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3 DRAFT VERSION NOVEMBER 13, 2017
4 Typeset using L^AT_EX **manuscript** style in AASTeX61
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8 A BINARY NATURE OF THE MARGINAL CP STAR SIGMA SCULPTORIS
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28 (Received November 9, 2017; Revised —; Accepted —)
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30 Submitted to PASP
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ABSTRACT

The A2 V star σ Scl was suspected to be a low-amplitude periodic variable of the Ap type by several authors. Aiming at deciding whether the star is a variable chemically peculiar (CP) star, we searched for the photometric and spectroscopic variability and determined chemical abundances of σ Scl. The possible variability was tested using several types of periodograms applied to the photometry from Long-Term Photometry of Variables project (LTPV) and Hipparcos. Sixty spectrograms of high quality were obtained and used for chemical analysis of the stellar atmosphere and for looking for spectral variability that is symptomatic for the CP stars. We neither find any signs of the light

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variability nor prominent chemical peculiarity, specific for the CP stars. The only exception is the abundance of scandium, which is significantly lower than the solar one. As a by-product of the analysis using 89 spectroscopic observations, we found that σ Scl is a single-lined spectroscopic binary with orbital period of 46.876(8) d. We argue that σ Scl is not an Ap star, but a marginal Am star in SB1 system. The spectral energy distribution of the binary reveals infrared excess due to circumstellar material.

Keywords: stars: chemically peculiar – stars: early type – stars: variables – binaries:
spectroscopic – stars: individual σ Scl

1. INTRODUCTION

The peculiarity of a part of the upper main sequence stars manifests itself through enhanced spectral lines of some chemical elements and weakened lines of helium, calcium and scandium. Even if this phenomenon is known since the end of the nineteenth century (Maury 1897), several fundamental questions remain unanswered up to now.

The peculiarity of stars with spectra were classified as „Ap” or „Bp”, was explained as a consequence of overabundance and deficit of chemical elements in the their atmospheres, which differ from the standard solar abundance (for example Preston 1974). It is accepted that such phenomenon is a consequence of a slow atomic diffusion or settling in calm upper layers of these stars, then designated as chemically peculiar „CP” stars (Michaud 1970). A further key feature of the classical Ap and Bp stars are periodic changes in the strength and radial velocity of spectral lines and a weak or moderate photometric variability of the same period.

Babcock (1947) discovered a global dipolar magnetic field in the star 78 Virginis and catalogized similar stars (Babcock 1958) in which also the variability of the field strength in many Ap and Bp stars – including even a reversal of magnetic polarity – was unveiled. Stibbs (1950) introduced the „Oblique Rotator” model representing a rigidly rotating star with a non-coincident magnetic and rotational axes. Due to the uneven distribution of chemical elements on the stellar surface, the spectral and the related photometric variabilities have been easily understood too. However, the expected strict relationship between the locations of spectral and photometric spots on the magnetic chemically peculiar (mCP) stars and their magnetic field geometry is documented in only an untold group of mCP stars.

Recently, Rusomarov et al. (2015) show that surface abundance maps of HD 24712 are inconsistent with the predictions of the current theory of atomic diffusion in the presence of magnetic fields. Krtićka et al. (2007) concluded that although the agreement between the observed and computed light curves seems satisfactory, other effects like e.g. NLTE may play a major role. Shulyak et al. (2010) summarize that the simulated light variability of ϵ UMa is in very close agreement with the observations, however, the theoretical variation of its β index is approximately one order of magnitude

smaller than the observed. [Krtička et al. \(2013\)](#) show that neither maximum overabundance of helium and silicon on the surface of the low amplitude CP star HD 64740 do not cause a substantial light variability.

For this study, we selected a bright star, σ Scl (HR 293), which is labelled in the SIMBAD database as an α^2 Canum Venaticorum (ACV) type star. This group of variable stars is known to have distinctive photometric variability with an amplitude of up to tenth of magnitude ([Hümmerich et al. 2016](#)). However, while [Rufener & Bartholdi \(1982\)](#) observed a variability with an amplitude of 0.06 mag in V filter, the revised version of the Geneva catalogue by [Rufener \(1988\)](#) claimed a smaller amplitude. [De Rosa et al. \(2014\)](#) included σ Scl in the Volume-limited A-STAR (VAST) survey and reported its duplicity.

In this study, we review all available photometric observations. We also obtained a series of high-resolution, high S/N ratio spectra to analyse the light variability and duplicity.

2. PHOTOMETRIC VARIABILITY

The star σ Scl (HD 6178, HIP 4852, and HR 293) is of A2V spectral type ([Abt & Morrell 1995](#)), Johnson's $B = 5.595$ mag, and $V = 5.509$ mag ([Wenger et al. 2000](#)). On the other hand, [Manfroid & Renson \(1994\)](#) analysed it photometrically in $uvby$ and concluded it as an Ap (mCP) star.

The star was observed within the years 1988-1993 in the Strömgren photometric system in the frame of the Long-Term Photometry of Variables project (LTPV, see [Sterken \(1983, 1986\)](#); [Manfroid et al. \(1991a,b\)](#); [Sterken et al. \(1993\)](#)), in which σ Scl served as a comparison star for α Scl. In total, 390 quaternions of $uvby$ measurements were published by [Manfroid & Renson \(1994\)](#), who confirmed the weak variability of σ Scl. They also indicated a few possible periods, with a most likely one $P = 2^d36975(20)$, while $P = 1^d37675(10)$ could be its possible alias.

We searched for a possible periodic variability using the LTPV data (1560 individual measurements) and H_p measurements by Hipparcos satellite (98 measurements, [ESA 1998](#)). We conclude that the amplitude of the star variability (if any) has to be evidently smaller than 0.006 mag. Especially, no

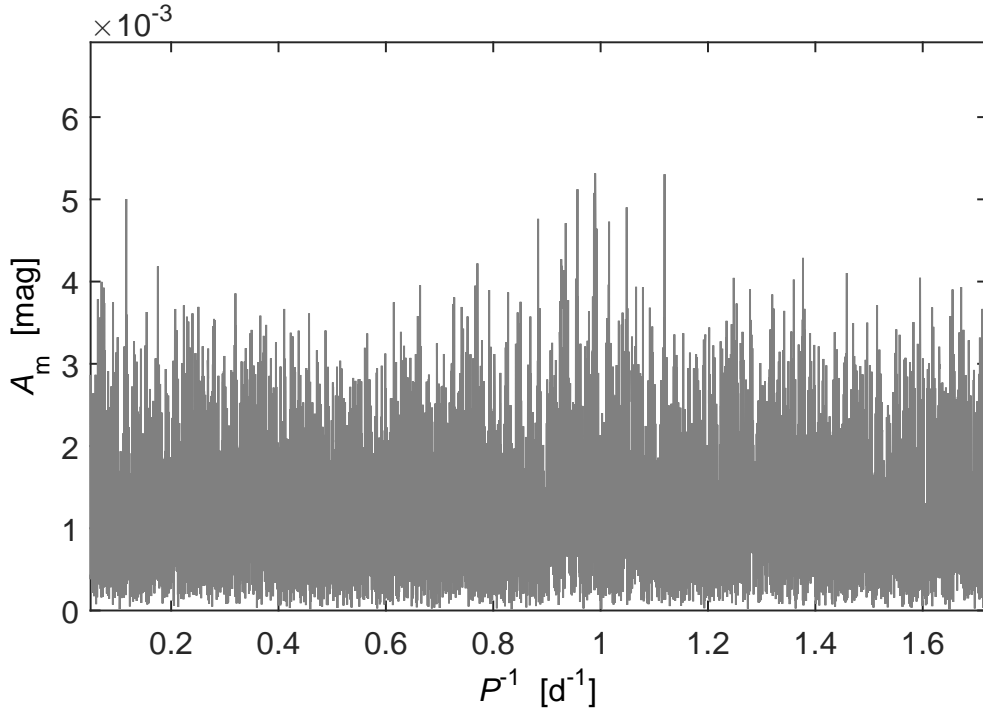


Figure 1. A periodogram based on the LTPV and H_p data where A_m is the amplitude of light variations in mag.

one from three types of periodograms used (Paunzen et al. 2013; Mikulášek et al. 2015), indicates any of the periods mentioned by Manfroid & Renson (1994).

The frequency spectrum of the star is almost featureless (see Fig. 1). The confidential bootstrap test (Paunzen et al. 2013; Mikulášek et al. 2015) approved that the significance of all relevant periods indicated using periodograms is smaller than 10 %, and thus the periodicity of photometric variability should be considered as unproven.

3. SPECTROSCOPIC VARIABILITY

As a next step, we investigated the spectroscopic characteristics of σ Scl.

3.1. Observations

The observations were performed with four different echelle spectrographs. First, we used the high dispersion echelle spectrograph FEROS ($R = 48\,000$) attached to the 2.2 MPG/ESO telescope at La Silla Observatory (ESO) in Chile. We collected 60 spectra during two observational runs (34

spectra in 2011 and 26 in 2012, respectively), with a total exposure time 530 minutes. The spectra **ar** listed in Table 1. Each individual spectrum has an integrating time 600 s reaching $S/N \approx 500$ with the exception of spectra for HJD=2 455 776.76378 – 2 455 776.96153, which have a half of the exposure time resulting in $S/N \approx 250$. The second instrument was the high dispersion echelle spectrograph GIRAFFE ($R = 23\,000$) attached to the 1.9-m telescope at South African Astronomical Observatory (SAAO). We obtained 3 spectra during two nights in 2011. Each individual spectrum has an integrating time of 1200 s reaching $S/N \approx 130$.

The third and fourth echelle spectrographs were REOSC ($R = 8\,600$) and high dispersion EBASIM ($R = 31\,000$) attached to the Jorge Sahade 2.15-m telescope at El Leoncito Observatory (CASLEO) in Argentina. The exposure times and S/N ratios are listed in Table 1.

The data were reduced with the standard IRAF¹ routines. The spectra were bias-corrected and flat-fielded. We used the Th-Ar-Ne comparison spectrum for a precise wavelength calibration, which was sensed through a second fiber simultaneously with the spectrum of the star through the first fiber in case for FEROS spectra. For the other instruments we used for calibration Th-Ar spectra which were measured before and after the stellar spectrum.

Table 1. List of spectra obtained using FEROS@La Silla (A), GIRAFFE@SAAO (B), REOSC@CASLEO (C), and EBASIM@CASLEO (D). HJD is the mean heliocentric Julian date reduced by 2 450 000. Phase φ is calculated according T_0 and P from Table 2.

HJD	φ	In.	Reg.[nm]	τ [s]	S/N	v_{rad} [km s ⁻¹]	HJD	φ	In.	Reg.[nm]	τ [s]	S/N	v_{rad} [km s ⁻¹]
5775.890	0.4087	A	370–920	600	500	-0.99 ± 0.92	6157.927	0.5585	A	370–920	600	500	-5.87 ± 0.73
5775.916	0.4093	A	370–920	600	500	-0.98 ± 0.69	6158.805	0.5772	A	370–920	600	500	-6.04 ± 0.78
5775.940	0.4098	A	370–920	600	500	-0.96 ± 0.84	6158.861	0.5784	A	370–920	600	500	-6.20 ± 0.77
5776.762	0.4273	A	370–920	300	250	-1.66 ± 0.86	6158.914	0.5796	A	370–920	600	500	-6.30 ± 0.78

Table 1 continued on next page

¹ IRAF is distributed by NOAO, which is operated by AURA, Inc., under cooperative agreement with the National Science Foundation.

Table 1 (*continued*)

HJD	φ	In.	Reg.[nm]	τ [s]	S/N	v_{rad} [km s ⁻¹]	HJD	φ	In.	Reg.[nm]	τ [s]	S/N	v_{rad} [km s ⁻¹]
5776.778	0.4277	A	370–920	300	250	-1.74 ± 0.87	6158.929	0.5799	A	370–920	600	500	-6.24 ± 0.75
5776.794	0.4280	A	370–920	300	250	-1.61 ± 0.82	6159.642	0.5951	A	370–920	600	500	-6.93 ± 0.70
5776.809	0.4283	A	370–920	300	250	-1.57 ± 0.70	6159.697	0.5963	A	370–920	600	500	-6.88 ± 0.65
5776.825	0.4287	A	370–920	300	250	-1.50 ± 0.75	6159.722	0.5968	A	370–920	600	500	-6.83 ± 0.63
5776.840	0.4290	A	370–920	300	250	-1.48 ± 0.78	6159.747	0.5973	A	370–920	600	500	-6.97 ± 0.74
5776.855	0.4293	A	370–920	300	250	-1.56 ± 0.81	6159.772	0.5979	A	370–920	600	500	-6.97 ± 0.70
5776.870	0.4296	A	370–920	300	250	-1.64 ± 0.90	6159.797	0.5984	A	370–920	600	500	-6.79 ± 0.69
5776.885	0.4300	A	370–920	300	250	-1.23 ± 0.73	6159.822	0.5989	A	370–920	600	500	-6.73 ± 0.63
5776.900	0.4303	A	370–920	300	250	-1.44 ± 0.77	6159.847	0.5995	A	370–920	600	500	-6.98 ± 0.78
5776.914	0.4306	A	370–920	300	250	-1.43 ± 0.65	6159.872	0.6000	A	370–920	600	500	-6.76 ± 0.70
5776.930	0.4309	A	370–920	300	250	-1.37 ± 0.79	6159.888	0.6003	A	370–920	600	500	-6.97 ± 0.72
5776.945	0.4312	A	370–920	300	250	-1.50 ± 0.75	6159.903	0.6007	A	370–920	600	500	-6.89 ± 0.66
5776.960	0.4316	A	370–920	300	250	-1.52 ± 0.82	6159.919	0.6010	A	370–920	600	500	-6.94 ± 0.65
5777.718	0.4477	A	370–920	600	500	-1.86 ± 0.78	6159.935	0.6013	A	370–920	600	500	-7.06 ± 0.73
5777.744	0.4483	A	370–920	600	500	-1.86 ± 0.76	7173.910	0.2319	C	536–782	1800	100	-2.32 ± 1.90
5777.770	0.4488	A	370–920	600	500	-1.67 ± 0.73	7200.847	0.8065	D	483–685	900	90	-17.83 ± 1.17
5777.794	0.4494	A	370–920	600	500	-1.94 ± 0.76	7200.860	0.8068	D	483–685	900	90	-17.28 ± 1.25
5777.819	0.4499	A	370–920	600	500	-1.97 ± 0.87	7201.900	0.8290	D	483–685	900	90	-19.95 ± 1.41
5777.845	0.4504	A	370–920	600	500	-1.95 ± 0.83	7201.911	0.8292	D	483–685	900	90	-19.89 ± 1.28
5777.870	0.4510	A	370–920	600	500	-1.88 ± 0.78	7204.904	0.8931	D	399–568	1800	90	-20.53 ± 0.86
5779.722	0.4905	A	370–920	600	500	-3.18 ± 0.92	7205.859	0.9134	D	399–568	1800	90	-21.29 ± 0.83
5779.747	0.4910	A	370–920	600	500	-3.20 ± 0.78	7584.874	0.9987	C	477–632	1200	300	-14.74 ± 1.14
5779.771	0.4915	A	370–920	600	500	-2.95 ± 0.71	7585.824	0.0190	C	477–632	1200	300	-17.14 ± 1.32
5779.785	0.4918	A	370–920	600	500	-3.08 ± 0.76	7585.883	0.0203	C	477–632	1200	300	-14.47 ± 1.32
5779.812	0.4924	A	370–920	600	500	-2.98 ± 0.81	7586.824	0.0403	C	477–632	1200	300	-15.34 ± 1.47
5779.840	0.4930	A	370–920	600	500	-3.18 ± 0.79	7586.864	0.0412	C	477–632	1200	300	-10.65 ± 1.55
5779.865	0.4935	A	370–920	600	500	-3.08 ± 0.79	7587.818	0.0615	C	477–632	1200	300	-16.46 ± 1.34
5779.895	0.4942	A	370–920	600	500	-3.18 ± 0.84	7587.858	0.0624	C	477–632	1200	300	-13.22 ± 1.48
5779.914	0.4946	A	370–920	600	500	-3.26 ± 0.86	7588.796	0.0824	C	477–632	1200	300	-20.08 ± 1.51
5779.946	0.4952	A	370–920	600	500	-3.36 ± 0.93	7588.839	0.0833	C	477–632	1200	300	-11.33 ± 1.57

Table 1 continued on next page

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Table 1 (*continued*)

HJD	φ	In.	Reg.[nm]	τ [s]	S/N	v_{rad} [km s $^{-1}$]	HJD	φ	In.	Reg.[nm]	τ [s]	S/N	v_{rad} [km s $^{-1}$]
5790.590	0.7223	B	418–682	1200	130	-10.00 ± 1.08	7652.744	0.4466	D	503–701	3600	90	-1.47 ± 0.65
5790.662	0.7239	B	418–682	1200	130	-13.33 ± 1.66	7652.880	0.4495	D	503–701	3600	90	-0.83 ± 0.65
5793.661	0.7878	B	418–682	1200	130	-14.47 ± 1.41	7653.690	0.4668	D	503–701	3600	90	-2.23 ± 0.98
6157.728	0.5543	A	370–920	600	500	-5.56 ± 0.65	7653.850	0.4702	D	503–701	3600	90	-2.08 ± 1.20
6157.753	0.5548	A	370–920	600	500	-5.91 ± 0.77	7655.708	0.5098	D	416–547	3600	100	-3.30 ± 0.98
6157.778	0.5553	A	370–920	600	500	-5.79 ± 0.75	7655.841	0.5126	D	416–547	3600	100	-4.69 ± 0.90
6157.815	0.5561	A	370–920	600	500	-5.86 ± 0.76	7656.654	0.5300	D	416–547	3600	100	-5.14 ± 0.87
6157.840	0.5567	A	370–920	600	500	-5.80 ± 0.67	7656.812	0.5334	D	416–547	3600	100	-4.75 ± 1.06
6157.865	0.5572	A	370–920	600	500	-5.74 ± 0.68	7658.742	0.5745	D	436–581	3600	100	-5.73 ± 0.99
6157.896	0.5579	A	370–920	600	500	-5.81 ± 0.66	7658.873	0.5773	D	436–581	3600	100	-4.79 ± 1.26
6157.911	0.5582	A	370–920	600	500	-5.77 ± 0.72							

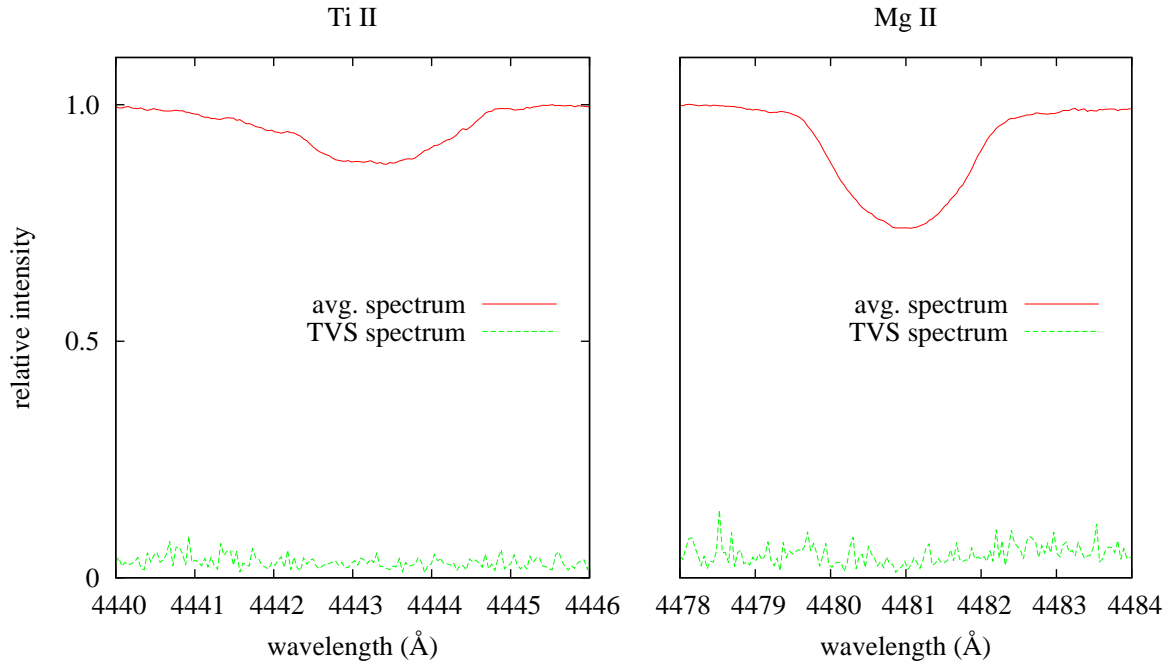


Figure 2. Averaged profiles of the spectral lines Ti II 4443.8007 Å (left) and Mg II 4481.327 Å (right) from night HJD 56159 and their TVS spectra.

3.2. Stellar parameters

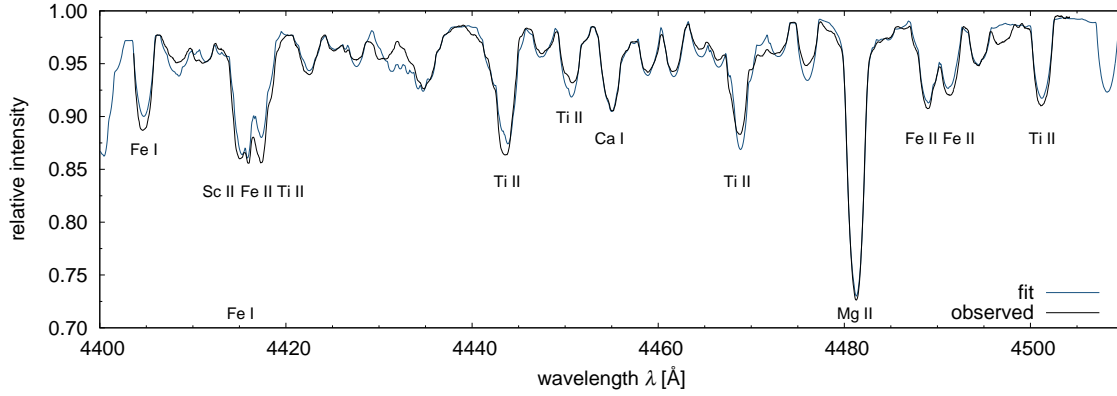


Figure 3. Comparison of the observed spectrum and the fit by the synthetic spectrum in a selected region of a FEROS spectrum.

Table 2. Stellar parameters of σ Scl

Orbital period	$P = 46^{\text{d}}.877(8)$
JD time of periastron passage	$T_0 = 2\,456\,225.5(7)$
Numerical eccentricity	$e = 0.195(25)$
γ velocity	$\gamma = -9.2(2) \text{ km s}^{-1}$
K_1 semiamplitude	$K_1 = -10.3(3) \text{ km s}^{-1}$
Argument of the periastron	$\omega = 45(5)^\circ$

The effective temperature and surface gravity values listed in Table 3 were derived using the codes UVBYBETA Moon & Dworetsky (1985) and TEFFLOGG Smalley & Dworetsky (1995) using the $wby\beta$ measurements (Paunzen 2015). The value of the projected rotational velocity was taken from Zorec & Royer (2012).

The evolutionary tracks for solar metallicity ($[Z] = 0.014$) from Bressan et al. (2012) and the above mentioned astrophysical parameters yield an age of about 500 Myr for σ Scl.

3.3. Line profile and radial velocity variability

We investigated a possible presence of line-profile variations as one of characteristic attributes of Ap stars. We used the temporal variance spectrum technique (TVS), devised by Fullerton et al. (1996).

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Table 3. Stellar parameters of σ Scl

Effective temperature T_{eff}	$8\,700 \pm 125$ K
Surface gravity $\log g$ (cgs)	4.11 ± 0.05
Projected rotational velocity $v_{\text{rot}} \sin i$	79 km s^{-1}

We calculated the self-TVS for last 13 spectra from the same night (HJD 56159) which are listed in Table 1. For two spectral lines in region $4400 - 4500 \text{ \AA}$ Ti II 4443.8007 \AA and Mg II 4481.327 \AA we obtained again a completely negative result (see Fig. 2).

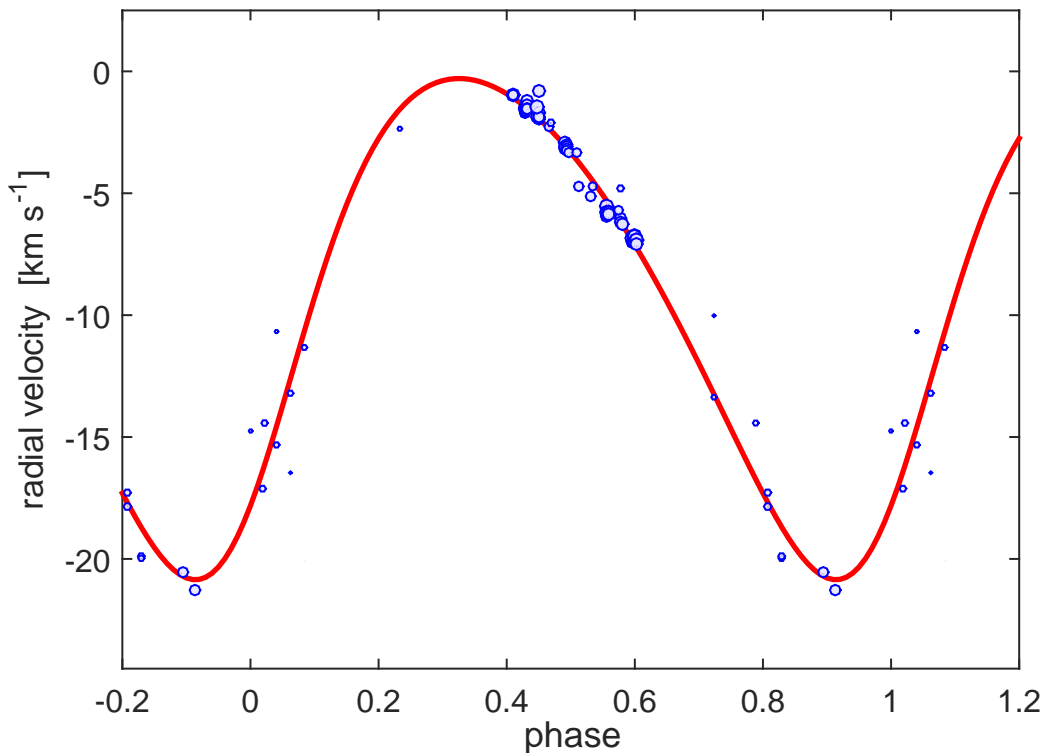


Figure 4. The radial velocity variations of σ Scl show that the star is a primary component of the single-lined spectroscopic binary with the orbital period $P = 46^{\text{d}}877(8)$. The size of circles corresponds to the weights of individual RV measurements.

We also measured the individual radial velocity for each of the 89 spectra (see Table 1) by means of a cross-correlation function (CCF) using the theoretical spectrum as a template (Zverko et al. 2007). We processed each spectral order of echelle spectra individually and calculated the mean value for

each observation. Our analysis based on inspection of 89 spectrograms from four spectrographs (see Table 1) confirmed the presence of the radial velocity variability reported in the literature (Neubauer 1930; Wilson 1953; Bond et al. 1971; Gontcharov 2006). We have found that the observed radial velocity variation is cyclic with the period $P = 46^{\text{d}}.877(8)$. The course of the RV phase curve (see Fig. 4) can be explained as the consequence of the orbital motion of the primary component bound in a spectroscopic double star with the parameters summarized in Table 2. The radial velocity curve was fitted by own robust regression eliminating the influence of outliers (for details see in Mikulášek & Zejda 2013). Any photometric variability connected with the orbital period of the star have not been revealed.

To detect assumptive spectral lines of the secondary component, we subtracted the average spectrum, which we have gained as sum of all spectra shifted by radial velocity of primary component, from each individual spectrum. Despite our best spectra were integrated to reach $S/N = 500$ no traces of spectral lines of the secondary component have been found. Therefore the secondary must be either a low luminosity star of a cool main-sequence one or a white/brown dwarf. To specify the possible main-sequence companion closely, we estimated its upper limit of the luminosity and mass. As an example, we consider the first spectrum introduced in Table 1. Provided that a signal safely identified in a noisy continuum amounts at least 3σ , the central depth of the weakest line visible in the spectrum is $D_c = 0.006$ in the scale of the continuum. The spectrum of a cool main-sequence star is covered by a plenty of metallic lines, the central depth of which reaches to $D_{\text{sec}} = 0.9$. Then the ratio of the lights $L_{\text{sec}}/L_{\text{prim}} = 0.006/0.9 = 0.0067$, that results in the magnitude difference between the A2V primary and the secondary $\Delta M = 5.44$ mag. Considering that primary star has absolute magnitude $M_p(V) = 1.3$ mag, secondary component should then have $M_s(V) = 6.8$ mag, that corresponds to a K3V main-sequence star (Drilling & Landolt 1999). The mass of it is $\approx 0.7 M_{\odot}$, and with the mass ratio of an A2V to K3V-star is $2.6/0.7 = 3.7$ and the radial velocity of the secondary in the orbital phase of the observed spectrum considered is $\approx -38 \text{ km s}^{-1}$. We do not identify lines of the secondary shifted with this value.

3.4. Chemical composition

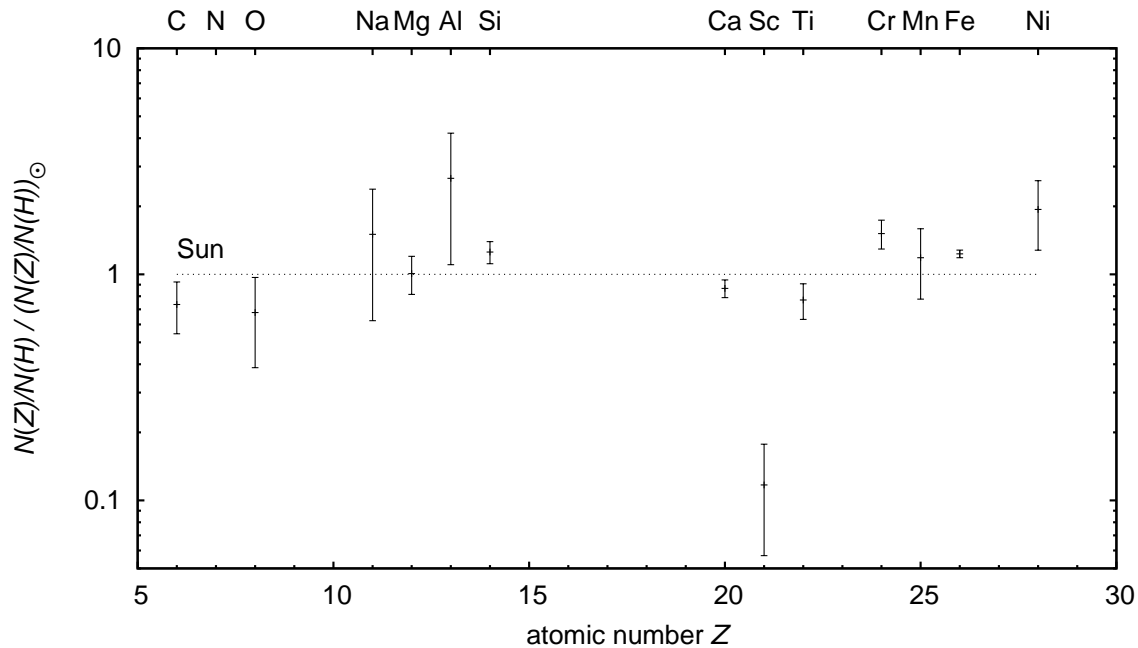
The chemical composition was derived using a χ^2 minimization of the synthetic spectra for each of the FEROS orders separately (see Fig. 3). Synthetic spectra were computed using the SYNSPEC code (Hubeny & Lanz 2011), ATLAS12 model atmospheres (Bischof 2005; Castelli 2005; Kurucz 2005), and atomic data from the VALD database (Piskunov et al. 1995; Kupka et al. 1999) and also from Lanz & Hubeny (2007). The observed spectra for abundance analysis were averaged over all observations. We used a simplex method with abundances of individual chemical elements as free parameters, and we fixed stellar parameters to that given in Table 3. The final abundances of σ Scl are given in Table 4 (see also Fig. 5). The abundances of individual elements were typically derived from numerous absorption lines present in the studied spectral range. For most elements, the lines were present at least in three orders of FEROS spectrum, and therefore we obtained at least three independent estimates of abundance for each element. The exception is Na and Al, for which the abundances were derived from only 5683 Å and 5688 Å lines in the case of Na and 4663 Å line in the case of Al. The abundances are consistent with solar chemical composition (Asplund et al. 2009) except a strong underabundance of scandium with a possible slight overabundance of chromium, iron, and nickel.

This may indicate that σ Scl is a marginal Am (CP1) star. This group of chemically peculiar stars does not have a significant global magnetic field and is preferably found in close binary systems (Adelman 1999; Hui-Bon-Hoa 2000). Almost all CP1 stars seem to be rather evolved with ages above 400 Myr (Künzli & North 1998). The derived age of σ Scl (about 500 Myr) is consistent with this conclusion.

The observed abundance pattern of CP1 stars is explained by the diffusion of elements together with the disappearance of the outer convection zone associated with the helium ionization because of gravitational settling of helium (Michaud et al. 1983). They predict a cut-off rotational velocity for such objects ($\approx 90 \text{ km s}^{-1}$), above which meridional circulation leads to a mixing in the stellar atmosphere. σ Scl is very close to this value which makes it very interesting for testing the models where the transit to marginal CP1 stars takes place.

Table 4. Derived σ Scl abundances

Elem.	$\log(N_{\text{el}}/N_{\text{H}}) + 12$		Elem.	$\log(N_{\text{el}}/N_{\text{H}}) + 12$	
	σ Scl	Sun		σ Scl	Sun
C	8.30 ± 0.11	8.43	Sc	2.22 ± 0.22	3.15
O	8.52 ± 0.19	8.69	Ti	4.83 ± 0.08	4.95
Na	6.42 ± 0.25	6.24	Cr	5.82 ± 0.06	5.64
Mg	7.60 ± 0.08	7.60	Mn	5.50 ± 0.15	5.43
Al	6.87 ± 0.25	6.45	Fe	7.59 ± 0.02	7.50
Si	7.61 ± 0.05	7.51	Ni	6.51 ± 0.15	6.22
Ca	6.28 ± 0.04	6.34			

**Figure 5.** Derived σ Scl abundances relative to the Sun.

3.5. Analysis of the Spectral Energy Distribution

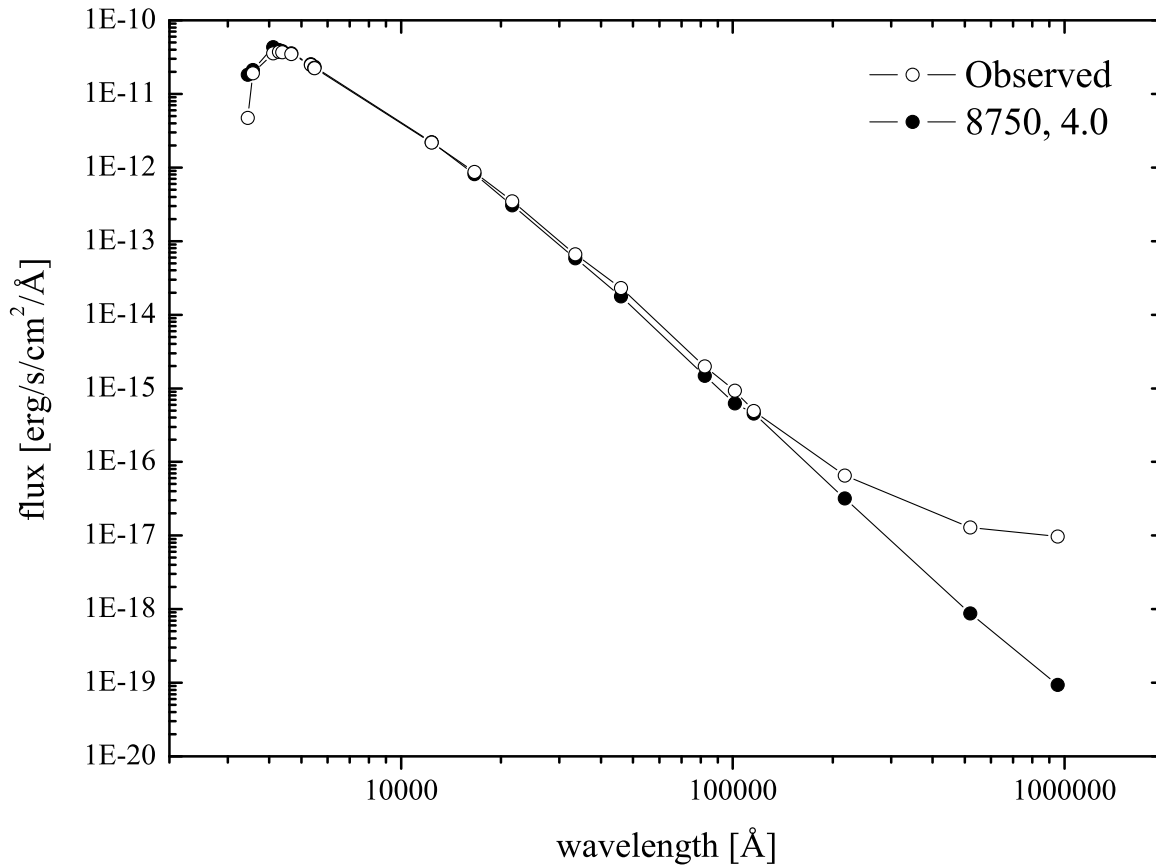


Figure 6. The SED of σ Scl fitted to all available photometry and the synthetic flux from an appropriate stellar atmosphere (Castelli & Kurucz 2004). A flux excess from about $1.5 \mu\text{m}$ is evident. The infrared excess beyond $20 \mu\text{m}$ is probable due to circumstellar material.

In order to put further constraints on the characteristics of the secondary component, we have investigated the Spectral Energy Distribution (SED) of σ Scl.

For this, the VO Sed Analyzer (VOSA v5.1, Bayo et al. 2008) tool, was used. All the available photometric data were included into the fitting process. We used the synthetic flux from a standard Kurucz model (Castelli & Kurucz 2004) with $T_{\text{eff}} = 8750 \text{ K}$, $\log g = 4.11$, and solar metallicity.

Figure 6 shows the observed and synthetic fluxes. We face two different phenomena. First, there is a flux excess visible which emerges with a continuous increase from $1.5 \mu\text{m}$. Therefore, we have calculated the difference of the fluxes and fitted a simple black body radiation curve. A temperature between 1500 and 2000 K reproduces the flux distribution well. The origin of this excess is unclear.

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4 The second, much stronger infrared excess, is visible beyond 20 μm which can be attributed cir-
5 cumstellar material within this system. It is well known that a significant number of stars in the solar
6 neighbourhood exhibit infrared excess due to circumstellar material (Cotten & Song 2016). Most of
7 these stars are not on the zero-age-main-sequence but already evolved on the main-sequence. The
8 origin of this circumstellar material is still unknown. σ Scl provides an excellent opportunity to
9 investigate its origin and characteristics within a binary system of a very low mass ratio.
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17 4. NACO DATA REDUCTION AND ANALYSIS

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19 For σ Scl (also HIP 4852) high-angular resolution observations are available, taken with the Adaptive
20 Optics camera NACO (for Nasmyth Adaptive optics system with the COude near-infrared imager
21 and spectrograph, Rousset et al. 2003; Lenzen et al. 2003) installed at the Nasmyth B focus of the
22 VLT/UT4 from 2001 through 2013.
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27 The data-set of σ Scl consists of one epoch, imaged in November 2004² with the IB_2.18 filter band
28 ($\lambda = 2.18 \pm 0.06 \mu\text{m}$) and employing the NACO S27 optics. The total integration time is 7.5 min.
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32 For the data reduction of the NACO images we used appropriate calibration data from the ESO
33 data-archive, i.e. darks with the same exposure time, and flats taken in the same filter and in the
34 same night as the science data. In order to reduce the imaging data we used ESO ECLIPSE/jitter³.
35 For each science frame a master-dark subtraction and flat-field correction was applied, and all images
36 were averaged using the provided shift+add procedure including measurement and subtraction of the
37 bright background of the sky in the near infrared.
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44 From the reduced the NACO IB_2.18 image of HIP4852, taken with the S27 optics, we measured the
45 achieved (S/N = 3) detection limit versus angular and projected separation. To measure the dynamic
46 range as a function of the angular separation, a one-dimensional sensitivity curve was created using
47 the following procedure. First, for each pixel on the detector the distance to the previously determined
48 position of the central star was calculated, and the corresponding measured flux F stored. Then, n
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56 ² Public data from the ESO data archive, taken in ESO program 074.D-0180(A)

57 ³ ESO C Library for an Image Processing Software Environment.
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annuli of equal width were selected within which the flux values were sorted, the upper and lower 10% excluded and those values rejected which deviate by more than 3σ from the median in order to remove the contribution of the stars contained in the annulus from the calculation. The standard deviation of the remaining values was then used as estimate of the background noise. The resulting 3σ contrast level as a function of the angular separation is finally evaluated by comparison of the estimated background noise to the peak flux of the central star. The procedure described above was also applied to the PSF-subtracted image. However, the differences between the two resulting curves were only marginal. The main reason for this is that the main source of the noise is the photon-noise which is superimposed with the measured signal. The removal of the signal by subtraction of the PSF does not significantly change the measured residual noise, and thus the computed sensitivity remains almost equal. The result is illustrated in 7.

No additional source could be detected within the fully covered field of view, i.e. at angular separations smaller than about 12 arcsec (~ 855 au of projected separation) around the central star σ Scl. In the background noise limited region a ($S/N = 3$) detection-limit of about 14 mag is reached at angular separations from HIP 4852 A beyond about 1.5 arcsec. At the derived age of σ Scl of about 500 Myr this allows the detection of low-mass stellar companions with masses down to 0.09 solar masses and projected separations from HIP 4852 A of more than about 107 au.

5. CONCLUSIONS

We reanalysed sets of photometric data available in the literature for suspected Ap star σ Scl to check them for periodicity. We do not confirm the periods indicated in the literature. We have not detected any significant light and line profile variations. The absence of line profile variations and mostly solar chemical composition support the classification of σ Scl as non-variable marginal Am main-sequence star. However, a detailed analysis of radial velocity variations revealed that σ Scl is a spectroscopic binary (SB1) with the period of 46.876(8) d. We did not find traces of the secondary component in the spectrum, but we found a flux excess from $1.5\mu\text{m}$ in spectral energy distribution. Yet larger flux excess is visible beyond $20\mu\text{m}$ and is caused probably due to circumstellar material.

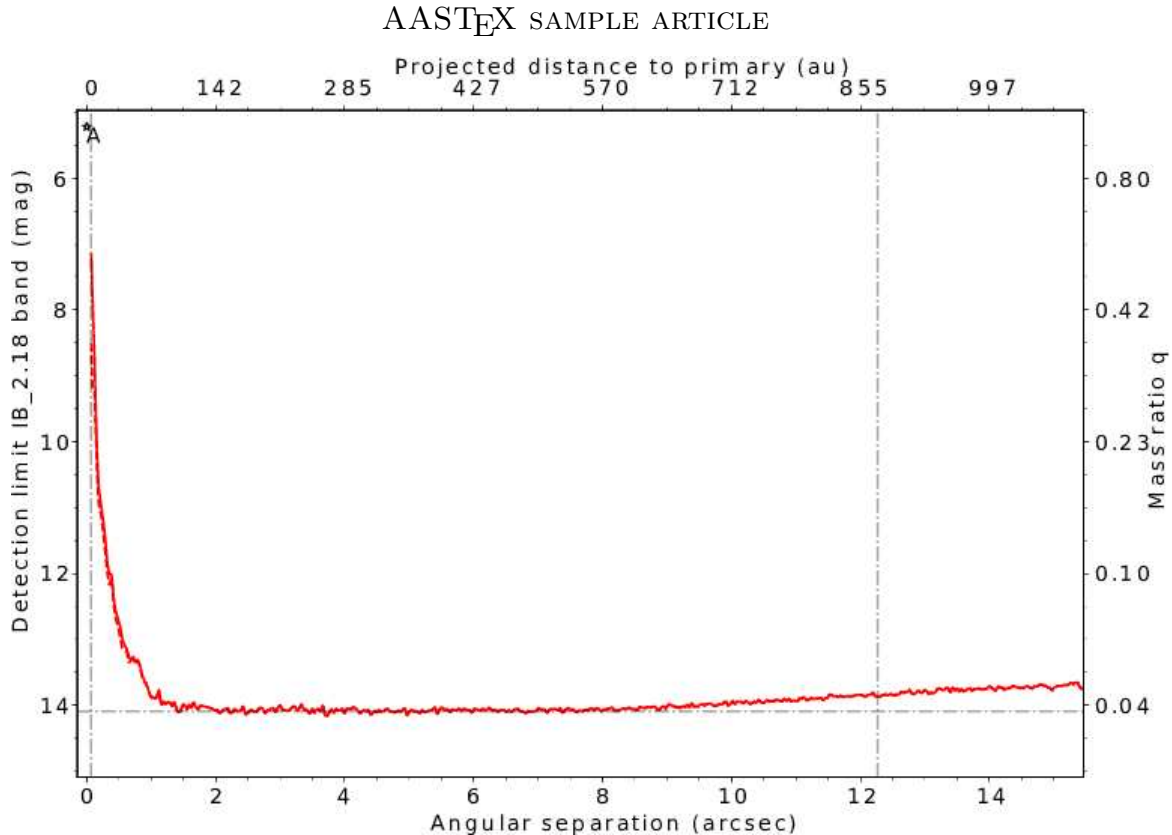


Figure 7. Dynamic range plot for σ Scl observed with NaCo in November 2004. We plot the achieved apparent magnitude (left y-axis) in the IB_2.18 filter band as function of angular (bottom x-axis) and projected separation (top x-axis), in au and assuming a distance of 71.2 ± 1.6 pc (McDonald et al. 2017). The position of the primary (A) is marked with an asterisk. The measured 3σ sensitivity curve is shown as full (red) line. Objects below this line cannot be detected. The measured sensitivity curve from the PSF-subtracted image is also shown as dashed (red) line. The detectable mass ratios q (right y-axis) obtained from evolutionary models by Bressan et al. (2012) are indicated as well as the diffraction limit (left vertical dash-dotted line) of the telescope and the closest image border to the central star (right vertical dash-dotted line) are given. The background-level remains constant at about 14 mag for separations outside of $1.5''$ ($\sim 0.09 M_{\odot}$ at 500 Myr). No visual companion were detected within these limits.

This work was supported by the grants GA ĀR 16-01116S as well as MŠMT 7AMB14AT015 and 7AMB17AT030. Olga Pintado is a Visiting Astronomer, Complejo Astronómico El Leoncito operated under agreement between the Consejo Nacional de Investigaciones Científicas y Técnicas de la República Argentina and the National Universities of La Plata, Córdoba and San Juan.

The research is based on observations obtained at the European Southern Observatory (ESO programmes 087.D-0099 and 089.D-0153), at the South African Astronomical Observatory (SAAO) and at the El Leoncito Astronomical Complex (CASLEO).

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