View metadata, citation and similar papers at core.ac.uk
brought to you by CORE
provided by Caltech Authors - Main

DRAFT VERSION MAY 22, 2018
Typeset using LaTeX **twocolumn** style in AASTeX62

arXiv:1804.00667v2 [astro-ph.CO] 21 May 2018

CLUMP-3D. Testing ΛCDM with galaxy cluster shapes

MAURO SERENO,[1,2] KEIICHI UMETSU,[3] STEFANO ETTORI,[1,4] JACK SAYERS,[5] I-NON CHIU,[3] MASSIMO MENEGHETTI,[1,4] JESÚS VEGA-FERRERO,[6,7] AND ADI ZITRIN[8]

[1]*INAF - Osservatorio di Astrofisica e Scienza dello Spazio di Bologna, via Piero Gobetti 93/3, I-40129 Bologna, Italy*
[2]*Dipartimento di Fisica e Astronomia, Università di Bologna, via Piero Gobetti 93/2, I-40129 Bologna, Italy*
[3]*Institute of Astronomy and Astrophysics, Academia Sinica, P. O. Box 23-141, Taipei 10617, Taiwan*
[4]*INFN, Sezione di Bologna, viale Berti Pichat 6/2, I-40127 Bologna, Italia*
[5]*Division of Physics, Math, and Astronomy, California Institute of Technology, Pasadena, CA 91125*
[6]*IFCA, Instituto de Física de Cantabria (UC-CSIC), Av. de Los Castros s/n, 39005 Santander, Spain*
[7]*Department of Physics and Astronomy, University of Pennsylvania, 209 S. 33rd St, Philadelphia, PA 19104, USA*
[8]*Physics Department, Ben-Gurion University of the Negev, P.O. Box 653, Be'er-Sheva 84105, Israel*

ABSTRACT

The ΛCDM model of structure formation makes strong predictions on concentration and shape of DM (dark matter) halos, which are determined by mass accretion processes. Comparison between predicted shapes and observations provides a geometric test of the ΛCDM model. Accurate and precise measurements needs a full three-dimensional analysis of the cluster mass distribution. We accomplish this with a multi-probe 3D analysis of the X-ray regular CLASH (Cluster Lensing And Supernova survey with Hubble) clusters combining strong and weak lensing, X-ray photometry and spectroscopy, and the Sunyaev-Zel'dovich effect. The cluster shapes and concentrations are consistent with ΛCDM predictions. The CLASH clusters are randomly oriented, as expected given the sample selection criteria. Shapes agree with numerical results for DM-only halos, which hints at baryonic physics being not so effective in making halos rounder.

Keywords: galaxies: clusters: general — gravitational lensing: weak — galaxies: clusters: intracluster medium

1. INTRODUCTION

Cold Dark Matter (CDM) and the cosmological constant Λ are well established in modern astrophysics (Peebles 2015). Their case rests on precise measurements on very large scales, notably the cosmic microwave background radiation (Planck Collaboration et al. 2016) and the baryon acoustic oscillations in the power spectrum of the matter distribution (Xu et al. 2013).

The ΛCDM paradigm has been very successful and occasional crises have been solved. At small scales, the properties of the innermost regions of DM halos (the cusp-core problem) or the Milky Way's dwarf galaxy satellites (the missing satellites problem) can be reconciled with ΛCDM by a better understanding of baryonic physics or deeper observations (Weinberg et al. 2015). The over-concentration problem and the very large Einstein rings in galaxy clusters can be due to selection effects and statistical biases (Oguri & Blandford 2009; Meneghetti et al. 2011; Sereno et al. 2015).

Corresponding author: Mauro Sereno
mauro.sereno@oabo.inaf.it

Here, we propose a new test of ΛCDM based on the shape of cluster-sized halos. A three-dimensional analysis of galaxy clusters beyond the usual spherical hypothesis is crucial in modern astrophysics (Limousin et al. 2013). Halo shapes are the products of structure formation and evolution over billions of years. The matter aggregation from large-scale perturbations determines the shape. The major axes of galaxies and clusters often share the same orientation as the surrounding matter distribution (West 1994; Jing & Suto 2002) possibly since very early epochs (West et al. 2017).

Furthermore, we have to account for halo shape and orientation to get unbiased estimates of mass and concentration (Oguri et al. 2005; Morandi et al. 2012; Sereno et al. 2013), which are needed in precision cosmology exploiting formation and growth of galaxy clusters.

Shape measurements are very challenging and demand precise and accurate analyses. Shapes of individual clusters can be recovered with multi-probe approaches (Fox & Pen 2002; Mahdavi & Chang 2011; Morandi et al. 2012; Tchernin et al. 2016). The CLUMP-3D (CLUster Multi-Probes in Three Dimensions) project exploits rich data-sets to infer unbiased measurements of mass and concentration together with the intrinsic shape and the equilibrium status of the cluster (Sereno et al. 2017; Umetsu et al. 2018; Chiu

et al. 2018). The full three-dimensional Bayesian analysis combines strong (SL) and weak lensing (WL), X-ray surface brightness and temperature, and the Sunyaev-Zel'dovich effect (SZe) (Sereno et al. 2017). In a nutshell, lensing constrains the 2D mass and concentration which are deprojected thanks to the shape and orientation information from X-ray and SZe.

Here, we apply the method to the CLASH (Cluster Lensing and Supernova Survey with Hubble) clusters (Postman et al. 2012) and test if the recovered concentrations, shapes and orientations are in agreement with the ΛCDM predictions. This offers a novel, geometric check of the structure formation and evolution.

The reference cosmological model we test is the concordance flat ΛCDM universe with matter density parameter $\Omega_{\rm M} = 1 - \Omega_\Lambda = 0.3$, Hubble constant $H_0 = 100h\ {\rm km\,s^{-1}Mpc^{-1}}$ with $h = 0.7$, and power spectrum amplitude $\sigma_8 = 0.82$. Notations and conventions follows Sereno et al. (2017).

2. THE CLASH SAMPLE

CLASH is a Multi-Cycle Treasury program with the *Hubble Space Telescope* (HST) complemented with high-quality, multi-wavelength data-sets (Postman et al. 2012). Twenty massive clusters are X-ray selected over the redshift range $0.2 \lesssim z \lesssim 0.9$ on the basis of their high temperature ($kT > 5$ keV), and symmetric and smooth X-ray emission. Five additional clusters are included for their lensing strength to find magnified high-z galaxies.

We extend the method first applied to M1206 in Sereno et al. (2017) to the 16 X-ray regular CLASH clusters with high quality, ground-based data for WL (Table 1). We do not consider the five lensing selected clusters, which are mostly merging or irregular systems. In fact, our modelling requires that matter and gas follow an ellipsoidal geometry and that gas and matter are aligned and co-centered. Modeling of complex distributions could require using finite mixture models as collections of ellipsoids that fit individual subclusters (Kuhn et al. 2014).

The data-sets have been comprehensively presented and detailed elsewhere. In the following, we provide the main references, and detail any change with respect to Sereno et al. (2017).

Weak lensing at large radii. The lensing analysis relies on ground-based data from the Suprime-Cam at Subaru Telescope or the ESO Wide Field Imager (Umetsu et al. 2014, 2016; Merten et al. 2015). We mainly refer to Umetsu et al. (2018), where projected mass maps are recovered from the joint analysis of shear and magnification bias. Our fitting analysis follows Sereno et al. (2017) but we do not limit the fit to a squared region of size of 4 Mpc$/h$, instead we consider the full field of view (Chiu et al. 2018). Accordingly, the noise from the large-scale structure is added to the uncertainty covariance matrix.

Strong and weak lensing in the cluster cores. Multiple image systems, shear in the *HST* field, and mass models of the inner cluster regions are presented in Zitrin et al. (2015). As in Sereno et al. (2017), we compute the mean convergence from the 'PIEMDeNFW' maps in equally spaced circular annuli. The innermost and the outermost radii are set to an angular scale of $5''$and two times the Einstein radius, $\theta_{\rm max,\,SL} = 2\theta_{\rm E}(z_{\rm s} = 2)$, respectively. The width of the annuli, $\Delta\theta_{\rm SL}$, is determined through the relation $N_{\rm im}\Delta\theta_{\rm SL}^2 \sim \pi\theta_{\rm E}^2$, where $N_{\rm im}$ is the number of images (Umetsu et al. 2016). We fix $N_{\rm im} = N_{\rm SL}/2$, where $N_{\rm SL}$ is the number of effective SL constraints (Zitrin et al. 2015, table 1). The bin size is rounded to have equally spaced annuli.

X-ray. Gas density and temperature profiles are measured from archival *Chandra* data, see Table 1. The data are analysed as in Sereno et al. (2017). Cleaned (by grade, status, bad pixels, and time intervals affected from flares in the background count rate) events file are prepared with the CIAO 4.8 software[1] and the calibration database CALDB 4.7.1. Backgrounds are extracted far from the cluster X-ray peak in circular regions of radius of $2'$.

Exposure-corrected images are produced in the [0.7-2] keV band. The point-sources are identified with the tool `wavedetct`, filtered out, and the corresponding regions filled with values of counts from surrounding background areas through the tool `dmfilth`.

We perform the 2D analysis of the number counts in the circular region enclosing 80% of the total source emission. Pixels are binned with a final resolution of $1.968''$. We excise the inner region of radius $5''$. Outside the 80% region, we examine the surface brightness profiles in circular annuli.

Spectra are extracted in circular annuli and analysed with the XSPEC v.12.9 software[2] with an absorbed thermal model `tbabs`, the Galactic absorption fixed by extrapolation from HI radio maps in Kalberla et al. (2005) and the `apec` emission spectrum with three free parameters (normalisation, temperature and metallicity). The same model with metallicity fixed to the median value is used for regression.

SZe. The CLASH clusters are part of the Bolocam X-ray SZ (BoXSZ) sample, with publicly available data from Bolocam. Details of the observational campaign and data reduction can be found in Sayers et al. (2011, 2016); Czakon et al. (2015). Additional data from the *Planck* all-sky survey are employed to set the mean signal estimates (Sereno et al. 2017).

Our analysis strictly follows Sereno et al. (2017). The integrated Compton parameter is computed from the unfiltered maps in five equally spaced annular bins up to a maximum radius of $5'$. The annulus width is set to $1'$, comparable to the PSF FWHM.

3. THE MODEL

The parametric joint analysis of the multi-probe data-sets is introduced in Sereno et al. (2017). The main assumption is that the total matter distribution is ellipsoidal. The halo shape is described by the two axis ratios: $q_{\rm mat,1}$ is the minor-

[1] http://cxc.harvard.edu/ciao/

[2] https://heasarc.gsfc.nasa.gov/xanadu/xspec/

Table 1. The cluster sample. Columns 4 and 5: right ascension and declination in degrees (J2000) of the associated brightest cluster galaxy (BCG), adopted as the cluster center. The X-ray data set is detailed in cols. 6–8. Col. 6: observation identification. Col. 7: nominal exposure time. Col. 8: Galactic absorption.

Name		z	RA	DEC	ObsID	$t_{\rm exp}$	$n_{\rm H}$
Full	Short					[ks]	$[10^{20}{\rm cm}^{-2}]$
ABELL 0383	A383	0.188	42.01409	−3.5292641	522/3579 (ACIS-I)	19.6	1.6
ABELL 0209	A209	0.206	22.96895	−13.6112720	524/2320 (ACIS-I)	28.2	3.5
ABELL 2261	A2261	0.225	260.61336	32.1324650	3194 (ACIS-S)	32.0	4.5
RX J2129+0005	R2129	0.234	322.41649	0.0892232	5007 (ACIS-I)	24.1	3.2
ABELL 0611	A611	0.288	120.23674	36.0565650	3257/3582/6108/7719 (ACIS-I)	71.3	4.6
MS 2137.3-2353	MS2137	0.313	325.06313	−23.6611360	3271 (ACIS-I)	22.6	4.3
RXC J2248.7-4431	R2248	0.348	342.18322	−44.5309080	3585/6111 (ACIS-I)	66.2	5.7
MACS J1115+0129	M1115	0.352	168.96627	1.4986116	3275/9375 (ACIS-I)	49.1	4.3
MACS J1931.8-2635	M1931	0.352	292.95608	−26.5758570	3277 (ACIS-I)	22.9	3.7
RX J1532.8+3021	R1532	0.363	233.22410	30.3498440	3280/6107 (ACIS-I)	51.9	3.4
MACS J1720.3+3536	M1720	0.391	260.06980	35.6072660	3282/9382 (ACIS-I)	111.2	8.3
MACS J0429.6-0253	M0429	0.399	67.40003	−2.8852066	5250 (ACIS-S)	18.9	3.8
MACS J1206.2-0847	M1206	0.44	181.55065	−8.8009395	3592/13516/13999 (ACIS-I)	149.8	4.6
MACS J0329.6-0211	M0329	0.45	52.42320	−2.1962279	14009 (ACIS-S)	85.8	2.3
RX J1347.5-1145	R1347	0.451	206.87756	−11.7526100	552/9370 (ACIS-I)	39.6	3.6
MACS J0744.9+3927	M0744	0.686	116.22000	39.4574080	4966/18611/18818 (ACIS-I)	123.1	1.2

to-major axis ratio, and $q_{\rm mat,2}$ is the intermediate-to-major axis ratio. The orientation is fixed by three Euler's angles: ϑ is the inclination angle between the major axis and the line-of-sight.

The gas distribution is assumed to be ellipsoidal too and co-centered and co-aligned with the total matter. The gas is taken to be more spherical than the total matter, as usual in regular systems, but we do not require that the gas is in equilibrium in the potential well. This modelling is supported by the analysis of the 2D maps (Umetsu et al. 2018), which show that gas and total matter have a negligible off-set and are aligned in projection, and that a constant matter ellipticity as a function of the radius provides a good description of the data.

WL and X-ray/SZ data probe dark matter and gas, respectively, on different scales. Whereas the matter shape is measured within the viral region, the gas shape is mostly sensitive to the inner regions (Sereno et al. 2017).

We fit the SL and WL convergence maps, the X-ray surface brightness and temperature, and the integrated Compton parameter. The total matter distribution, the gas density, and the gas temperature are modelled with flexible parametric ellipsoidal 3D profiles. Since we do not require hydrostatic equilibrium, the determinations of the matter or the gas density profiles are largely independent apart from an overall normalization related to the orientation (Sereno et al. 2017). The more elongated the cluster is along the line-of-sight, as mainly inferred from X-ray and SZe measurements, the smaller the central gas density, and the smaller the total mass and concentration. Here, we are interested in the global shape and we do not discuss the gas properties, which are simultaneously fitted in the same regression procedure. To conservatively deal with parameter degeneracy, we focus on 1D probability distributions obtained after marginalization of the remaining parameters.

The total mass (DM plus galaxies plus gas) is described as a Navarro-Frenk-White (NFW) density profile (Navarro et al. 1996),

$$\rho_{\rm NFW} = \frac{\rho_{\rm s}}{(\zeta/\zeta_{\rm s})(1+\zeta/\zeta_{\rm s})^2}, \tag{1}$$

where ζ is the ellipsoidal radius and $\zeta_{\rm s}$ is the scale radius. In our notation, M_Δ is the mass within the ellipsoid of semi-major axis ζ_Δ,

$$M_\Delta \equiv (4\pi/3)\Delta\, \rho_{\rm cr}(z)\, q_{\rm mat,1} q_{\rm mat,2} \zeta_\Delta^3, \tag{2}$$

where $\rho_{\rm cr}(z)$ is the critical density. The concentration is $c_\Delta \equiv \zeta_\Delta/\zeta_{\rm s}$. For comparison with numerical simulations, we also consider the mass $M_{\rm sph,\Delta}$ measured in spherical regions, and the concentration $c_{\rm sph,\Delta}$ computed by fitting the spherically averaged NFW profile.

The relation between $\rho_{\rm s}$ and the concentration takes the same form in the spherical or ellipsoidal model, whereas $r_{\rm sph,s} \sim (q_{\rm mat,1} q_{\rm mat,2})^{1/3}\zeta_{\rm s}$. Then, $c_{\rm sph,\Delta} < c_\Delta$.

For our Bayesian analysis, we adopt priors spanning large parameter regions. For mass and concentration, priors are uniform distributions in logarithmically spaced intervals, as suitable for positive parameters (Sereno & Covone 2013): $p_{\rm prior}(M_{200}) \propto 1/M_{200}$ and $p_{\rm prior}(c_{200}) \propto 1/c_{200}$ in the allowed ranges $0.01 \le M_{200}/(10^{15}h^{-1}M_\odot) \le 10$ and

$1 \leq c_{200} \leq 10$, or null otherwise. We assume a flat prior for the matter shape (q-flat), i.e. the probability $p_{\rm prior}(q_{\rm mat,1})$ and the conditional probability $p_{\rm prior}(q_{\rm mat,2}|q_{\rm mat,1})$ are constant. The minimum axis ratio is $q_{\rm min} = 0.1$, and $q_{\rm min} \leq q_{\rm mat,1} \leq q_{\rm mat,2} \leq 1$. This prior allows either very triaxial clusters ($q_{\rm mat,1}, q_{\rm mat,2} \ll 1$) or spherical systems ($q_{\rm mat,1} \lesssim 1$), which are preferentially excluded by N-body simulations (Jing & Suto 2002). A priori, the cluster orientation is random, i.e. $p_{\rm prior}(\cos\vartheta) = 1$.

4. THEORETICAL PREDICTIONS

Numerical simulations can picture the cluster properties in the ΛCDM scenario. We consider two kinds of simulated samples: i) halos selected as actual CLASH clusters ('MUSIC2-CLASH'); ii) a complete population of relaxed massive halos ('ΛCDM-rel').

MUSIC2-CLASH. Simulated clusters mimicking the CLASH sample are presented in Meneghetti et al. (2014), which study nearly 1400 halos over $0.25 \leq z \leq 0.67$ from the MUSIC-2 N-body/hydrodynamical simulations[3]. These halos are mass-limited ($> 10^{15} M_\odot/h$ at $z = 0$) and are re-simulated by adding baryons to the DM distributions (Sembolini et al. 2013). Here, we consider the runs not including radiative processes.

Meneghetti et al. (2014) classified halos as regular or relaxed, the two conditions being non-equivalent. Regular clusters are so in their X-ray features. They show small centroid shift, small ellipticity, and small third- and fourth-order power ratios of the X-ray morphology in the soft-energy band, but large X-ray surface-brightness concentrations. On the other hand, clusters are classified as relaxed according to their center of mass displacement, their virial ratio, and their substructure mass fraction (Neto et al. 2007). Regular clusters can be unrelaxed. The X-ray morphology is mostly evaluated in the inner regions, whereas relaxation is evaluated on scales up to the virial radius.

Meneghetti et al. (2014) find simulated avatars of the CLASH X-ray-selected clusters by matching the X-ray morphology. The association does not use gas temperatures or X-ray luminosities, whose physical processes are more challenging to simulate. Masses and redshifts have to be compatible too.

Only the simulated halos closest to each individual CLASH cluster in the morphological parameter space are used. The total number of associations analysed here is 166, with from 2 to 26 associations per cluster. CLASH clusters are found to be prevalently relaxed and only modestly affected by the strong lensing bias. The regularity of the matched clusters is not extreme (Meneghetti et al. 2014). The fraction of $\sim$ 70 per cent of relaxed halos among X-ray regular clusters is larger than in the full sample, and the average concentration is boosted.

We measure the shapes and orientations of the total matter distribution, i.e. DM plus gas particles, by mimicking our real measurement process (Bonamigo et al. 2015). We compute the mass tensor of the particles selected inside the ellipsoid, centred in the most bound particle, that encloses an overdensity $\Delta = 200$. The procedure was reiterated until both $q_{\rm mat,1}$ and $q_{\rm mat,2}$ converge within a 0.5% of error.

ΛCDM-rel. Relaxed galaxy clusters in N-body simulations are well represented as a population of ellipsoidal, co-aligned, triaxial halos (Jing & Suto 2002; Allgood et al. 2006; Bonamigo et al. 2015; Vega-Ferrero et al. 2017). Results from different groups agree if methodological differences on how the halo shape is measured are taken into account (Vega-Ferrero et al. 2017).

We base the ΛCDM prediction for halo shape on Bonamigo et al. (2015), who analyse the relaxed halos from the Millennium XXL simulation and provide statistically significant predictions for masses above $3 \times 10^{14} M_\odot/h$. They measure the shape of the ellipsoid enclosing an overdensity equal to the virial one. Unrelaxed clusters are removed by selecting only halos whose offset between the most bound particle and the centre of mass of the particles enclosed by the ellipsoid is less than 5 per cent of their virial radius. This criterion is not very stringent and a fraction of unrelaxed clusters might still be included.

Theoretical estimates provide a consistent picture of the halo concentrations (Bhattacharya et al. 2013; Dutton & Macciò 2014; Ludlow et al. 2016). As reference prediction, we follow Meneghetti et al. (2014), who measure the mass-concentration relation under different selection criteria and for either projected or 3D concentrations and masses. The MUSIC-2 halos follow an intrinsic concentration-mass relation with a slightly larger normalisation than other recent results, but with the usual weak redshift evolution. The more sensible comparison to our analysis is with the NFW fitting in 3D of the relaxed sample.

We assume that the unbiased population of relaxed clusters is randomly oriented.

5. RESULTS

Cluster masses and concentrations, and shape and orientation parameters are listed in Table 2. We fitted all the CLUMP-3D parameters (Sereno et al. 2017, table 1), including the gas parameters, which are not reported here and will be discussed separately.

Halo parameters are usually measured assuming spherical symmetry and by fitting the projected maps. Masses and concentrations so determined are called 2D. Here, we measure shape and concentration in 3D and we average the ellipsoidal mass profile in spherical regions. Our spherical masses are then unbiased.

The sample distributions are shown in Fig. 1. The theoretical ΛCDM prediction for each cluster is computed based on the observed mass distribution. For the MUSIC2-CLASH simulated clusters, we consider random samples of 16 associations, one per cluster. The expected distribution for each cluster is convolved with a Gaussian whose dispersion is equal to the statistical uncertainty on the one-point estimate.

[3] http://music.ft.uam.es/

Table 2. Halo properties. $M_{\rm sph,200}$ (col. 2) and $c_{\rm sph,200}$ (col. 3) refer to mass and concentration measured in spheres. M_{200} (col. 4) and c_{200} (col. 5) refer to ellipsoids. Masses are in units of $10^{15}\ M_\odot/h$. We quote the bi-weighted estimators of the marginalised posterior distributions.

Cluster	$M_{\rm sph,200}$	$c_{\rm sph,200}$	M_{200}	c_{200}	$q_{\rm mat,1}$	$q_{\rm mat,2}$	$\cos\vartheta$
A383	0.56 ± 0.11	5.87 ± 0.96	0.63 ± 0.14	6.75 ± 1.20	0.30 ± 0.06	0.63 ± 0.14	0.25 ± 0.15
A209	0.80 ± 0.28	1.66 ± 0.45	1.02 ± 0.25	1.83 ± 0.42	0.28 ± 0.12	0.51 ± 0.24	0.97 ± 0.03
A2261	1.63 ± 0.28	5.13 ± 0.61	2.06 ± 0.37	6.94 ± 0.92	0.17 ± 0.04	0.71 ± 0.17	0.40 ± 0.19
R2129	0.49 ± 0.13	4.55 ± 0.92	0.55 ± 0.16	5.17 ± 1.09	0.32 ± 0.08	0.68 ± 0.13	0.39 ± 0.17
A611	0.62 ± 0.15	2.91 ± 0.74	0.75 ± 0.18	3.47 ± 0.88	0.28 ± 0.12	0.74 ± 0.21	0.94 ± 0.03
MS2137	0.66 ± 0.16	3.60 ± 0.92	0.71 ± 0.17	3.78 ± 1.07	0.51 ± 0.14	0.85 ± 0.12	0.53 ± 0.19
R2248	1.16 ± 0.26	4.38 ± 1.06	1.30 ± 0.29	4.93 ± 1.33	0.34 ± 0.08	0.58 ± 0.16	0.18 ± 0.14
M1115	1.32 ± 0.27	3.77 ± 0.67	1.63 ± 0.34	4.64 ± 0.90	0.22 ± 0.04	0.58 ± 0.13	0.20 ± 0.13
M1931	0.53 ± 0.14	5.51 ± 1.71	0.57 ± 0.16	5.94 ± 1.68	0.47 ± 0.13	0.72 ± 0.19	0.71 ± 0.26
R1532	0.51 ± 0.13	6.00 ± 1.08	0.58 ± 0.16	7.12 ± 1.39	0.30 ± 0.09	0.77 ± 0.16	0.25 ± 0.21
M1720	0.73 ± 0.17	5.84 ± 1.45	0.80 ± 0.18	6.62 ± 1.74	0.36 ± 0.09	0.69 ± 0.14	0.41 ± 0.21
M0429	0.51 ± 0.10	7.03 ± 1.56	0.54 ± 0.11	7.58 ± 1.69	0.45 ± 0.09	0.85 ± 0.13	0.39 ± 0.25
M1206	1.20 ± 0.15	6.14 ± 0.60	1.33 ± 0.19	6.86 ± 0.84	0.33 ± 0.07	0.69 ± 0.13	0.20 ± 0.14
M0329	0.83 ± 0.16	4.34 ± 0.89	0.93 ± 0.22	4.92 ± 0.97	0.33 ± 0.10	0.58 ± 0.19	0.65 ± 0.22
R1347	2.45 ± 0.46	4.61 ± 0.82	3.23 ± 0.61	6.34 ± 1.33	0.15 ± 0.03	0.66 ± 0.24	0.50 ± 0.19
M0744	1.23 ± 0.38	4.36 ± 1.09	1.76 ± 0.53	6.55 ± 1.69	0.11 ± 0.01	0.46 ± 0.17	0.16 ± 0.12

Table 3. Test hypothesis. Measurements are compared to the theoretical predictions, either MUSIC2-CLASH (rows 1-5) or ΛCDM-rel (rows 6-10). For $q_{\rm mat,1}$, we also report results for the subsample with $\bar{q}_{\rm mat,1} > 0.2$. The reported p-values are computed with either the Kolmogorov-Smirnov or the Pearson χ^2 test. We report the mean and the standard deviation of the p-value distribution (cols. 2 and 4), accounting for finite sample size and observational uncertainties, and the upper limit of the 2σ-confidence region (cols. 3 and 5).

Parameter	Kolmogorov-Smirnov		Pearson χ^2	
	MUSIC2-CLASH			
$c_{\rm sph,200}$	0.63 ± 0.29	$\lesssim 1$	0.52 ± 0.25	0.97
$q_{\rm mat,1}$	0.01 ± 0.03	0.09	0.21 ± 0.20	0.73
$q_{\rm mat,2}$	0.59 ± 0.31	$\lesssim 1$	0.50 ± 0.25	0.97
$\cos\vartheta$	0.40 ± 0.31	0.95	0.39 ± 0.24	0.93
Subsample $\bar{q}_{\rm mat,1} > 0.2$				
$q_{\rm mat,1}$	0.07 ± 0.13	0.59	0.32 ± 0.23	0.87
	ΛCDM-rel			
$c_{\rm sph,200}$	0.45 ± 0.33	$\lesssim 1$	0.41 ± 0.25	0.94
$q_{\rm mat,1}$	0.08 ± 0.13	0.43	0.24 ± 0.19	0.73
$q_{\rm mat,2}$	0.53 ± 0.31	$\lesssim 1$	0.45 ± 0.25	0.96
$\cos\vartheta$	0.59 ± 0.31	$\lesssim 1$	0.47 ± 0.25	0.97
Subsample $\bar{q}_{\rm mat,1} > 0.2$				
$q_{\rm mat,1}$	0.25 ± 0.27	0.90	0.38 ± 0.24	0.87

The sample distributions are finally obtained by averaging the distributions of the single clusters. Measurements agree with ΛCDM predictions. For a quantitative assessment, we performed the Kolmogorov-Smirnov and the Pearson χ^2 tests, see Table 3.

Triaxial analyses facilitate the agreement of concentrations measured in massive lensing clusters with theoretical predictions (Oguri et al. 2005; Sereno & Zitrin 2012). In fact, the 3D analysis does not suffer from the orientation bias that can affect clusters preferentially elongated along the line-of-sight, whose concentration is over-estimated under the spherical hypothesis.

The CLASH clusters show a triaxial shape in good agreement with theoretical predictions. The distribution of $q_{\rm mat,1}$ slightly exceeds expectations at small values but the excess is not significant. The excess is mostly driven by three possibly unrelaxed clusters with $q_{\rm mat,1} \lesssim 0.2$, where the 3D combined analysis might experience some problems. M0744 and R1347 show evidence for shocks in high resolution *MUSTANG* SZe data (Korngut et al. 2011). R1347 hosts a radio mini-halo. Diffuse radio emission was suspected in A2261 (Giacintucci et al. 2017).

Adiabatic contraction and radiative cooling can make the total mass distribution rounder but AGN (Active Galactic Nuclei) feedback can mitigate the effect of cooling and make the final shape more similar to results in DM-only simulations (Suto et al. 2017). If spherically averaged profiles are considered, the baryon physics is important only within less than 10% of the virial radius. On the other hand, the non-sphericity of the DM distribution can be affected by the baryon physics operating in the central region up to half the virial radius (Suto et al. 2017).

The MUSIC2-CLASH shapes are measured in non-radiative simulations and may differ from simulations accounting for feedback processes. In fact, the agreement is better for DM-only simulations.

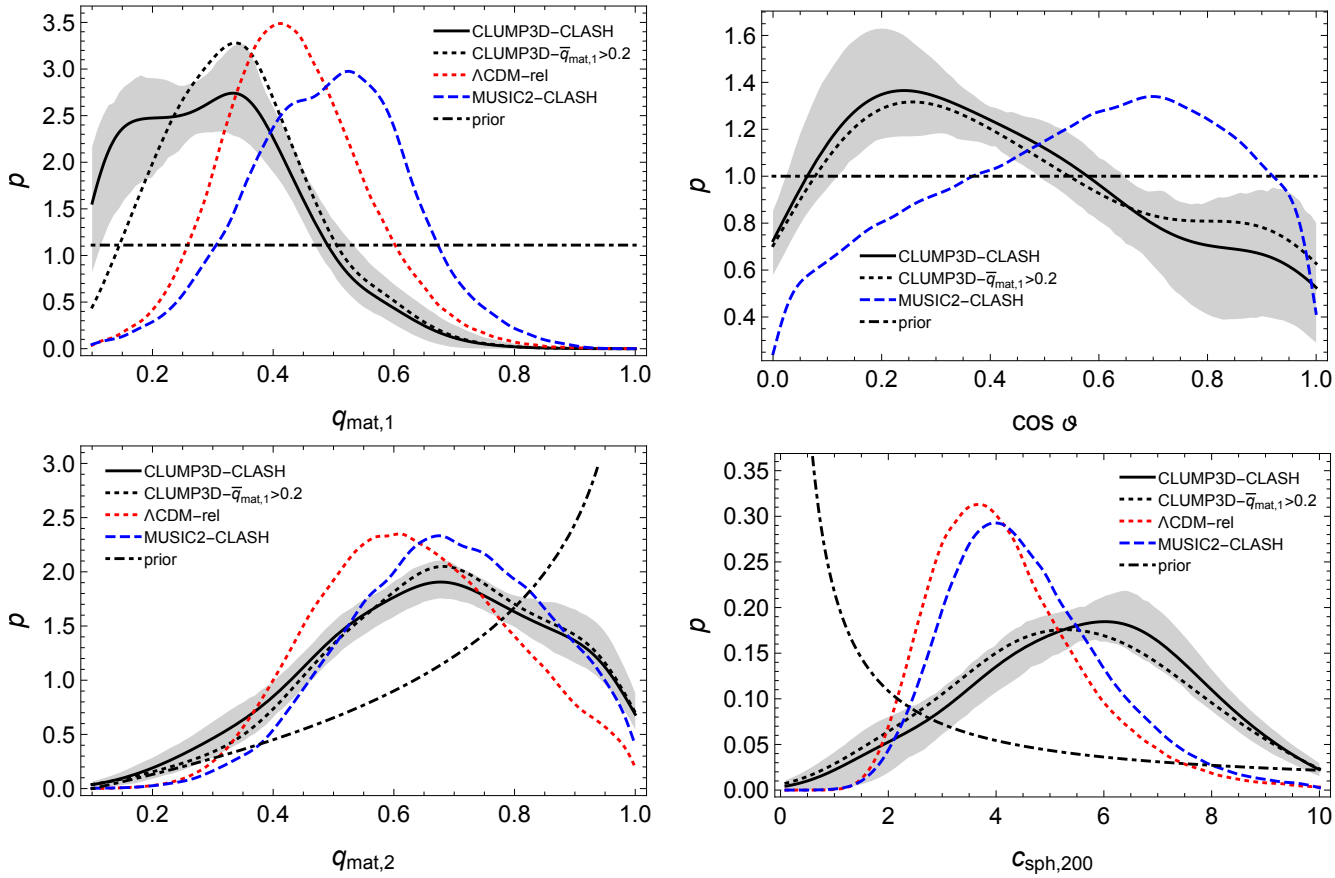

Figure 1. Probability density functions: minor-to-major axis ratio (top-left panel), intermediate-to-major axis ratio (bottom-left), inclination angle $\cos\vartheta$ (top-right), and concentration (bottom-right) of the total matter distribution of the CLASH clusters. The full-black lines show the measured distribution. The shadowed regions include the 1-σ region as obtained from a bootstrap resampling of the marginalized distributions. The black dotted lines show the distributions for the subsample with $\bar{q}_{\rm mat,1}$>0.2. The red-dotted lines show the ΛCDM-rel predictions, estimated from Bonamigo et al. (2015) given the inferred mass distribution smoothed for the observational uncertainties for the axial ratios, or the scattered mass-concentration relation of relaxed clusters from Meneghetti et al. (2014) for the concentration. The blue-dashed line is the expected distribution for the MUSIC2-CLASH simulated halos. The dot-dashed black lines show the priors. For visualisation purposes, the measured and the MUSIC2-CLASH distribution are smoothed by the Silverman's scale of the measured one-point estimates (Vio et al. 1994).

X-ray regular clusters may suffer from orientation bias: prolate clusters whose major axis is aligned with the line-of-sight show round X-ray isophotes and can be preferentially included. The observed distribution of inclination angles is consistent with random orientations. The secondary peak of the distribution at high values is due to a couple of very elongated clusters as A209 which is a known merger system along the line-of-sight (Cassano et al. 2010). Our algorithm can efficiently recover the orientation even in this peculiar system.

For the CLASH sample, we find $\cos\vartheta = 0.42 \pm 0.31$ or $\vartheta = 65 \pm 23\,\mathrm{deg}$. The mean angle for random orientations (MUSIC2-CLASH) is $60\,\mathrm{deg}$ or $\cos\vartheta = 0.5$ ($54\,\mathrm{deg}$ or $\cos\vartheta = 0.57$), consistent with our results and giving no evidence for any orientation bias.

6. CONCLUSIONS

We compared shapes and concentrations of X-ray selected CLASH clusters to the ΛCDM paradigm of structure formation. We performed a full three-dimensional analysis of the cluster mass profiles exploiting lensing, X-ray, and SZe measurements. We could then measure unbiased concentrations and masses and recover the halo shapes and orientations. We found that shapes and concentrations of the CLASH clusters are consistent with theoretical predictions giving a further validation of the ΛCDM paradigm.

Even though results from simulations are consistent with our measurements, we still lack a comprehensive analytical model of halo formation. The conventional ellipsoidal collapse mode and the simulations differ after the turn-around epoch (Suto et al. 2016). While simulated massive halos are more spherical initially, they gradually become less spherical after the turn-around epoch. This tendency is opposite to the analytical prediction (Suto et al. 2016).

Furthermore, the role of gas physics in halo shapes is yet to be fully understood (Suto et al. 2017). In DM-only simulations, the inner regions are less spherical than the outer ones (Allgood et al. 2006; Suto et al. 2016; Vega-Ferrero et al. 2017; Despali et al. 2017). Internal parts retain memory of the violent formation process keeping the major axis oriented towards the preferential direction of the infalling material while the outer regions become rounder due to continuous isotropic merging events (Despali et al. 2017).

However, baryonic physics can significantly affect the non-sphericity of the DM distribution up to the half of the virial radius (Suto et al. 2017). Radiative cooling makes DM halos more spherical (Kazantzidis et al. 2004) but AGN feedback can counterbalance. As a result, the radial trend of asphericity can be opposite to that predicted in DM-only simulations (Suto et al. 2017).

Our results suggest that baryonic physics is not effective in making cluster rounder.

The authors thank Rossella Cassano and Gianfranco Brunetti for useful discussions, and Gustavo Yepes for supplying the MUSIC-2 simulations performed at the Barcelona Supercomputing Center, Spain. MS and SE acknowledge support from the contracts ASI-INAF I/009/10/0, NARO15 ASI-INAF I/037/12/0, ASI 2015-046-R.0 and ASI-INAF n.2017-14-H.0. KU acknowledges support from the Ministry of Science and Technology of Taiwan (grants 103-2112-M-001-030-MY3 and 106-2628-M-001-003-MY3) and from the Academia Sinica Investigator Award. JS was supported by NSF/AST-1617022. MM acknowledges support from MAECI and contracts ASI-INAF/I/023/12/0 and ASI n. I/023/12/0. JV-F was supported by AYA2015-64508-P (MINECO/FEDER).

REFERENCES

Allgood, B., Flores, R. A., Primack, J. R., et al. 2006, MNRAS, 367, 1781, doi: 10.1111/j.1365-2966.2006.10094.x

Bhattacharya, S., Habib, S., Heitmann, K., & Vikhlinin, A. 2013, ApJ, 766, 32, doi: 10.1088/0004-637X/766/1/32

Bonamigo, M., Despali, G., Limousin, M., et al. 2015, MNRAS, 449, 3171, doi: 10.1093/mnras/stv417

Cassano, R., Ettori, S., Giacintucci, S., et al. 2010, ApJL, 721, L82, doi: 10.1088/2041-8205/721/2/L82

Chiu, I., Umetsu, K., Sereno, M., et al. 2018, arXiv:1804.00676. https://arxiv.org/abs/1804.00676

Czakon, N. G., Sayers, J., Mantz, A., et al. 2015, ApJ, 806, 18, doi: 10.1088/0004-637X/806/1/18

Despali, G., Giocoli, C., Bonamigo, M., Limousin, M., & Tormen, G. 2017, MNRAS, 466, 181, doi: 10.1093/mnras/stw3129

Dutton, A. A., & Macciò, A. V. 2014, MNRAS, 441, 3359, doi: 10.1093/mnras/stu742

Fox, D. C., & Pen, U.-L. 2002, ApJ, 574, 38, doi: 10.1086/340897

Giacintucci, S., Markevitch, M., Cassano, R., et al. 2017, ApJ, 841, 71, doi: 10.3847/1538-4357/aa7069

Jing, Y. P., & Suto, Y. 2002, ApJ, 574, 538, doi: 10.1086/341065

Kalberla, P. M. W., Burton, W. B., Hartmann, D., et al. 2005, A&A, 440, 775, doi: 10.1051/0004-6361:20041864

Kazantzidis, S., Kravtsov, A. V., Zentner, A. R., et al. 2004, ApJ, 611, L73, doi: 10.1086/423992

Korngut, P. M., Dicker, S. R., Reese, E. D., et al. 2011, ApJ, 734, 10, doi: 10.1088/0004-637X/734/1/10

Kuhn, M. A., Feigelson, E. D., Getman, K. V., et al. 2014, ApJ, 787, 107, doi: 10.1088/0004-637X/787/2/107

Limousin, M., Morandi, A., Sereno, M., et al. 2013, SSRv, 177, 155, doi: 10.1007/s11214-013-9980-y

Ludlow, A. D., Bose, S., Angulo, R. E., et al. 2016, MNRAS, 460, 1214, doi: 10.1093/mnras/stw1046

Mahdavi, A., & Chang, W. 2011, ApJL, 735, L4, doi: 10.1088/2041-8205/735/1/L4

Meneghetti, M., Fedeli, C., Zitrin, A., et al. 2011, A&A, 530, A17, doi: 10.1051/0004-6361/201016040

Meneghetti, M., Rasia, E., Vega, J., et al. 2014, ApJ, 797, 34, doi: 10.1088/0004-637X/797/1/34

Merten, J., Meneghetti, M., Postman, M., et al. 2015, ApJ, 806, 4, doi: 10.1088/0004-637X/806/1/4

Morandi, A., Limousin, M., Sayers, J., et al. 2012, MNRAS, 425, 2069, doi: 10.1111/j.1365-2966.2012.21196.x

Navarro, J. F., Frenk, C. S., & White, S. D. M. 1996, ApJ, 462, 563

Neto, A. F., Gao, L., Bett, P., et al. 2007, MNRAS, 381, 1450, doi: 10.1111/j.1365-2966.2007.12381.x

Oguri, M., & Blandford, R. D. 2009, MNRAS, 392, 930, doi: 10.1111/j.1365-2966.2008.14154.x

Oguri, M., Takada, M., Umetsu, K., & Broadhurst, T. 2005, ApJ, 632, 841, doi: 10.1086/452629

Peebles, P. J. E. 2015, Proceedings of the National Academy of Science, 112, 12246, doi: 10.1073/pnas.1308786111

Planck Collaboration, Ade, P. A. R., Aghanim, N., et al. 2016, A&A, 594, A13, doi: 10.1051/0004-6361/201525830

Postman, M., Coe, D., Benítez, N., et al. 2012, ApJS, 199, 25, doi: 10.1088/0067-0049/199/2/25

Sayers, J., Golwala, S. R., Ameglio, S., & Pierpaoli, E. 2011, ApJ, 728, 39, doi: 10.1088/0004-637X/728/1/39

Sayers, J., Golwala, S. R., Mantz, A. B., et al. 2016, ApJ, 832, 26, doi: 10.3847/0004-637X/832/1/26

Sembolini, F., Yepes, G., De Petris, M., et al. 2013, MNRAS, 429, 323, doi: 10.1093/mnras/sts339

Sereno, M., & Covone, G. 2013, MNRAS, 434, 878, doi: 10.1093/mnras/stt1086

Sereno, M., Ettori, S., Meneghetti, M., et al. 2017, MNRAS, 467, 3801, doi: 10.1093/mnras/stx326

Sereno, M., Ettori, S., Umetsu, K., & Baldi, A. 2013, MNRAS, 428, 2241, doi: 10.1093/mnras/sts186

Sereno, M., Giocoli, C., Ettori, S., & Moscardini, L. 2015, MNRAS, 449, 2024, doi: 10.1093/mnras/stv416

Sereno, M., & Zitrin, A. 2012, MNRAS, 419, 3280, doi: 10.1111/j.1365-2966.2011.19968.x

Suto, D., Kitayama, T., Nishimichi, T., Sasaki, S., & Suto, Y. 2016, PASJ, 68, 97, doi: 10.1093/pasj/psw088

Suto, D., Peirani, S., Dubois, Y., et al. 2017, PASJ, 69, 14, doi: 10.1093/pasj/psw118

Tchernin, C., Eckert, D., Ettori, S., et al. 2016, A&A, 595, A42, doi: 10.1051/0004-6361/201628183

Umetsu, K., Zitrin, A., Gruen, D., et al. 2016, ApJ, 821, 116, doi: 10.3847/0004-637X/821/2/116

Umetsu, K., Medezinski, E., Nonino, M., et al. 2014, ApJ, 795, 163, doi: 10.1088/0004-637X/795/2/163

Umetsu, K., Sereno, M., Tam, S.-I., et al. 2018, arXiv:1804.00664. https://arxiv.org/abs/1804.00664

Vega-Ferrero, J., Yepes, G., & Gottlöber, S. 2017, MNRAS, 467, 3226, doi: 10.1093/mnras/stx282

Vio, R., Fasano, G., Lazzarin, M., & Lessi, O. 1994, A&A, 289, 640

Weinberg, D. H., Bullock, J. S., Governato, F., Kuzio de Naray, R., & Peter, A. H. G. 2015, Proceedings of the National Academy of Science, 112, 12249, doi: 10.1073/pnas.1308716112

West, M. J. 1994, MNRAS, 268, 79, doi: 10.1093/mnras/268.1.79

West, M. J., de Propris, R., Bremer, M. N., & Phillipps, S. 2017, Nature Astronomy, 1, 0157, doi: 10.1038/s41550-017-0157

Xu, X., Cuesta, A. J., Padmanabhan, N., Eisenstein, D. J., & McBride, C. K. 2013, MNRAS, 431, 2834, doi: 10.1093/mnras/stt379

Zitrin, A., Fabris, A., Merten, J., et al. 2015, ApJ, 801, 44, doi: 10.1088/0004-637X/801/1/44