

Extending the Calibration of C IV-based Single-epoch Black Hole Mass Estimators for Active Galactic Nuclei*

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Abstract

We provide an updated calibration of C IV $\lambda 1549$ broad emission line–based single-epoch (SE) black hole (BH) mass estimators for active galactic nuclei (AGNs) using new data for six reverberation-mapped AGNs at redshift $z = 0.005\text{--}0.028$ with BH masses (bolometric luminosities) in the range $10^{6.5}\text{--}10^{7.5} M_{\odot}$ ($10^{41.7}\text{--}10^{43.8} \text{ erg s}^{-1}$). New rest-frame UV-to-optical spectra covering $1150\text{--}5700 \text{ \AA}$ for the six AGNs were obtained with the *Hubble Space Telescope* (*HST*). Multicomponent spectral decompositions of the *HST* spectra were used to measure SE emission-line widths for the C IV, Mg II, and $H\beta$ lines, as well as continuum luminosities in the spectral region around each line. We combine the new data with similar measurements for a previous archival sample of 25 AGNs to derive the most consistent and accurate calibrations of the C IV-based SE BH mass estimators against the $H\beta$ reverberation-based masses, using three different measures of broad-line width: full width at half maximum (FWHM), line dispersion (σ_{line}), and mean absolute deviation (MAD). The newly expanded sample at redshift $z = 0.005\text{--}0.234$ covers a dynamic range in BH mass (bolometric luminosity) of $\log M_{\text{BH}}/M_{\odot} = 6.5\text{--}9.1$ ($\log L_{\text{bol}}/\text{erg s}^{-1} = 41.7\text{--}46.9$), and we derive the new C IV-based mass estimators using a Bayesian linear regression analysis over this range. We generally recommend the use of σ_{line} or MAD rather than FWHM to obtain a less biased velocity measurement of the C IV emission line, because its narrow-line component contribution is difficult to decompose from the broad-line profile.

Key words: galaxies: active – galaxies: nuclei – methods: statistical

1. Introduction

Understanding the cosmic growth of the supermassive black hole (BH) population and the coevolution of BHs with their host galaxies is now recognized to be one of the essential ingredients for a complete picture of galaxy formation and evolution (see Ferrarese & Ford 2005 and Kormendy & Ho 2013). To probe the high-redshift BH population and the evolution of BH-galaxy scaling relations over cosmic time, it is essential to have reliable methods to determine BH masses in distant active galactic nuclei (AGNs; Shen 2013).

The rest-frame UV C IV $\lambda 1549$ broad emission line is commonly used for BH mass estimates in high-redshift AGNs (i.e., $2 \lesssim z \lesssim 5$) when single-epoch (SE) optical spectra are available. The method of deriving SE mass estimates based on broad-line widths and continuum luminosities in quasar spectra relies on reverberation-mapped (RM) AGNs for its fundamental calibration.¹² Achieving an accurate calibration of C IV-based SE

BH mass estimators using the most reliable AGN BH mass estimates obtained from RM is thus important for improving the precision and accuracy of SE mass estimates for AGNs. Due to a lack of direct C IV RM measurements, however, C IV SE calibration has been performed against the $H\beta$ RM-based BH masses, which is the best practical approach at present. Note that $H\beta$ is, so far, the most studied and understood emission line in RM studies, with many reliable $H\beta$ -based RM results; it can thus be arguably regarded as the most reliable line for AGN BH mass measurements (see Shen 2013 for a related discussion).

Previously, Vestergaard & Peterson (2006, hereafter VP06) provided a calibration of C IV-based BH mass estimators using a sample of low-redshift AGNs for which both $H\beta$ RM measurements and rest-frame UV spectra were available. Since then, the number of AGNs with BH mass estimates from RM has increased, as has the number of AGNs for which UV spectra have been obtained by the *Hubble Space Telescope* (*HST*). Park et al. (2013, hereafter P13) revisited the calibrations of C IV-based BH mass estimators by taking advantage of high-quality *HST* UV spectra for the RM AGN sample and using improved measurement methods. The P13 sample included 25 AGNs, of which six have estimated $M_{\text{BH}} < 10^{7.5} M_{\odot}$ and only one has $M_{\text{BH}} < 10^{7.0} M_{\odot}$. In order to improve the calibration of SE BH masses at the low end of the AGN mass range ($M_{\text{BH}} \lesssim 10^{7.5} M_{\odot}$), it is important to further expand the sample of AGNs with both RM measurements and *HST* UV spectroscopy. Similarly, a calibration of BH masses

* Based on observations made with the NASA/ESA *Hubble Space Telescope*, obtained at the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS 5-26555. These observations are associated with program GO-12922.

¹⁰ EACOA fellow.

¹¹ NASA Einstein fellow.

¹² See a recent review by Bentz (2016) of the current status and future prospects for RM studies.

based on the broad Mg II $\lambda 2798$ emission line, another commonly used rest-frame UV line at intermediate redshifts, has also been performed (see, e.g., McLure & Jarvis 2002; Wang et al. 2009). As with C IV, there is also much room for improvement in the calibration of Mg II-based BH masses, and extending the calibration to a larger sample of AGNs over a wider dynamic range in BH mass is a high priority.

There have been several efforts in the literature to improve the calibration of C IV-based SE BH mass estimators, e.g., by taking advantage of the ratio of UV to optical continuum luminosities (color dependence; Assef et al. 2011), the ratio of full width at half maximum (FWHM) to σ_{line} of C IV (line shape; Denney 2012), the peak flux ratio of the $\lambda 1400$ feature to C IV (Eigenvector 1; Runnoe et al. 2013a; Brotherton et al. 2015), and the C IV blueshift (Shen & Liu 2012; Coatman et al. 2017).

As an extension of our previous work (P13), this paper presents new *HST* UV and optical spectra of six RM AGNs with BH masses of $10^{6.5}$ – $10^{7.5} M_{\odot}$. High-quality spectra, quasi-simultaneously covering the C IV to H β spectral regions with a consistent aperture size and slit width, were obtained with the Space Telescope Imaging Spectrograph (STIS). The new data enable a consistent comparison between the broad emission lines while minimizing measurement systematics due to time variability or aperture effects.

Using the new spectra, we provide updated calibrations of C IV-based SE BH mass estimators for three different measures of broad-line width: the FWHM and the line dispersion (σ_{line}), which have been commonly used in previous work on SE mass estimates, and the mean absolute deviation (MAD), which was recently suggested by Denney et al. (2016b) to be a useful line width measure for virial mass estimation.

We use a Bayesian linear regression method, which is independently implemented for this work, to carry out the calibration of the C IV virial mass relation. Our method follows the work of Kelly et al. (2012; see also Kelly 2007) using the Stan probabilistic programming language (Stan Development Team 2015a). The Bayesian methodology and model specifications for the linear regression analysis will be described in detail in a forthcoming paper (D. Park 2017, in preparation).

This paper is organized as follows. In Section 2, the calibration sample, *HST* observations, and data reduction procedures are described. In Section 3, we present measurements of the C IV, Mg II, and H β emission lines and comparisons of their profiles. The new calibration of the SE virial mass estimators based on the FWHM, σ_{line} , and MAD of the C IV line profile are presented in Section 4 with a comparison to previous calibrations and a test of methodological differences in the linear regression analysis. We also systematically compare our updated calibration with the corrected prescriptions in the literature for the C IV BH mass calibration described above. We summarize this work and provide discussion in Section 5. The following standard cosmological parameters were adopted to calculate distances: $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.30$, and $\Omega_{\Lambda} = 0.70$, which is the same as used by P13.

2. Sample, *HST* Observations, and Data Reduction

The sample for this work is based on the sample of 25 AGNs (BH mass $\log M_{\text{BH}}/M_{\odot} = 7.0$ – 9.1 , bolometric luminosity¹³ $\log L_{\text{bol}}/\text{erg s}^{-1} = 43.2$ – 46.9 , redshift $z = 0.009$ – 0.234)

from P13, supplemented by six new AGNs at redshift $z = 0.005$ – 0.028 that have low-mass BHs (i.e., $\log M_{\text{BH}}/M_{\odot} = 6.5$ – 7.5) from H β -based RM measurements and low bolometric luminosities (i.e., $\log L_{\text{bol}}/\text{erg s}^{-1} = 41.7$ – 43.8). The P13 sample contains RM AGNs with available archival *HST* spectra selected by taking into account data quality, spectral coverage, and contamination of C IV by absorption features. The enlarged dynamic range in mass for the expanded sample enables us to calibrate the C IV SE virial relationship over almost three orders of magnitude in BH mass. The new targets have been selected from recent RM programs. These include Arp 151, Mrk 1310, NGC 6814, and SBS 1116+583A from the Lick AGN Monitoring Project 2008 campaign (Bentz et al. 2009b; Park et al. 2012b), Mrk 50 from the Lick AGN Monitoring Project 2011 campaign (Barth et al. 2011b), and the *Kepler*-field AGN Zw 229-015 (Barth et al. 2011a). Table 1 summarizes the properties of the P13 AGN sample and the six new objects presented in this work. Note that the virial factor f with its uncertainty is adopted from Park et al. (2012a) and Woo et al. (2010; see also Woo et al. 2013, 2015) and applied to all RM BH masses (i.e., $\log f = 0.71 \pm 0.31$), which is consistent with previous measurements and direct measurements by Pancoast et al. (2012, 2014). The f represents the dimensionless scale factor of order unity that depends on the detailed geometry, kinematics, and inclination of the broad-line region (BLR), which is thus used to convert the measured virial product into actual BH mass ($M_{\text{BH}} = f \times \text{VP}_{\text{BH}}$). The adopted uncertainty (0.31 dex) for the virial factor is derived from the scatter of the AGN $M_{\text{BH}} - \sigma_*$ relation (0.43 dex), which gives an upper limit of random scatter of the virial factor itself after subtracting in quadrature the assumed intrinsic scatter (0.3 dex) of the relation (see also the related discussion in Park et al. 2012b). Note that the virial factor uncertainty is the dominant portion of the error budget for the RM masses, since the measurement uncertainty propagated from the reverberation lags and H β line widths is substantially smaller than this 0.31 dex uncertainty for individual AGNs (see Table 1).

For the six new AGNs, we obtained UV spectra with the STIS as part of *HST* program GO-12922 (PI: Woo). In addition to the UV data, optical spectra were obtained quasi-simultaneously (during the same *HST* visit) with a consistent slit width and aperture size. Note that the temporal gaps between the end of the optical exposures and the start of the UV exposures were less than ~ 6 minutes. Individual exposures in and between UV gratings were obtained within a maximum temporal gap of ~ 50 minutes. The ability to obtain nearly simultaneous UV and optical spectra through a consistent aperture is a unique capability of the STIS instrument and is essential in order to minimize possible systematic biases from AGN variability and different amounts of host galaxy and narrow-line region contributions.

We used the G140L, G230L, and G430L gratings with the 52×0.2 slit (i.e., a long slit of width $0''.2$) to acquire a spectrum covering the Ly α , C IV, Mg II, and H β emission lines for each target. The consistent and small spectroscopic aperture has the benefit of minimizing the contamination from host-galaxy starlight. For the CCD G430L observations, we used the EI aperture position to minimize losses due to the imperfect charge transfer efficiency, as recommended in the STIS instrument handbook. Total integrations of 1170–1464 s for G140L, 627–1471 s for G230L, and 120–200 s for G430L were split into two or three exposures, depending on the grating, and dithered along the slit for optimal cleaning of cosmic-ray hits and bad pixels. The observations are summarized in Table 2.

¹³ The bolometric luminosity is computed as $L_{\text{bol}} = 3.81 \times \lambda L_{1350}$ (see Shen et al. 2008 and references therein).

Table 1
Optical Spectral Properties from H β Reverberation Mapping

Object	z	τ_{cent} (H β) (days)	σ_{rms} (H β) (km s $^{-1}$)	$\log(M_{\text{BH}}/M_{\odot})$ (RM)	References
(1)	(2)	(3)	(4)	(5)	(6)
Sample Presented in P13 ^a					
3C 120	0.03301	27.2 $^{+1.1}_{-1.1}$	1514 \pm 65	7.80 \pm 0.31	6
3C 390.3	0.05610	23.60 $^{+6.45}_{-6.45}$	3105 \pm 81	8.43 \pm 0.33	1
Ark 120	0.03230	39.05 $^{+4.57}_{-4.57}$	1896 \pm 44	8.14 \pm 0.32	1
Fairall 9	0.04702	17.40 $^{+3.75}_{-3.75}$	3787 \pm 197	8.38 \pm 0.32	1
Mrk 279	0.03045	16.70 $^{+3.90}_{-3.90}$	1420 \pm 96	7.51 \pm 0.33	1
Mrk 290	0.02958	8.72 $^{+1.21}_{-1.02}$	1609 \pm 47	7.36 \pm 0.32	4
Mrk 335	0.02578	14.1 $^{+0.4}_{-0.4}$	1293 \pm 64	7.37 \pm 0.31	6
Mrk 509	0.03440	79.60 $^{+5.75}_{-5.75}$	1276 \pm 28	8.12 \pm 0.31	1
Mrk 590	0.02638	24.23 $^{+2.11}_{-2.11}$	1653 \pm 40	7.65 \pm 0.32	1
Mrk 817	0.03145	19.05 $^{+2.45}_{-2.45}$	1636 \pm 57	7.66 \pm 0.32	1
NGC 3516	0.00884	11.68 $^{+1.02}_{-1.53}$	1591 \pm 10	7.47 \pm 0.31	4
NGC 3783	0.00973	10.20 $^{+2.80}_{-2.80}$	1753 \pm 141	7.44 \pm 0.32	1
NGC 4593	0.00900	3.73 $^{+0.75}_{-0.75}$	1561 \pm 55	6.96 \pm 0.32	2
NGC 5548	0.01717	4.18 $^{+0.86}_{-1.30}$	3900 \pm 266	7.80 \pm 0.34	3, 5
NGC 7469	0.01632	4.50 $^{+0.75}_{-0.75}$	1456 \pm 207	7.05 \pm 0.31	1
PG 0026+129	0.14200	111.00 $^{+26.20}_{-26.20}$	1773 \pm 285	8.56 \pm 0.33	1
PG 0052+251	0.15500	89.80 $^{+24.30}_{-24.30}$	1783 \pm 86	8.54 \pm 0.32	1
PG 0804+761	0.10000	146.90 $^{+18.85}_{-18.85}$	1971 \pm 105	8.81 \pm 0.31	1
PG 0953+414	0.23410	150.10 $^{+22.10}_{-22.10}$	1306 \pm 144	8.41 \pm 0.32	1
PG 1226+023	0.15830	306.80 $^{+79.70}_{-79.70}$	1777 \pm 150	8.92 \pm 0.32	1
PG 1229+204	0.06301	37.80 $^{+21.45}_{-21.45}$	1385 \pm 111	7.83 \pm 0.38	1
PG 1307+085	0.15500	105.60 $^{+41.30}_{-41.30}$	1820 \pm 122	8.61 \pm 0.33	1
PG 1426+015	0.08647	95.00 $^{+33.50}_{-33.50}$	3442 \pm 308	9.08 \pm 0.34	1
PG 1613+658	0.12900	40.10 $^{+15.10}_{-15.10}$	2547 \pm 342	8.42 \pm 0.38	1
PG 2130+099	0.06298	12.8 $^{+1.2}_{-0.9}$	1825 \pm 65	7.63 \pm 0.31	6
New Sample Presented Here					
Arp 151	0.02109	3.99 $^{+0.49}_{-0.68}$	1295 \pm 37	6.83 \pm 0.32	3, 5
Mrk 1310	0.01956	3.66 $^{+0.59}_{-0.61}$	921 \pm 135	6.50 \pm 0.34	3, 5
Mrk 50	0.02343	10.64 $^{+0.82}_{-0.93}$	1740 \pm 101	7.50 \pm 0.32	7
NGC 6814	0.00521	6.64 $^{+0.87}_{-0.90}$	1697 \pm 224	7.28 \pm 0.34	3, 5
SBS 1116+583A	0.02787	2.31 $^{+0.62}_{-0.49}$	1550 \pm 310	6.74 \pm 0.38	3, 5
Zw 229-015	0.02788	3.86 $^{+0.69}_{-0.90}$	1590 \pm 47	6.99 \pm 0.32	8

Note. Col. (1) Name. Col. (2) Redshifts are from the NASA/IPAC Extragalactic Database (NED). Col. (3) Rest-frame H β time lag measurements. Col. (4) Line dispersion (σ_{line}) measured from rms spectra. Col. (5) The M_{BH} estimates from reverberation mapping: $M_{\text{BH}}(\text{RM}) = fVP_{\text{BH}} = fc\tau_{\text{cent}}\sigma_{\text{rms}}^2/G$, where the virial factor f with its uncertainty is adopted from Park et al. (2012a) and Woo et al. (2010; i.e., $\log f = 0.71 \pm 0.31$). Col. (6) References: 1. Peterson et al. (2004); 2. Denney et al. (2006); 3. Bentz et al. (2009b); 4. Denney et al. (2010); 5. Park et al. (2012b); 6. Grier et al. (2012); 7. Barth et al. (2011b); 8. Barth et al. (2011a).

^a Note that the sample and measurements are from P13. One difference here is that the adopted uncertainty for the virial factor (i.e., 0.31 dex) has been added in quadrature to the final RM BH mass uncertainties, although this homoscedastic uncertainty addition to the dependent variables does not alter any of the calibration results in this work, except for the values of intrinsic scatter term and slight changes in the constrained uncertainty ranges of regression coefficients.

Note that the slit position angle (PA) was not constrained, in order to maximize the *HST* scheduling opportunities. But the three grating data for each object were obtained in a single *HST* visit with the same orientation.

While we used the fully reduced data provided by the *HST* STIS pipeline for the UV gratings, we performed a custom reduction for the optical grating data from the raw science and reference files in order to improve the cleaning of cosmic-ray charge transfer trails in the raw images from the badly degraded STIS CCD. Based on the standard reduction of the STIS pipeline, an additional cosmic-ray removal step was added to the

processes employing the LA_COSMIC (van Dokkum 2001) routine following the approach described by Walsh et al. (2013). The raw data for the optical G430L grating were first calibrated with the BASIC2D task, including trimming the overscan region, bias and dark subtraction, and flat-fielding. Cosmic rays and hot pixels were then cleaned with LA_COSMIC, and wavelength calibration was performed.

The dithered individual exposures for each grating were then aligned and combined using the IMSHIFT and IMCOMBINE PYRAF tasks. After that, one-dimensional spectra from each grating were extracted with the X1D task and joined together to

Table 2
Summary of *HST*/STIS Observations for the Six New AGNs

Object	Observation Date	Slit PA (deg)	Total Exposure Time		
			G140L (s)	G230L (s)	G430L (s)
Arp 151	2013 Apr 29	97.7	3801	2639	495
Mrk 1310	2013 Jun 07	70.8	2624	1255	360
Mrk 50	2012 Dec 12	-110.8	2624	1255	360
NGC 6814	2013 May 07	-149.6	2848	1299	540
SBS 1116 +583A	2013 Jul 12	28.9	3714	2648	600
Zw 229-015	2013 Jul 23	117.6	4302	2942	600

produce a final single spectrum by taking into account the flux and noise levels in overlapping regions around ~ 1700 and ~ 3100 Å. Following P13, we corrected the spectra for Galactic extinction using the values of $E(B - V)$ from Schlafly & Finkbeiner (2011), as listed in the NASA/IPAC Extragalactic Database (NED), and the reddening curve of Fitzpatrick (1999). Figure 1 shows the fully reduced and calibrated rest-frame spectra of the six AGNs.

3. Spectral Measurements

To measure the broad emission-line widths and the continuum luminosity adjacent to each broad line, we carried out a multicomponent spectral decomposition analysis of the spectral region surrounding C IV $\lambda 1549$, Mg II $\lambda 2798$, and H β $\lambda 4861$. A combination of these two observables, line width and continuum luminosity measured from an SE spectrum, is commonly used to estimate BH masses via SE BH mass estimators because they can be adopted as reasonable proxies for the velocity of the broad-line gas clouds and the size of the BLR (Kaspi et al. 2000, 2005; Bentz et al. 2006, 2009a, 2013), respectively. Following the standard approach that has been adopted in previous works (e.g., Shen et al. 2008, 2011), we measured monochromatic continuum luminosities at 1350, 3000, and 5100 Å to compute the SE virial masses from C IV, Mg II, and H β , respectively.

Our fits are based on a local decomposition of the spectral region around each broad line, rather than a global decomposition of the entire UV-optical spectrum. Owing to the complexity of the spectra and the large number of emission-line and continuum components that are present, we found that local decompositions are able to achieve a more precise fit to the data around each line than would be possible in a simultaneous, global fit to the full STIS spectrum (see Section 3.5 for a discussion of the global versus local fits). The local spectral decomposition technique employed here is based on those by P13 and Park et al. (2015) and slightly updated and modified for the STIS data and the spectral region in question. Our spectral-modeling method consists of separate procedures for continuum fitting and line emission fitting applied independently to the C IV, Mg II, and H β regions of the data. During fitting, model parameters are optimized using `mpfit` (Markwardt 2009) in Interactive Data Language. The model components and fitting details for each of the H β , Mg II,

and C IV line regions are described in the following subsections, and the decomposition results are given in Figure 2.

3.1. H β

We used the multicomponent spectral decomposition code developed by Park et al. (2015) for modeling the H β region of our STIS data. In brief, the code works by first simultaneously fitting a pseudocontinuum that consists of a single power law, an Fe II template, and a host-galaxy template in the surrounding continuum regions of 4430–4770 and 5080–5450 Å and then fitting the H β emission line complex with Gauss–Hermite series functions (van der Marel & Franx 1993; Cappellari et al. 2002) for one broad emission component (H β) and three narrow emission components (H β and [O III] $\lambda\lambda 4959, 5007$) and two Gaussian functions for the nearby blended He II $\lambda 4686$ emission line after subtracting the best-fit pseudocontinuum model (see Park et al. 2015 and references therein for details of the measurement procedure; see also Woo et al. 2006; Bennert et al. 2015; Runco et al. 2016). The H β line widths, $\text{FWHM}_{\text{H}\beta}$ and $\sigma_{\text{H}\beta}$, are measured from the best-fit broad-line model (i.e., the Gauss–Hermite series function), and the continuum luminosity at 5100 Å, $\lambda L_{5100 \text{ Å}}$, is measured from the best-fit power-law model.

Note that there are two differences between the method adopted here and the approach given by Park et al. (2015), specifically in the model components used for the Fe II emission and host-galaxy starlight. The template for host-galaxy starlight is excluded in this work because stellar absorption features, which are critical to achieving reliable host-galaxy template fits, are not observable in the small-aperture STIS spectra. The minimal contribution of host-galaxy light and the relatively low signal-to-noise ratio (S/N) and spectral resolution of the STIS optical data make it difficult to detect any host-galaxy features. Moreover, the fits did not converge when we included the host-galaxy starlight component in the model. As a rough check, we provided a crude estimate of an upper limit for host-galaxy starlight contribution to the STIS spectra using the object SBS 1116+583A, which shows the highest host-galaxy fraction in the ground-based spectroscopic observations (see Park et al. 2012b; Barth et al. 2015) from our STIS sample. The host-galaxy flux in the STIS spectrum can be roughly estimated by subtracting the AGN flux at 5100 Å, which is obtained by subtracting the *HST* imaging-based galaxy flux at 5100 Å (Bentz et al. 2013) from the ground-based spectroscopic total flux at 5100 Å (Park et al. 2012b), from the total flux at 5100 Å of the STIS spectrum. The resulting host-galaxy fraction in the STIS spectrum is found to be $\sim 31\%$. Note that the other AGNs will have much lower contributions due to the lower host-galaxy fractions shown by Park et al. (2012b) and Barth et al. (2015).

Available Fe II templates for the H β region include empirically constructed monolithic templates by Boroson & Green (1992) and Véron-Cetty et al. (2004), a theoretical template by Bruhweiler & Verner (2008), and a semi-empirical multicomponent template by Kovačević et al. (2010). After performing extensive tests using each of the templates and a linear combination of them for our STIS data, we opted to use the template of Kovačević et al. (2010) based on its overall performance as quantified by the χ^2 -statistics and residuals of the fits (see also Barth et al. 2013, 2015). As expected, the multicomponent template generally performed better than the monolithic templates, particularly for the objects showing

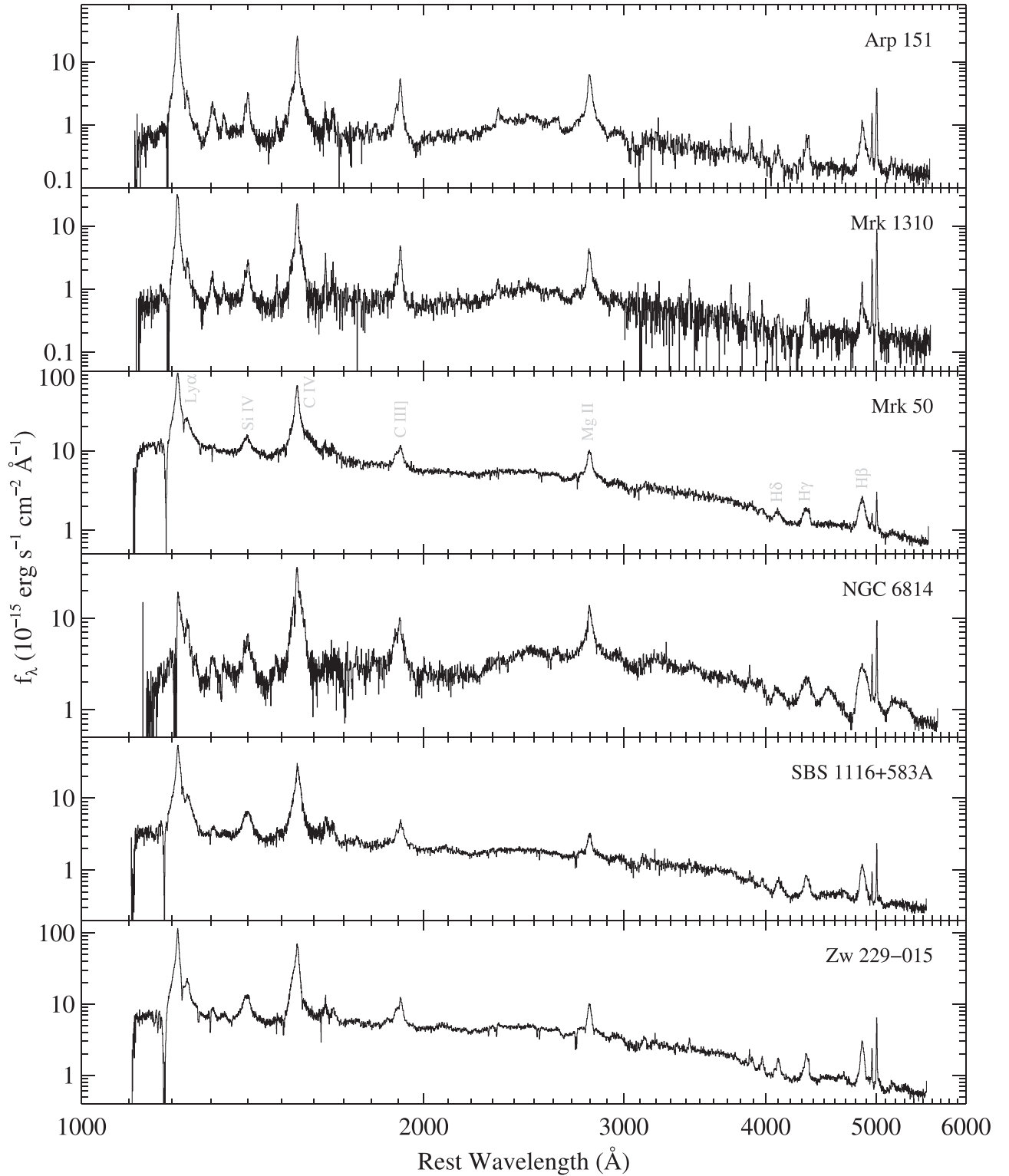


Figure 1. Final fully reduced and combined STIS spectra for our sample of the six low-mass AGNs.

strong Fe II emission. The Kovačević et al. (2010) template appears to be the best currently available for accurately fitting diverse Fe II emission blends in AGNs by allowing for different relative intensities between five Fe II multiplet subgroups. To sum up, we followed the method described by Park et al. (2015), except that we used the template of Kovačević et al. (2010) instead of that of Boroson & Green (1992) for Fe II

emission, and we omitted the host-galaxy starlight template from the fits.

3.2. Mg II

For the Mg II spectral region, we first fit a pseudocontinuum model in the surrounding continuum regions of 2450–2750 and

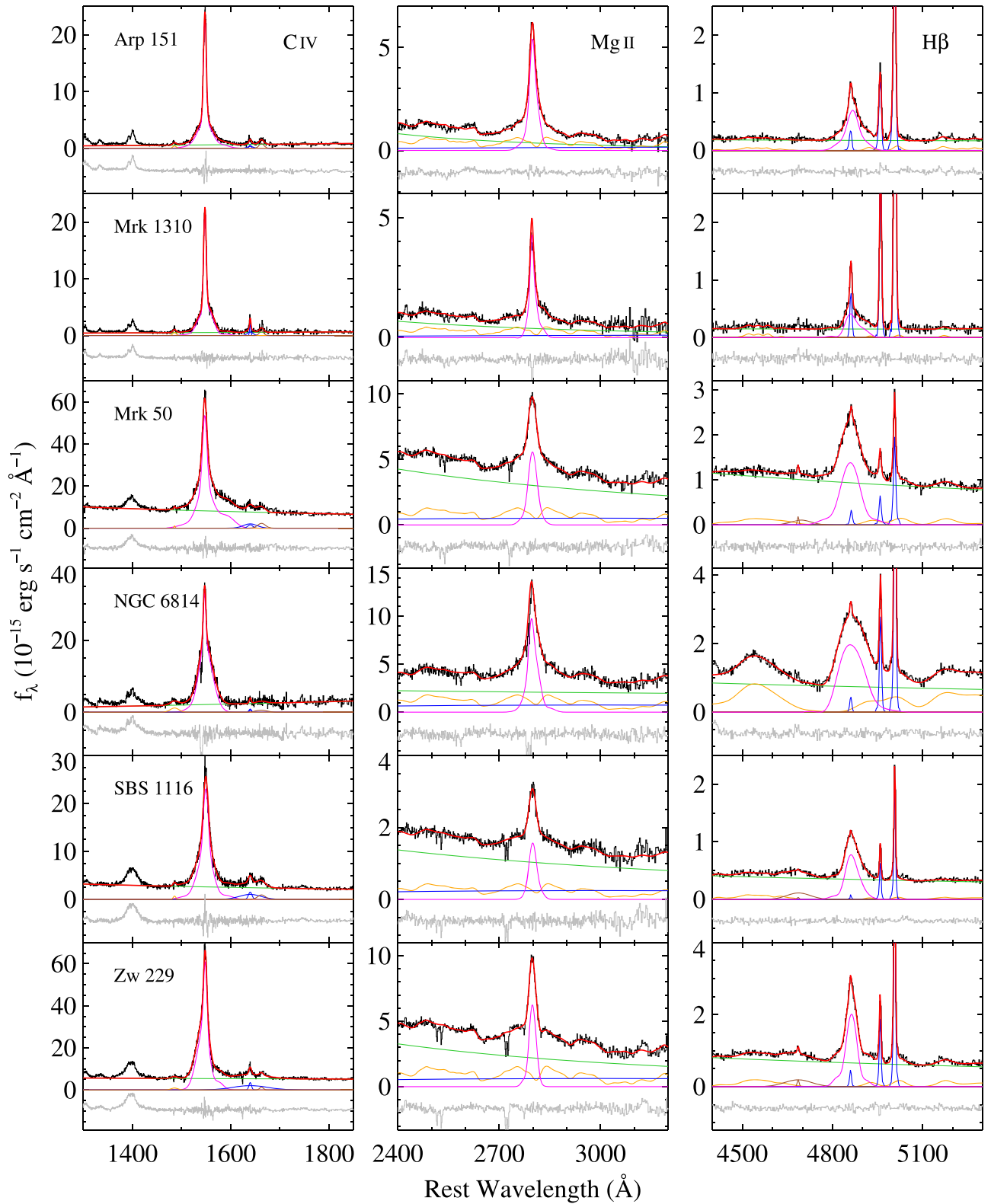


Figure 2. Multicomponent spectral decompositions in the spectral regions of three major broad emission lines—C IV $\lambda 1549$, Mg II $\lambda 2798$, and H β $\lambda 4861$ —for our six AGNs. In each panel, the observed spectrum (black) is decomposed into various components. Left panels (C IV): power-law continuum (green), C IV $\lambda 1549$ (magenta), and other nearby blended lines, including N IV] $\lambda 1486$ (orange), He II $\lambda 1640$ (blue), and O III] $\lambda 1663$ (brown). Middle panels (Mg II): power-law continuum (green), Fe II template (orange), Balmer continuum (blue), and Mg II $\lambda 2798$ (magenta). Right panels (H β): power-law continuum (green), Fe II template (orange), three narrow emission lines (H β and [O III] $\lambda\lambda 4959, 5007$ blue), broad H β (magenta), and broad and narrow He II $\lambda 4686$ components (brown; only included if blended with H β). The red lines in each panel indicate the full model combining all the best-fit model components. The gray lines at the bottom of each panel represent the residual, i.e., data–model, shifted downward for clarity.

2850–3100 Å. The pseudocontinuum model is composed of a single power-law function representing the AGN featureless continuum, an Fe II emission template, and an empirical model for the Balmer continuum. We adopted the UV Fe II template, which is made from observations of I Zw 1, from Tsuzuki et al. (2006). Using the template of Tsuzuki et al. (2006) is arguably better for modeling the Mg II line region than using that of Vestergaard & Wilkes (2001) because it contains semi-empirically constrained Fe II contribution underneath the Mg II line, while the template by Vestergaard & Wilkes (2001) has no Fe II flux at all under the Mg II line due to the difficulty of decomposing this spectral region.

Based on the investigations of Grandi (1982) and Wills et al. (1985; see also Malkan & Sargent 1982 for the first practical measurement of the Balmer continuum shape), Dietrich et al. (2002, 2003) described a practical procedure for Balmer continuum modeling in high- z quasar spectra that has become a standard practice for fitting the Balmer continuum. This empirical model assumes that the Balmer continuum is generated from partially optically thick gas clouds with a uniform effective temperature ($T_e = 15,000$ K) as

$$F_{\lambda}^{\text{BaC}}(A, T_e, \tau_{\text{BE}}) = AB_{\lambda}(T_e)(1 - e^{-\tau_{\text{BE}}(\lambda/\lambda_{\text{BE}})^3}), \lambda \leq \lambda_{\text{BE}}, \quad (1)$$

where A and τ_{BE} are the normalized flux density and optical depth at the Balmer edge ($\lambda_{\text{BE}} = 3646$ Å), respectively, and $B_{\lambda}(T_e)$ is the Planck function at the electron temperature T_e . At $\lambda > \lambda_{\text{BE}}$, higher-order Balmer lines using the relative intensity calculations from Storey & Hummer (1995) are used to represent the smooth rise to the Balmer edge. Many studies (e.g., Kurk et al. 2007; Wang et al. 2009; Greene et al. 2010; De Rosa et al. 2011, 2014; Ho et al. 2012; Shen & Liu 2012; Kokubo et al. 2014) have used variants of this method with slightly different ways of constraining the model parameter ranges based on the available data quality and spectral coverage. We found that if we treated all three parameters in the model (A, T_e, τ_{BE}) as free parameters during fitting, as was done by Wang et al. (2009) and Shen & Liu (2012), they were very poorly constrained due to the degeneracy with the power-law continuum and Fe II emission blends.

Recently, Kovačević et al. (2014) suggested an improved way to constrain the normalization A : by taking into account the fact that A can be obtained by calculating the sum of all intensities of higher-order Balmer lines at the Balmer edge. We followed this procedure with slight modifications. The value of A was separately determined from the intensity calculation using high-order Balmer lines (up to 400) with the template line profile adopted from the best-fit $\text{H}\beta$ emission-line model obtained above. The other parameters, T_e and τ_{BE} , were then fitted simultaneously with the Fe II template and power-law function during pseudocontinuum modeling. However, it should be noted that, following the Dietrich approach, the temperature was finally fixed to be 15,000 K, and the optical depth was allowed to vary between 0.1 and 2. We also independently checked that the constrained Balmer continuum component only exhibited marginal changes over temperatures ranging from 10,000 to 30,000 K and optical depths varying from 0.1 to 2. Note that the resulting continuum luminosity estimates are consistent with each other within ~ 0.04 dex scatter.

After subtracting the best-fit pseudocontinuum model, the Mg II emission line was fitted using a linear combination of a

sixth-order Gauss–Hermite series and a single Gaussian function to account for its full line profile, typically showing a more peaky core (i.e., narrower and sharper line peak) and more extended wings than a Gaussian profile, in the spectral region ~ 2700 – 2900 Å. We use the full line profile without a decomposition of narrow and broad components for the line width measurements for UV lines in this work (the same approach adopted by P13), in contrast to $\text{H}\beta$. This is because no reliable and clear distinction between broad and narrow components in the UV lines is usually possible, and sometimes no narrow components of the UV lines are seen at all; their presence is still uncertain and under debate. Thus, the Mg II line widths, $\text{FWHM}_{\text{Mg II}}$ and $\sigma_{\text{Mg II}}$, are measured from the best-fit full line profile, and continuum luminosity at 3000 Å, $\lambda L_{3000 \text{ Å}}$, is measured from the best-fit power-law function. During fitting, Galactic absorption lines such as Fe II $\lambda\lambda 2586, 2600$, Mg II $\lambda\lambda 2796, 2803$, and Mg I $\lambda 2852$ (cf. Savaglio et al. 2004) are masked out with exclusion windows.

3.3. C IV

Spectral measurements for the C IV line region in the archival sample of the local 25 RM AGNs were described by P13. Here, we focus on analysis of the six objects with newly obtained STIS data. We used the same methods as in P13 for consistency and to avoid additional systematic biases. We fit the AGN featureless continuum with a single power-law function, and we chose to omit a UV iron template (e.g., Vestergaard & Wilkes 2001) from the fits because no clear contribution of iron emission over the C IV region was observed. Although we performed a test by including the UV iron template in the model as in P13, its contribution was too small to be constrained accurately with the template, at least in our sample (see also Shen et al. 2008, 2011).

After the best-fit continuum model was subtracted, the C IV emission line was fitted with a linear combination of a sixth-order Gauss–Hermite series and a single Gaussian function. The contaminating nearby blended emission lines (e.g., N IV] $\lambda 1486$, Si II $\lambda 1531$, He II $\lambda 1640$, and O III] $\lambda 1663$) were fitted simultaneously as well using up to two Gaussian functions for each line. Again, we used the combined model of one Gauss–Hermite function and one Gaussian function to fit the full line profile of C IV without decomposing it into broad and narrow components. The C IV line widths, $\text{FWHM}_{\text{C IV}}$ and $\sigma_{\text{C IV}}$, were measured from the best-fit full line profile, and continuum luminosity at 1350 Å, $\lambda L_{1350 \text{ Å}}$, was measured from the best-fit power-law function. Narrow absorption spikes were masked out using a 3σ clipping threshold during fitting, and broad absorption features around the line center were masked out manually with exclusion windows (see P13 and references therein for more details of the C IV measurement method and results for the archival sample).

3.4. Measurement Uncertainty Estimation

Uncertainties for the above spectral measurements are estimated with the Monte Carlo method used by Park et al. (2012b) and P13 (see also Shen et al. 2011; Shen & Liu 2012). For each spectral region, 1000 mock spectra are generated by resampling the original spectra with the addition of Gaussian random noise based on the error spectrum for each object. We then measure line widths and luminosities from each of the mock spectra using the same measurement methods and take

Table 3
Ultraviolet Spectral Properties from C IV SE Estimates

Object	Telescope/Instrument	Date Observed	S/N (1450 or 1700 Å) (pix ⁻¹)	$E(B - V)$ (mag)	$\log(\lambda L_{\lambda}/\text{erg s}^{-1})$ (1350 Å)	FWHM _{SE} (C IV) (km s ⁻¹)	σ_{SE} (C IV) (km s ⁻¹)	MAD _{SE} (C IV) (km s ⁻¹)	$\rho(W, X)$	$\rho(W, Y)$	$\rho(W, Z)$
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)
Sample Presented in P13 ^a											
3C 120	<i>IUE</i> /SWP	1994 Feb 19, 27; 1994 Mar 11	12	0.263	44.399 ± 0.021	3093 ± 291	3106 ± 157	2280 ± 74	-0.05	-0.35	-0.50
3C 390.3	<i>HST</i> /FOS	1996 Mar 31	18	0.063	43.869 ± 0.003	5645 ± 202	6154 ± 65	4674 ± 42	-0.11	-0.30	-0.32
Ark 120	<i>HST</i> /FOS	1995 Jul 29	17	0.114	44.400 ± 0.005	3471 ± 108	3219 ± 53	2345 ± 31	0.01	-0.30	-0.37
Fairall 9	<i>HST</i> /FOS	1993 Jan 22	24	0.023	44.442 ± 0.004	2649 ± 77	2694 ± 20	1987 ± 13	0.04	-0.21	-0.17
Mrk 279	<i>HST</i> /COS	2011 Jun 27	9	0.014	43.082 ± 0.004	4093 ± 388	2973 ± 53	2202 ± 18	0.01	-0.04	-0.11
Mrk 290	<i>HST</i> /COS	2009 Oct 28	24	0.014	43.611 ± 0.002	2052 ± 36	3531 ± 32	2544 ± 14	-0.00	-0.13	-0.20
Mrk 335	<i>HST</i> /COS	2009 Oct 31; 2010 Feb 08	29	0.032	43.953 ± 0.001	1772 ± 14	1876 ± 12	1311 ± 7	0.03	-0.02	0.02
Mrk 509	<i>HST</i> /COS	2009 Dec 10, 11	107	0.051	44.675 ± 0.001	3872 ± 18	3568 ± 9	2601 ± 6	0.02	-0.07	-0.03
Mrk 590	<i>IUE</i> /SWP	1991 Jan 14	17	0.033	44.094 ± 0.007	5362 ± 266	3479 ± 165	2630 ± 139	-0.09	-0.12	-0.05
Mrk 817	<i>HST</i> /COS	2009 Aug 04; 2009 Dec 28	38	0.006	44.326 ± 0.001	4580 ± 48	3692 ± 23	2756 ± 14	-0.01	-0.25	-0.21
NGC 3516	<i>HST</i> /COS	2010 Oct 04; 2011 Jan 22	20	0.038	42.615 ± 0.002	2658 ± 34	4006 ± 49	2864 ± 29	-0.03	0.07	0.06
NGC 3783	<i>HST</i> /COS	2011 May 26	29	0.105	43.400 ± 0.001	2656 ± 444	2774 ± 91	2014 ± 9	-0.01	-0.01	-0.17
NGC 4593	<i>HST</i> /STIS	2002 Jun 23, 24	10	0.022	43.761 ± 0.005	2952 ± 166	2946 ± 162	2135 ± 33	-0.01	-0.00	-0.02
NGC 5548	<i>HST</i> /COS	2011 Jun 16, 17	36	0.018	43.822 ± 0.001	1785 ± 82	4772 ± 80	3528 ± 102	0.02	-0.11	-0.02
NGC 7469	<i>HST</i> /COS	2010 Oct 16	32	0.061	43.909 ± 0.001	2725 ± 66	2849 ± 237	2060 ± 15	-0.02	-0.01	-0.14
PG 0026+129	<i>HST</i> /FOS	1994 Nov 27	25	0.063	45.236 ± 0.005	1604 ± 50	4965 ± 113	3610 ± 77	-0.04	-0.11	-0.22
PG 0052+251	<i>HST</i> /FOS	1993 Jul 22	21	0.042	45.292 ± 0.004	5380 ± 87	4648 ± 50	3463 ± 30	-0.12	-0.54	-0.61
PG 0804+761	<i>HST</i> /COS	2010 Jun 12	34	0.031	45.493 ± 0.001	3429 ± 23	2585 ± 20	1932 ± 13	0.04	0.14	0.17
PG 0953+414	<i>HST</i> /FOS	1991 Jun 18	18	0.012	45.629 ± 0.005	3021 ± 74	3448 ± 55	2472 ± 35	-0.02	-0.43	-0.48
PG 1226+023	<i>HST</i> /FOS	1991 Jan 14, 15	93	0.018	46.309 ± 0.001	3609 ± 29	3513 ± 29	2595 ± 19	-0.18	-0.52	-0.60
PG 1229+204	<i>IUE</i> /SWP	1982 May 01, 02	28	0.024	44.609 ± 0.009	4023 ± 163	2621 ± 90	1989 ± 62	-0.29	-0.48	-0.49
PG 1307+085	<i>HST</i> /FOS	1993 Jul 21	14	0.030	45.113 ± 0.006	3604 ± 111	4237 ± 80	3010 ± 54	-0.13	-0.57	-0.58
PG 1426+015	<i>IUE</i> /SWP	1985 Mar 01, 02	45	0.028	45.263 ± 0.004	4220 ± 258	4808 ± 305	3549 ± 95	-0.10	-0.23	-0.42
PG 1613+658	<i>HST</i> /COS	2010 Apr 08, 09, 10	37	0.023	45.488 ± 0.001	6398 ± 51	4204 ± 17	3286 ± 13	-0.01	-0.10	-0.04
PG 2130+099	<i>HST</i> /COS	2010 Oct 28	22	0.039	44.447 ± 0.001	2147 ± 18	2225 ± 47	1554 ± 21	0.02	-0.06	-0.07
New Sample Presented Here											
Arp 151	<i>HST</i> /STIS	2013 Apr 29	6	0.012	41.791 ± 0.017	1489 ± 26	2900 ± 61	1864 ± 35	-0.03	-0.38	-0.46
Mrk 1310	<i>HST</i> /STIS	2013 Jun 07	5	0.027	41.715 ± 0.025	1434 ± 78	2447 ± 108	1603 ± 54	0.00	-0.25	-0.31
Mrk 50	<i>HST</i> /STIS	2012 Dec 12	19	0.015	43.213 ± 0.003	2807 ± 63	4443 ± 160	3140 ± 115	0.02	-0.13	-0.10
NGC 6814	<i>HST</i> /STIS	2013 May 07	6	0.164	41.105 ± 0.021	2651 ± 264	2804 ± 103	2096 ± 59	0.02	-0.07	-0.11
SBS 1116+583A	<i>HST</i> /STIS	2013 Jul 12	13	0.010	42.867 ± 0.005	3253 ± 302	3315 ± 231	2302 ± 121	-0.04	-0.13	-0.15
Zw 229-015	<i>HST</i> /STIS	2013 Jul 23	17	0.064	43.129 ± 0.007	2573 ± 71	2608 ± 56	1891 ± 33	-0.05	-0.20	-0.18

Note. Col. (1) Name. Col. (2) Telescope/instrument from which archival UV spectra were obtained. Note that the new COS spectra were obtained after 2009. Col. (3) Observation date for combined spectra. Col. (4) S/N per pixel at 1450 or 1700 Å in the rest frame. Col. (5) Value of $E(B - V)$ from the NED based on the recalibration of Schlafly & Finkbeiner (2011). Col. (6) Continuum luminosity measured at 1350 Å. Col. (7) FWHM measured from SE spectra. Col. (8) Line dispersion (σ_{line}) measured from SE spectra. Col. (9) MAD (mean absolute deviation around weighted median) measured from SE spectra. Col. (10) Correlation coefficient between measurement errors of W and X , where $W = \log \lambda L_{\lambda}$ and $X = \log \text{FWHM}_{\text{SE}}$. Col. (11) Correlation coefficient between measurement errors of W and Y , where $Y = \log \sigma_{\text{SE}}$. Col. (12) Correlation coefficient between measurement errors of W and Z , where $Z = \log \text{MAD}_{\text{SE}}$.

^a Note that the sample and measurements are from P13. One difference here is that measurements for MAD and error correlations have been included.

the standard deviation of the distribution of the measurements as the estimate of the measurement uncertainty.

Typical uncertainty levels of line widths for all objects are found to be $\sim 2\%$ – 4% with a maximum of $\sim 17\%$ due to the high quality of the *HST* spectra. For continuum luminosity, we derive uncertainties of $\sim 1\%$ – 2% with a maximum of $\sim 6\%$. These are small compared to the overall systematic mass uncertainty of ~ 0.4 dex in the SE virial method. Covariances between the measurement uncertainties of the line widths and luminosities for each object in a logarithmic scale are also estimated from the resulting distributions of the Monte Carlo simulations, which are given as $\text{cov}(\log \lambda L_\lambda, \log \text{FWHM}) = \rho \sigma(\log \lambda L_\lambda) \sigma(\log \text{FWHM})$ and $\text{cov}(\log \lambda L_\lambda, \log \sigma_{\text{line}}) = \rho \sigma(\log \lambda L_\lambda) \sigma(\log \sigma_{\text{line}})$, where cov , ρ , and σ are the covariance, correlation coefficient, and measurement uncertainty of the logarithms of the luminosities and line widths, respectively. Table 3 lists the line widths and luminosities for our sample, along with the measurement uncertainties and their error correlation coefficients.

3.5. Continuum Luminosities and Emission-Line Widths

There are several issues in regard to measuring continuum luminosities and line widths accurately. It is important to take into account the Balmer continuum over the Mg II line region in order to accurately decompose the power-law continuum for the luminosity measurements. Based on our investigation of the STIS data, the $\lambda L_{3000\text{\AA}}$ values will be overestimated by ~ 0.14 dex, on average, if the Balmer continuum component is not accounted for. This is consistent with the investigation of Shen & Liu (2012), who found a ~ 0.12 dex systematic offset. This bias will then be propagated into the final M_{BH} estimates by as much as a ~ 0.07 dex ($\sim 17\%$) systematic offset if the Balmer continuum model is not included properly.

If the original Vestergaard & Wilkes (2001) Fe II template is used, the FWHM (σ_{line}) estimates are overestimated by ~ 0.03 (0.07) dex, on average, compared with the results derived using the template of Tsuzuki et al. (2006). This will again be propagated into the M_{BH} estimates by up to ~ 0.06 (0.14) dex of systematic offset, which is consistent with the result of Nobuta et al. (2012).

To make the maximum use of the wide spectral coverage of our STIS data, we also performed extensive tests of global continuum fits covering C IV to H β simultaneously. Using a more flexible double power-law model to represent the AGN featureless continuum, we fit a pseudocontinuum model including the Balmer continuum model and Fe II emission to many line-free continuum regions (see also, e.g., Shen & Liu 2012 and Mejía-Restrepo et al. 2016 for related recent work). The global continuum fits produced results consistent with the local continuum fits described in the previous sections, except that the global fits failed to constrain the Fe II emission on the red side of the Mg II regions for Arp 151 and Mrk 1310. This is probably because the double power-law model is not flexible enough to properly describe the steep local slope changes around Mg II for these two objects that are coming from intrinsic changes of spectral shapes and/or from strong internal reddening, along with the incompleteness of currently available UV Fe II templates across the regions. In any case, there is no significant improvement of the global fits compared to the local fits.

The simultaneous coverage of our STIS data also makes it possible to consistently compare the major UV and optical emission lines (C IV, Mg II, and H β) without biases from intrinsic variability (Figure 3). The C IV profile shows, on

average, more peaky cores with extended wings than those of the Mg II and H β lines (see also Wills et al. 1993; Brotherton et al. 1994). There is no significant velocity offset between the line peaks of the three emission lines. Figure 4 compares FWHM and σ_{line} line width measurements among the emission lines. Although it is hard to draw a clear picture due to small-number statistics, we use only our STIS sample of six AGNs in order to perform a consistent comparison between the line widths in quasi-simultaneously observed data. We find that the FWHM of C IV is, on average, smaller than that of H β (and Mg II), which may indicate that the FWHM of C IV is not a good proxy for virial BLR velocity, probably due to contamination from a nonreverberating C IV core component (Denney 2012). This contamination would be one of the biases correlated with Eigenvector 1 (EV1; Boroson & Green 1992), as discussed by Runnoe et al. (2013a, 2014) and Brotherton et al. (2015). They investigated and used the peak flux ratio of the $\lambda 1400$ feature to C IV as a UV indicator of EV1 to correct for the C IV-based BH masses. However, the interpretation is not straightforward. It could also be the case that the BLR geometry is different for the regions emitting these three lines, resulting in different individual virial factors (f) for each line (see also Runnoe et al. 2013b). By contrast, $\sigma_{\text{C IV}}$ is, on average, larger than $\sigma_{\text{H}\beta}$, which is consistent with the simple virial expectation of stratified BLR structure and the shorter reverberation lags of C IV (Peterson & Wandel 1999; Kollatschny 2003), thus corroborating the use of σ_{line} over FWHM for C IV-based BH mass estimates. More detailed intercomparisons and systematic investigation of multiline properties including more objects from the literature will be presented in a forthcoming paper.

4. Bayesian Calibration of C IV-Based M_{BH} Estimators

Now that we have the continuum luminosity and line width measurements from the SE spectra, we can perform a calibration of the C IV-based SE BH mass estimators against the H β RM-based BH masses as a fiducial baseline. We assume that the BH masses from the H β RM are the most reliable mass estimates available for these galaxies, and our goal is to find the combination of SE measurements that most closely reproduces the RM mass scale using the following equation:

$$\log\left(\frac{M_{\text{BH}}^{\text{RM}}}{M_{\odot}}\right) = \alpha + \beta \log\left(\frac{\lambda L_{1350\text{\AA}}^{\text{SE}}}{10^{44} \text{ erg s}^{-1}}\right) + \gamma \log\left(\frac{\Delta V_{\text{C IV}}^{\text{SE}}}{1000 \text{ km s}^{-1}}\right), \quad (2)$$

where $\Delta V_{\text{C IV}}^{\text{SE}} = \text{FWHM}_{\text{C IV}}^{\text{SE}}$ or $\sigma_{\text{C IV}}^{\text{SE}}$. This equation essentially expresses the virial relation $M_{\text{BH}} \sim r_{\text{BLR}} V^2 / G$, assuming that the BLR radius scales with AGN luminosity according to $r_{\text{BLR}} \propto L^{\beta}$ and allowing for the virial exponent γ to differ from the physically expected value of 2 in order to achieve the best fit.

Note that calibrations of the BH mass estimators based on the emission lines that have no direct reverberation measurements, e.g., C IV and Mg II, have been performed indirectly against the H β -based reverberation results for the same objects, if available (e.g., VP06 and P13). Although there are a few direct (or, in some cases, tentative) reverberation results for the UV emission lines (e.g., Kaspi et al. 2007; Trevese et al.

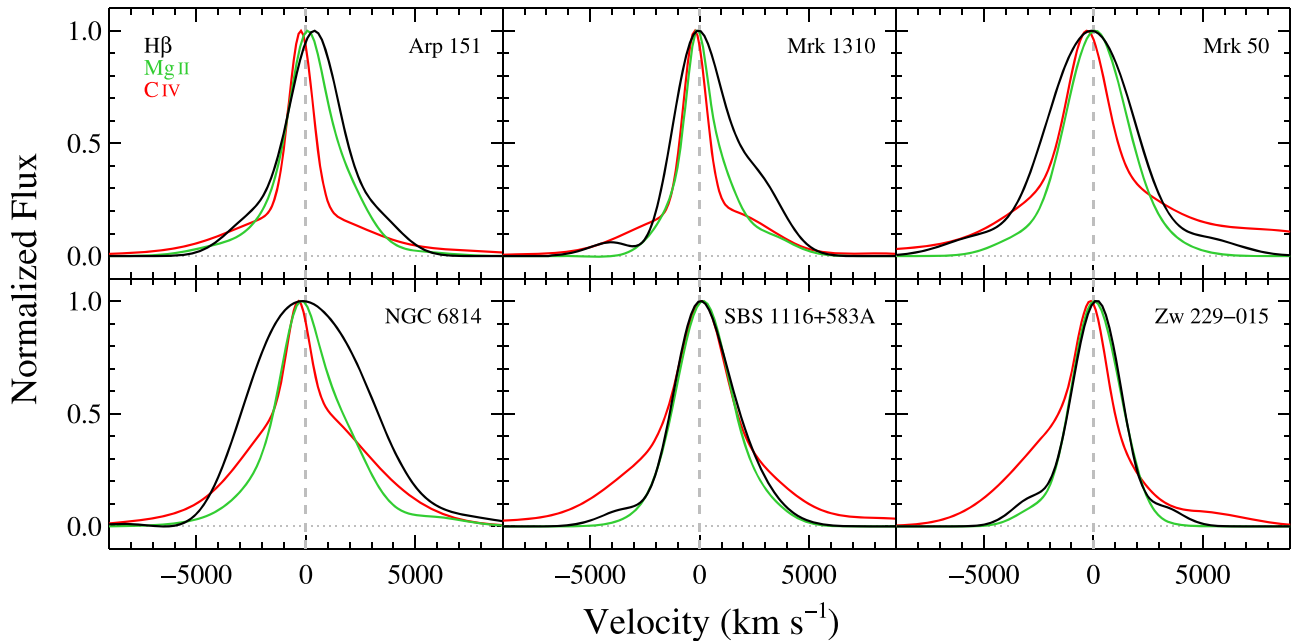


Figure 3. Comparison of the modeled emission-line profiles, normalized by each peak, of C IV (red), Mg II (green), and H β (black) for our STIS sample.

2014; see also Shen et al. 2016b), most of the available reverberation studies have been done with the H β line, which gives the most reliable AGN BH masses at present.

To perform the calibration, we adopt a Bayesian approach to linear regression analysis. An advantage of the Bayesian method over the traditional χ^2 -based method is that, by obtaining probability density functions (PDFs) for parameters of interest instead of just calculating a point estimate, it provides more reliable uncertainty estimates, incorporating all the error sources modeled and simply marginalizing over nuisance parameters. It is also easy to explore the covariance between parameters from resulting joint probability distributions. The Bayesian linear regression method outperforms other classical methods, especially when the measurement error of the independent variables is large and/or the sample size is small (see Kelly 2007).

For the full Bayesian inference, we use the Stan probabilistic programming language (Stan Development Team 2015a), which contains an adaptive Hamiltonian Monte Carlo (HMC; Neal 2012; Betancourt & Girolami 2013) No-U-Turn sampler (NUTS; Hoffman & Gelman 2014) as its sampling engine. This provides a simple implementation for specifying complex hierarchical Bayesian models and achieves good computational efficiency. We set up the Bayesian hierarchical model following Kelly et al. (2012) and implement it by referring to Gelman et al. (2013), Kruschke (2014), and the Stan Development Team (2015b). The practical details of the sampler and the model specification will be described in a separate paper discussing the methodology (Park 2017, in preparation). As a brief summary, the t -distribution is adopted to obtain an outlier-robust statistical inference following the investigation of Kelly et al. (2012), with an additional improvement of treating degrees of freedom in the t model as a free parameter, instead of fixing it to be a preselected constant. Thus, the likelihood function, which specifies the measurement, regression, and covariate distribution models, is built with t -distributions, and the prior distributions are specified based on suggestions by Barnard et al. (2000), Gelman et al. (2013), and Kruschke (2014). Our Markov chain Monte Carlo

(MCMC) simulations have been run via the PyStan package (v2.9.0; Stan Development Team 2016) with careful assessment of the convergence of the MCMC chains.

4.1. The SE Mass Calibration

Figure 5 shows the results of the calibration of C IV-based SE BH mass estimators against the H β -based RM masses using the full sample of 31 local RM AGNs with Equation (2), while Figure 6 presents the corresponding marginal projections of each pair of parameters of interest with one-dimensional marginalized distributions from the full posterior distribution, from which parameter covariances are simply identifiable. We take the best-fit values and uncertainties of the parameters of interest (α , β , γ , and σ_{int}) from posterior median estimates and 68% posterior credible intervals, as recommended by Kelly (2007) and Hogg et al. (2010).

In each panel of Figure 5, the two mass estimates (RM and SE) are fairly consistent. The overall scatter of the SE BH masses based on the calibrated equation using the C IV FWHM (σ_{line}) compared to the RM BH masses are at the level of 0.37 dex (0.33 dex). This indicates generally quite good consistency, given the unavoidable object-to-object scatter of the virial factor f (0.31 dex; Woo et al. 2010), since we are adopting a single ensemble average value of the f factor for all objects.

To assess the resulting model fit to the data, we present the posterior predictive distribution (blue shaded contour), which is generated from simulations according to the fit parameters of the model, to check whether the posterior prediction replicates the original observed data distribution reasonably well. As can be seen, the 95% credible region depicted by the light blue shaded contour describes most of the data distribution well, except for one outlier: NGC 6814. We have also calculated posterior predictive p -values by following the method of Chevillard & Charlot (2016; see also Gelman et al. 2013 and the PyMC User’s Guide¹⁴). Using χ^2 deviance as a discrepancy

¹⁴ <https://pymc-devs.github.io/pymc/modelchecking.html>

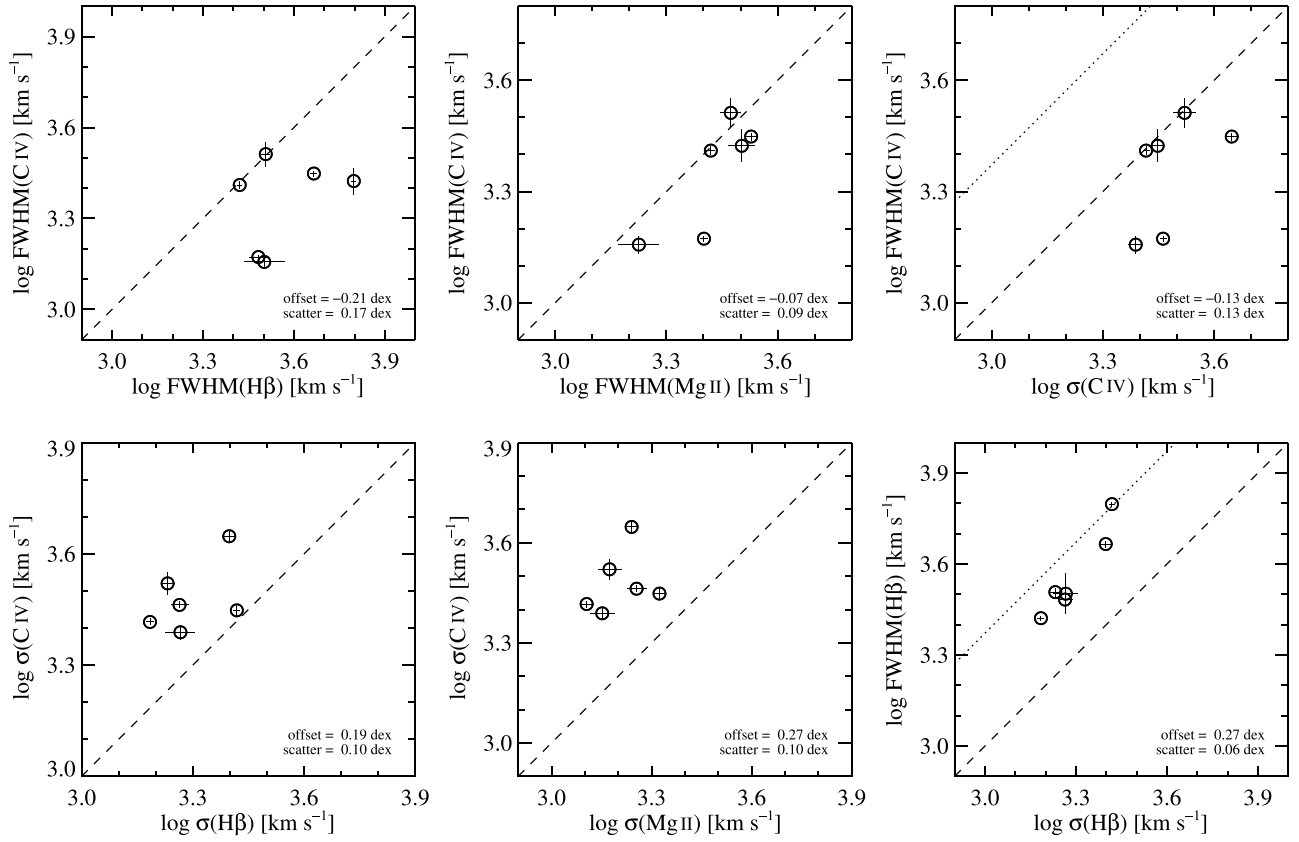


Figure 4. Intercomparison of line width measurements, FWHM and σ_{line} , of C IV, Mg II, and H β for our STIS sample. The dashed line in each panel represents a one-to-one relation. The dotted lines show the ratio of FWHM to σ_{line} for a Gaussian profile. The mean offset (i.e., average of line width differences) and 1σ scatter (i.e., standard deviation of line width differences) are given in the lower right corner of each panel.

measure, the Bayesian p -value estimates are mostly ~ 0.2 , ranging from 0.14 to 0.26 in this work, indicating successful model fits to the data. Note that there is a problem (e.g., misfit or inadequacy in the descriptive model) if the p -value is extreme, i.e., < 0.05 or > 0.95 . All the calibration results performed in this work for various cases are listed in Table 4. Note that the presence of the single outlier, which is not very extreme, does not alter any of the conclusions, and virtually the same result is obtained for the given uncertainties if the outlier is removed.

The luminosity slope, β , is consistent with the photoionization expectation (i.e., 0.5) within the uncertainties for both FWHM-based and σ_{line} -based estimators. This may also imply that the size–luminosity relation for the UV 1350 Å continuum has a slope consistent with the relation for the optical 5100 Å continuum (e.g., $0.533^{+0.035}_{-0.033}$ by Bentz et al. 2013). The velocity slope, γ , for the σ_{line} -based estimator is close to the virial expectation (i.e., 2) within the uncertainties, while this is not the case for the FWHM-based estimator, whose γ value is consistent with zero, given the uncertainty interval. This is generally consistent with the calibration result based on C IV FWHM by Shen & Liu (2012), who found a much smaller slope (0.242) than 2, although their luminosity dynamic range probed is much higher than ours. And the σ_{line} -based SE masses show overall less intrinsic and total scatter than the FWHM-based masses. Thus, this indicates that the C IV σ is a better tracer of the BLR velocity field than the C IV FWHM because it is closer to the virial relation and shows less scatter in mass estimates. These results are in overall agreement with those of Denney (2012), who found that the C IV FWHM is much more

affected by a contaminating nonvariable C IV core component (see also P13 and Denney et al. 2013 for related discussion and interpretation). Velocity slopes that are shallower (0.50, 1.66) than the virial expectation (2) would also be expected in part due to an additional dispersion of the measured line-of-sight velocities stemming from orientation dependence (see Shen & Ho 2014).

Our final best fits are as follows (see also Table 4):

$$\log \left[\frac{M_{\text{BH}}(\text{SE})}{M_{\odot}} \right] = 6.73^{+0.07}_{-0.07} + 0.43^{+0.06}_{-0.06} \log \left(\frac{\lambda L_{1350 \text{ \AA}}}{10^{44} \text{ erg s}^{-1}} \right) + 2 \log \left[\frac{\sigma(\text{C IV})}{1000 \text{ km s}^{-1}} \right], \quad (3)$$

with the overall scatter against RM masses of 0.33 dex, which is defined by the standard deviation of mass residuals $\Delta = \log M_{\text{BH}}(\text{RM}) - \log M_{\text{BH}}(\text{SE})$, and

$$\log \left[\frac{M_{\text{BH}}(\text{SE})}{M_{\odot}} \right] = 7.54^{+0.26}_{-0.27} + 0.45^{+0.08}_{-0.08} \log \left(\frac{\lambda L_{1350 \text{ \AA}}}{10^{44} \text{ erg s}^{-1}} \right) + 0.50^{+0.55}_{-0.53} \log \left[\frac{\text{FWHM}(\text{C IV})}{1000 \text{ km s}^{-1}} \right], \quad (4)$$

with the overall scatter against RM masses of 0.37 dex. For the case of the σ_{line} -based estimator, note that the value of γ is fixed at 2 (i.e., consistent with the virial expectation) in our final analysis because it is consistent with 2 within the uncertainty

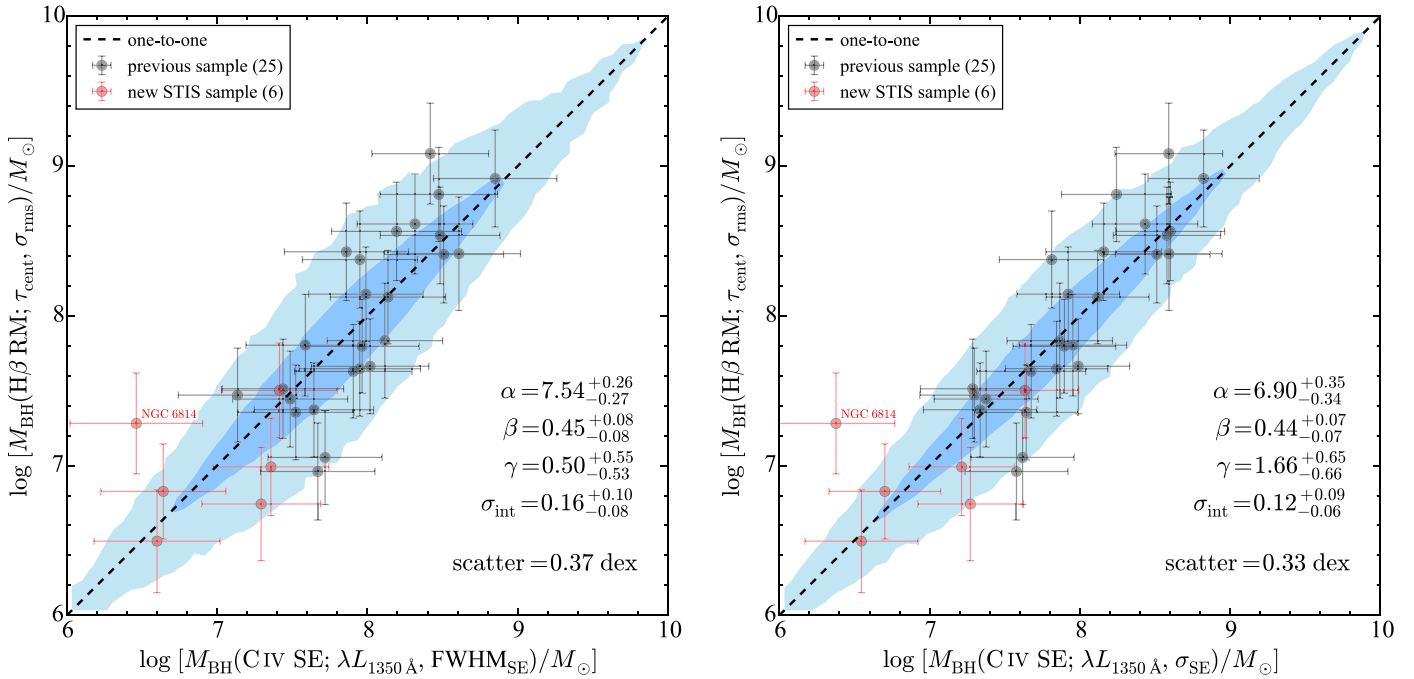


Figure 5. Calibration results of C IV-based SE BH mass estimators using FWHM (left) and σ_{line} (right). The newly added low-mass AGN sample (six objects) is indicated with red circles, while the black circles represent the previous archival sample of 25 objects from P13. Blue shaded contours, whose dark blue (light blue) area corresponds to the 68% (95%) credible region, show the posterior predictive distributions under the fitted model as a self-consistency check (see text). The resulting regression parameters with the uncertainty estimates and overall scatter are given in the lower right corner of each panel. Note that the error bars in the y-axis (RM mass) represent the quadratic sum of the propagated RM measurement uncertainty for the virial product (VP) and the adopted uncertainty for the virial factor, and those in the x-axis (SE mass) represent the quadratic sum of the propagated SE calibration uncertainty and the resulting overall scatter.

estimate when we treat it as a free parameter. Fixing γ to the physically motivated value helps to avoid possible object-by-object biases and systematics due to small-number statistics and reduces the uncertainties of the other resulting regression parameters (see also P13). The overall scatter of the σ_{line} -based calibration is virtually unchanged if we fix $\gamma = 2$, while that of the FWHM-based calibration increases considerably if γ is fixed to 2 (see Table 4).

In Figure 7, we show the calibration results in a three-dimensional space of luminosity, velocity, and mass for clarity. There is a strong dependence of mass on luminosity, while there is a much weaker dependence of mass on velocity, partly due to the small dynamic range of broad-line velocity in our sample. Especially for FWHM, the change of mass as a function of velocity is only marginal. As expected from the very small value of γ for the FWHM-based estimator, it seems that the C IV FWHM velocity characterization does not significantly add useful information to mass estimates, which is consistent with the result by Shen & Liu (2012; see also discussions on the importance of line width information by Croom 2011 and Assef et al. 2012). It would thus be possible to achieve a comparable level of accuracy in mass estimates using the luminosity information only, at least for the present sample. This much-weaker dependence of FWHM on mass than that of σ_{line} can also be observed from the projected RM mass-SE velocity plane in Figure 8, which reinforces that σ_{line} is a better velocity width measurement for the C IV line than FWHM. We generally recommend using the σ_{line} -based C IV SE M_{BH} estimator instead of the FWHM-based estimator, since it is closer to virial relation and shows a better correlation and less scatter against the RM masses.

In this work (i.e., Equations (3) and (4) and Table 4), we have used the virial factor, $\log f = 0.71$, taken from Park et al. (2012a) and Woo et al. (2010) for a consistent comparison with the result of P13. If one wants to use a more recent value of the virial factor for BH mass estimates, e.g., the recently updated virial factor, $\log f = 0.65$, from Woo et al. (2015), one can simply subtract the difference (0.06) between the adopted virial factors from the normalization, α , of all the calibration results in this work. Our results can similarly be rescaled to any other adopted value for virial factor f .

4.2. Comparison with P13

The final best-fit calibrated equations (Equations (3) and (4)) are very similar to those of our previous work (see Equations (2) and (3) of P13). If we compare the two mass estimates based on both estimators using the same measurements of our sample, there are very small offsets (0.02–0.03 dex) with small scatters of ~ 0.09 dex for both FWHM-based and σ_{line} -based masses, which are mostly coming from slight differences in the slopes between P13 and this work. Although the changes in the adopted virial relation are modest, it is worth noting that this work directly extends the applicability of the C IV-based M_{BH} estimators toward lower BH masses ($\sim 10^{6.5} M_{\odot}$) than were present in the P13 sample.

4.3. Comparison with Other Linear Regression Methods

The advantages of the adopted statistical model using Stan in this work are (1) using the outlier-robust t -distribution as an alternative to the normal distribution for error distributions and (2) modeling the intrinsic distribution of covariates explicitly with a multivariate t -distribution. To check the performance of

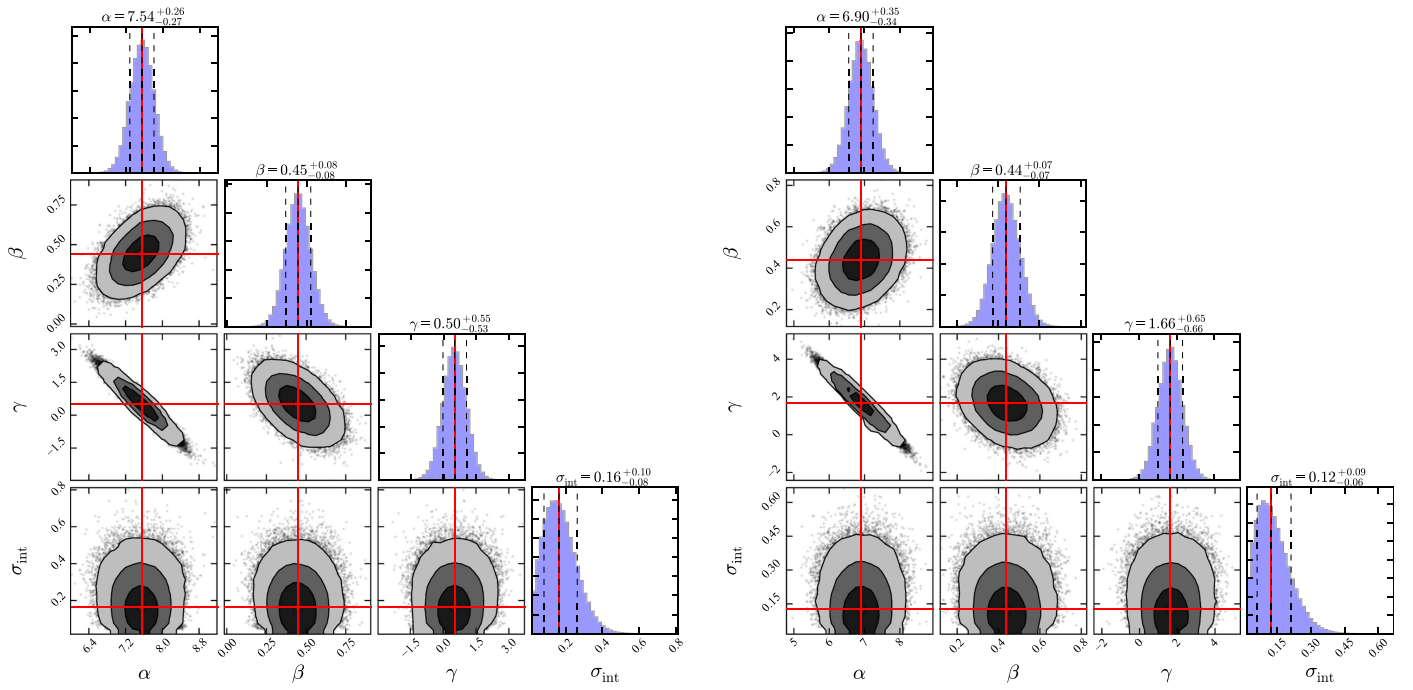


Figure 6. Posterior distributions of the resulting parameters of the calibrations in Figure 5 for FWHM (left) and σ_{line} (right). The red solid lines indicate the posterior median estimate, and the black dashed lines mark the uncertainty ranges (i.e., 16% and 84% posterior quantiles). The 2D marginal distributions (black) of the parameter pairs are shown in the off-diagonal panels, while the 1D marginalized histograms (blue) for each parameter from the posterior sample are given in the diagonal panels. This figure was made using the `corner.py` (<https://github.com/dfm/corner.py>).

our model, we provide a comparison of results by performing the same regression work with other available methods (i.e., `mlinmix_err` and `FITEXY`).

Figure 9 compares the resulting posterior distributions of parameters obtained from three different regression methods for the same data (i.e., calibration of the FWHM-based estimator using the full sample of 31 local RM AGNs as obtained above). The `Stan` Bayesian model implemented for this work (Park 2017, in preparation) uses the Student t -distributions for measurement errors, intrinsic scatter, and the covariate distribution model. The `mlinmix_err` method, a Bayesian linear regression code developed by Kelly (2007), employs a normal mixture model for covariate distribution and assumes Gaussian distributions for measurement errors and intrinsic scatter. The `FITEXY` method, a widely used traditional χ^2 -based linear regression method (Tremaine et al. 2002; see also Park et al. 2012a and references therein), also uses Gaussian distributions for measurement errors and intrinsic scatter, but it has no model specified for the covariate distribution and does not take into account possible correlations between measurement errors.

The left panel in Figure 9 shows an overall consistency of results between the Bayesian methods, `Stan`, and `mlinmix_err`, except for the distributions of intrinsic scatter. The quite strong difference of the σ_{int} distributions is expected, because `Stan` uses t -distributed intrinsic scatter while `mlinmix_err` uses normally distributed intrinsic scatter. By definition, the σ_{int} of the t -distribution is smaller than that of the Gaussian distribution due to the broader tails of the t -distribution (see Kelly et al. 2012 and D. Park 2017, in preparation). Another noticeable difference between the posterior distributions is that the widths of the probability distributions for the regression parameters obtained from `Stan` are slightly wider than those from `mlinmix_err`. Although not

significant, this seems to indicate more reliable uncertainty estimates with `Stan`, probably due to the flexibility of the adopted t -distributions with degrees-of-freedom parameters. Note that the t -distribution ranges widely from the Cauchy distribution to the normal distribution with a varying degrees-of-freedom parameter, but the number of Gaussian components for the normal mixture model used in `mlinmix_err` is fixed as a constant (e.g., 3 by default, although a few Gaussians are usually enough to obtain a reasonable description of the observed distributions of many astronomical samples and data, as is the case in this work).

The right panel in Figure 9 also shows an overall consistency between the resulting distributions of the Bayesian `Stan` and the χ^2 -based `FITEXY` method. However, underestimates of the parameter uncertainties from `FITEXY` are a bit more noticeable, possibly due to the absence of the covariate model description and not accounting for correlations between measurement errors in `FITEXY` estimates. The parameter distributions from `FITEXY` are obtained with a bootstrapping method, so it may not be a consistent comparison with the Bayesian posterior distributions. Many zero values in the σ_{int} distribution of `FITEXY` are also noticeable, which indicates that many realizations of bootstrap samples are optimized without the addition of intrinsic scatter. This behavior is one of the downsides of the χ^2 -based `FITEXY` estimator, which employs a somewhat ad hoc iterative procedure to determine σ_{int} because it cannot be constrained simultaneously with the regression parameters (see Kelly 2011 and Park et al. 2012a).

The adopted best-fit parameters and uncertainties from the three methods are listed in Table 5 for comparison. Again, there is no significant difference between the parameter estimates; they are basically consistent with each other within the uncertainties. The primary reason for this consistency is that the measurement uncertainties for the covariates (line widths

Table 4
C IV M_{BH} Estimator Calibration Results $\log[M_{\text{BH}}(\text{RM})/M_{\odot}] = \alpha + \beta \log(L_{1350 \text{ \AA}}/10^{44} \text{ erg s}^{-1}) + \gamma \log[\Delta V(\text{C IV})/1000 \text{ km s}^{-1}]$

ΔV (IV)	α	β	γ	σ_{int}	Mean Offset (dex)	1σ Scatter (dex)	References
Previous Calibrations							
σ_{line}	6.73 ± 0.01	0.53	2	0.33	VP06
FWHM	6.66 ± 0.01	0.53	2	0.36	VP06
σ_{line}	6.71 ± 0.07	0.50 ± 0.07	2	0.28 ± 0.04	0.00	0.295	P13
FWHM	7.48 ± 0.24	0.52 ± 0.09	0.56 ± 0.48	0.35 ± 0.05	0.00	0.347	P13
This Work							
σ_{line}	$6.90^{+0.35}_{-0.34}$	$0.44^{+0.07}_{-0.07}$	$1.66^{+0.65}_{-0.66}$	$0.12^{+0.09}_{-0.06}$	0.01	0.33	...
FWHM	$7.54^{+0.26}_{-0.27}$	$0.45^{+0.08}_{-0.08}$	$0.50^{+0.55}_{-0.53}$	$0.16^{+0.10}_{-0.08}$	0.00	0.37	Best fit ^a
MAD	$7.15^{+0.24}_{-0.25}$	$0.42^{+0.07}_{-0.07}$	$1.65^{+0.61}_{-0.62}$	$0.12^{+0.09}_{-0.06}$	0.00	0.33	...
This Work (Fixing $\gamma = 2$)							
σ_{line}	$6.73^{+0.07}_{-0.07}$	$0.43^{+0.06}_{-0.06}$	2	$0.12^{+0.09}_{-0.06}$	0.01	0.33	Best fit ^a
FWHM	$6.84^{+0.09}_{-0.09}$	$0.33^{+0.07}_{-0.07}$	2	$0.22^{+0.11}_{-0.10}$	-0.01	0.43	...
MAD	$7.01^{+0.07}_{-0.07}$	$0.41^{+0.06}_{-0.06}$	2	$0.12^{+0.09}_{-0.06}$	0.00	0.33	Best fit ^a
This Work (Fixing $\beta = 0.5$)							
σ_{line}	$6.99^{+0.34}_{-0.34}$	0.5	$1.49^{+0.63}_{-0.62}$	$0.12^{+0.09}_{-0.06}$	0.01	0.34	...
FWHM	$7.62^{+0.23}_{-0.23}$	0.5	$0.31^{+0.46}_{-0.45}$	$0.16^{+0.10}_{-0.08}$	0.01	0.38	...
MAD	$7.23^{+0.24}_{-0.24}$	0.5	$1.41^{+0.57}_{-0.58}$	$0.12^{+0.09}_{-0.06}$	0.01	0.34	...
This Work (Fixing $\beta = 0.5$ and $\gamma = 2$)							
σ_{line}	$6.72^{+0.07}_{-0.07}$	0.5	2	$0.12^{+0.08}_{-0.06}$	0.01	0.35	...
FWHM	$6.82^{+0.09}_{-0.09}$	0.5	2	$0.26^{+0.11}_{-0.11}$	0.00	0.47	...
MAD	$7.00^{+0.07}_{-0.07}$	0.5	2	$0.12^{+0.09}_{-0.06}$	0.01	0.35	...

Note. The mean offset and 1σ scatter for our calibrations are measured from the average and standard deviations of mass residuals between the RM masses and calibrated SE masses, $\Delta = \log M_{\text{BH}}(\text{RM}) - \log M_{\text{BH}}(\text{SE})$. Note that the apparent big difference in σ_{int} estimates between the previous calibrations and this work is mostly due to the differences in the adopted RM mass error and statistical model. The uncertainty of $\log f$ (i.e., 0.31 dex) is added in quadrature to the uncertainties of the RM BH masses in this work. The standard deviation (σ) of the t -distribution is by definition different (larger) from that of the Gaussian distribution due to the heavy tail when the degrees-of-freedom parameter is small. In this case, the σ_{int} parameter of the t -distribution model is not the same as the data spread (σ) of the t -distribution. Boldface font indicates best-fit values.

^a We suggest these calibrations as the best M_{BH} estimators.

and luminosities) are very small in this work (i.e., only a few percent, on average, due to high-quality *HST* spectra). The resulting covariances between measurement errors are consequently very small as well, thus leading to virtually no effect of error correlations on the regression parameter estimates, even though there are correlations between measurement errors (see Table 3). To summarize, all three methods (*Stan*, *mlin-mix_err*, and *FITEXY*) produce consistent results in this work, except for arguably more reliable parameter uncertainty estimates when using *Stan*, given the small measurement errors of the covariates. Although the difference is marginal in this work, more flexible t -distributed errors, as well as an explicit covariate model description, are generally recommended to get a correct central trend against outliers by avoiding the effects of possible unaccounted systematic errors (see Park 2017, in preparation, for details).

4.4. MAD-Based Calibration

Although we prefer line dispersion (σ_{line}) to FWHM in measuring C IV line width, as investigated above, one downside

of using σ_{line} is that it requires high-S/N data to accurately fit the line wings. Noisy data can lead to biases in line width measurements, especially when the line profile has very extended wings, as is typical of C IV lines (Denney et al. 2013; see also Fine et al. 2010).

Recently, Denney et al. (2016b) suggested another way of measuring line width: the MAD around the flux-weighted median wavelength. They suggested it as the most reliable method of line width measurement for low-quality data. The MAD is by definition less affected by the core and wing parts of the profile. Instead, the middle portions of the velocity profile (relative to the median velocity) contribute primarily to the determination of line width. The lower sensitivity of the MAD to the line core in comparison with FWHM is quite useful in order to obtain the least-biased line width measurement when there is a nonvarying core component in the C IV line profile (see Denney 2012). Such components are very hard to identify and remove without using multi-epoch RM data. Additionally, the MAD has the useful property of being less sensitive to high-velocity line wings than line dispersion (i.e., absolute deviation versus squared deviation as weights). This is

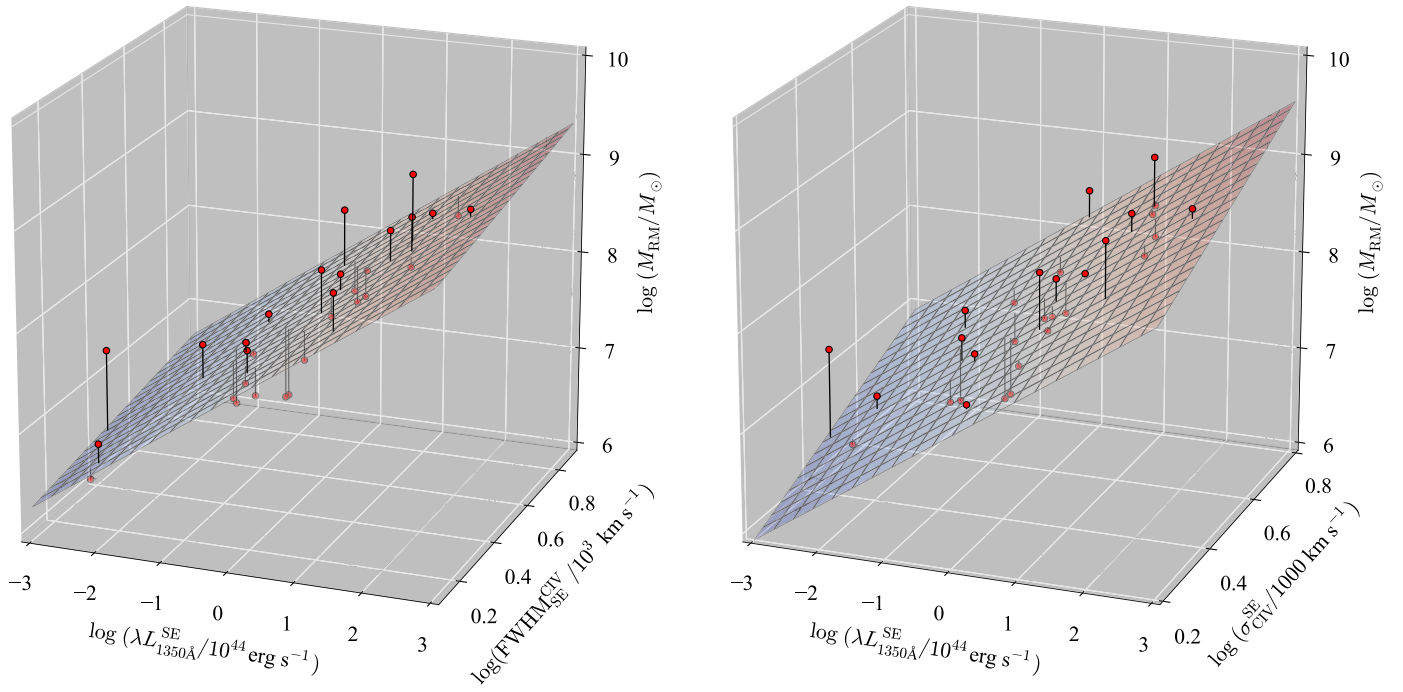


Figure 7. Three-dimensional representation of the calibration results of Figure 5 for clarification. Red circles indicate observed data. Colored tilted planes represent the resulting calibration with Equation (2). Black vertical lines connecting the data points to the fitted plane show the mass deviation between the observed RM mass and calibrated SE mass.

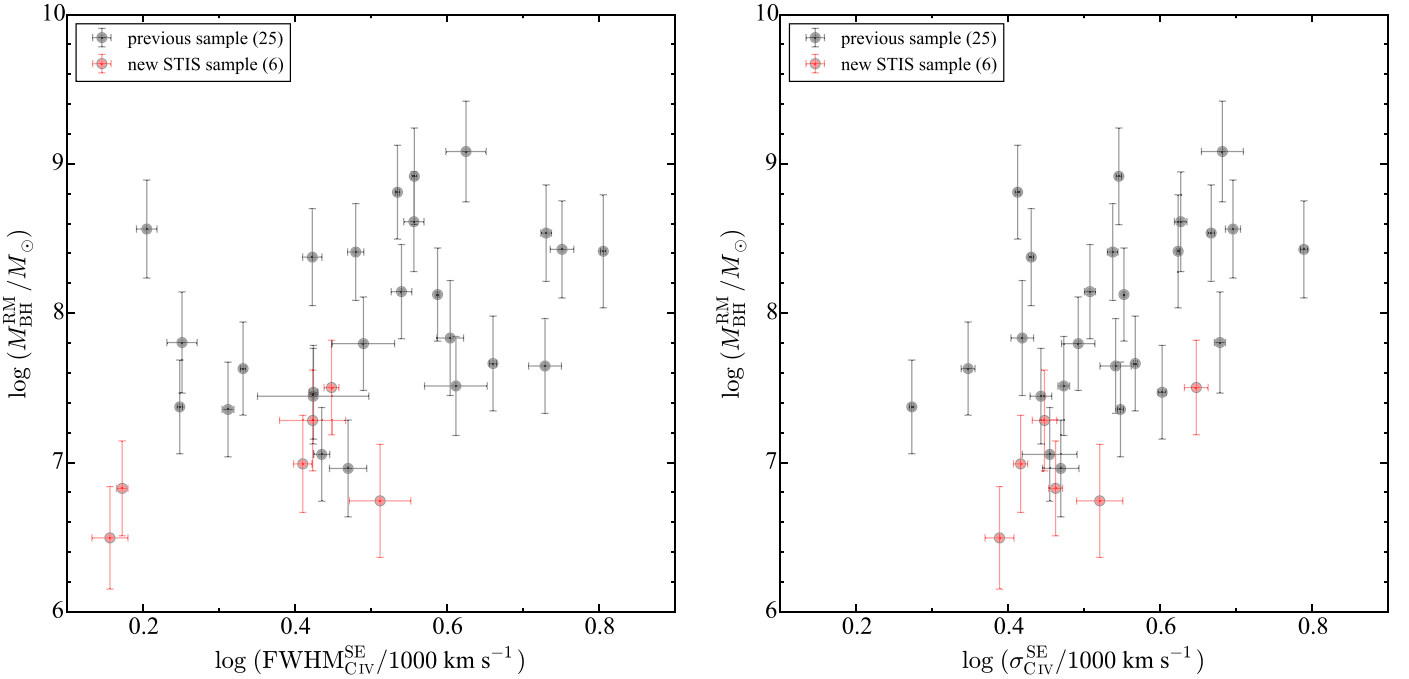


Figure 8. Comparison of the SE C IV velocity width measurements, FWHM (left) and σ_{line} (right), to the observed RM masses.

important when using low-S/N data, which makes accurate characterization of line wings very difficult.

Thus, the MAD inherits some of the practical merits of both σ_{line} and FWHM and possibly works better in low-quality data. We have carried out MAD measurements for the broad lines in our sample, and we find good consistency between the MAD and σ_{line} measurements (Figure 10). The two measurements are very nicely correlated with a marginal scatter, while a poor

correlation with a large scatter is observed between the FWHM and MAD. In this regard, the MAD may be the best line width measurement method for the C IV emission line when using survey-quality spectra, as advocated by Denney et al. (2016b).

As it is for the case of σ_{line} , the γ of the MAD is also consistent with 2 for the given uncertainty estimates if left as a free parameter (see Table 4). Fixing the virial slope to $\gamma = 2$, we find the following best-fit calibration of the SE C IV mass

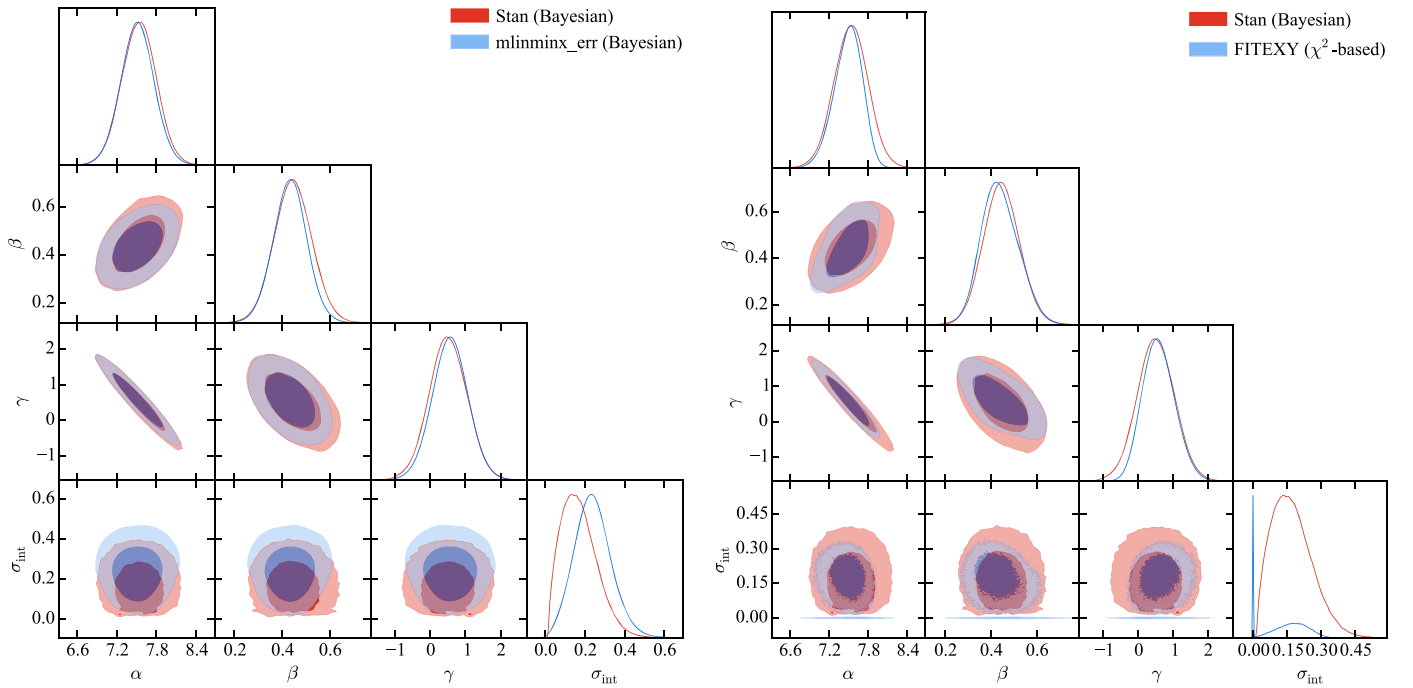


Figure 9. Comparisons of the resulting posterior distributions using Stan to those using `mlinmix_err` (left) and `FITEXY` (right). The one- and two-dimensional distributions of a parameter (diagonal panels) and parameter pairs (off-diagonal panels) are shown with the kernel density estimate using the `GetDist` (<https://github.com/cmbant/getdist>) python package. Note that some amount of smoothing has been applied for clarity of comparison between the distributions.

estimator based on the MAD as the measure of C IV line width:

$$\log \left[\frac{M_{\text{BH}}(\text{SE})}{M_{\odot}} \right] = 7.01_{-0.07}^{+0.07} + 0.41_{-0.06}^{+0.06} \log \left(\frac{\lambda L_{1350 \text{ \AA}}}{10^{44} \text{ erg s}^{-1}} \right) + 2 \log \left[\frac{\text{MAD}(\text{C IV})}{1000 \text{ km s}^{-1}} \right]. \quad (5)$$

In this case, the overall scatter against RM masses is 0.33 dex. The resulting MAD-based calibration and posterior distributions, which are not shown here, are very similar to the σ_{line} -based, except for a slight difference in the intercept α (see Table 4).

4.5. Possible Biases Due to C IV Blueshift

In our calibration sample (i.e., local RM AGNs), C IV blueshifts are basically insignificant (see Richards et al. 2011; Shen 2013), so our calibration based on the local RM AGNs is relatively free of possible biases stemming from the effect of a large blueshift. However, the applicability of this calibration to high- z quasars may be uncertain, because large C IV blueshifts are known to be common in high- z , high-luminosity quasars (see, e.g., Richards et al. 2002). Available C IV-based M_{BH} estimators have been used for measuring the BH masses of a statistical sample of such high- z AGNs, simply based on assumption and extrapolation without a direct test. The best way of investigating and possibly correcting for the effect of C IV blueshifts on BH mass estimates would be to use direct C IV RM data (see Denney 2012 for the case of local AGNs). The number of AGNs with direct C IV RM observations is, however, very limited, due to the major practical difficulties of obtaining RM measurements for high- z , high-luminosity AGNs

and the difficulty of obtaining space-based UV monitoring data for low- z AGNs.

Instead, Coatman et al. (2016, 2017; see also Shen & Liu 2012) recently provided a new empirical correction to C IV FWHM-based BH mass estimators as a function of C IV blueshift by comparing SE C IV measurements to SE H α measurements.

In Figure 11, we compare the overall distributions of C IV FWHM-based BH mass estimates as a function of C IV blueshift using the spectral measurements of DR9 BOSS quasars.¹⁵ The blue shaded contour presents BH masses computed from the blueshift-corrected formula from Coatman et al. (2017), while the red and green shaded contours show those calculated with the updated recipe in this work and the original VP06 equation, respectively. As can be seen, at a large blueshift ($\gtrsim 2000 \text{ km s}^{-1}$), our estimator, which does not take into account the C IV blueshift, produces a similar mass distribution to the blueshift-corrected distribution. Note that the overestimated BH masses from the VP06 estimator are reduced by correcting for the blueshift effect on the C IV FWHM (Coatman et al. 2017). Thus, our locally calibrated FWHM-based M_{BH} estimator is applicable to a sample of high- z , high-luminosity quasars with high C IV blueshifts (e.g., $\gtrsim 2000 \text{ km s}^{-1}$), giving a consistent mass scale on average.

At a small blueshift range ($\sim 0\text{--}1000 \text{ km s}^{-1}$), where our calibration sample is distributed, the overall mass scale from our estimator is smaller than those of VP06 and Coatman et al. (2017) in the high-mass regime ($\gtrsim 10^{8.5} M/M_{\odot}$) and larger in the low-mass regime ($\lesssim 10^{8.5} M/M_{\odot}$). This trend has been described in detail by P13. Arguably, our calibration in this work (and P13) has resulted in a better agreement (overall mass scale) with RM masses than that of VP06 in terms of intrinsic

¹⁵ Provided by Yue Shen at http://quasar.astro.illinois.edu/BH_mass/dr9.htm.

Table 5
Comparing Calibration Results with Other Linear Regression Methods

Method	α	β	γ	σ_{int}	Mean Offset (dex)	1 σ Scatter (dex)
Stan (Bayesian)	$7.54^{+0.26}_{-0.27}$	$0.45^{+0.08}_{-0.08}$	$0.50^{+0.55}_{-0.53}$	$0.16^{+0.10}_{-0.08}$	0.00	0.37
m1inmix_err (Bayesian)	$7.51^{+0.25}_{-0.25}$	$0.43^{+0.07}_{-0.07}$	$0.57^{+0.50}_{-0.51}$	$0.24^{+0.09}_{-0.09}$	-0.00	0.37
FITEXY (χ^2 -based)	7.50 ± 0.22	0.43 ± 0.08	0.59 ± 0.46	0.20 ± 0.10	-0.00	0.37

Note. For a consistent comparison, the exact same methodology of FITEXY used by P13 is applied.

scatter and using the higher-quality data set of the updated sample, at least in the mass range ($10^{6.5-9.1} M/M_{\odot}$) where the calibration has been performed. Note that neither our calibration nor that of VP06 (also Coatman et al. 2017, which is based on VP06’s calibration) is directly confirmed in the very high-mass BH regime ($\gtrsim 10^9 M/M_{\odot}$, where most of the Coatman et al. sample is) due to a lack of high-mass RM AGNs in the calibration samples. All of these trends are shown in Figure 13 of Coatman et al. (2017).

It is also worth noting that the striking mass increase toward the negative blueshift from using the equation of Coatman et al. (2017) is obviously unreliable, as already discussed by Coatman et al. (2017), due to their insufficient dynamic range of blueshift. This C IV blueshift-corrected recipe should therefore not be used for objects with a negative C IV blueshift.

As shown by Coatman et al. (2016), the H α line seems to also be systematically changing as a function of C IV blueshift, although the H α line measurements may not be very accurate due to a very low S/N for the H α spectral region (mostly $\lesssim 10$ per resolution element; see their Table 1). If this is true, calibrating the SE C IV line to the SE H α (or H β) line, as done by Coatman et al. (2016), would be flawed. In other words, correcting the C IV FWHM as a function of blueshift against the H α FWHM would still be biased, since the H α FWHM is also correlated with the C IV blueshift. As an ultimate goal, calibrating the SE C IV mass estimators against direct C IV RM data (or indirectly against the RM Balmer line if C IV RM data is unavailable) for a much larger sample including high-luminosity, high-C IV blueshift AGNs will be the best way to improve the SE mass method. However, given the difficulty of obtaining many direct C IV RM measurements for both low- and high- z AGNs and determining accurate blueshifts (and systemic redshifts; see, e.g., Denney et al. 2016a; Shen et al. 2016a), our simple calibration of SE C IV-based M_{BH} estimators will still be useful when estimating BH masses from C IV observations of AGNs over a wide range of redshift and luminosity.

4.6. Comparison to Other Prescriptions

In Figure 12, we compare the H β RM-based BH masses to the C IV FWHM-based SE BH masses from the corrected prescriptions presented by Denney (2012, their Equation (1)) and Runnoe et al. (2013a, their Equation (3)). Note that we use our sample of the local RM AGNs, except for four objects (PG 0026+129, PG 0052+251, PG 1226+023, and PG 1307+085) that do not have enough spectral coverage to measure the $\lambda 1400$ feature (see Figure 1 of P13). The peak flux of the emission-line blend of Si IV + O IV] (i.e., the $\lambda 1400$ feature) has been measured by fitting it with a local power-law continuum and multi-Gaussian functions following the same method as Runnoe et al. (2013a; see also Shang et al. 2007). As

a direct comparison, we also show the C IV FWHM-based masses using our new calibration, Equation (4).

The SE BH masses using the C IV line shape (FWHM/ σ_{line})-based correction by Denney (2012) show an overall scatter of 0.39 dex, which is the same as that of our calibration. However, this corrected prescription is not practically useful because it requires a σ_{line} measurement, as well as FWHM, to obtain the shape measurement for the correction to C IV masses. One can use σ_{line} directly, if it is available, rather than using FWHM. The slightly larger scatter of 0.43 dex is observed for the case of the peak flux ratio ($\lambda 1400/\text{C IV}$)-based correction by Runnoe et al. (2013a). The effect of the correction is less pronounced for our sample, which is not surprising based on the investigation by Brotherton et al. (2015), who found that the peak flux ratios measured for the RM AGN sample did not correlate with the difference between the H β and C IV velocity widths. Our simple calibration is a useful practical tool in situations in which such additional measurements are not available.

5. Summary and Discussion

We have updated the calibration of C IV-based SE M_{BH} estimators based on an enlarged AGN sample with high-quality *HST* UV spectra and using Bayesian linear regression analysis. As an extension of the work of P13, there are several improvements over the previous calibration: the sample now covers masses down to $\sim 10^{6.5} M_{\odot}$ with measurements from high-quality and quasi-simultaneous UV-to-optical STIS spectra, and we have used a Bayesian linear regression method to perform outlier-robust inference and take into account covariate distributions and possible correlations between measurement errors.

The results presented in this work are consistent with those of our previous work (P13) and are also in line with those of Denney (2012) and Denney et al. (2013). We generally recommend use of the σ_{line} -based or MAD-based C IV M_{BH} estimators when the measurement are available, since they are better proxies for the BLR velocity field (close to the virial relation) and show less scatter in mass estimates than the FWHM-based measurements. Using σ_{line} or MAD rather than FWHM for C IV line width measurement is supported by the fact that accurately decomposing and removing a C IV narrow component, if any, is difficult to accomplish with SE spectra. Thus, to avoid possible biases due to a possible C IV core component (Denney 2012), using the line width measurement that is least affected by an uncertain line core (i.e., σ_{line} or MAD) appears to be the best approach at present. Measuring σ_{line} requires high-quality data to accurately characterize the line wings, while the MAD is less sensitive to high-velocity wings. The C IV-based SE M_{BH} estimators are commonly applied to survey-quality data (e.g., SDSS quasars), where σ_{line}

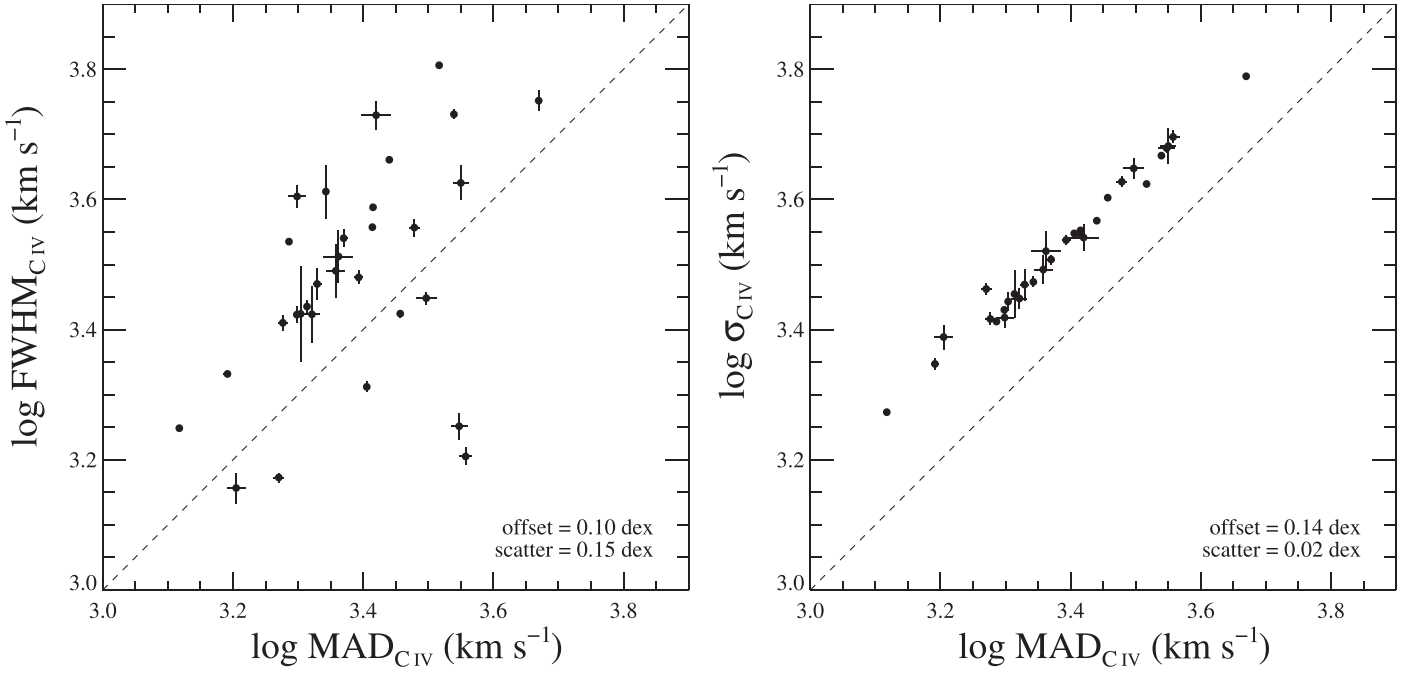


Figure 10. Comparison of MAD to FWHM (left) and line dispersion (σ_{line} ; right) measurements for our sample of all 31 AGNs. The dashed lines show a one-to-one relation. The mean offset and 1σ scatter are given in the lower right corner of each panel.

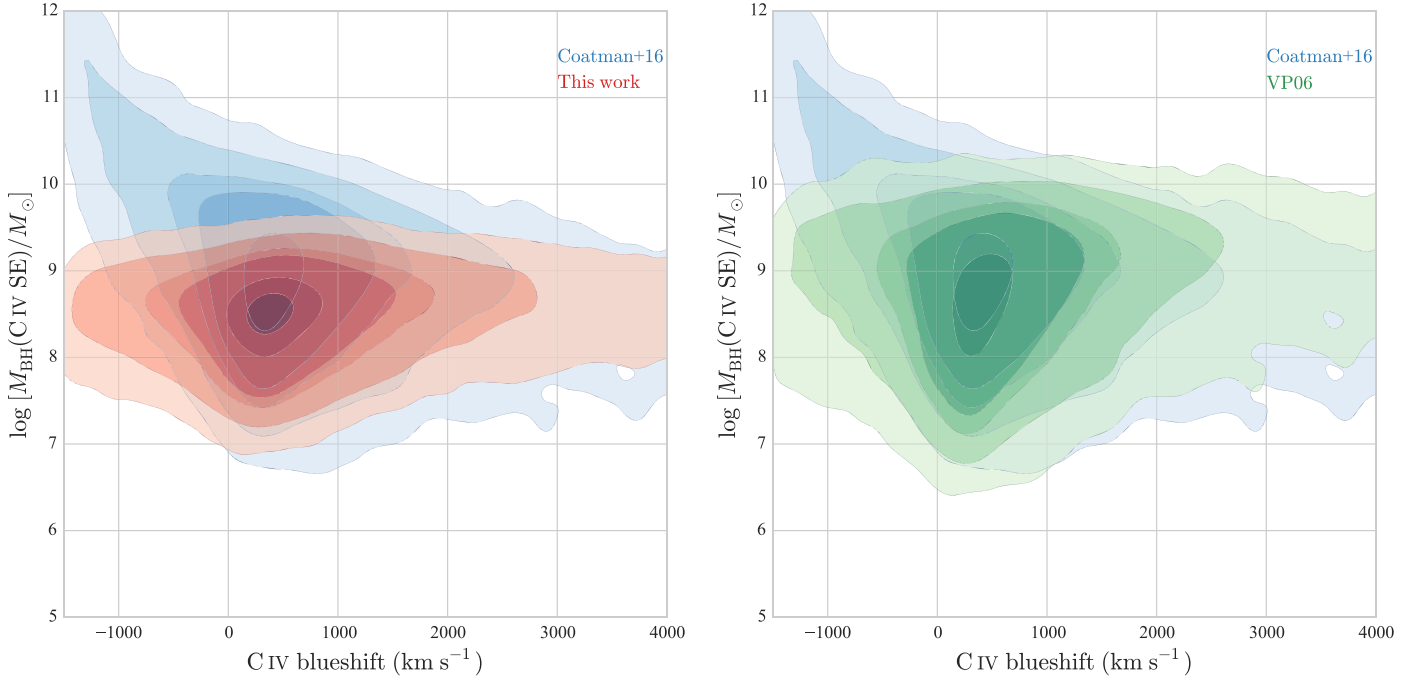


Figure 11. Distributions of C IV FWHM-based M_{BH} estimates as a function of C IV blueshift for SDSS quasars from the DR9 BOSS quasar catalog. Most of them are in the redshift range $2 \lesssim z \lesssim 3$. The left panel compares BH masses computed from the new calibration in this work (red) to those using the blueshift-corrected recipe from Coatman et al. (2017) (blue), while BH masses using the calibration of VP06 (green) are compared in the right panel. This figure was made using the Seaborn (<http://seaborn.pydata.org/>) python package.

might not be robustly measured. The FWHM is relatively straightforward to measure even in low-S/N data, and FWHM measurements are usually provided as the primary measure of line width in survey catalogs (see, e.g., Shen et al. 2011; Pâris et al. 2017). However, this does not mean that the C IV FWHM provides an unbiased estimate of the C IV virial velocity in low-quality data. Furthermore, Denney et al. (2013) showed that high-quality data do not improve C IV FWHM-based BH mass

estimates, and the best-quality C IV-based BH masses are obtained using σ_{line} values measured from high-quality data.

All the calibrations presented in this work and similar works from the literature are, however, subject to sample biases from incompleteness of the calibration samples. Although we were able to expand the BH mass range to lower masses compared with previous work, there is still a lack of calibration objects at very high masses ($\gtrsim 10^9 M_{\odot}$) and in the regime of strong

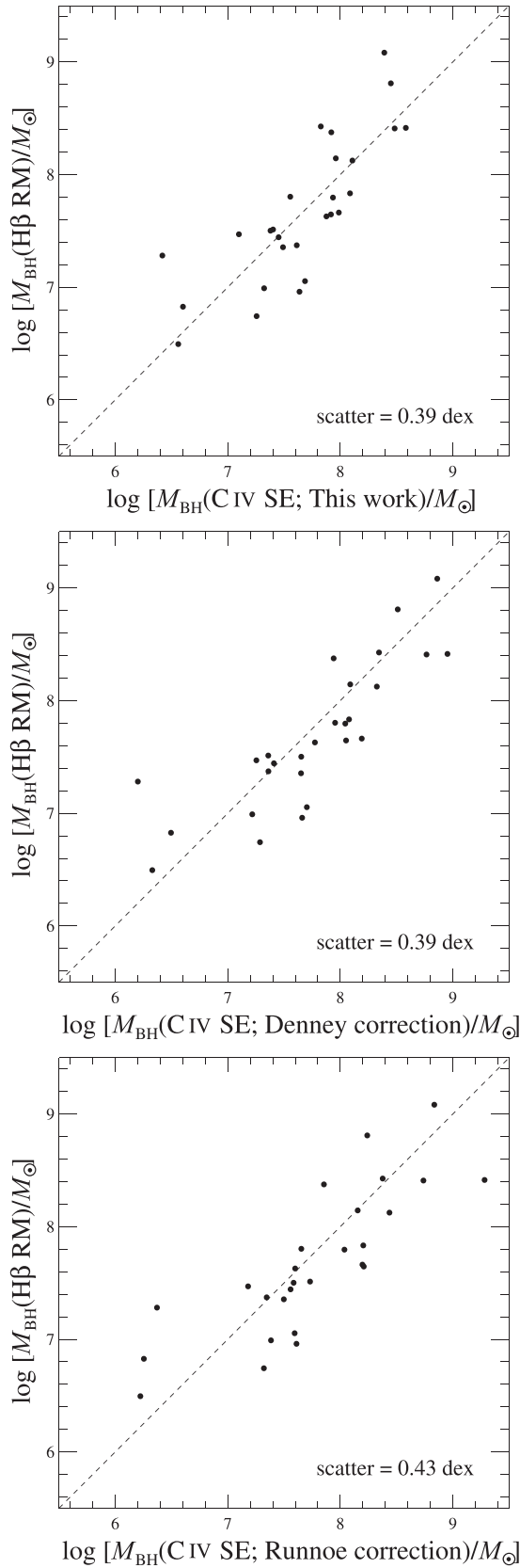


Figure 12. Comparisons of C IV FWHM-based SE M_{BH} estimates using our new calibration (top), the C IV line shape-based correction by Denney (2012) (middle), and the $\lambda 1400$ feature-based correction by Runnoe et al. (2013a) (bottom) to the H β RM-based BH masses for our sample of local RM AGNs. The dashed lines show a one-to-one relation, and the 1σ scatter is given in the lower right corner of each panel.

blueshifts. It is thus important to conduct direct tests of the reliability of the extrapolation of this calibration toward high-redshift, high-luminosity quasars, which commonly have high BH masses and/or strong C IV blueshift, as discussed by Richards et al. (2011) and Shen (2013). Current and future multi-object RM programs, including Sloan Digital Sky Survey Reverberation Mapping (Shen et al. 2015), the Australian spectroscopic Dark Energy Survey (King et al. 2015), and the Maunakea Spectroscopic Explorer (McConnachie et al. 2016), will help to improve this situation by providing direct reverberation measurements for the rest-frame UV lines in large numbers of quasars. There have been other efforts to improve the calibration of the C IV SE mass scale by taking into account C IV blueshifts (e.g., Shen & Liu 2012; Coatman et al. 2016, 2017) and by making use of other measured quantities, including UV-to-optical color, line shape, and nearby line peak flux ratio in the calibration of the SE method (Assef et al. 2011; Denney 2012; Runnoe et al. 2013a; Brotherton et al. 2015). However, the most fundamental and best way of achieving an accurate calibration of C IV-based BH mass estimators will be direct C IV reverberation mapping of significant samples of AGNs.

A spectrum with broad wavelength coverage observed simultaneously is essential in order to accurately investigate the rest-frame UV-to-optical continuum, the rest-frame UV-to-optical continuum, emission lines, and velocity offsets between the emission lines without suffering from systematics due to intrinsic AGN variability (see, e.g., Ho et al. 2012; Capellupo et al. 2015; and Mejía-Restrepo et al. 2016 for such data sets). However, even with such data at hand, it is difficult to achieve good continuum fits over the region from ~ 3100 to ~ 4000 Å if fitting the entire spectral region at once, due to the incompleteness and limitation of the currently used AGN Fe II emission templates. No available template covers the full UV/optical range, which is essential to constrain the Balmer continuum and Fe II emission accurately and continuously. There is a need for further improvement in Fe II templates, and an ideal data set for construction of a new template would consist of complete UV and optical spectra at high S/N observed with a small spectroscopic aperture to minimize starlight and narrow emission-line components. New *HST* observations are currently planned that will enable the construction of a new Fe II template using quasi-simultaneous UV and optical STIS data for the nearby Seyfert 1 galaxy Mrk 493 (program GO-14744; PI: Park).

Calibration of Mg II-based SE BH mass estimators using available *HST* archival spectra, as well as our STIS data, will be presented in a future paper. As an extension of the work of Wang et al. (2009), our STIS sample will provide an expanded BH mass range to calibrate the Mg II virial relationship, and we will provide σ_{line} - and MAD-based calibrations and updated FWHM-based calibrations using the uniform measurement and analysis methods for spectral decompositions, uncertainty estimates, and Bayesian linear regression presented in this work.

We thank the anonymous referee for a prompt and constructive report that has improved the paper. Support for *HST* program GO-12922 was provided by NASA through a grant from the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS 5-26555. D.P. acknowledges support through the EACOA Fellowship from

the East Asian Core Observatories Association, which consists of the National Astronomical Observatories, Chinese Academy of Science (NAOC), National Astronomical Observatory of Japan (NAOJ), Korean Astronomy and Space Science Institute (KASI), and Academia Sinica Institute of Astronomy and Astrophysics (ASIAA). Research by A.J.B. is supported in part by National Science Foundation (NSF) grant AST-1412693. JHW acknowledges support by the National Research Foundation of Korea grant funded by the Korean government (No. 2010-0027910 and No. 2016R1A2B3011457). T.T. acknowledges support by the Packard Foundation through a Packard Research Fellowship and by the NSF through grant AST-1412315. V.N.B. gratefully acknowledges assistance from NSF Research at Undergraduate Institutions (RUI) grant AST-1312296. Note that the findings and conclusions do not necessarily represent the views of the NSF. R.J.A. was supported by FONDECYT grant number 1151408. We thank Jonelle L. Walsh for help with the STIS data reduction.

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