was interpreted as evidence for the plasmoid instability for the onset of fast reconnection. In space and astrophysical environments highly elongated current sheets are subject to a “plasmoid instability” that leads to their breakup and gives rise to an increase in the reconnection rate (collisional plasma regime: Biskamp, 1986; Loureiro et al., 2007, 2012, 2013; Bhattacharjee et al., 2009; Cassak et al., 2009; Daughton et al., 2009; Huang & Bhattacharjee, 2010; Huang et al., 2011; Murphy et al., 2013; Ni et al., 2015; Comisso et al., 2015; Comisso et al., 2016; Comisso & Grasso, 2016, and collisionless plasma regime: Daughton et al., 2006; Retinò et al., 2008; Daughton et al., 2011a; Baalrud et al., 2012; Huang & Bhattacharjee, 2013).

FTEs subjected to an external and/or internal driving source evolve dynamically with time. During this evolution, the physical parameters, which characterize the state of the system, become a function of time. Each FTE is observed at a different point in its evolution, and thus, the dynamics of growth and expansion can be studied on a statistical basis. In section 3.1, the cross-sectional profiles of the magnetic and plasma characteristics of the 55 FTEs are investigated. Inherently, no forces exist inside perfectly “force-free” structures (i.e., $\mathbf{J} \times \mathbf{B} \sim \nabla \mathbf{P} \sim 0$). However, residual magnetic and thermal forces are expected inside “quasi-force-free” FTEs described here by a goodness of fit (i.e., $\chi^2 \leq 0.1$; Akhavan-Tafti et al., 2018). In fact, the residual forces may contribute to FTEs' temporal evolution. In section 3.2, the MHD forces are investigated inside the quasi-force-free FTEs. In order to seek evidence for FTE growth and the underlying mechanisms, we evaluate the stress balance inside of these FTE-type flux ropes and their pressure balance with the surrounding magnetosheath. The thermodynamics of FTE growth is studied. In section 3.3, the force imbalance across the FTE boundary and the extent to which it may contribute to the expansion of FTEs is discussed. Adiabatic expansion of flux ropes, as suggested by Kumar and Rust (1996), requires plasma pressure inside ion-scale FTEs to drop rapidly with increasing FTE size. However, reconnection slows the rate at which plasma pressure drops with increasing FTE diameter via plasma heating. We conclude that FTE growth in the subsolar region is most likely driven by continuous reconnection at adjacent X-lines and coalescence. This result is important because it underpins the interplay between FTE evolution and adiabatic (e.g., betatron acceleration) and nonadiabatic (e.g., reconnection, turbulence, wave-particle interaction) energization of plasmas at the magnetopause (e.g., Hwang et al., 2016).

2. Experimental Approach

2.1. Methods

Subsolar FTEs are selected from 55 near-equatorial orbital passes (3 November 2015 to 28 December 2015). High spatial and temporal MMS fields (Torbert et al., 2016) and plasma suites (Pollock et al., 2016) are used to investigate the structure and dynamics of FTEs. The four MMS spacecraft were maintained at a tetrahedron formation at an average separation of 10 km (Burch et al., 2016).

FTE identification procedure is done via applying the minimum variance analysis (MVA) on magnetic field measurements where the FTE axis lies in the M (intermediate) direction while N (minimum) and L (maximum) signify the normal and tangential components, respectively (Xiao et al., 2004). Our event selection criteria only allow FTEs with (1) small impact parameter ($IP < 0.5$) and (2) cylindrical symmetry (force-free model goodness of fit parameter, $\chi^2 < 0.1$; cf. Akhavan-Tafti et al., 2018). A complete list of all 63 burst-mode FTEs used in this study and the force-free flux rope model results are provided in Table S1 in the supporting information.

Taylor expansion of parameters is enabled by the MMS's multi-dimensional measurements and is used here to compute the barycentric value:

$$\mathbf{V} = \mathbf{V}_\alpha + (\nabla \mathbf{V}_\alpha) \cdot (\bar{\mathbf{r}} - \mathbf{r}_\alpha) \quad (1)$$

where α indicates the satellite's number ($\alpha = 1, \dots, 4$); $\bar{\mathbf{r}}$ denotes the barycentric position and is the average of all four spacecraft position vectors, \mathbf{r}_α ; \mathbf{V} is an n -dimensional vector (scalar: $n = 1$ and 3-D vector: $n = 3$); and

$$(\nabla \mathbf{V})_{ij} = \frac{1}{N^2} \left[\sum_{\alpha, \beta \neq \alpha} (\mathbf{V}_{\alpha i} - \mathbf{V}_{\beta i}) (\mathbf{r}_{\alpha k} - \mathbf{r}_{\beta k}) \right] \mathbf{R}_{kj}^{-1} \quad (2)$$

where the symbol $\sum_{\alpha \neq \beta}$ indicates summation over all $N(N-1)/2$ independent terms with $\alpha \neq \beta$. Here r and N signify the position and the number of spacecraft, respectively, and \mathbf{R}_{kj}^{-1} is the inverse of the *volumetric tensor*:

$$\mathbf{R}_{kj} = \frac{1}{N} \sum_{\alpha=1}^N r_{\alpha k} r_{\alpha j}, \quad (3)$$

where $i, j, k = 1, \dots, n$ (Harvey, 1998). Equation (2) is also used to calculate the Lorentz force and electron and ion pressure terms in the momentum equation:

$$\rho \frac{d\mathbf{u}}{dt} = \frac{1}{\mu_0} (\mathbf{B} \cdot \nabla) \mathbf{B} - \nabla \left(\frac{B^2}{2\mu_0} \right) - \nabla p \quad (4)$$

where $\frac{d\mathbf{u}}{dt}$ on the left-hand side of the equation is a total convective time derivative. The terms on the right-hand side are magnetic curvature force, magnetic pressure gradient, and (electron and ion) thermal pressure gradient forces, respectively. The thermal pressure gradients are calculated assuming an isotropic (scalar) pressure, p .

2.2. Assumptions

This study relies on three fundamental assumptions to identify FTEs and to obtain a simple empirical model to investigate FTE structures, dynamics, and evolution. First, FTEs are modeled as force-free flux ropes in order to determine their scale size, impact parameter, and cylindricity (goodness of fit: $\chi^2 < 0.1$). Second, FTEs are assumed axisymmetric to obtain radial profiles of their magnetic and plasma characteristics. Third, FTE diameter is used as a proxy to determine their stage of evolution. To accurately study FTE evolution, individual FTEs should be tracked and studied as a function of time and/or distance from their place of origin. However, this is unattainable with present in situ observational capabilities (i.e., small number of closely spaced satellites). To address this, Akhavan-Tafti et al. (2018) used FTE diameter as a proxy to determine FTEs' stage of evolution and, therefore, to study how different physical parameters change. The present study will further investigate how FTE characteristics and dynamics evolve as a function of FTE diameter.

Table 1
Thermal (P_{th}) and Magnetic (P_m) Pressure Averages Inside and Outside FTEs Observed in the Magnetosheath (MSH) and at the Magnetopause (MP)

	Inside		Outside		Ratio	
	MSH	MP	MSH	MP	MSH	MP
P_{th} (nPa)	1.95	1.6	2.02	1.56	0.97	1.03
P_m (nPa)	0.79	0.93	0.57	0.69	1.39	1.35

The ratio of pressures inside and outside FTEs (ratio = $P_{inside}/P_{outside}$) is determined for each category.

In this study, it is further assumed that the internal properties remain relatively constant along the axis of the FTEs. To test this assumption, the 55 subsolar quasi-force-free FTEs are divided into two categories. The first category (35 events) includes only the FTEs that are observed inside the magnetosheath and are labeled as MSH. The second category (20 events) involves FTEs that are observed at or within <10 s from a magnetopause crossing and are labeled as MP. Table 1 summarizes the average thermal, P_{th} , and magnetic, P_m , pressures inside and outside each FTE category. The ratio of average pressures is also determined (ratio = $P_{inside}/P_{outside}$), where ratio >1 indicates that the average pressure inside is larger than the average pressure outside the FTEs. It is shown that thermal pressure averages are nearly similar

(i.e., $P_{th, inside} = P_{th, outside}$) inside and outside both MSH and MP FTE categories. However, the average magnetic pressure is found to be larger inside both FTE categories (i.e., $P_{m, inside} > P_{m, outside}$). More importantly, it is revealed that the internal magnetic and thermal properties of an FTE do not necessarily change depending on where along the FTEs' axes observations are made. This suggests that FTEs are closed structures and, therefore, their internal properties remain relatively constant along the axis of the structure making them independent of the spacecraft location. We take advantage of this conclusion to study the properties and the evolution of the 55 FTEs irrespective of where along the axis of each FTE, in the magnetosheath or the magnetosphere, the spacecraft traversed.

3. Analysis and Results

3.1. FTE Physical Properties

Assuming cylindrical symmetry, the constant- α force-free flux rope model is used to infer the size, the core magnetic field strength, the magnetic flux content, and the spacecraft trajectories through the FTEs. Having knowledge of the trajectory of the spacecraft through the FTE enables statistical study of the structure of quasi-force-free flux ropes as a function of distance from the central axis (Burlaga, 1988; Lepping et al., 1990; Slavin, 2003). The model defines impact parameter (IP) as the spacecraft distance from the axis of the flux rope at the closest approach (see Akhavan-Tafti et al., 2018, Figure 11). Figures 2–4 provide single and multispacecraft cross-sectional views of the magnetic and plasma characteristics of the 55 FTEs. IP=0 indicates the central region of the FTE and IP=1 denotes the FTE “edge,” beyond which the magnetic field connectivity changes and magnetic fields are no longer bound to the structure (Farrugia et al., 1988, 2016; Rijnbeek et al., 1987). The assumption here is that variables inside the FTE are axisymmetric.

Physical variables are binned (bin width; BW = 0.1) and averaged across all 55 FTEs. Environmental variability such as location can change average parameters inside FTEs. In order to take these variabilities into account, every data point within each FTE was normalized with an assigned “weight” (i.e., the ratio of FTE-average of a variable to the average of the variable between all 55 FTEs). The detailed procedure is summarized in Appendix A. The error bars indicate the normalized variations of parameters inside individual bins and is known as the standard error, $\sigma_{mean} \equiv \frac{\sigma}{\sqrt{n}}$, where σ and n are the standard deviation and the number of events in each bin, respectively. Furthermore, the contribution from the Fast Plasma Investigation instrumental measurement errors are negligible compared to the reported plasma moments as suggested by the cross-comparison of magnetic field and plasma moments at the spacecraft barycenter (cf. Akhavan-Tafti et al., 2018, Figure 9).

Figures 2a–2c show how plasma moments vary across the FTEs. As shown in Figure 2a, plasma density is lower closer to the central region of the flux rope. The plasma density dip suggests that flux ropes are open-ended (three-dimensional) allowing plasma to evacuate upon strengthening magnetic field due to the twisting of the field lines (Akhavan-Tafti et al., 2018; Ma et al., 1994; Øieroset et al., 2016; Zhang et al., 2010). The dashed line indicates that the average plasma density inside FTEs at the subsolar magnetopause is $\langle \bar{N} \rangle \pm \sigma_{\bar{N}} = 28 \pm 3 \text{ cm}^{-3}$. Figure 2b indicates that electron and ion velocities, $\langle \bar{V}_e \rangle = 190 \text{ km/s}$ and $\langle \bar{V}_i \rangle = 230 \text{ km/s}$, respectively, are independent of the distance from the FTE's central axis. Electron and ion temperatures vary similarly as a function of IP. Figure 2c shows that the temperatures are lowest

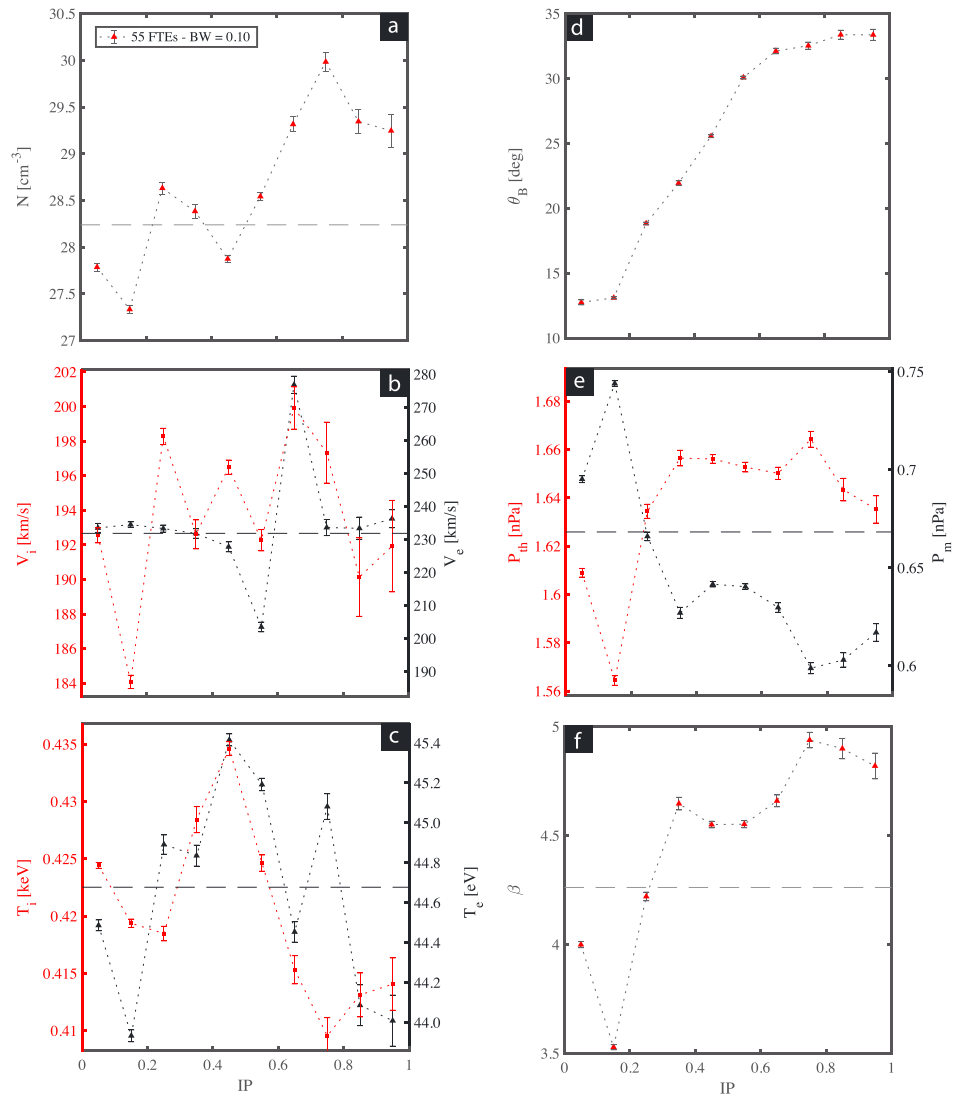


Figure 2. Cross-sectional profiles of (a) plasma density; (b) ion (left axis; red markers) and electron (right axis; black markers) bulk velocity; (c) ion (left axis; red markers) and electron (right axis; black markers) temperature; (d) magnetic pitch angle, $\theta_B = \text{atan}(B_L/B_M)$, where B_L and B_M represent tangential and axial magnetic field components, respectively; (e) thermal, P_{th} (left axis; red markers), and magnetic, P_m , pressures (right axis; black markers); and (f) plasma beta. All data points across the 55 FTEs are grouped and averaged inside bins (bin width; BW = 0.1). The error bar denotes the standard error σ_{mean} and is defined as the normalized variations of a given parameter inside each individual bin, $\sigma_{\text{mean}} \equiv \frac{\sigma}{\sqrt{n}}$, where σ and n are the standard deviation and the number of events in each bin, respectively. Note that the primary and secondary y axes are different in range and, in the case of plasma temperature, unit.

at the outer edge of the FTE. Electron and ion temperatures reach their maxima at IP=0.5 before dropping to their averages, $\langle T_e \rangle = 45$ eV and $\langle T_i \rangle = 420$ eV, respectively, near the central axis.

Figure 2d shows the twisting of the field lines inside an FTE. The magnetic field twist, θ_B , is a function of the ratio of the tangential, B_L , to the axial, B_M , components of magnetic field inside the FTE, $\theta_B = \text{atan}\left(\frac{B_L}{B_M}\right)$. In particular, it indicates that the axial component of the field lines becomes dominant near the central axis. The tangential component of the magnetic field becomes larger near the edges.

Figure 2e demonstrates the cross-sectional variability of both magnetic, P_m , and thermal, P_{th} , pressures inside quasi-force-free subsolar FTEs. Magnetic pressure, $\langle P_m \rangle = 0.65$ nPa, intensifies near the central axis where total magnetic field peaks. In contrast, thermal pressure, $\langle P_{th} \rangle = 1.6$ nPa, reaches a minimum near

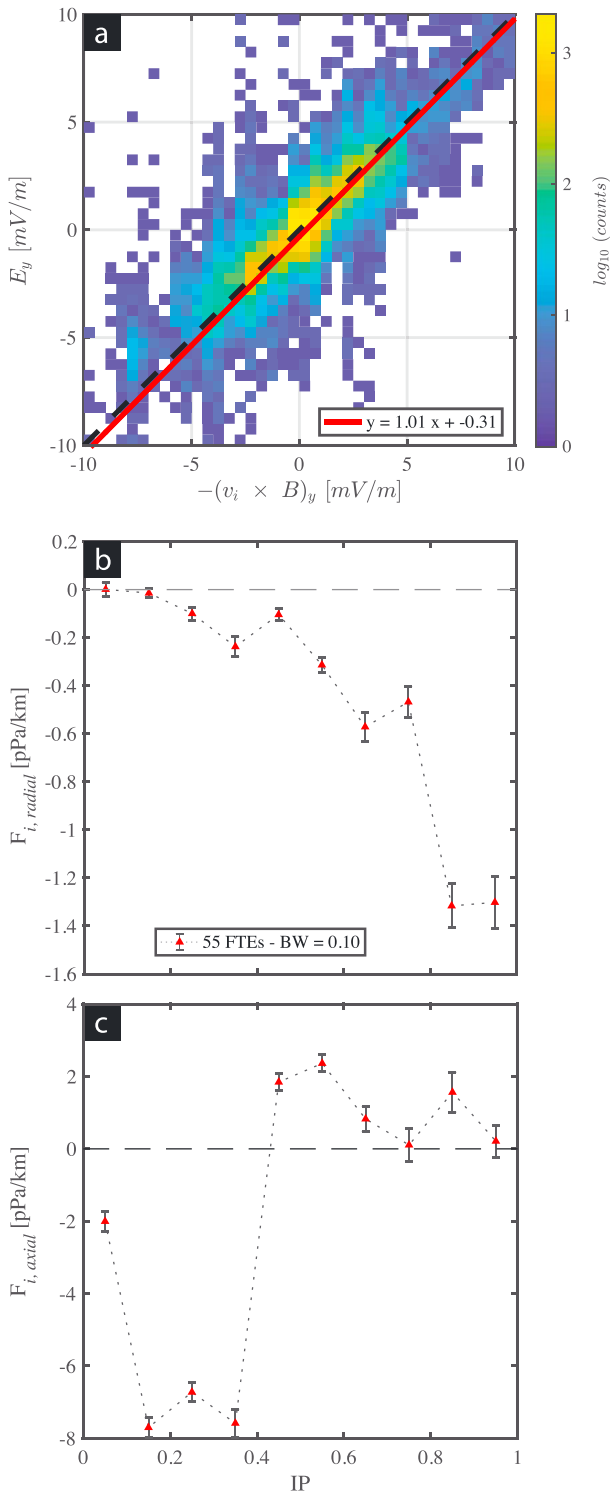


Figure 3. (a) Cross-comparison of the y_{GSE} components of $-v_i \times B$ and the electric field measurements grouped and counted inside bins (BW = 0.5 mV/m). The solid red line represents the orthogonal linear regression fit. The cross-sectional profile of the (b) radial component of the net force and (c) axial component of the net force experienced by ions inside the 55 quasi-force-free FTEs. All data points across the 55 FTEs are grouped and averaged inside bins (BW = 0.1). The error bar denotes the standard error inside each individual bin.

the central axis corresponding to the dip in plasma density. The thermal pressure profile suggests that plasma density variation, not temperature, is the dominant factor determining the plasma pressure profile inside FTEs. More importantly, the pressures are shown to be inversely correlated. A first-order approximation of the pressure gradients inside FTEs (i.e., $\nabla P \sim \frac{\Delta P}{R_{FR}}$, where R_{FR} is the FTE radius) also suggests that the magnetic and thermal pressure gradients point in opposite directions and nearly balance each other.

Figure 2f further shows that plasma beta, which is the ratio of thermal and magnetic pressures, $\beta = P_{th} / P_m$, dips below its average value of $\langle \beta \rangle = 4.3$ near the FTE's central axis. The average plasma beta is larger than those reported inside typical FTEs. This is consistent with Akhavan-Tafti et al. (2018), where newly-born FTEs were shown to contain large plasma density. Over time, plasma will be depleted while magnetic field enhances inside the FTE resulting in lower plasma beta.

During (continuous and secondary) reconnection additional magnetic flux and plasma are added to the outer layers of FTEs. Internal magnetic and thermal forces aim to release this stress and bring the structure back to equilibrium. We first investigate the “frozen-in” condition (i.e., $E \sim -v_i \times B$) for ions inside the 55 FTEs. The y_{GSE} components of both the electric field measurements (from the Electric Dual Probes (EDP) instrument (Ergun et al., 2016; Lindqvist et al., 2016; Torbert et al., 2016)) and the cross product of ion bulk velocity (from the Fast Plasma Investigation (FPI) instrument (Pollock et al., 2016)) and magnetic field measurements (from the fluxgate magnetometer (FGM) instrument (Russell et al., 2016)) are bin-averaged (BW = 0.5 mV/m) and are shown in Figure 3a. The dashed black line represents a line of slope unity for reference ($y = x$). The solid red line represents the fit ($y = 1.01x - 0.31$) determined using orthogonal linear regression indicating that ions remain frozen-in inside our quasi-force-free FTEs.

Second, the average net force components are investigated across our quasi-force-free FTEs. Here net force refers to the vector sum of the magnetic, $J \times B$, and thermal pressure, ∇P , forces. Figure 3b shows the bin-averaged (BW = 0.1) values of the radial component of the net force on ions, $F_{i, radial}$, across the cross sections of the 55 FTEs. The detailed coordinate transformation procedure is summarized in Appendix B. Here a negative value of $F_{i, radial}$ indicates inward force. We find that the net force is radially inward. Farther away from the FTE axis, the magnitude of the radially inward net force enhances. The bin-averaged values of the axial component of the net force, $F_{i, axial}$, across the cross section of the FTEs are further shown in Figure 3c. In contrast with the radial component, the axial component of the net force is found to enhance significantly near the FTE's central axis, $IP < 0.5$. In conclusion, the results found in Figures 3a–3c suggest that the “frozen-in” magnetic field lines are compressed inward with the plasma giving rise to strengthening core magnetic field. However, due to the strengthening magnetic field the radially inward net force is reduced near the FTE's central axis while the axial net force enhances. These findings indicate that plasmas are accelerated toward the FTE's central axis, though closer to the FTE's core region the radial force is reduced and the plasma is instead accelerated along the axial direction causing the plasma to escape the stressed region (e.g., Ma et al., 1994).

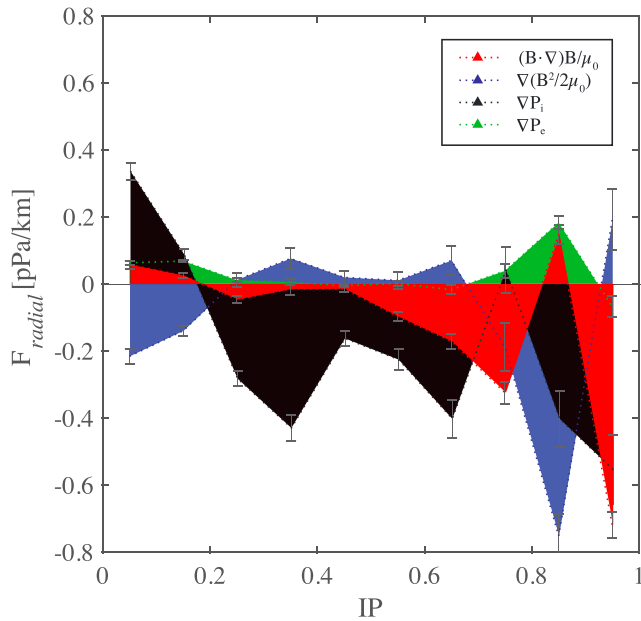


Figure 4. The radial components of the magnetic and thermal forces as a function of IP. The magnetic forces include the magnetic curvature force (shown in red) and the magnetic pressure gradient (shown in blue). The thermal forces are the ion (black shade) and electron (green shade) thermal pressure gradients forces. All data points across the 55 FTEs are grouped and averaged inside bins ($BW = 0.1$). The error bar denotes the standard error, σ_{mean} , and is defined as the normalized variations of a given parameter inside each individual bin, $\sigma_{\text{mean}} \equiv \frac{\sigma}{\sqrt{n}}$, where σ and n are the standard deviation and the number of events in each bin, respectively. The dotted lines further show the boundaries of each of the shaded areas.

3.2. FTE Boundaries

In this section, the physical characteristics and dynamics are studied inside and just outside the 55 quasi-force-free FTEs. As shown in Figure 6, the term “inside” refers to the cross section of the flux rope (shaded red) that meets the constant-alpha force-free flux rope model criteria (Akhavan-Tafti et al., 2018; Burlaga, 1988; Lepping et al., 1990), while “outside,” shaded as light blue, indicates the outer perimeter of the flux rope (i.e., the draping region (Farrugia et al., 1988; Rijnbeek et al., 1987). The thickness of the outer perimeter was selected at 10 local ion-inertial lengths, d_i . This thickness is nearly one third of the nominal FTE diameter at the dayside magnetopause (Akhavan-Tafti et al., 2018) and is sufficiently large to include reconnection-associated kinetic and MHD effects (Smith, 1977; Tóth et al., 2017).

Table 2 lists the relative contribution, that is, percentage of total, of the force terms. The force terms are averaged inside and outside the 55 FTEs. The shading in each row corresponds to the relative contribution of individual force terms with the darkest shading representing the largest contribution. The table indicates that magnetic pressure gradient and ion thermal pressure gradient are the dominant force terms contributing to the total force inside and outside FTEs. Magnetic forces (magnetic curvature force and magnetic pressure gradient) are comparable in magnitude to the ion thermal pressure gradient inside FTEs. The thermal pressure gradient becomes dominant outside the FTEs.

Figure 7 compares plasma beta, current density, and net force vector components inside and outside FTEs. Observations outside the 55 FTEs are bin-averaged ($BW = 1 d_i$, where $d_i = c/\omega_{pi} = 2.28 \times 10^7 (N_i)^{-1/2}$ cm, is the average local ion inertial length and is a function of average ion density, N_i , outside each FTE). Figure 7a indicates that plasma beta varies quite remarkably across the FTE boundary. This is partly due to the enhanced magnetic pressure inside FTEs. Figure 7b shows that current density is, on average, enhanced by nearly 50% inside FTEs. Figures 7c–7e present the net force vector components in the cylindrical coordinate system. The radial component of the vector sum of the magnetic and thermal forces remains relatively small outside the FTEs, suggesting that the radial magnetic and thermal forces are balanced. Furthermore,

Figure 4 further provides the cross-sectional profiles of the radial components of the magnetic and thermal forces across the FTE. The error bars indicate the standard error inside each individual bin. The radial profiles indicate that inside FTEs magnetic pressure gradient (blue shade) and ion thermal pressure gradient (black shade) are dominant near the FTE's central axis. It is also found that magnetic curvature force (red shade) is radially inward and becomes more significant near the FTE's edge. As expected, electron thermal pressure gradient (green shade) is small throughout the structure. On the other hand, ion thermal pressure gradient remains large, relatively steady, and radially inward across the structure.

Figures 5a–5d compare the spatial profiles of perpendicular and parallel components of plasma current density, net force, and electron and ion temperatures. In Figure 5a, the ratio of perpendicular to parallel components of the current density is shown to be smaller than unity (i.e., $R_J = J_{\perp} / J_{\parallel} < 1$; Akhavan-Tafti et al., 2018). This is a result of our event selection criteria where only quasi-force-free FTE candidates were chosen. Figure 5b shows that the majority of the net force inside the FTE accelerates plasmas perpendicular to the local magnetic field. The relatively large parallel force is consistent with the observed enhanced axial component of the net force inside FTEs. The parallel net force can facilitate the escape of the stressed plasma through accelerating them along the magnetic field. The parallel net force may also affect the helical magnetic structure of the FTEs (Zhao et al., 2016). Temperature anisotropy inside FTEs is evident in Figures 5c and 5d. Parallel electron temperature, $T_{e\parallel}$, and perpendicular ion temperature, $T_{i\perp}$, are observed to be the larger components of the electron and ion temperatures, respectively.

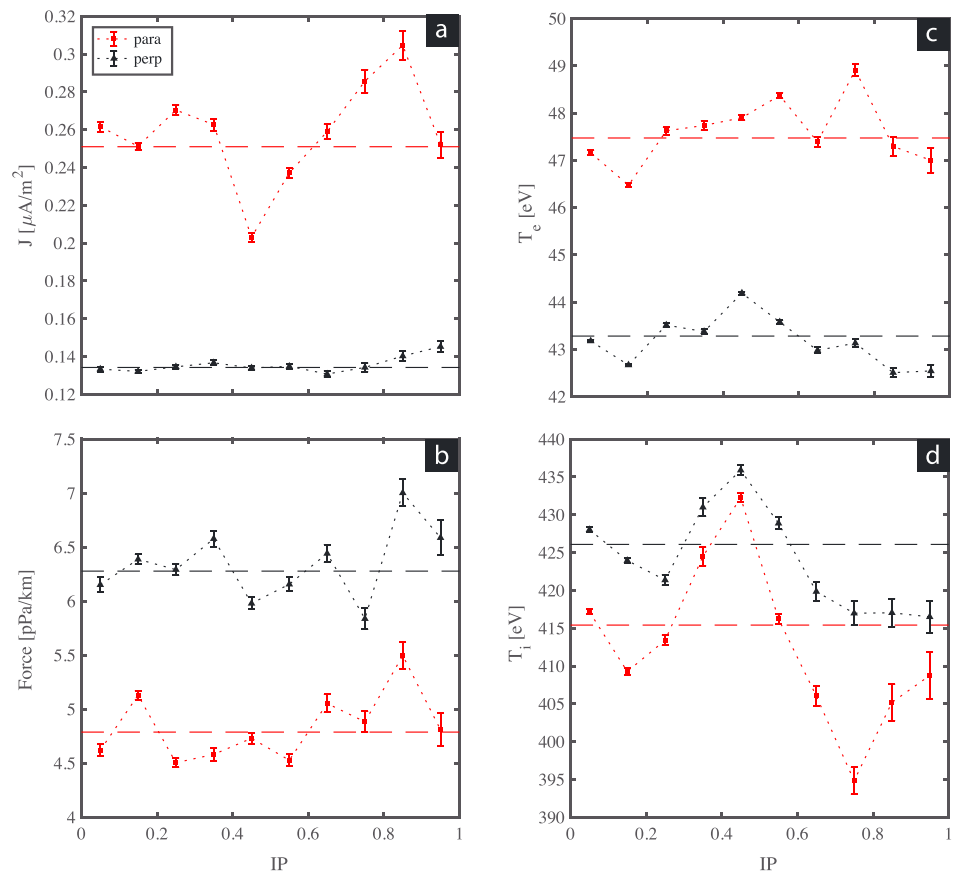


Figure 5. Cross-sectional profiles of the parallel (red) and perpendicular (black) components of current density, net force, electron, and ion temperatures. All data points across the 55 FTEs are grouped and averaged inside bins ($BW = 0.1$). The error bar denotes the standard error inside each individual bin. The dashed line represents the average value of each parameter across all FTEs.

the magnitude of the tangential component of the net force, defined as $F_{\text{tangential}} = |\mathbf{F} - (F_{\text{radial}} \hat{\mathbf{r}} + F_{\text{axial}} \hat{\mathbf{M}})|$, enhances near the FTE edge. Finally, the axial component of the net force remains steady and non-zero outside the FTE. The residual tangential and axial net force components may impact the twist of the field lines at the outer layers of the FTEs (e.g., Zhao et al., 2016).

3.3. Case Study

The magnetic and thermal force magnitudes inside the 55 quasi-force-free FTEs are nonzero (i.e., $\mathbf{J} \times \mathbf{B} \neq 0$ and $\nabla P \neq 0$). Here we provide two quasi-force-free FTEs observed in the magnetosheath with enhanced net force components. Figure 8 summarizes the magnetic and plasma moment observations. The panels in the left column include the measurements by MMS2: (a) total magnetic field, (b) magnetic field vector components in the GSE coordinate system, (c) plasma density, (d) ion velocity, (e) plasma current density, and (f) ion temperature. The magnetic and thermal force components are further included in the right column: (g) total magnetic field, (h) magnetic and thermal pressure components, (i) magnetic curvature force, (j) magnetic pressure gradient force, (k) ion thermal pressure gradient force, and (l) net force which is defined as the vector sum of the magnetic and thermal forces. The vector components are presented with respect to the local magnetic field. The component of each vector parallel to \mathbf{B} is shown in red, while the magnitude of the perpendicular component of the vector is shown in black. The red gradient and the

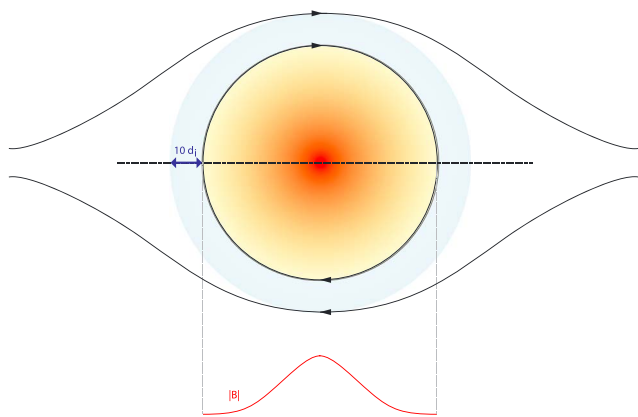


Figure 6. The red gradient indicates the magnetic field magnitude inside the FTE, while the light blue shade denotes the outer perimeter of the FTE with a thickness of 10 ion inertial lengths.

Table 2

The Relative Contribution, That Is, the Percentage of Total, of the Force Terms Inside (Blue Shade) and Outside (Red Shade) FTEs

	$ (\mathbf{B} \cdot \nabla)\mathbf{B}/\mu_0 $	$ \nabla(\mathbf{B}^2/2\mu_0) $	$ \nabla P_i $	$ \nabla P_e $
Inside FTE	15%	33%	43%	9%
Outside FTE	13%	29%	48%	11%

Note. The shading in each row corresponds to the relative contribution to the total force averaged inside all 55 FTEs with the darkest shading representing the largest contribution. FTE = flux transfer event.

light-blue solid color shadings represent the areas “inside” and “outside” each FTE. The width of the blue shaded area is 10 ion inertial lengths which is determined using the average local ion density, $N_i = 40 \text{ cm}^{-3}$.

Zhou et al. (2017) reported the presence of an ion jet reversal ($v_{iz \text{ GSM}} = \pm 200 \text{ km/s}$; Figure 8c) during a magnetopause crossing (17 November 2015—02:13:40–02:14:40 UT) corresponding to a “dissipative” interaction of two neighboring FTEs. Shortly after this crossing, at the peak of the southward ion velocity, a chain of small-scale FTEs is observed, labeled as “FTE 1” and “FTE 2.” In particular, two FTEs centered at 02:20:48 UT and 02:21:24 UT are observed surrounded by two southward v_{iz} peaks (02:17:30 UT and 02:22 UT) which met our selection criteria (Akhavan-Tafti et al., 2018).

Plasma density is shown to dip by 50% at the center of FTE 1 and by 25% inside FTE 2. Parallel current density is further found to enhance inside the two events. Ion temperature anisotropy is observed across both FTEs; however, the parallel ion temperature, $T_{i\parallel}$, is smaller than the perpendicular component, $T_{i\perp}$, inside FTE 1, $T_{i\perp} > T_{i\parallel}$, while $T_{i\perp} < T_{i\parallel}$ inside FTE 2. Magnetic pressure enhances inside both FTEs resulting in increased magnetic pressure gradient force. $\beta = P_{th}/P_m \sim 4$ outside the two FTEs. Plasma beta is reduced inside both structures. Ion thermal pressure gradient force is found dominant inside the two FTEs.

4. Discussion

In this study, we investigate the structural characteristics of the 55 subsolar quasi-force-free FTEs identified and modeled by Akhavan-Tafti et al. (2018). The determination of the impact parameter of the MMS spacecraft trajectories through modeling FTEs as force-free flux ropes with cylindrical geometry allows the study of cross-sectional profiles of physical parameters as a function of distance from the FTEs’ central axis. It is revealed that

1. Plasma density and plasma beta increase with increasing distance from the FTE’s central axis.
2. The radial component of the net force (solid black arrow in Figure 9a) at the outer layers of the flux rope points inward. The axial component of the force, shown as an into-the-plane arrow, becomes significant near the FTE’s central axis.

As illustrated in Figure 9a, the continuous supply of magnetic flux and plasma to the FTEs’ outer layers, shown as grey-shaded area, by magnetic reconnection at adjacent X-lines and/or coalescence with the neighboring FTEs, shown as X points, drives FTEs to “grow” (e.g., Akhavan-Tafti et al., 2018; Hasegawa et al., 2010) while also contributing to

1. The observed average decrease in plasma density near the FTE’s central axis. Cross-sectional profile of plasma density is represented as a blue and red-shaded area. As discussed above, plasma density drops below average, white-shaded area, near the FTE’s core region.
2. The observed radially-inward net force at the outer layers of the FTEs as well as the enhanced axial net force closer to the FTE’s central axis.

These observations are consistent with the continuous supply of magnetic flux and plasma heating by magnetic reconnection. As suggested by the radially inward net force, the additional magnetic flux and plasma, in the reconnection exhaust, is compressed and accelerated, respectively, while being transported toward the FTE’s core region. Closer to the FTE’s central axis, where B_{axial} enhances, the axial net force component

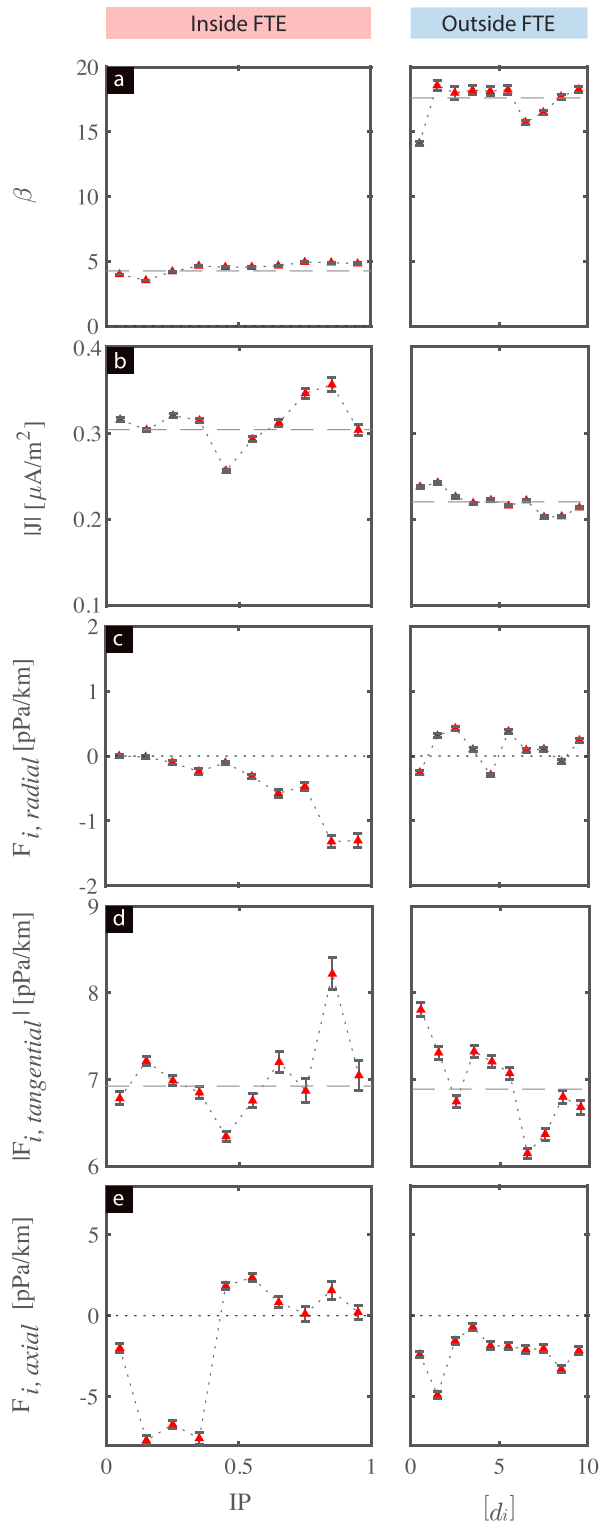


Figure 7. The cross-sectional profiles of (a) plasma beta, (b) current density, (c) the radial component of the net force, (d) the magnitude of the tangential component of the net force and (e) the axial component of the net force inside and outside (10 local ion inertial lengths from the FTE edge) the 55 FTEs. All data points inside (outside) the 55 FTEs are grouped and averaged inside bins ($BW = 1 d_i$). The error bar denotes the standard error inside each individual bin.

strengthens. The strengthening axial force can then accelerate plasmas along the FTE's axis and cause the “trapped” plasma population to escape the stressed region (e.g., Akhavan-Tafti et al., 2018; Ma et al., 1994; Sun et al., 2019).

1. The tangential component, B_L , of the magnetic field increases farther away from the FTE's central axis. Figure 9b shows the profile of magnetic pitch angle, $\theta_B = \text{atan}\left(\frac{B_L}{B_M}\right)$, across the cross section of a typical quasi-force-free FTE. The tangential component of the magnetic field enhances at the FTE's outer layers. The FTEs' magnetic pitch angle may impact the coalescence process (e.g., Yamada et al., 1990).
2. Ion and electron temperatures are found to be anisotropic throughout the cross section of FTE-type flux ropes observed both inside the magnetosheath and at the magnetopause.
3. We observe that, on average, $T_{e\parallel} > T_{e\perp}$ throughout the cross section of subsolar FTEs. Firehose, $T_{e\parallel} > T_{e\perp}$, and whistler, $T_{e\perp} > T_{e\parallel}$, instabilities can arise due to electron temperature anisotropy (e.g., Cuperman, 1981; Drake et al., 2010; Gary & Madland, 1985; Hollweg & Völk, 1970a, 1970b; Kennel & Petschek, 1966; Pilipp & Völk, 1971; Wang et al., 2019; Zhao et al., 1996). Zhao et al. (1996) note that the former condition is rarely observed in space plasmas and, hence, infrequently reported in the literature, but our analysis suggests that it may occur in FTEs.
4. We also observe that, on average, $T_{i\perp} > T_{i\parallel}$ inside small-scale subsolar FTEs. In theory, anisotropic ion temperature, $T_{i\perp} > T_{i\parallel}$, can act as a source of free energy for various low-frequency electromagnetic instabilities in the magnetosheath, such as the ion cyclotron anisotropy and instability and the mirror instability (e.g., Gary et al., 1976; Lee et al., 1988; Midgley & Davis, 1963; Remya et al., 2013; Zwan & Wolf, 1976). A future MMS study will examine anisotropies in the plasma moment distributions and the associated instabilities inside FTEs.

The profiles of the magnetic and plasma properties are further analyzed outside FTEs to gain insight into the role of environmental conditions in impacting the evolution, that is, growth, of FTEs. In particular, it is shown that, compared to the average properties inside FTEs,

1. Plasma beta enhances significantly, by a factor of >4 , outside FTEs. This may be the result of enhanced thermal pressure, and reduced magnetic pressure, at the adjacent X-lines.
2. Current density is reduced outside FTEs.
3. The radial component of the net force remains steady and negligible outside FTEs which may be a result of our selection process, that is, quasi-force-free FTEs. However, the tangential and the axial components of the net force are found to be significant outside FTEs. These observations suggest that the structure of FTEs, in particular, the magnetic field twist at the FTE's outer layers, $\theta_B = \text{atan}\left(\frac{B_L}{B_M}\right)$, may be affected by environmental net forces (e.g., Zhao et al., 2016).

The observed internal and external magnetic and plasma characteristics suggest that FTEs evolve with time. Previously, the FTEs at the subsolar magnetopause have been found to range between 10^2 and 10^4 km in diameter. Akhavan-Tafti et al. (2018) suggested that FTEs are generated at ion-scale at the subsolar magnetopause and grow as they convect away

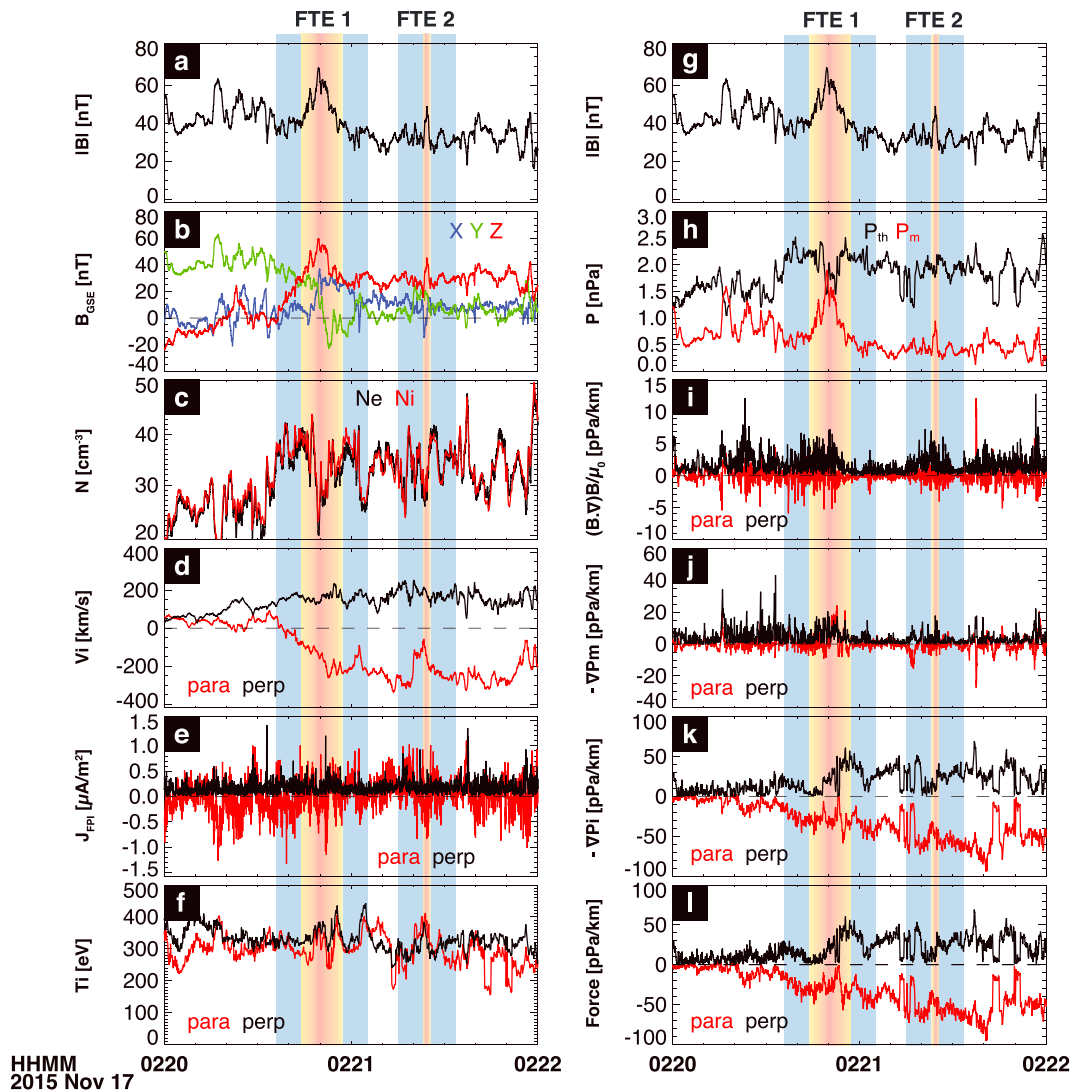


Figure 8. The 17 November 2015 magnetopause boundary crossing. (a and g) Magnetic field magnitude, (b) magnetic field components in the GSE coordinates, (c) ion (red) and electron (black) density, (d) ion velocity parallel (red) and perpendicular (black) components, (e) plasma current density parallel and perpendicular components, (f) ion temperature components, (h) ion thermal and magnetic (black) pressures, (i) magnetic curvature force parallel and perpendicular components, (j) magnetic pressure gradient force parallel and perpendicular components, (k) ion thermal pressure gradient force parallel and perpendicular components, and (l) net force parallel and perpendicular components. Two quasi-force-free FTEs, labeled as “FTE 1” and “FTE 2,” are shown with gradient and solid color shadings corresponding to the observations made “inside” and “outside” the FTEs, respectively. The width of the blue solid color shading is 10 ion inertial lengths, $d_i \sim 1,400$ km, from an average of $N_i = 40 \text{ cm}^{-3}$.

toward flanks and high-latitude magnetopause where the magnetosheath pressure is much lower (Eastwood et al., 2012; Sibeck et al., 1990). From the three FTE growth mechanisms discussed in Figure 1, a radially outward net force across the FTE boundary is required to cause the “expansion” of FTEs as a function of distance from the subsolar region. However, our observations indicate that the net force inside FTEs is radially inward and that the radial component of the net force is negligible, that is, force balance $J \times B - \nabla P \sim 0$, outside FTEs. To further confirm or deny expansion as an FTE growth mechanism, we examine the plasma pressure inside FTEs as a function of FTE diameter.

In the absence of magnetic reconnection, the irreversible adiabatic expansion of FTEs requires, cf. Appendix C, that the thermal pressure inside an ion-scale FTE to drop by at least 2 orders of magnitude, that is, $p < 10^{-4}$ nPa, while the FTE grows macroscale, $> 1 R_E$ in diameter. Figure 10a shows histograms of thermal ($P_{th} = Nk_B T$, where $T = T_i + T_e$ and k_B is the Boltzmann constant) and magnetic pressures

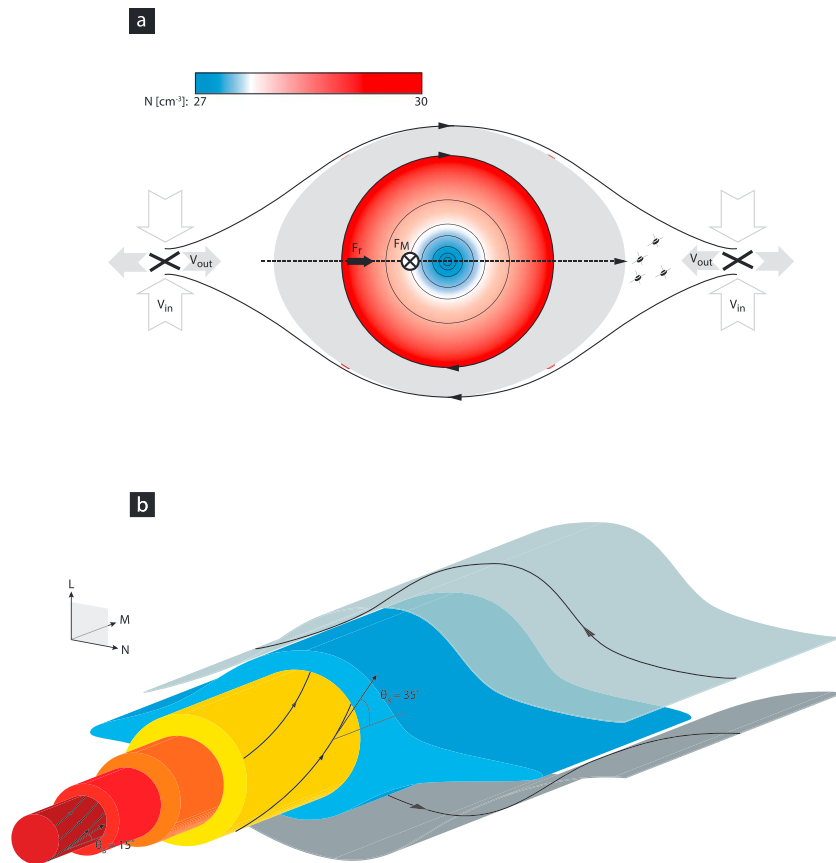


Figure 9. Cartoon representation of cross-sectional profiles of (a) plasma density profile and the radial (solid black arrow) and the axial (into-the-plane arrow) components of the net force and (b) magnetic pitch angle, $\theta_B = \text{atan}(B_L / B_M)$, where B_L and B_M represent tangential and axial magnetic field components, respectively. Plasma density colorbar corresponds to the bin-averaged values reported in this study. The grey shaded area around the flux rope illustrates the magnetic flux and plasma (grey arrows labelled as “ V_{out} ”) continuously added to the FTEs’ outer layers by reconnection (shown as “X”).

measured inside the 55 FTEs at the subsolar magnetopause. The thermal and magnetic pressure distributions are nearly symmetric about their means at 1.5 nPa ($0.3 < P_{th} \text{ [nPa]} < 5.0$) and 0.5 nPa ($0.03 < P_m \text{ [nPa]} < 3$), respectively. Therefore, our observations do not confirm FTE growth via expansion.

We further examine the rate at which pressure varies as a function of FTE size. Figure 10b shows the change in average thermal, P_{th} , and magnetic pressures, P_m , as a function of FTE diameter. Here pressures are averaged across the cross section of individual FTEs and plotted with respect to the FTE diameter. The 55 FTEs are grouped into bins (BW = 500 km) and the error bars indicate standard error of values inside each bin ($\sigma_{\text{mean}} \equiv \frac{\sigma}{\sqrt{n}}$, where n is the number of events in each bin). The solid curves represent power law fits to the data with shaded areas signifying the 95% confidence intervals. The dependence of thermal pressure as a function of the FTE diameter, λ , is measured on average as $\lambda^{-0.24}$ compared to $\lambda^{-3.3}$, shown as a grey dashed line, expected for adiabatic expansion of FTEs. This observed slow decay of the thermal pressure is consistent with plasma undergoing significant heating (Drake et al., 2006), further indicating that adiabatic expansion alone cannot explain the observed FTE range of diameters at the magnetopause. This result suggests that adiabatic FTE expansion in the subsolar magnetosheath is accompanied by processes that add (i) magnetic flux and (ii) thermal pressure to the outer layers of FTEs. Hence, FTE growth can occur via three mechanisms:

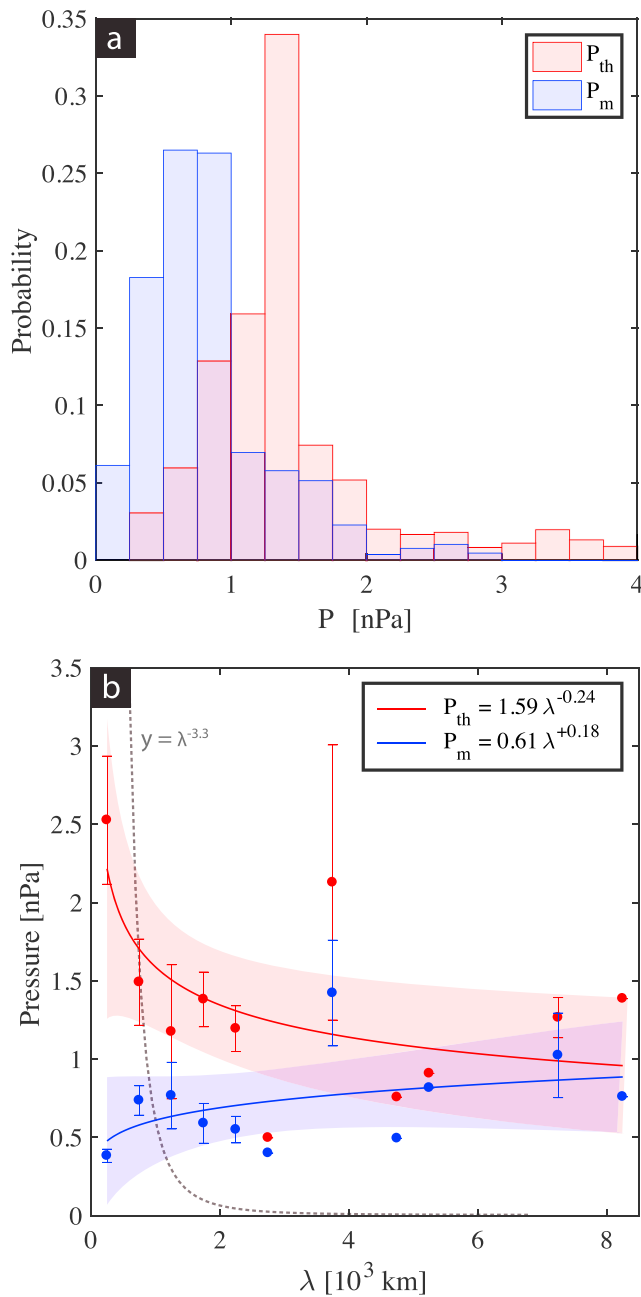


Figure 10. (a) Histograms of thermal and magnetic pressure data points inside the 55 FTEs. (b) Average thermal pressure and magnetic pressure profiles as a function of FTE diameter across the 55 FTEs. The circles indicate the bin-averaged (BW = 500 km) values of pressures.

1. *Continuous reconnection:* Ongoing multiple X-line reconnection can add magnetic layers to an existing FTE as predicted by simulations (Chen et al., 2017; Lee & Fu, 1985). The resulting increase in the structures' magnetic flux and plasma content can explain also the exponential distribution of magnetic flux inside subsolar FTEs (Akhavan-Tafti et al., 2018; Fermo et al., 2011). Further, plasma is heated in the exhausts of the primary X-line and plasma pressure is expected to increase in the outer layers of these flux ropes as new layers of helical magnetic flux are added and the FTE grows in size.
2. *Coalescence:* PIC simulations suggest that numerous small-scale flux ropes form in thin current sheets (e.g., Drake & Lee, 1977a, 1977b; Galeev et al., 1978; Hesse et al., 1999; Pritchett, 2001, 2005; Ricci et al., 2004; Daughton et al., 2011b; Nakamura et al., 2016; Huang et al., 2016; Lapenta et al., 2018). Neighboring small-scale flux ropes then reconnect with each other and, as a result, grow into larger merged FTEs (e.g., Alm et al., 2017; Wang et al., 2016; Zhou et al., 2017). Through coalescence, the magnetic flux content of the merged FTE is equal to or greater than the larger of the two (i.e., Fermo et al., 2011). However, plasma is added to the outer layers of the merged FTE and heating may occur as a result of the reconnection between the outer layers of the two merging FTEs.
3. *Adiabatic expansion due to the pressure imbalance across the FTE boundary:* Upon formation at the subsolar magnetopause, FTEs advect with the magnetosheath flow toward the flanks and high-latitude regions where the magnetosheath thermal pressure is lower. The pressure imbalance across the FTE boundary created as they are carried tailward may lead to the expansion of FTEs (Eastwood et al., 2012). The expansion will continue until the internal magnetic forces and plasma pressure gradients are again in balance. FTEs expand to relieve the stress most likely produced by reconnection-driven heating of their outer layers.

5. Conclusions

In summary, we have used the high temporal and spatial-resolution MMS magnetic and plasma measurements to investigate the structure and dynamics of 55 subsolar FTEs. The force-free flux rope model was initially used to derive the relative distance of the spacecraft with respect to the FTE's central axis, that is, the impact parameter. Having knowledge of the impact parameter facilitates the statistical study of cross-sectional profiles of physical parameters inside and outside FTEs. In particular, we observe that plasma density and plasma beta decrease near the FTE's central axis. We also find that the sum of the magnetic and thermal forces points radially inward at the outer layers of FTEs, suggesting that plasmas and the frozen-in magnetic flux are accelerated toward the center of the FTE causing the field lines to compress. However, closer to the FTE's central axis, the axial net force is found to dominate and to accelerate the plasma along the FTE's axis and out of the region of compressing field lines.

Magnetic and plasma characteristics are also studied outside FTEs to gain insight into how changing external conditions may impact FTEs. It is observed that plasma beta outside the FTEs is, on average, significantly larger than the values inside. Similarly, the tangential and axial net force components are also found to enhance outside FTEs. Since magnetic field pitch angle is a function of the ratio of the tangential and axial magnetic field components, the enhanced tangential and axial net force components can likely impact the twist of magnetic field lines at the FTEs' outer layers.

We also find that the average thermal pressure inside FTEs decreases with increasing FTE diameter. If the 55 FTEs had started out with similar diameters and external conditions, the observed rate of thermal pressure decrease with increasing FTE diameter is shown to be slower than the rate expected for the adiabatic expansion of FTEs. This result suggests that the adiabatic expansion of FTEs at the subsolar magnetosheath is accompanied by processes, such as magnetic reconnection, that add (i) magnetic flux and (ii) thermal pressure to the outer layers of FTEs. Magnetic reconnection at adjacent X-lines and/or coalescence with the neighboring FTEs is most likely the external source of magnetic flux and plasma. This continuous addition of magnetic flux and plasma, while causing an FTE to grow in size, drives a net radially inward force which leads to the compression of the magnetic field lines and the acceleration of plasmas inside the FTE. This process leads naturally to the low-beta core region of the FTE as it evolves toward its final relaxed, force-free state, that is, the Taylor state (Taylor, 1986).

Appendix A

An analytical weight determines that the i_{th} observation comes from a sub-population with a particular mean and variance, σ^2/w_i , where σ^2 is a common variance and w_i is the weight of the i_{th} observation. Analytical weights are commonly used in meta-analyses where each *observation* is the mean of a sample (Kish, 1992). In this study, each observed data point, $X_{i,n}$, is weighted with

$$w_n = \langle X \rangle / \bar{X}_n$$

where w_n is the analytical weight assigned to data points inside the n_{th} FTE. Here $\langle X \rangle \equiv \frac{1}{P} \sum_{n=1}^P \bar{X}_n$, where 55 and is the total number of FTEs used in this study. \bar{X}_n indicates the average value of variable X across the cross section of the n_{th} FTE; hence, it is defined as

$$\bar{X}_n \equiv \frac{1}{N} \sum_{i=1}^N X_{i,n}$$

where subscript $i = 1, \dots, N$ denotes individual data points inside an FTE with total N measurements.

Appendix B

In order to determine the radial component of a vector in a force-free flux rope, the vector is projected from the LMN coordinates into the cylindrical coordinates. Figure B1 demonstrates that the vector, A , is first projected into the flux rope cross-section plane (LN plane). The vector can then be written in terms of the in-plane components, $A(\hat{N}, \hat{L}) = A_N \hat{N} + A_L \hat{L}$, where unit vector \hat{N} points along the spacecraft trajectory and $\hat{L} = \hat{M} \times \hat{N}$. The transformation into the cylindrical coordinates requires the knowledge of angle θ which is a function of (1) the distance of the spacecraft location along its trajectory ($0 < X < d$, where d indicates the distance between the two flux rope edges) from the flux rope's central axis and (2) the closest approach ($d_{closest\ approach} = r_{flux\ rope} \times IP$):

$$\theta = \tan^{-1} \left(\frac{\left(\frac{X-d/2}{r_{flux\ rope}} \right)}{IP} \right)$$

where $r_{flux\ rope}$ is the flux rope radius as determined by the force-free flux rope model. The radial component of the vector can be written as

$$A_r = A_N \sin(\theta).$$

It is important to note that for statistical analyses, all flux ropes should be rotated into the same LMN coordinates. To achieve this, all flux ropes are rotated with respect to the \hat{N} axis for their axes, \hat{M} , to point in the same (e.g., dawn-dusk) direction. Because we are only interested in the radial

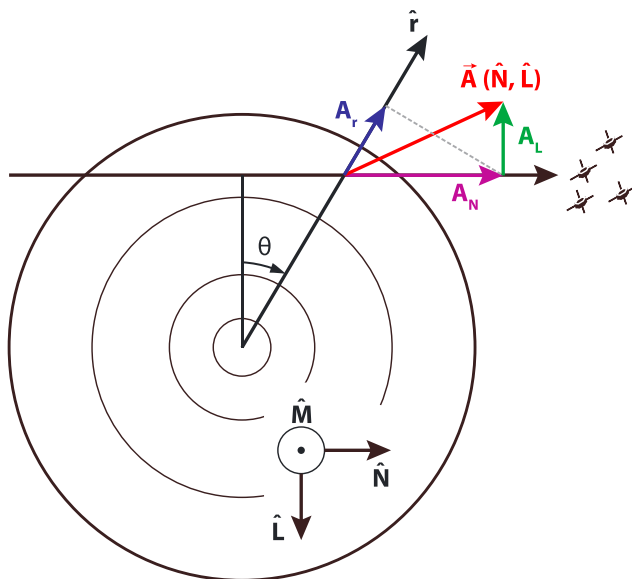


Figure B1. Coordinate transformation between flux rope's LMN and cylindrical coordinate systems.

component of the vector the different flux rope helicities (e.g., right-hand helicity) will not impact the results.

Appendix C

The second law of thermodynamics states that the entropy inside an irreversibly expanding FTE must not decrease:

$$p_2 V_2^\gamma \geq p_1 V_1^\gamma$$

where pressure $p \propto N T$, where N and T are the zeroth (density) and second (temperature) plasma moments, respectively. V denotes the flux rope volume, hence, proportional to the square of the FTE diameter ($V \propto \lambda^2 H$, where λ is the FTE diameter and H is the FTE length, assuming cylindrical symmetry; e.g., Lingam & Comisso, 2018). The length, H , of the FTE is assumed to remain constant. It is important to note that recent observational evidence suggest that reconnection sites can expand, though primarily in local time along the boundary, which may result in a change in H (Zou et al., 2018). The subscripts indicate the stage of evolution (t_1 and $t_2 \gg t_1$) and $\gamma=5/3$ is the adiabatic specific heat ratio for monoatomic gases. Thereby, the equation above can be written as

$$p_2/p_1 \geq \left(\lambda_2/\lambda_1\right)^{-3.3}$$

where λ is the FTE diameter.

The role of magnetic field in the above approximation can also be included by using the double-adiabatic approximation (e.g., Baumjohann & Treumann, 2012). This approximation takes into account the different variations in the parallel and perpendicular pressures in the presence of a sufficiently strong magnetic field. The perpendicular thermal pressure P_\perp is a linear function of plasma density n and magnetic field magnitude B as suggested by the conservation of magnetic moment, $\mu = P_\perp/nB$, in systems wherein plasma gyration takes longer than temporal variations of the plasma. In the absence of any dissipative mechanisms and heat transfer, the parallel component of thermal pressure is also found to be a function of both plasma density and magnetic field:

$$\begin{aligned} P_\perp &\propto n B \\ P_\parallel &\propto n^3 B^{-2} \end{aligned}$$

It is further shown that the parallel γ_\parallel and the perpendicular γ_\perp adiabatic indices are defined in the double-adiabatic equations of state as

$$\begin{aligned} \gamma_\perp &= 1 + \frac{\ln(B_2/B_1)}{\ln(n_2/n_1)}, \text{ and} \\ \gamma_\parallel &= 3 - 2 \frac{\ln(B_2/B_1)}{\ln(n_2/n_1)} \end{aligned}$$

Therefore, inside an adiabatically growing flux rope, wherein no heat transfer is assumed, the thermal pressure becomes increasingly anisotropic, that is, $P_\perp/P_\parallel \neq 1$. Plasma pressure components can then be written as

$$p_{\perp 2}/p_{\perp 1} \geq \left(\lambda_2/\lambda_1\right)^{-2 \gamma_\perp}, \text{ and } p_{\parallel 2}/p_{\parallel 1} \geq \left(\lambda_2/\lambda_1\right)^{-2 \gamma_\parallel}$$

Given the MMS observations, Akhavan-Tafti et al. (2018) found that plasma density of subsolar FTEs is, to a first-order approximation, a linear function of the flux rope core magnetic field, $n [\text{cm}^{-3}] = -0.36|B|[\text{nT}] + 43.81$. Here we find that

$$P_{\perp 2} / P_{\parallel 1} \geq \left(\lambda_2 / \lambda_1 \right)^{-5.9}, \text{ and } P_{\parallel 2} / P_{\parallel 1} \geq \left(\lambda_2 / \lambda_1 \right)^{+1.9}$$

which suggests that for small-scale FTEs to grow macroscale via adiabatic expansion their internal pressure must become dominantly parallel. This is not physically possible since the strong anisotropy will trigger plasma instabilities, that is, firehose instability, to return the plasma population to a lower energy state. Dissipative mechanisms and/or plasma energization mechanisms including magnetic reconnection can further thermalize plasma and prevent anisotropy.

Acknowledgments

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