



UKIRT-2017-BLG-001Lb: A Giant Planet Detected through the Dust

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Abstract

We report the discovery of a giant planet in event UKIRT-2017-BLG-001, detected by the United Kingdom Infrared Telescope (UKIRT) microlensing survey. The mass ratio between the planet and its host is $q = 1.50_{-0.14}^{+0.17} \times 10^{-3}$, about 1.5 times the Jupiter/Sun mass ratio. The event lies $0^{\circ}35'$ from the Galactic center and suffers from high extinction of $A_K = 1.68$. Therefore, it could be detected only by a near-infrared (NIR) survey. The field also suffers from large spatial differential extinction, which makes it difficult to estimate the source properties required to derive the angular Einstein radius. Nevertheless, we find evidence suggesting that the source is located in the far disk. If correct, this would be the first source star of a microlensing event to be identified as belonging to the far disk. We estimate the lens mass and distance using a Bayesian analysis to find that the planet's mass is $1.28_{-0.44}^{+0.37} M_J$, and it orbits a $0.81_{-0.27}^{+0.21} M_{\odot}$ star at an instantaneous projected separation of $4.18_{-0.88}^{+0.96}$ au. The system is at a distance of $6.3_{-2.1}^{+1.6}$ kpc, and so likely resides in the Galactic bulge. In addition, we find a non-standard extinction curve in this field, in agreement with previous results toward high-extinction fields near the Galactic center.

Key words: binaries: general – Galaxy: bulge – gravitational lensing: micro – planetary systems

Supporting material: data behind figure

1. Introduction

Gravitational microlensing is unique in its ability to probe relatively untapped reservoirs of exoplanet parameter space (Gaudi 2012). These include planets at all masses near the “snowline” where gas and ice giants are likely to form (Ida & Lin 2005), hence probing the frequency and mass function of snowline planets (Cassan et al. 2012; Shvartzvald et al. 2016; Suzuki et al. 2016); planetary systems around all types of stars throughout the Galaxy, thus allowing measurement of the Galactic distribution of planets (Calchi Novati et al. 2015a; Penny et al. 2016; Zhu et al. 2017); and free-floating planets (Sumi et al. 2011; Mróz et al. 2017). However, currently only 51 planetary systems hosting 53 planets¹⁰ have been discovered using microlensing. While new ground-based microlensing surveys with global networks of telescopes and high cadences (e.g., KMTNet; Kim et al. 2016) are expected to detect a few tens of planets per year (Henderson et al. 2014), it is clear that a space-based survey is required in order to significantly increase the number of detected systems and to detect planets with masses substantially less than that of the Earth, thus fully exploiting the microlensing potential (Bennett & Rhie 2002; Penny et al. 2013).

The proposed *Wide Field InfraRed Survey Telescope* (WFIRST) flagship mission (Spergel et al. 2015), which is planned to launch in mid-2020s, would dedicate $\sim 25\%$ of its lifetime to a microlensing survey. This survey is predicted to

discover thousands of exoplanets near or beyond the snowline via their microlensing light curve signatures (M. T. Penny et al. 2018, in preparation), enabling a *Kepler*-like statistical analysis of planets ~ 1 – 10 au from their host stars and potentially revolutionizing our understanding of planet formation. In addition to the superb photometry, high cadence, and continuous observations, the survey will be conducted in the near-infrared (NIR). An NIR microlensing survey suffers from less extinction than traditional optical surveys, enabling observations closer to the Galactic plane and center, where the stellar surface density of sources and lenses, and thus the microlensing event rate, is highest. However, until recently no dedicated NIR microlensing survey has been conducted, and so the event rate in the NIR has not been measured, which is crucial for WFIRST field optimization (Yee et al. 2014).

During the 2015 and 2016 seasons, we conducted the first NIR microlensing survey with the United Kingdom Infrared Telescope (UKIRT) as support for the *Spitzer* and *Kepler* microlensing campaigns. From examination of 5% of the 2016 UKIRT fields (overlapping with optical survey fields), we discovered five highly extinguished, low-Galactic latitude microlensing events (Shvartzvald et al. 2017). These events were not detected by optical surveys, likely due to the high extinction. Combining these detections with additional events that were also detected by optical surveys in these fields, we found evidence that the event rate is indeed higher in this region, closer to the Galactic plane. In 2017, we initiated a full NIR microlensing survey with UKIRT, covering all potential WFIRST fields including the Galactic plane and center (see Figure 1), which are inaccessible to optical surveys due to the

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¹⁰ <https://exoplanetarchive.ipac.caltech.edu>

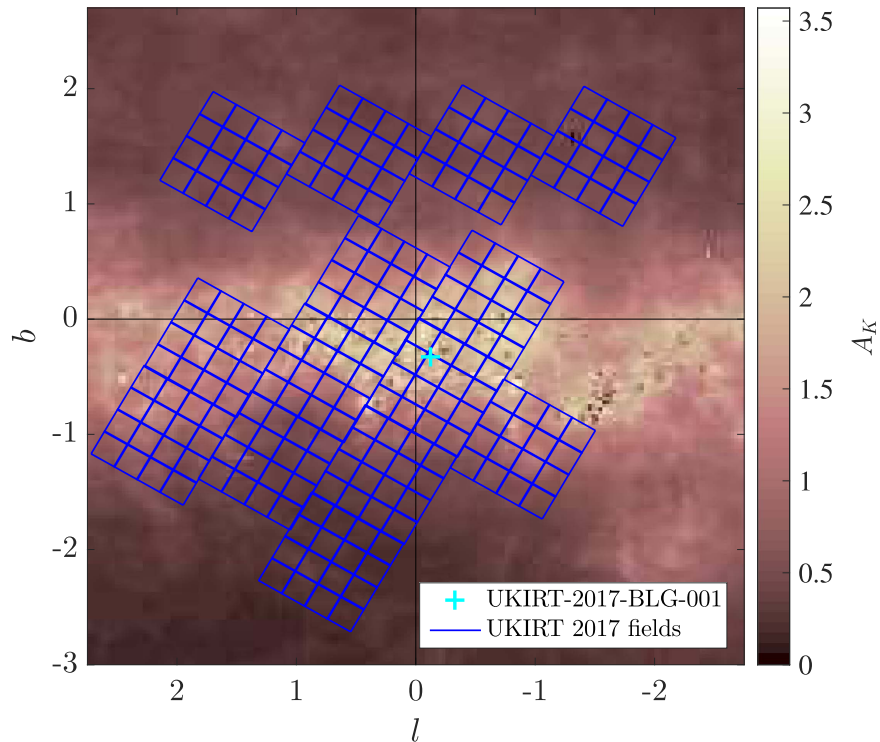


Figure 1. UKIRT-2017 microlensing survey fields (blue), plotted over the A_K extinction map from Gonzalez et al. (2012). The cyan plus marks the location of the event UKIRT-2017-BLG-001.

high extinction. The fields are observed with a daily cadence, allowing us to easily detect microlensing events with their typical timescale of approximately 20 days (Shvartzvald et al. 2016), and with a cadence of three epochs/night in the central fields, which are expected to have an excess of short-timescale bulge–bulge events (Gould 1995).

Bound microlensing planets are discovered through their anomalous signature on the otherwise smooth single-lens light curve. These anomalies occur when the source passes near or over a caustic. There are three types of caustics (for a thorough review see Gaudi 2012): central caustics (associated with the host), planetary caustics (associated with the planet), and resonant caustics, which occur when the central and planetary caustics merge into a single caustic. The location and shape of the caustics are determined by the mass ratio q between the host and the planet, and the scaled (with respect to the Einstein radius θ_E) instantaneous projected separation s . The size of the caustics is mostly determined by the mass ratio q (as the separation s changes by only a factor 2 in the “lensing zone” (Griest & Safizadeh 1998)), and thus for events due to a star and a planet (and thus small q) the caustics are usually small. Anomalies due to planets typically last $\lesssim 1$ day for planetary or central caustics, and thus require high cadence in order for them to be detected and well characterized. However, resonant caustics are always larger than either the planetary or central caustics (for a given mass ratio q) and thus have a larger cross section (Dominik 1999), as well as longer anomaly duration, and therefore can be detected even with a daily cadence. In fact, about 40% of the microlensing planetary systems to date have been detected through resonant caustic perturbations.¹¹ These include the first microlensing planet (Bond et al. 2004), the first

planet simultaneously observed from ground and space (Udalski et al. 2015), and the two multiple-planet systems (Gaudi et al. 2008; Han et al. 2013). All of these events also had long timescales (60–150 days), thus the anomalies lasted several days.

In this Letter, we report the discovery of the planetary event UKIRT-2017-BLG-001. This is the first planet to be detected solely by the UKIRT survey. The planetary perturbation is caused by a resonant caustic of a giant planet ($q \sim 10^{-3}$) and the event has a long timescale of ~ 100 days. Thus, the survey cadence was sufficient for the detection and characterization of the planet. The event lies close to the Galactic center ($0^\circ 35'$) and in a field with significant total and spatially variable extinction, thus it could not be detected by optical microlensing surveys.

We describe the UKIRT observations and event detection in Section 2. In Section 3 we present the best-fit microlensing model of the event. In Section 4 we derive the source properties by analyzing the color–magnitude diagram (CMD) of the event, finding indications that suggest it is a far-disk source. In Section 5 we estimate the physical properties of the planetary system using a Bayesian analysis. An analysis of the multi-band extinction toward this field is presented in Section 6. Finally, we summarize and discuss our results in Section 7.

2. Observations

The UKIRT microlensing survey uses the wide-field NIR camera (WFCAM) at the UKIRT 3.8 m telescope on Mauna Kea, Hawaii. The 2017 fields cover the northern bulge ($b > 0$), the Galactic center, and the southern bulge (see Figure 1). Observations of the central fields are primarily done with the K_S -band, with a nominal cadence of three epochs/night. In addition, H -band observations are taken once every three

¹¹ We also consider events with degenerate solutions when at least one of the solutions is due to a resonant caustic.

nights. The northern and southern bulge fields are observed once per night with the H -band and once every five nights with the K_S -band. Each epoch is composed of 16 5-s co-added dithered exposures (two co-adds, two jitter points, and 2×2 microsteps). The UKIRT dithered images are reduced, astrometrically calibrated, and stacked by the Cambridge Astronomy Survey Unit (CASU; Irwin et al. 2004). The light curves of all sources are then extracted using two methods—(a) Two Micron All Sky Survey (2MASS)-calibrated soft-edge aperture photometry (standard CASU individual image catalogs; Hodgkin et al. 2009), and (b) 2MASS-calibrated point-spread function (PSF) photometry using SExtractor (Bertin & Arnouts 1996) and PSFEx (Bertin 2011). The latter is better for detecting and measuring faint objects in our crowded bulge fields.

We searched the full 2017 data set for microlensing events using the new event detection algorithm of Kim et al. (2018). Among the events found (the full analysis will be presented in a subsequent paper), we identified UKIRT-2017-BLG-001. The event lies inside one of our central fields, at equatorial coordinates (R.A., decl.)_{J2000.0} = (17:46:36.98, -29:12:40.9), and Galactic coordinates (l, b) = (-0.12, -0.33).

After the discovery of the event and the anomaly over its peak, we extended the observations of the field covering the event for additional two weeks beyond the main 2017 campaign. However, these observations were only with one pointing (toward the specific field) unlike the standard observational sequence of a full tile (covering a continuous area with four pointings). This results in a different procedure for deriving the sky frames that are used to correct for additive artifacts such as scattered light, residual reset anomaly, and illumination-dependent detector artifacts (for more details see the UKIRT/WFCAM technical website¹²). The sky estimation for single pointings leaves some residual object flux and thus introduces a small systematic bias that progressively increases at faint magnitudes. In order to incorporate this systematic effect in the PSF photometry light curve, we measured the distribution of flux offset using all of the sources near the event ($<2'$) for each epoch in the additional two weeks compared to the mean flux during the main survey. We then estimate the systematic error per epoch using the flux offset dispersion after subtracting the mean flux offset dispersion of the main survey. Finally, we add (in quadrature) this systematic error to the reported PSF flux errors.

For the modeling of the event (Section 3) we use the PSF photometry of the full data set, while for the CMD that is used to derive the source properties (Section 4), we use the PSF catalog from the main survey only. Based on the rms distribution of the PSF photometry for all of the sources in our field, we add (in quadrature) a minimum error, $e_{\min} = 0.015$, to the pipeline reported errors, in order to compensate for unrealistically small Poisson flux errors when the event is bright. We note that this is an empirical correction and is not driven by the fit for the model.

3. Light Curve Modeling

The light curve (Figure 2) has a clear and long anomaly (~ 8 days) just after the peak of a moderate-magnification microlensing event. The anomaly starts with a small deviating

rise above the expected single-lens model, followed by a sharp fall. Then the event rises again and continues to fall approximately as expected by a single-lens microlensing event. The combination of these features suggests a source crossing over a planetary resonant caustic, exiting at the “back” side of the caustic opposite to the planet (and thus the sharp fall below the single-lens model), followed by a cusp approach (see Figure 3).

We model the event using seven geometric parameters to calculate the magnification, $A(t)$, of a binary-lens system. These are the three point-lens parameters (t_0, u_0, t_E) (Paczynski 1986), and three parameters for the companion: the mass ratio q , the instantaneous scaled projected separation s , and the angle α (measured counter-clockwise) between the source trajectory and the binary axis in the lens plane. The seventh parameter is the scaled angular source size $\rho = \theta_*/\theta_E$. To calculate the model magnifications near and during the anomaly we employ contour integration (Gould & Gauchere 1997) with 10 annuli to allow for limb darkening. We adopt linear coefficients $u_H = 0.3895$ and $u_K = 0.3324$ (Claret & Bloemen 2011), based on the source type derived in Section 4.¹³ Far from the anomaly we employ limb-darkened multipole approximations (Gould 2008; Pejcha & Heyrovský 2009). Finally, each data set has two flux parameters representing the source ($f_{s,H/K_S}$) and any additional blend ($f_{b,H/K_S}$): $F_{H/K_S}(t) = f_{s,H/K_S} A(t) + f_{b,H/K_S}$.

To identify initial possible solutions we search over a grid of mass ratios ranging $q = 0.0001$ – 1.0 and separations $s = 0.1$ – 1.4 , fully containing the relevant parameter space of resonant caustics. For each grid point we initiate a set of Markov-chain Monte-Carlo (MCMC) chains with all possible angles α . We find an isolated single minimum centered on $(s, q) \sim (1.02, 0.001)$ (see Figure 4) and initiate from this point a full MCMC analysis (with all parameters free) to find the best-fit model.

The source of the event is faint, with $K_{S,s} \gtrsim 16$, which is very close to the survey limiting magnitude (see Figure 6). Moreover, our data do not completely cover the baseline of the event. Yee et al. (2012) showed that some of the standard microlensing parameters, in particular t_E and f_s , can be poorly constrained for such events with faint sources. However, they introduced the microlensing invariants ($t_0, t_{\text{eff}} \equiv u_0 t_E, t_* \equiv \rho t_E, q t_E, f_{\text{lim}} \equiv f_s/u_0$), which are well-constrained from the region near the peak and anomaly of the event. Table 1 gives the values and uncertainties of the best-fit parameters and the microlensing invariants.

We further try to include in the model microlensing parallax or orbital motion of the binary-lens system. The inclusion of these higher-order effects does not significantly improve the fit (which requires two additional parameters for each effect) with $\Delta\chi^2 = 5$ when including only microlens parallax, $\Delta\chi^2 = 9$ when including only orbital motion, and $\Delta\chi^2 = 14$ when including both effects. We do not consider either of these as a detection because they all require significant negative blending (which is unphysical) and systematics at that level are well known in microlensing. Nevertheless, we can set a conservative upper limit ($\Delta(\chi^2 - \chi_{\text{best}}^2) < 30$) on the microlensing parallax of $\pi_E < 0.7$. In Section 5 we use this limit when deriving the physical properties of the system.

¹² <http://casu.ast.cam.ac.uk/surveys-projects/wfcam/technical/sky-subtraction>

¹³ We note that the best fit is insensitive to the exact limb darkening coefficient for giant stars.

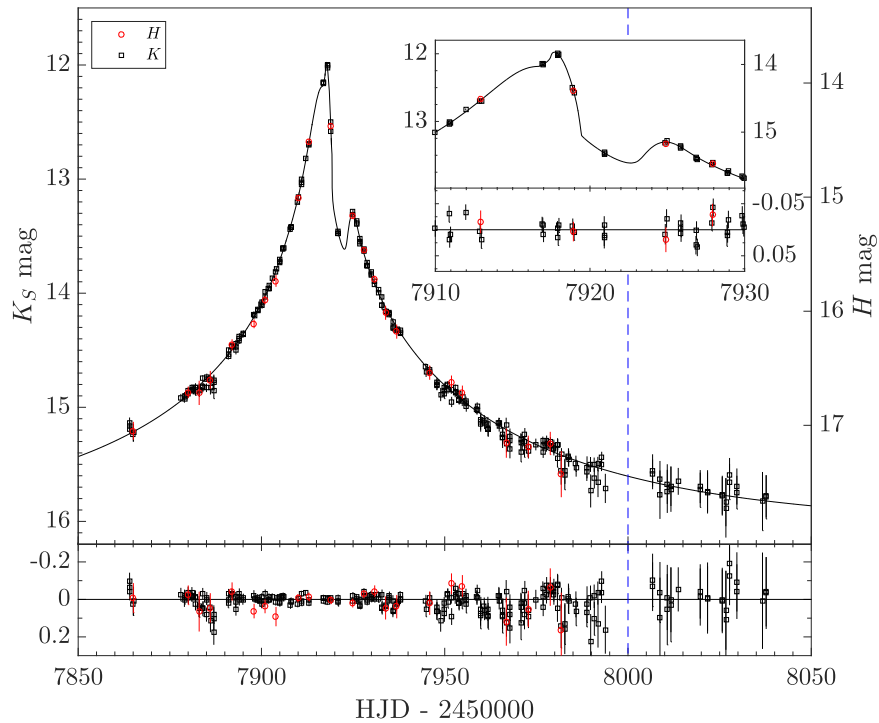


Figure 2. Light curve of UKIRT-2017-BLG-001 in K_S (black squares and left horizontal axis) and H (red circles and right horizontal axis). The best-fit planetary model is shown in black. The anomaly (inset) over the peak is clear and covered sufficiently well with our three epochs/night cadence. The vertical blue dashed line indicates the end of the main survey, beyond which we include systematic errors due to poor sky estimations. The data used to create this figure are available.

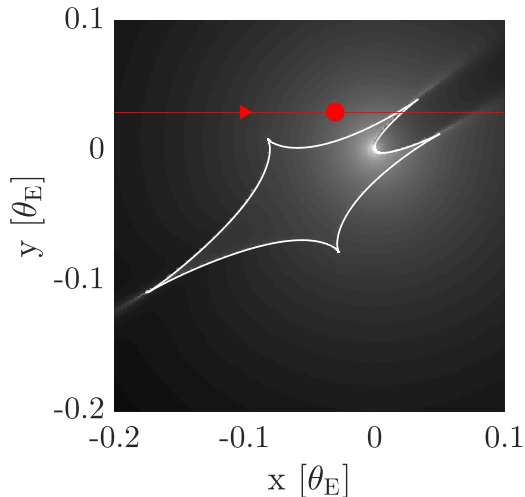


Figure 3. Magnification map of the event. The source trajectory (red line) over the resonant caustic (white curve). The red circle indicates the source size. The length of the caustic (long axis) is larger than $0.2\theta_E$, showing the large cross section of giant planet resonant caustics. The width is $\sim 0.1\theta_E$ and the event timescale is $t_E \approx 100$ days, suggesting a typical anomaly duration of ~ 10 days.

4. Source Properties

A standard way to derive the intrinsic source properties is by measuring the offset between the source position on a CMD and the centroid of the red clump (RC; Yoo et al. 2004). A fundamental assumption for this method is that the source and the RC are both behind the same dust column. The CMD is thus usually constructed using stars near ($\lesssim 2'$) the target, as spatial extinction variations toward the bulge are usually on larger scales. However, in the case of our event there are several challenges with this method.

The first challenge is the observed source properties. The entire range of $(H - K_S)$ colors for giants (which our source likely is, as we show below) is narrow, with only 0.24 mag difference between G0 to M7 giants. Even for dwarfs the range is only 0.38 mag between B8 and M6 dwarfs (Bessell & Brett 1988).¹⁴ This is because $(H - K_S)$ basically probes the Rayleigh–Jeans tail for most stars. Therefore, a precise measurement of the source color is essential in order to derive the source properties. We determine the model-independent source $(H - K_S)_s$ color from regression of H versus K_S flux as the source magnification changes (Gould et al. 2010), and find $(H - K_S)_s = 1.87 \pm 0.01$. (This color is in excellent agreement with the color derived from the model of $(H - K_S)_s = 1.878 \pm 0.004$). Each H epoch was taken in between two K_S observations on the same night, minimizing systematics and allowing us to achieve the required precision. Next, the source is faint and, as discussed in Section 3, we do not have coverage of the event’s baseline. Therefore, the uncertainty on the source magnitude is relatively large. The K_S source magnitude as inferred from the microlensing model is $K_{S,s} = 16.07^{+0.09}_{-0.11}$. The model also indicates a fainter blend flux, with 5σ upper limit of $K_{S,b} > 17.7$. We use this measurement later as an upper limit on the lens flux when we estimate the physical properties of the planetary system (see Section 5). For completeness, we note that no centroid shift as a function of magnification was detected.

The second and more prominent challenge is the determination of the extinction and reddening using the RC centroid. There are clear dust clouds in the region around our event (see Figure 5), which cause large spatial reddening variations on

¹⁴ We convert Bessell & Brett (1988) $(H - K)_{\text{BB}}$ colors to 2MASS-calibrated $(H - K_S)_{\text{2MASS}}$ colors using the relations from Carpenter (2001).

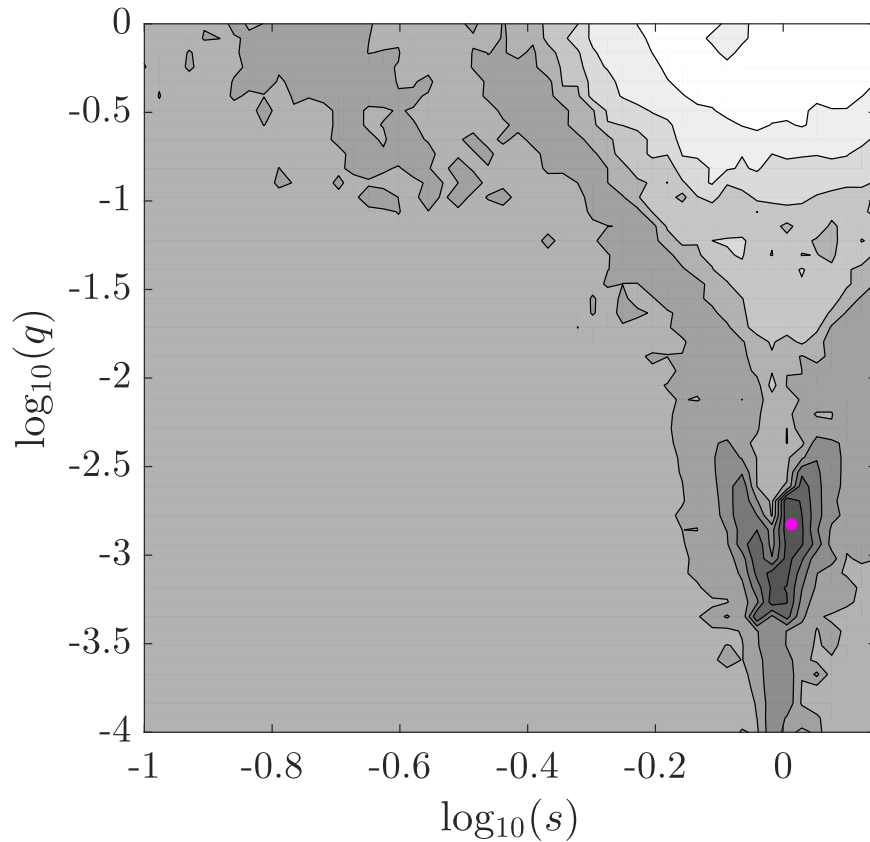


Figure 4. Results from the grid search on the s - q (scaled separation, mass ratio) plane. The values of the best-fit solution are marked as a magenta circle. A single isolated minimum is clearly detected. The contours indicate a steep surface, where all results with $\Delta\chi^2 < 100$ compared to the best fit are within the inner contour and results with $\Delta\chi^2 < 500$ and $\Delta\chi^2 < 1000$ are within the second and third contours, respectively.

Table 1
Microlensing Model

Parameter	
t_0 [HJD']	7916.243 ± 0.022
u_0	$0.0303^{+0.0036}_{-0.0026}$
t_E [day]	$101.0^{+8.2}_{-9.6}$
ρ [10^{-3}]	$6.64^{+0.75}_{-0.56}$
α [rad]	$3.7036^{+0.0084}_{-0.0076}$
s	$1.0318^{+0.0032}_{-0.0045}$
q [10^{-3}]	$1.50^{+0.17}_{-0.14}$
$K_{S,s}$	$16.07^{+0.09}_{-0.11}$
$K_{S,b}$	$>17.7^a$
t_{eff} [day]	3.062 ± 0.050
t_* [day]	0.671 ± 0.016
qt_E [day]	0.1517 ± 0.0041
f_{lim}	198.2 ± 2.0

Note. HJD' = HJD-2450000.

^a 5σ limit.

scales smaller than the standard $\lesssim 2'$ region. We avoid obvious dust stripes and use stars in a 1.6 arcmin^2 box around the event, as marked in Figure 5, to construct the CMD (Figure 6). We measure the RC centroid following the procedure described in Nataf et al. (2013) and find $(H - K_S, K_S)_{\text{cl}} = (1.63, 14.63)$. The number of RC stars used, as derived from the fit, is $N_{\text{RC}} = 197 \pm 14$, which is sufficient for a reliable measurement of the RC. By comparing the apparent clump centroid to

the intrinsic centroid of $(H - K_S, K_S)_{\text{cl},0} = (0.15, 12.95)$ toward the event's location (assuming a distance of 8.17 kpc; Nataf et al. 2013, 2016; Hawkins et al. 2017) we find that the mean extinction and reddening toward the event are $\langle A_{K_S} \rangle = 1.68$ and $\langle E_{H-K_S} \rangle = 1.48$, respectively. We note that the reddening-to-extinction ratio that we find is non-standard, and we investigate this further in Section 6.

However, both the color and magnitude dispersions of the RC stars are large, suggesting large extinction and reddening dispersions (and thus large uncertainties on the intrinsic source properties) as well as a possibly wide distance dispersion of clump stars. The observed clump color dispersion is $\sigma_{(H-K_S)\text{cl}} = 0.16$ (as derived from the clump centroid fit; see Nataf et al. 2013 for details), which combines the intrinsic clump color dispersion and reddening dispersion. In order to estimate the intrinsic dispersion we apply the RC centroid procedure to several of our low-Galactic-latitude UKIRT fields that have low and uniform extinction. From the clump color dispersion on these fields we find that the intrinsic clump color dispersion is only 0.04, and thus the uncertainty on the reddening (e.g., the reddening dispersion) in the event's field is $\sigma_{E_{H-K_S}} = 0.15$. This large uncertainty, if taken as is, implies that the intrinsic color of the source cannot be well constrained, as the full range of colors for giant stars is within $2\sigma_{E_{H-K_S}}$ and the full dwarf color range is within $2.5\sigma_{E_{H-K_S}}$. The observed clump magnitude dispersion, $\sigma_{K_{S,\text{cl}}} = 0.36$, is even higher. Intrinsic dispersion, extinction dispersion, and distance modulus dispersion all contribute to the total magnitude dispersion (the dispersion due to photometric noise is negligible). The

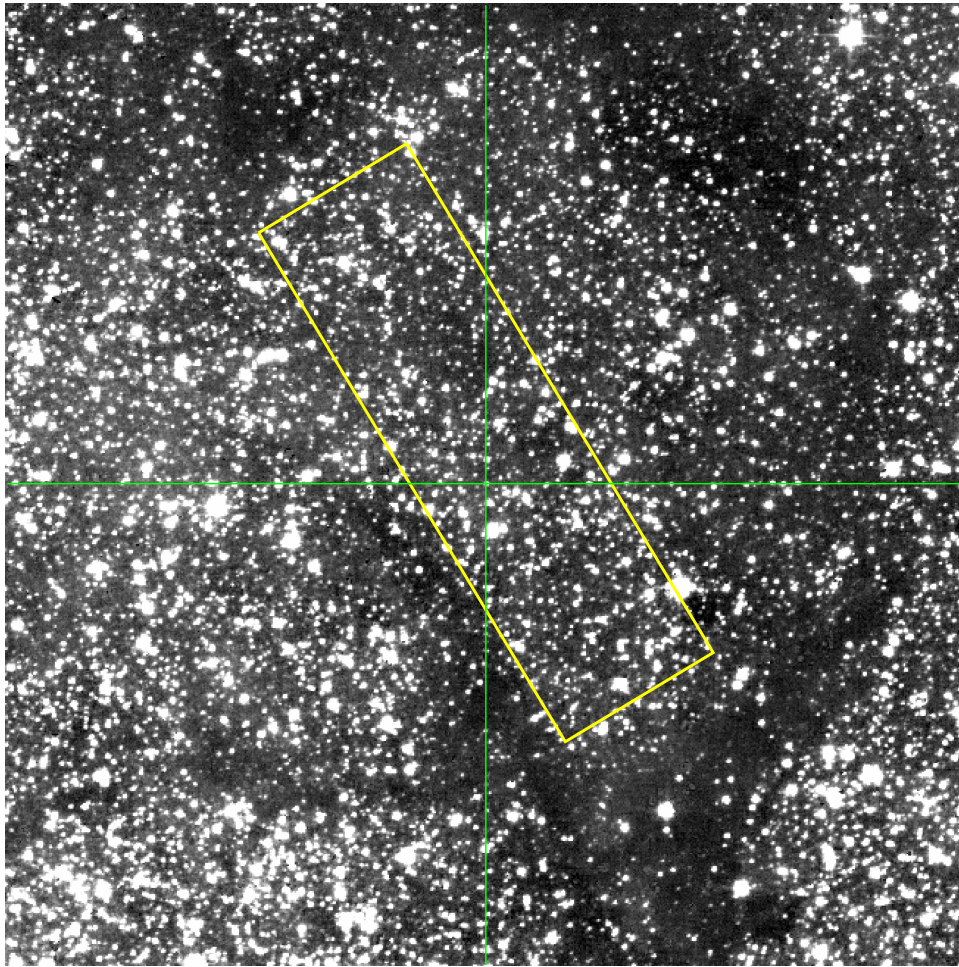


Figure 5. UKIRT $3/8 \times 3/8$ H -band image of the field around the event (marked with a cross-hair). Dust “stripes” are clearly seen around the target, indicating the high differential reddening in the field. The yellow box marks the region used to construct the CMD.

intrinsic magnitude dispersion is 0.17 (Hawkins et al. 2017), and the extinction dispersion can be estimated from the reddening dispersion and the mean extinction-to-reddening ratio to be $\sigma_{A_{K_S}} = 0.17$. Thus, the distance modulus dispersion to the clump is 0.28. While this value is larger than typical toward Baade’s window (Nataf et al. 2013), the event is located very close to the Galactic center (and thus the Galactic plane), implying a large geometrical dispersion due to the fact that there are significant contributions from both bulge and far disk stars (see more details in Section 4.1 below).

The large differential reddening implies that the basic assumption that the source is behind the same dust as the clump is not reliable, as the extinction itself is not uniform across the field. Moreover, assuming that the reddening and extinction to the source are the mean reddening and extinction in the field, $\langle E_{H-K_S} \rangle = 1.48$ and $\langle A_{K_S} \rangle = 1.68$ yield intrinsic source properties of $[(H - K_S), K_S]_{s,“0”} = (0.39, 14.39)$. While the magnitude suggests a giant source, the color is 0.05 redder than an M7 giant (which is already very rare, as suggested by the small intrinsic dispersion of clump star colors). This implies that the source is suffering from higher reddening than the mean clump. While we cannot use the usual method to derive the source properties with small uncertainties, we can use it to set possible limits on the source properties. We assume that the extinction and reddening toward the source are $\langle A_{K_S} \rangle + \Delta A_{K_S}$ and $\langle E_{H-K_S} \rangle + \Delta E_{H-K_S}$, respectively. In order

to estimate the boundaries of the possible source angular size range, we take two extreme cases of giant sources. As the reddest limit, we assume an M7 giant with $(H - K_S) = 0.34$, thus suffering from additional reddening $\Delta E_{H-K_S} = 0.05$ and correspondingly additional extinction $\Delta A_{K_S} = 0.06$. As the bluest limit, we assume a G0 giant with $(H - K_S) = 0.10$, thus suffering from $\Delta E_{H-K_S} = 0.29$ and $\Delta A_{K_S} = 0.33$. For each case, we derive the source intrinsic magnitude by accounting for the total extinction. We then use the $(V - K)$ colors (Bessell & Brett 1988) corresponding to the source spectral type and the surface brightness to angular source size relation of Kervella et al. (2004) to infer that the source angular size is in the range $3.3 < \theta_*/\mu\text{as} < 8.5$.

4.1. Far Disk Source?

The large differential reddening and the “too red” source color imply a non-uniform extinction toward each individual star, and in particular the source. These can be due to either spatial differential reddening on the scale of a few arcseconds (even at the apparently uniform box selected), a different accumulated dust column along the line of sight toward sources at different distances, or a combination of both. We first investigate the possibility of variations of dust columns due to source distances, which we find suitable to fully explain both problems. Nevertheless, in Section 4.2 we try to examine the possible spatial differential reddening.

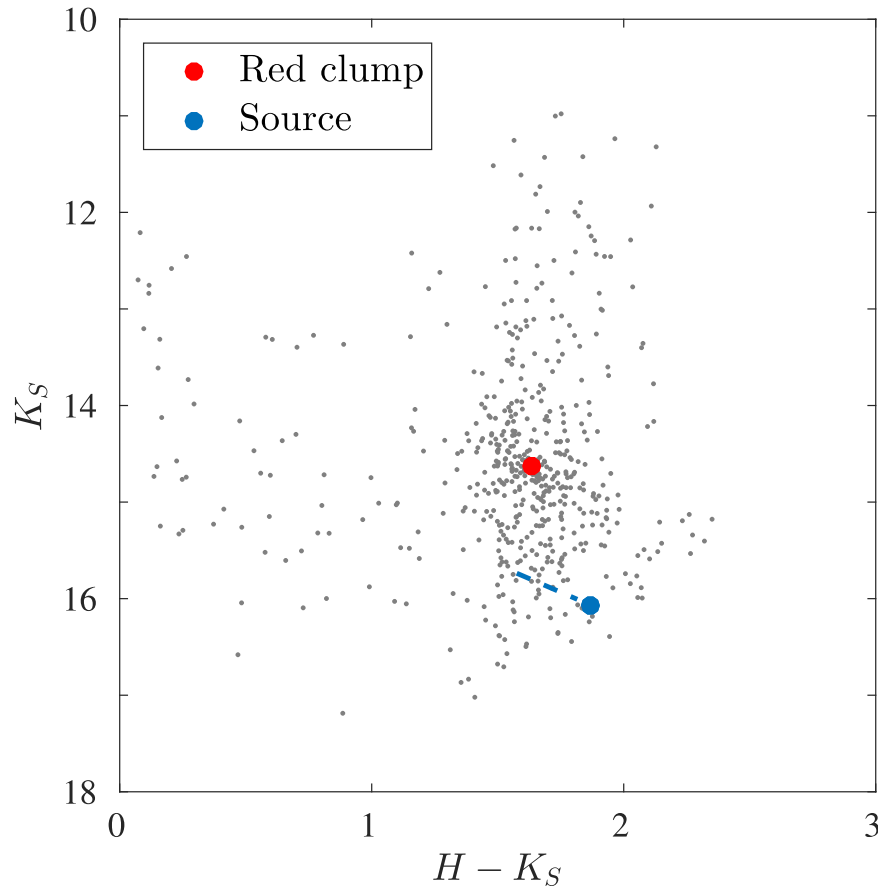


Figure 6. CMD of stars around the event. The red circle indicates the centroid of the RC, with the relatively large dispersion around it indicating the high differential reddening. The large color offset between the source (blue) and the clump suggests that the source suffers from additional reddening. The blue dashed line indicates the source position as if it were behind the same dust column as the clump for the range of possible giant colors.

The event is located at Galactic latitude $b = -0^{\circ}.33$, corresponding to ~ 50 pc below the Galactic plane at the distance of the Galactic bulge (~ 8 kpc), or 115 pc at the extreme far disk (20 kpc). Therefore, the Galactic thin disk population can have a significant contribution to the number of observed RC stars on the CMD, because the scale height of the Galactic thin disk is 300 pc (Jurić et al. 2008). This can partially explain the observed scatter. In addition, the dust scale height is 120 pc (Jones et al. 2011). First, this can explain the large differential reddening within the bulge. Second, stars in the far disk will be significantly more extinguished. The Marshall et al. (2006) 3D Galactic interstellar extinction model, which combines 2MASS data (Cutri et al. 2003) and the Besançon Galactic model (Robin et al. 2003), suggests an additional extinction toward the event of $\Delta A_{K_S} \approx 0.4$ behind the bulge (between 8.5 and 12 kpc where they are limited by the data). The Green et al. (2015) 3D Galactic dust map, which combines the 2MASS and Pan-STARRS1 (Kaiser et al. 2010) data sets, suggests a $\sim 15\%$ increase in reddening (corresponding to $\Delta E_{H-K_S} \approx 0.25$) between 8 and 16 kpc. However, both of these maps (Marshall et al. 2006; Green et al. 2015) run out of stars before or around the distance to the Galactic center, and thus they probably underestimate the true amount of additional extinction on the far side of the Galaxy. In conclusion, the combined effects of the significant thin disk population and the accumulated dust result in the large scatter of the apparent magnitude of the RC stars in the field.

The source distance and position on the CMD can similarly be well explained by the above arguments. First, we estimate the general probability of a far-disk source with $15 < K_{S,s} < 17$ toward the events coordinates using a new Galactic population synthesis model (M. J. Huston et al. 2018, in preparation)¹⁵ and Monte-Carlo integration of the event rate (Penny et al. 2013; Awiphan et al. 2016). We find that the contribution of far-disk sources, at distances $D_s > 11$ kpc, is 50% of the overall general source probability. Next, as we found above, the source is $\Delta(H - K_S) = 0.25$ redder than the bulge mean clump. If, for example, the source is a typical RC star at the far disk, then this difference is due to additional reddening $\Delta E_{H-K_S} = 0.25$. Using the mean extinction-to-reddening ratio derived by the mean clump position, this corresponds to $\Delta A_{K_S} \approx 0.28$. The remaining difference between the source magnitude and the mean clump from the change in distance modulus is thus 1.16. This implies a source that is 1.71 times more distant than the mean clump, corresponding to $D_s \approx 14$ kpc. As we showed above, Galactic dust models and observations suggest that, at such distance, the source is likely to suffer from the observed additional

¹⁵ The model assumes a bulge with an E3 density distribution from Dwek et al. (1995), with parameters estimated by Cao et al. (2013) using OGLE-III clump giants. The bar has an angle of 29° to the line of sight, and kinematics that match the Bulge Radial Velocity Assay (BRAVA) survey (Howard et al. 2009). The disk density and age distribution follows that of Robin et al. (2003). Stellar properties are derived from MIST version 1.1 solar metallicity isochrones (Choi et al. 2016; Dotter 2016).

extinction and reddening. The distance could be somehow less than 14 kpc, if the source is intrinsically fainter than the clump or if the dust on the other side of the Galaxy has an extinction curve that is closer to the standard law than the dust between the Sun and the RC. Later, in our Bayesian analysis (Section 5) we find that the source is most likely in the far disk ($D_S = 11.2^{+3.6}_{-2.6}$ kpc), and thus we conclude that this is probably the correct explanation.

4.2. Spatial Differential Reddening

The field around the event is severely impacted by dust stripes. While we constructed our CMD by selecting a box without any obvious dust clouds, it is still possible that the large differential reddening is due to local spatial variations on the scale of a few arcminutes. We tried to examine this possibility by extracting multi-color photometry of the stars around the event. By combining extinction ratios (either by using the RC stars for each band or through color-color relations) and the colors of stars, we can estimate the extinction toward individual stars. This will allow, in principle, the derivation of a spatial reddening map.

We obtained *grizY* bands catalog of stars near the event using data from the DECam Plane Survey (Schlafly et al. 2018), with corresponding AB limiting magnitudes of 23.9, 23.0, 22.6, 21.9, and 21.2. In addition, we extracted a *Spitzer* 3.6 μm catalog of stars near the event from an image taken as part of the Galactic Legacy Infrared Mid-Plane Survey Extraordinaire (GLIMPSE) Proper project (Benjamin et al. 2015) using the new Infrared Array Camera (IRAC) crowded-field photometry algorithm (Calchi Novati et al. 2015b), and calibrating it to Ramírez et al. (2008). The limiting magnitude is 15.4. Finally, we extracted VISTA Variables in the Vía Láctea (VVV; Minniti et al. 2010) *J*-band catalog,¹⁶ with a limiting magnitude of 19.4. Unfortunately, most of the stars in our field are not detected in the DECam *grizY* and VVV *J* catalogs (see Figure 8). In particular, the bulge mean RC is not fully covered in all of these bands. Therefore, we cannot derive the spatial reddening map using the method suggested above, and thus further observations are required in order to explore this viable explanation, as we discuss in Section 7. Nevertheless, we use the limiting magnitudes of these catalogs to study the extinction curve toward this field in Section 6.

5. Planetary System Physical Properties

The angular Einstein radius and the relative geocentric proper motion between the source and the lens can be derived from the light curve model and CMD analysis, $\theta_E = \theta_*/\rho$ and $\mu_{\text{geo}} = \theta_E/t_E$. Applying the limits derived above on θ_* then gives

$$0.44 < \theta_E [\text{mas}] < 1.39 \quad 1.5 < \mu_{\text{geo}} [\text{mas yr}^{-1}] < 5.5. \quad (1)$$

Unfortunately, even using these limits we cannot determine directly the physical properties of the planetary system, as the microlens parallax was not detected. We therefore estimate them using a Bayesian analysis that incorporates the limits on θ_E and μ_{geo} into a Galactic model. We follow the procedures described in Shvartzvald et al. (2014) and adopt the Galactic model of Han & Gould (1995, 2003), which reproduces well the

observed statistical distribution of properties of microlensing events. For the source probability, we apply conservative color and magnitude limits of $15 < K_{S,s} < 17$ and $(H - K_S)_s > 1.3$. We furthermore set two limits on the mass–distance relation of the lensing system. The first is from the microlens parallax

$$\pi_E \equiv \sqrt{\kappa M \pi_{\text{rel}}} < 0.7 \quad \kappa \equiv \frac{4G}{c^2 \text{au}} \simeq 8.14 \frac{\text{mas}}{M_\odot}. \quad (2)$$

Here $\pi_{\text{rel}} = \text{au}(D_L^{-1} - D_S^{-1})$ is the lens–source relative parallax. The second limit is an upper limit on the lens flux by using the 5σ upper limit on the blend flux that was derived from the model, $K_{S,l} \geq K_{S,b} > 17.7$. We use 5 Gyr Padova isochrones (Bressan et al. 2012; Marigo et al. 2017) and assume (as a conservative limit) that the lens is behind the overall dust column toward the bulge to convert the flux limit to a mass–distance limit. This limit alone already gives an upper limit on the planet mass (if all the blend flux is attributed to the lens), because at this age the maximal host mass is $\lesssim 1.25 M_\odot$, and thus the planet mass will be $< 2.7 M_J$. We note that if the host is a bulge star, with age of ≈ 10 Gyr, the limits on the host and planet masses are $\lesssim 1.05 M_\odot$ and $< 2.2 M_J$, respectively. However, as the event is within one scale height of the thin disk throughout the Galaxy, we use the 5 Gyr isochrones to derive the limit on the lens flux for all distances, as there is a non-negligible probability that the lens might be a part of the disk population even at 6–10 kpc.

The 2D posterior distribution of the host mass and distance is shown in Figure 7, as well as the limits on θ_E , π_E , and the lens flux. We infer that the host is a $0.81^{+0.21}_{-0.27} M_\odot$ dwarf, likely in the Galactic bulge at $6.3^{+1.6}_{-2.1}$ kpc. Using the mass ratio and scaled instantaneous projected separation from the model, we find that the planet mass is $1.28^{+0.37}_{-0.44} M_J$, and it orbits its host beyond the snowline at a projected separation of $4.18^{+0.96}_{-0.88}$ au. The estimated physical properties are summarized in Table 2.

One of the main uncertainties in our Galactic model is the source distance. However, we note the the posterior probability for the lens mass (and thus the planet mass) is weakly dependent on the source distance. For a source at 8 kpc the mass is only 10% lower than for a source at 15 kpc, well within the range of our uncertainty. The lens distance, and consequently the projected separation between the planet and its host, is more sensitive to the source distance, with a difference of 45% between a bulge source and a far disk source. Future observations can resolve this and give a better estimation of the source distance, as we discuss below in Section 7.

6. Extinction Law

Recent studies of dust properties in the inner bulge suggest deviations from the standard extinction law (Nataf et al. 2016; Alonso-García et al. 2017). The extinction coefficient measured in our field, $\langle A_{K_S} \rangle / \langle E_{H-K_S} \rangle = 1.14$, is lower than the value of 1.37 from Nishiyama et al. (2009). This offset is a disconcertingly large ~ 0.4 mag in K_S . The offset relative to the value of $A_{K_S} / E_{H-K_S} \sim 2.07$ predicted by Fitzpatrick (1999) is even larger—about 1.4 mag in K_S .

The magnitude of the RC can also be predicted in the *YJ* bands by assuming the color–color relations from Nataf et al. (2016) and the extinction coefficients from Fitzpatrick (1999).

¹⁶ We note that our UKIRT data are deeper than the *H* and K_S VVV catalogs.

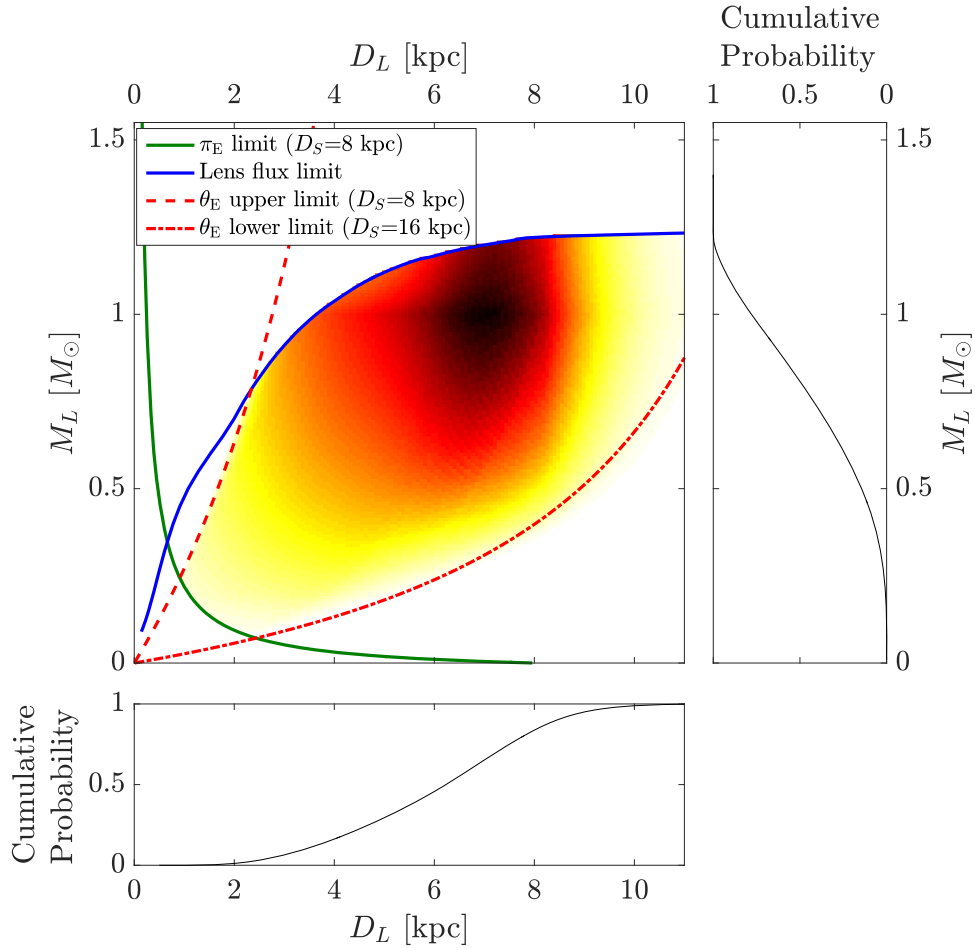


Figure 7. Posterior probability distributions of the host mass and distance derived from the Bayesian analysis. Also shown are the limits on the lens flux (blue), on θ_E for either bulge or far disk sources (red), and on the microlens parallax for a bulge source (green).

Table 2
Physical Properties

$M_1 [M_\odot]$	$0.81^{+0.21}_{-0.27}$
$m_2 [M_J]$	$1.28^{+0.37}_{-0.44}$
r_\perp [au]	$4.18^{+0.96}_{-0.88}$
D_L [kpc]	$6.3^{+1.6}_{-2.1}$
D_S [kpc]	$11.2^{+3.6}_{-2.6}$
μ_{geo} [mas yr $^{-1}$]	$2.34^{+0.84}_{-0.42}$

These are $Y_{RC} \approx 14.63 + 1.63 + (7.35-3.12) \approx 20.49$ and $J_{RC} \approx 14.63 + 1.63 + (5.44-3.12) \approx 18.58$, where we add the corresponding intrinsic color offsets and reddening offsets, respectively, to the observed clump H -band magnitude.

Given that the limiting magnitudes for the DECam and VVV data sets are, respectively, $Y \sim 21.2$ and $J \sim 19.4$ (see Section 4.2 above), the RC should be clearly detected in each of these bandpasses by a wide margin of 0.7–0.8 mag, respectively. However, the photometry in both cases barely covers the RC. Figure 8 shows the HK_s CMD of our field (gray points), with the sources detected in Y shown as yellow circles in the left panel, and the points detected in J shown as purple in the right panel. We conclude that the extinction toward this field is steeper than standard (as defined by Fitzpatrick 1999).

7. Discussion

We have presented the discovery of a roughly Jupiter/Sun mass ratio planet. The system likely resembles a version of the Jupiter/Sun system, with the host being a G/K dwarf slightly less massive than the Sun, located in the Galactic bulge. The event was detected as part of our UKIRT microlensing survey, which has the primary goal of deriving the NIR event rate toward the *WFIRST* target field region. As such, the survey is designed to have a \sim daily cadence, which in principle is not optimal for the detection of short planetary anomalies. However, the perturbation was due to a resonant planetary caustic, and the event had a long timescale, so the cadence was sufficient to detect the planetary anomaly. This event contributes to the set of interesting planetary microlensing events that were discovered through resonant caustic perturbations. The resonant caustic parameter space is relatively wide for giant planets. For $q = 10^{-3}$, resonant caustics exist for instantaneous projected separations of $0.93 < s < 1.15$, corresponding to 2.6–3.3 au for a typical Einstein radius. However, the range in projected separations for which resonant caustics exist scales as $q^{1/3}$, and thus resonant caustics are less important for lower mass ratios q .

The field around the event suffers from high and differential extinction that creates a challenge for deriving the intrinsic source properties, and thus the physical properties of the

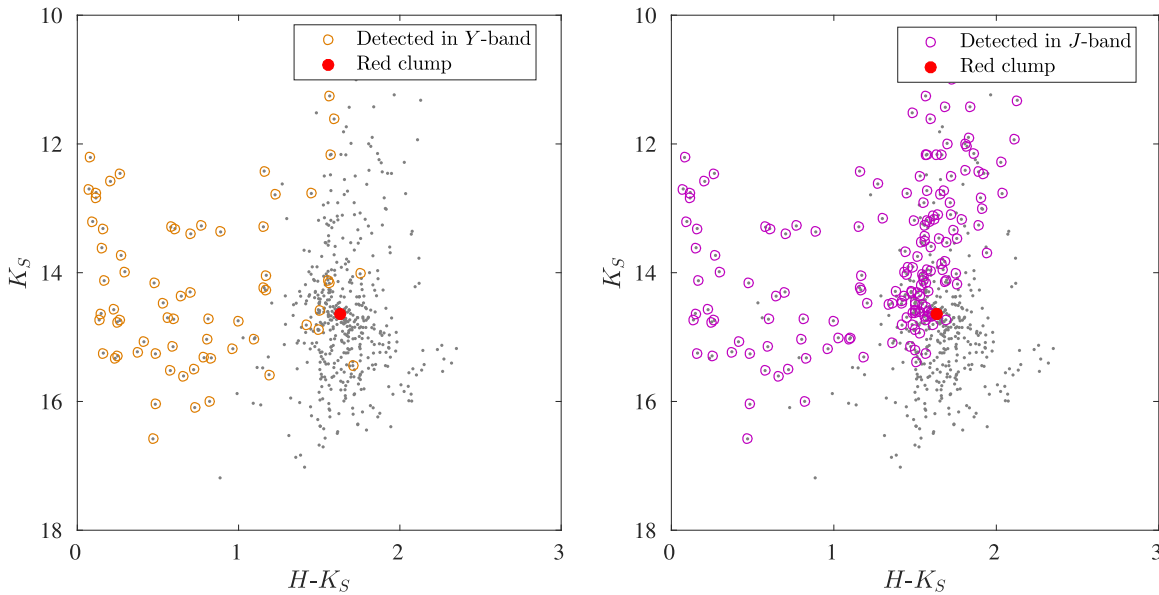


Figure 8. Detection limits of DECam Y -band data (left) and VVV J -band data (right). These limits are brighter than expected using the extinction and reddening of the UKIRT HK_S data (underlying gray points) and standard extinction curve, suggesting a non-standard and steeper extinction law.

planetary system. Follow-up observations of UKIRT-2017-BLG-001Lb can reduce these uncertainties in several ways. First, deep high-angular resolution imaging of the field in the infrared would enable an accurate measurement of the extinction toward each star (including the source) using the Rayleigh–Jeans color excess method (Majewski et al. 2011). This can be done from the ground with Keck in the NIR or, preferably, with *JWST* in mid- and near-infrared bands allowing for a wider spectral range. Second, *JWST* could also measure a medium-resolution spectrum of the source star in the NIR to estimate its temperature and improve the estimate of the angular Einstein radius (the upgraded Near Infrared Spectrograph (NIRSPEC) on Keck will potentially have the sensitivity for a low-resolution spectrum of the source in H -band). Third, the high-resolution image will resolve out possible unrelated sub-arcsecond blend stars around the target (see, e.g., Beaulieu et al. 2018), allowing the measurement of any excess flux from the target above the measured source flux. However, the blend flux from the microlensing model ($K_{S,b} > 17.7$) already sets an upper limit on the excess flux. For such faint excess the probability for a significant contribution from stars that are not the lens (i.e., companion to lens, companion to source, or ambient star) is high (see, e.g., Koshimoto et al. 2017). Finally, a second high-resolution epoch of UKIRT-2017-BLG-001Lb could resolve the lens and the source and directly measure the relative lens-source proper motion and the lens flux (e.g., Batista et al. 2015). Given that the relative proper motion is in the range $1.5 < \mu[\text{mas yr}^{-1}] < 5.5$ (with Bayesian estimate of $\mu_{\text{geo}} \sim 2.5 \text{ mas yr}^{-1}$) the source and the lens will be sufficiently separated in 10–50 years. However, centroid shifts due to the relative lens-source proper motion (e.g., Bennett et al. 2015; Bhattacharya et al. 2017) could be detected earlier.

Future NIR microlensing surveys, such as the planned *WFIRST* microlensing program or the one proposed with *Euclid* (Penny et al. 2013), may consider avoiding such regions with large spatial extinction variations. On the other hand, our current results (Shvartzvald et al. 2017) suggest that the NIR event rate is highest close to the Galactic center, and thus the

field selection should balance between the high differential extinction and the high event rate. We note that the currently planned *WFIRST* fields (M. T. Penny et al. 2018, in preparation) lie at lower Galactic latitudes ($b \approx -2.05$ to -0.45), in regions of lower extinction and lower differential extinction than UKIRT-2017-BLG-001Lb.¹⁷


The multi-band analysis of our field suggests a non-standard (steeper) extinction law. This supports previous suggestions of interstellar extinction law variations toward the inner bulge (Nataf et al. 2016; Alonso-García et al. 2017). Our ongoing UKIRT survey will enable the creation of extinction and reddening maps of the Galactic center and bulge. These, combined with deep optical surveys, will allow us to confirm and fully constrain these variations. These results can also have important implications for observational cosmology, as they suggest that, if non-standard extinction laws occur in external galaxies, they may lead to systematic errors in the distances derived from studies of SNe Ia and Cepheids, if those studies adopt standard extinction laws.

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¹⁷ Based on the extinctions of RC stars measured by Nidever et al. (2012), the current *WFIRST* fields have quartile K_S extinctions of 0.38 and 0.64.

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