

# *NuSTAR* catches the unveiling nucleus of NGC 1068

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## ABSTRACT

We present a *NuSTAR* and *XMM-Newton* monitoring campaign in 2014/2015 of the Compton-thick Seyfert 2 galaxy, NGC 1068. During the 2014 August observation, we detect with *NuSTAR* a flux excess above 20 keV ( $32 \pm 6$  per cent) with respect to the 2012 December observation and to a later observation performed in 2015 February. We do not detect any spectral variation below 10 keV in the *XMM-Newton* data. The transient excess can be explained by a temporary decrease of the column density of the obscuring material along the line of sight (from  $N_{\text{H}} \simeq 10^{25} \text{ cm}^{-2}$  to  $N_{\text{H}} = 6.7 \pm 1.0 \times 10^{24} \text{ cm}^{-2}$ ), which allows us for the first time to unveil the direct nuclear radiation of the buried active galactic nucleus in NGC 1068 and to infer an intrinsic 2–10 keV luminosity  $L_{\text{X}} = 7_{-4}^{+7} \times 10^{43} \text{ erg s}^{-1}$ .

**Key words:** galaxies: active – galaxies: individual: NGC 1068 – galaxies: Seyfert.

## 1 INTRODUCTION

Since Antonucci & Miller (1985) proposed the unification scheme for type-1 and type-2 active galactic nucleus (AGN), it has been commonly thought that highly absorbed (i.e. Compton-thick, with  $N_{\text{H}} \geq 1.5 \times 10^{24} \text{ cm}^{-2}$ ) Seyfert 2s are obscured by neutral gaseous matter embedded in a thick molecular torus located at parsec distances from the central X-ray source (see Netzer 2015, for a recent review). Reflection from the torus reveals itself through a very intense neutral iron K  $\alpha$  emission line at 6.4 keV, with

equivalent widths of  $\sim 1$  keV, and a prominent Compton hump peaking at  $\sim 20$  keV (Ghisellini, Haardt & Matt 1994).

Recently, both the size and the distance of this thick screen have been questioned by a number of observations that have measured significant column density variability of the innermost absorber over time-scales of days or even hours in nearby bright sources such as NGC 1365 (Risaliti et al. 2005; Rivers et al. 2015a), NGC 4388 (Elvis et al. 2004), NGC 4151 (Puccetti et al. 2007) and NGC 7582 (Bianchi et al. 2009; Rivers et al. 2015b). On the other hand, spatially resolved iron K  $\alpha$  line emission, extended on scales of hundreds of parsecs, has been detected in the brightest Compton-thick objects such as NGC 1068 (Young, Wilson & Shopbell 2001; Brinkman et al. 2002), NGC 4945 (Marinucci et al. 2012a) and Mrk 3 (Guainazzi et al. 2012). These measurements suggest that

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different absorbers/reflectors, located on a variety of spatial scales, may contribute (see e.g. Bianchi, Maiolino & Risaliti 2012) to the absorption and reprocessing.

NGC 1068 ( $D_L = 14.4$  Mpc; Tully 1988) is one of the best studied Seyfert 2 galaxies. Indeed, the unification model was first proposed to explain the presence of broad optical lines in its polarized light. In X-rays, it was first studied by *Ginga*, which detected a strong (EW  $\sim 1.3$  keV) neutral iron line (Koyama et al. 1989), an unambiguous sign that we are observing reflected, rather than direct radiation (Matt, Brandt & Fabian 1996). This result was later confirmed by *ASCA* (Ueno et al. 1994; Iwasawa, Fabian & Matt 1997) which resolved the iron line into neutral and ionized components. *BeppoSAX* (Matt et al. 1997) found no evidence for transmitted radiation up to 100 keV, implying a column density of the absorbing material in excess of  $10^{25}$  cm $^{-2}$ . Matt et al. (2004) and Pounds & Vaughan (2006) studied the *XMM-Newton*/EPIC spectra, and found evidence for an iron overabundance with respect to the solar value.

Recently, Bauer et al. (2015) analysed the multi-epoch X-ray spectra of NGC 1068 using different observatories, including 3–79 keV data from *NuSTAR*. They interpreted the broad-band cold reflected emission of NGC 1068 as originating from multiple reflectors with three distinct column densities. The highest  $N_H$  component ( $N_{H,1} \simeq 10^{25}$  cm $^{-2}$ ) is the dominant contribution to the Compton hump, while the lowest  $N_H$  component ( $N_{H,2} \sim 1.5 \times 10^{23}$  cm $^{-2}$ ) produces much of the line emission. The authors also confirm that almost 30 per cent of the neutral Fe K  $\alpha$  line flux arises from regions outside the central 140 pc.

Guainazzi et al. (2000) found evidence for variability, in the 3–10 keV band, comparing two *BeppoSAX* observations performed about one year apart. Later on, comparing *ASCA*, *RossixTE* and *BeppoSAX* spectra taken at different epochs spanning a few months, Colbert et al. (2002) claimed variations in both the continuum and He-like iron line flux on time-scales as short as four months, using the 2–10 keV energy band. Matt et al. (2004), comparing an *XMM-Newton* observation with *BeppoSAX* observations performed a few years earlier, found possible evidence for flux variability of both the cold and the ionized reflectors.

We present a joint *XMM-Newton* and *NuSTAR* monitoring campaign of NGC 1068, from 2014 July until 2015 February, and report on the discovery of a transient excess above 20 keV.

We adopt the cosmological parameters  $H_0 = 70$  km s $^{-1}$  Mpc $^{-1}$ ,  $\Omega_\Lambda = 0.73$  and  $\Omega_m = 0.27$ , i.e. the default ones in *XSPEC* 12.8.1 (Arnaud 1996). Errors correspond to the 90 per cent confidence level for one interesting parameter ( $\Delta\chi^2 = 2.7$ ), if not stated otherwise.

## 2 OBSERVATIONS AND DATA REDUCTION

*NuSTAR*: NGC 1068 was observed by *NuSTAR* with its two co-aligned X-ray telescopes five times. The first three times were in 2012 December: those data are discussed in Bauer et al. (2015). After that, NGC 1068 was the target of a monitoring campaign with *XMM-Newton* composed of four observations, from 2014 July until 2015 February. *NuSTAR* observed the source simultaneously with the third and fourth *XMM-Newton* pointings. The Level 1 data products were processed with the *NuSTAR* Data Analysis Software (NUSTARDAS) package (v. 1.3.0). Cleaned event files (level 2 data products) were produced and calibrated using standard filtering criteria with the *NUPIPELINE* task and the latest calibration files available in the *NuSTAR* calibration data base (CALDB 20150316). Since no spectral variation is found within each observation, we decided to use time-averaged spectra for each epoch. Background light curves are constant within each observation and do not present any flares

due to spurious emission. The background levels are perfectly consistent between the three epochs. We co-added data taken in 2012 December, since no variation was found between those three pointings. The extraction radii for the source and background spectra were 1.5 arcmin each. Net exposure times, after this process, can be found in Table 1, for both focal plane modules A and B. The two pairs of *NuSTAR* spectra were binned in order to oversample the instrumental resolution by at least a factor of 2.5 and to have a signal-to-noise ratio (SNR) greater than 5 in each spectral channel.

*XMM-Newton*: the monitoring campaign of NGC 1068 with *XMM-Newton* was composed of four  $\sim 40$  ks observations, starting on 2014 July 10 with the EPIC CCD cameras, the pn (Strüder et al. 2001) and the two MOS (Turner et al. 2001), operated in small window and thin filter mode. Details of the *XMM-Newton* observations and analysis can be found in Bianchi et al. (in preparation). The resulting net exposure times can be found in Table 1 for the EPIC-pn. Spectra were binned in order to oversample the instrumental resolution by at least a factor of 3 and to have no less than 30 counts in each background-subtracted spectral channel. Cross-calibration constants between the *NuSTAR*-FPMA/B and EPIC-pn are within 10 per cent, in agreement with values presented in Madsen et al. (2015).

NGC 1068 hosts a strong Ultra Luminous X-ray source (ULX) near its nucleus already studied in Matt et al. (2004). The source (at a distance of  $\sim 28$  arcmin from the AGN) is present in our 2014/2015 *XMM* data and we find no differences with respect to the properties discussed in Matt et al. (2004): its contribution to the 4–10 keV spectrum is constrained to be  $\leq 5$  per cent.

## 3 SPECTRAL ANALYSIS

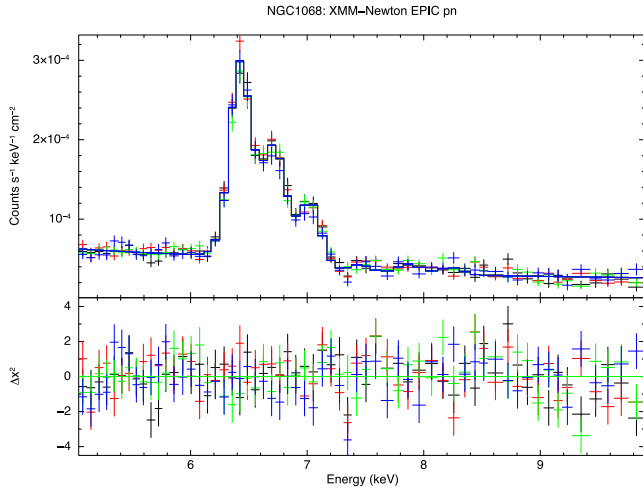
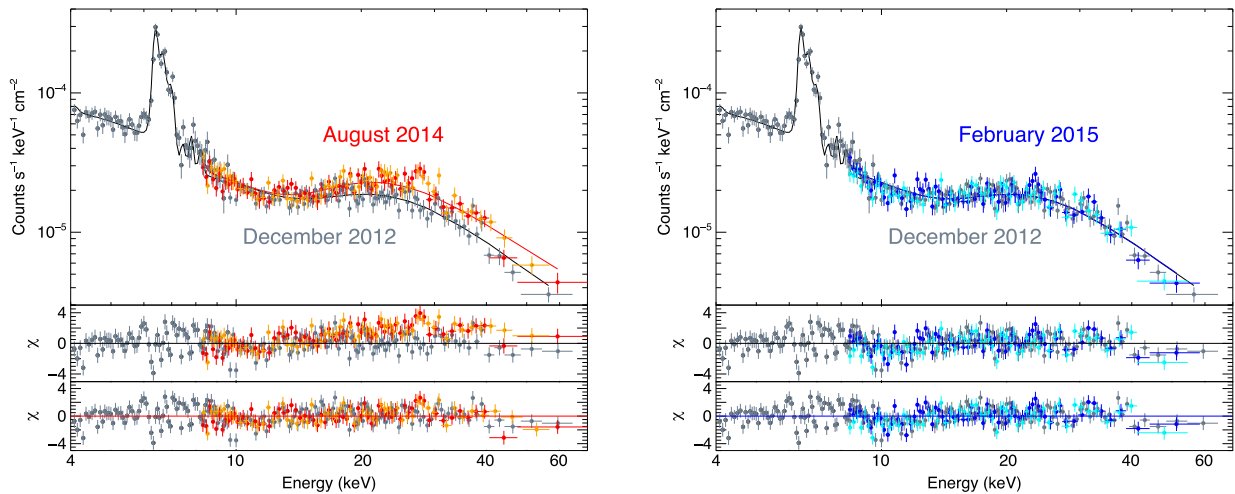
We start our analysis by checking for variability in the four *XMM-Newton* spectra obtained between 2014 July and 2015 February. In our analysis, we only consider data above 4 keV due to the strong contribution at lower energies from distant photoionized, extra-nuclear emission, which will be discussed in Bianchi et al. (in preparation). No differences are found between the spectra, nor do we find differences between the new data set and the one taken in 2000 July, we refer to Bianchi et al. (in preparation) for details. Applying the model discussed in Matt et al. (2004) for the old *XMM* observation to the new 5–10 keV pn spectra, we infer that the flux of the narrow core of the iron K  $\alpha$  line is constant within 5 per cent. The best fit of the four spectra is shown in Fig. 1, no spectral or flux variations are apparent.

Since no significant changes are found between the *XMM* spectra, we used only the observation simultaneous to the high-energy transient event (ObsId 0740060401: Table 1, Fig. 2) and then included *NuSTAR* FPMA and FPMB spectra from observations performed in 2012 December, 2014 August and 2015 February. The final data set is therefore comprised of seven spectra: one *XMM* spectrum taken in 2014 August and three pairs of *NuSTAR* spectra taken in 2012, 2014 and 2015. We only considered *NuSTAR* data above 8 keV because *XMM* spectra have higher spectral resolution and higher SNR in the Fe K  $\alpha$  energy range.

We then apply the best-fitting model discussed in Bauer et al. (2015) to our *XMM+NuSTAR* 4–79 keV data set. This model fits data from 1996 until 2012 from multiple X-ray observatories. The authors found, using *MYTORUS* tables (model M2d), that the reflecting material is composed of three distinct components with  $N_{H,1} \simeq 10^{25}$  cm $^{-2}$ ,  $N_{H,2} = (1.5 \pm 0.1) \times 10^{23}$  cm $^{-2}$  and  $N_{H,3} = (5.0_{-1.9}^{+4.5}) \times 10^{24}$  cm $^{-2}$ .  $N_{H,1}$  is the absorbing column density along the line of sight. *Chandra* observations show that the

**Table 1.** Observation log for the *NuSTAR* and *XMM-Newton* monitoring of NGC 1068.

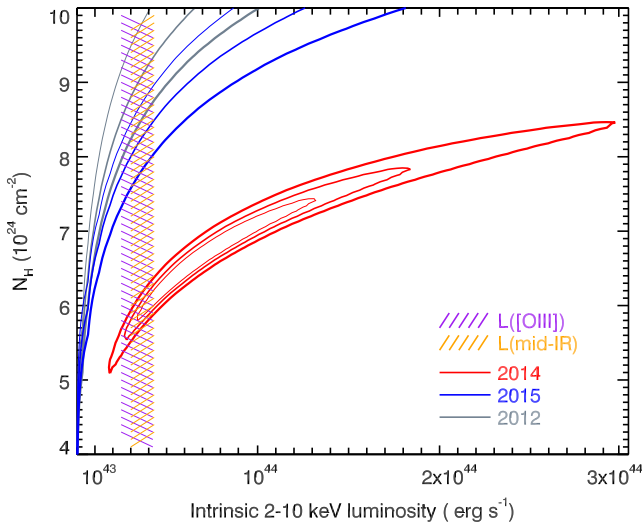
Obs. ID	Date	Exp. time (ks)	4–20 keV count rate (cts s <sup>-1</sup> )		20–80 keV count rate (cts s <sup>-1</sup> )	
			FPMA	FPMB	FPMA	FPMB
60002030002	2012-12-18	57	0.1579 ± 0.0011	0.1494 ± 0.0011	0.0324 ± 0.0006	0.0286 ± 0.0006
60002030004	2012-12-20	48				
60002030006	2012-12-21	19				
60002033002	2014-08-18	52	0.1556 ± 0.0018	0.1456 ± 0.0018	0.0417 ± 0.0010	0.0387 ± 0.0010
60002033004	2015-02-05	53	0.1573 ± 0.0018	0.1444 ± 0.0017	0.0329 ± 0.0009	0.0285 ± 0.0008
4–10 keV Count rate (cts s <sup>-1</sup> )						
EPIC-pn						
0740060201	2014-07-10	44	0.1593 ± 0.0022			
0740060301	2014-07-18	39	0.1540 ± 0.0021			
0740060401	2014-08-19	37	0.1537 ± 0.0021			
0740060501	2015-02-03	37	0.1625 ± 0.0025			

**Figure 1.** Best fit of the four EPIC-pn spectra, with residuals. Black, red, green and blue data points indicate observations performed on 2014 July 10, July 18, August 19 and 2015 February 3, respectively.**Figure 2.** Best-fitting models, spectra and residuals are shown. The 2014 August EPIC-pn spectrum and the 2012 FPMA/B spectra are plotted as grey circles. The 2014 FPMA and FPMB data are overlotted in the top-left panel in red and orange, respectively. The new 2015 FPMA and FPMB data are overlotted in the top-right panel in blue and cyan, respectively. Middle panels show residuals of the two 2014 and 2015 data set to the 2012 best-fitting model (black solid line). Bottom panels show residuals to a model in which the column density and the nuclear flux are left free in the new observations (red and blue solid lines). See the text for details.

spectral features attributed to the  $N_{H,1}$  and  $N_{H,2}$  components arise from the central 2 arcsec only, while  $N_{H,3}$  corresponds to regions outside the central 2 arcsec.  $N_{H,1}$  and  $N_{H,3}$  contribute primarily to the Compton hump, while  $N_{H,2}$  and  $N_{H,3}$  provide dominant contributions to the Fe K lines (Bauer et al. 2015).

For this analysis, the normalizations of the absorbed, scattered and line emission in MYTORUS tables were kept tied together (coupled reprocessor solution; Yaqoob 2012): further details about this solution can be found in Bauer et al. (2015). Applying their best-fitting model to our data set, we find  $\chi^2/\text{dof} = 1166/896 = 1.30$  because the model does not reproduce well the 2014 August *NuSTAR* data (Fig. 2, left-middle panel). Indeed, Fig. 2 and Table 1 show that there is a  $32 \pm 6$  per cent flux increase above 20 keV and a clear additional spectral feature in the *NuSTAR* observation taken in 2014 August.

We first model this excess by leaving the column density  $N_{H,1}$  (along the line of sight) free to vary in the 2014 spectra, which results in a significant improvement of the fit ( $\Delta\chi^2 = 177$  for one additional free parameter). A marginal improvement is found when we also leave the normalization of the primary component free to vary ( $\Delta\chi^2 = 10$  for one additional free parameter), for a final  $\chi^2/\text{dof} = 979/894 = 1.09$ ; no strong residuals are seen



**Figure 3.** Contour plots between the column density along the line of sight and the intrinsic 2–10 keV luminosity. Grey, red and blue lines represent contours for 2012, 2014 and 2015 data sets, respectively. Boldest to thinnest contours indicate 99 per cent, 90 per cent and 68 per cent confidence levels. Vertical orange and purple stripes indicate the 2–10 keV intrinsic luminosities inferred from the mid-IR and [O III] observed luminosities with their relative dispersions from Gandhi et al. (2009) and Lamastra et al. (2009).

throughout the whole 4–79 keV energy band (Fig. 2, bottom panels). Best-fitting values for the column density along the line of sight and for the nuclear component normalization are  $N_{\text{H},1} = (6.7 \pm 1.0) \times 10^{24} \text{ cm}^{-2}$  and  $A_{\text{nuc}} = 0.9_{-0.5}^{+1.0} \text{ ph cm}^{-2} \text{ s}^{-1} \text{ keV}^{-1}$  at 1 keV, respectively. This normalization leads to an unabsorbed 2–10 keV luminosity  $L_X = 7_{-4}^{+7} \times 10^{43} \text{ erg s}^{-1}$ , which is consistent with the value presented in Bauer et al. (2015), within the error bars (Fig. 3). The intrinsic luminosity presented in Bauer et al. (2015) is  $L_X = 2.2 \times 10^{43} \text{ erg s}^{-1}$  and indeed, the authors state that it is a factor of  $\sim 1.6$  lower than the one derived from the mid-IR to X-ray relation in Gandhi et al. (2009) (which is the orange vertical stripe in Fig. 3). However, we note that at high column densities, the derived intrinsic luminosity is highly dependent on the (unknown) geometry of the absorber (e.g. Matt, Pompilio & La Franca 1999). The fit does not improve if we leave the normalization and column density of the 2015 spectra free to vary, indicating that there is no difference between the 2012 and 2015 data sets. Residuals around 25–30 keV in both 2014 August and 2015 February observations (Fig. 2, middle and bottom panels) may be ascribed to residual instrumental features in the NuSTAR ARFs (see figs 7 and 8; Madsen et al. 2015). If we include the instrumental background emission lines between 22 and 35 keV, no significant variations in the best-fitting parameters are found. This effect represents  $\sim 2$  per cent of the total 20–80 keV flux in the 2014 August observation and  $\sim 7$  per cent of the observed flux excess, in the 20–80 keV energy band. Fig. 3 shows the contour plots between the column density along the line of sight (the one directly obscuring the primary continuum) and the intrinsic 2–10 keV nuclear luminosity extrapolated from the de-absorbed best fit from the primary continuum for the three NuSTAR data sets (colours as in Fig. 2). The inferred X-ray luminosity  $L_X \sim 10^{43}–10^{44} \text{ erg s}^{-1}$  is almost four orders of magnitude greater than usually observed in ULXs (Swartz et al. 2004; Walton et al. 2011): the lack of variations below 10 keV and the sharp cutoff in ULX spectra above 20 keV (Bachetti et al. 2013; Walton et al. 2013, 2014, 2015) lead us to conclude that this transient excess cannot be attributed to the ULX in the NGC 1068 FOV.

We then considered the possibility that the high-energy excess in the 2014 August observation might be due to a rise in the reflected emission only, due to matter with  $N_{\text{H}} \sim 10^{25} \text{ cm}^{-2}$ . The only way to have a variation in the Compton hump without an associated variation in the iron line is for the reflector to be almost completely self-obscured. Indeed, fits show that the inclination angle of the reflector would have to be larger than  $87^\circ$ , assuming a toroidal configuration. We therefore conclude that this interpretation, while not impossible, is unlikely.

## 4 DISCUSSION

We interpret the high-energy excess detected in the 2014 August NuSTAR spectra as the first unveiling event ever observed in NGC 1068, in which there is a drop in the column density along our line of sight. If we take into account the mid-IR and [O III] luminosities as proxies of the intrinsic nuclear luminosity, we have additional pieces of information to add to the contour plots shown in Fig. 3. The vertical orange lines represent the intrinsic 2–10 keV luminosity inferred from the mid-IR luminosity (Gandhi et al. 2009). Purple lines indicate the intrinsic 2–10 keV luminosity calculated from the extinction-corrected [O III] luminosity (Marinucci et al. 2012b) using the [O III]–X-ray relation from Lamastra et al. (2009). Contour plots show that the intrinsic X-ray luminosity for the three observations is consistent with those inferred using other proxies, and all the spectral difference can be attributed to a change in the absorbing column density, from  $N_{\text{H}} > 8.5 \times 10^{24} \text{ cm}^{-2}$  in the 2012 observation to  $N_{\text{H},1} = (5.9 \pm 0.4) \times 10^{24} \text{ cm}^{-2}$  in 2014 (Fig. 3, using 90 per cent confidence level for two interesting parameters).

Assuming the bolometric correction from Marconi et al. (2004), we infer  $L_{\text{bol}} = 2.1_{-1.4}^{+3.2} \times 10^{45} \text{ erg s}^{-1}$ , in agreement with Hönic, Prieto & Beckert (2008). The black hole mass of NGC 1068 is estimated to be  $\simeq 1 \times 10^7 M_{\odot}$  (Greenhill et al. 1996; Lodato & Bertin 2003). For consistency with the mid-IR and [O III] luminosities (Fig. 3), we take the lower value  $L_{\text{bol}} = 7 \times 10^{44} \text{ erg s}^{-1}$ , leading to an accretion rate  $\eta_{\text{edd}} \sim 0.55$ , confirming the highly accreting nature of the source.

Absorption variability is common when observations performed months to years apart are compared (Risaliti, Elvis & Nicastro 2002), and has been found on time-scales of hours to days in several sources. However, even in the so-called changing-look AGN (sources that switched from the Compton-thick to the Compton-thin state and vice versa) an eclipsing/unveiling event affecting only the spectrum above 10 keV has never been observed: we emphasize that this is the first time that a Compton-thick unveiling event of this kind has been reported. We note that this is different from the intrinsic variability recently reported for the Compton-thick AGN NGC 4945 (Puccetti et al. 2014). Our finding is supporting a clumpy structure of the obscuring material along the line of sight (Nenkova et al. 2008).

In this scenario, we do not have a single, monolithic obscuring wall, but the total column density along the line of sight is the sum of the contributions from a discrete number of clouds. The NuSTAR sensitivity above 10 keV allowed us to infer only a lower limit on the column density variation ( $\Delta N_{\text{H}} \simeq 2.5 \times 10^{24} \text{ cm}^{-2}$ ) but greater changes could have occurred (top-left corner of Fig. 3: the parameter space with  $N_{\text{H}} > 8.5 \times 10^{24} \text{ cm}^{-2}$ ) but were not measurable with our data. Further monitoring of NGC 1068 could provide constraints on the number of clouds and their distance from the illuminating source.



## 5 CONCLUSIONS

We presented a spectral analysis of the 4–79 keV *NuSTAR* and *XMM-Newton* monitoring campaign of NGC 1068 obtained between 2014 July and 2015 February. We found a clear transient excess above 20 keV in the 2014 August *NuSTAR* observation, while no variations are found in the *XMM* data below 10 keV. The most plausible explanation is an unveiling event, in which for a short while the total absorbing column, probably composed by a number of individual clouds, became less thick so as to permit to the nuclear radiation to pierce through it. Our result provides further evidence that the obscuring material along our line of sight is clumpy, and enables us to infer a 2–10 keV intrinsic luminosity of  $L_X = 7_{-4}^{+7} \times 10^{43} \text{ erg s}^{-1}$ .

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## REFERENCES

- Antonucci R. R. J., Miller J. S., 1985, *ApJ*, 297, 621  
 Arnaud K. A., 1996, in Jacoby G. H., ed., *ASP Conf. Ser. Vol. 101, Astronomical Data Analysis Software and Systems V*. Astron. Soc. Pac., San Francisco, p. 17  
 Bachetti M. et al., 2013, *ApJ*, 778, 163  
 Bauer F. E. et al., 2015, *ApJ*, 812, 116  
 Bianchi S., Piconcelli E., Chiaberge M., Bailón E. J., Matt G., Fiore F., 2009, *ApJ*, 695, 781  
 Bianchi S., Maiolino R., Risaliti G., 2012, *Adv. Astron.*, 2012, 17  
 Brinkman A. C., Kaastra J. S., van der Meer R. L. J., Kinkhabwala A., Behar E., Kahn S. M., Paerels F. B. S., Sako M., 2002, *A&A*, 396, 761  
 Colbert E. J. M., Weaver K. A., Krolik J. H., Mulchaey J. S., Mushotzky R. F., 2002, *ApJ*, 581, 182  
 Elvis M., Risaliti G., Nicastro F., Miller J. M., Fiore F., Puccetti S., 2004, *ApJ*, 615, L25  
 Gandhi P., Horst H., Smette A., Hönic S., Comastri A., Gilli R., Vignali C., Duschl W., 2009, *A&A*, 502, 457  
 Ghisellini G., Haardt F., Matt G., 1994, *MNRAS*, 267, 743  
 Greenhill L. J., Gwinn C. R., Antonucci R., Barvainis R., 1996, *ApJ*, 472, L21  
 Guainazzi M., Molendi S., Vignati P., Matt G., Iwasawa K., 2000, *New Astron.*, 5, 235  
 Guainazzi M., La Parola V., Miniutti G., Segreto A., Longinotti A. L., 2012, *A&A*, 547, A31  
 Hönic S. F., Prieto M. A., Beckert T., 2008, *A&A*, 485, 33  
 Iwasawa K., Fabian A. C., Matt G., 1997, *MNRAS*, 289, 443  
 Koyama K., Inoue H., Tanaka Y., Awaki H., Takano S., Ohashi T., Matsuoka M., 1989, *PASJ*, 41, 731  
 Lamastra A., Bianchi S., Matt G., Perola G. C., Barcons X., Carrera F. J., 2009, *A&A*, 504, 73  
 Lodato G., Bertin G., 2003, *A&A*, 398, 517  
 Madsen K. K. et al., 2015, *ApJS*, 220, 8  
 Marconi A., Risaliti G., Gilli R., Hunt L. K., Maiolino R., Salvati M., 2004, *MNRAS*, 351, 169  
 Marinucci A., Risaliti G., Wang J., Nardini E., Elvis M., Fabbiano G., Bianchi S., Matt G., 2012a, *MNRAS*, 423, L6  
 Marinucci A., Bianchi S., Nicastro F., Matt G., Goulding A. D., 2012b, *ApJ*, 748, 130  
 Matt G., Brandt W. N., Fabian A. C., 1996, *MNRAS*, 280, 823  
 Matt G. et al., 1997, *A&A*, 325, L13  
 Matt G., Pompilio F., La Franca F., 1999, *New Astron.*, 4, 191  
 Matt G., Bianchi S., Guainazzi M., Molendi S., 2004, *A&A*, 414, 155  
 Nenkova M., Sirocky M. M., Nikutta R., Ivezić Ž., Elitzur M., 2008, *ApJ*, 685, 160  
 Netzer H., 2015, *ARA&A*, 53, 365  
 Pounds K., Vaughan S., 2006, *MNRAS*, 368, 707  
 Puccetti S., Fiore F., Risaliti G., Capalbi M., Elvis M., Nicastro F., 2007, *MNRAS*, 377, 607  
 Puccetti S. et al., 2014, *ApJ*, 793, 26  
 Risaliti G., Elvis M., Nicastro F., 2002, *ApJ*, 571, 234  
 Risaliti G., Elvis M., Fabbiano G., Baldi A., Zezas A., 2005, *ApJ*, 623, L93  
 Rivers E. et al., 2015a, *ApJ*, 804, 107  
 Rivers E. et al., 2015b, preprint ([arXiv:1511.01951](https://arxiv.org/abs/1511.01951))  
 Strüder L. et al., 2001, *A&A*, 365, L18  
 Swartz D. A., Ghosh K. K., Tennant A. F., Wu K., 2004, *ApJS*, 154, 519  
 Tully R. B., 1988, *Nearby Galaxies Catalog*. Cambridge Univ. Press, Cambridge  
 Turner M. J. L. et al., 2001, *A&A*, 365, L27  
 Ueno S., Mushotzky R. F., Koyama K., Iwasawa K., Awaki H., Hayashi I., 1994, *PASJ*, 46, L71  
 Walton D. J., Roberts T. P., Mateos S., Heard V., 2011, *MNRAS*, 416, 1844  
 Walton D. J. et al., 2013, *ApJ*, 779, 148  
 Walton D. J. et al., 2014, *ApJ*, 793, 21  
 Walton D. J. et al., 2015, *ApJ*, 806, 65  
 Yaqoob T., 2012, *MNRAS*, 423, 3360  
 Young A. J., Wilson A. S., Shopbell P. L., 2001, *ApJ*, 556, 6

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