

Open Research Online

The Open University's repository of research publications and other research outputs

Protoplanetary disks in the hostile environment of Carina

Journal Item

How to cite:

Mesa-Delgado, A.; Zapata, L.; Henney, W. J.; Puzia, T. H. and Tsamis, Y. G. (2016). Protoplanetary disks in the hostile environment of Carina. *The Astrophysical Journal*, 825(1), article no. L16.

For guidance on citations see [FAQs](#).

© 2016 The American Astronomical Society

Version: Version of Record

Link(s) to article on publisher's website:

<http://dx.doi.org/doi:10.3847/2041-8205/825/1/L16>

Copyright and Moral Rights for the articles on this site are retained by the individual authors and/or other copyright owners. For more information on Open Research Online's data [policy](#) on reuse of materials please consult the policies page.

oro.open.ac.uk



PROTOPLANETARY DISKS IN THE HOSTILE ENVIRONMENT OF CARINA

A. MESA-DELGADO¹, L. ZAPATA², W. J. HENNEY², T. H. PUZIA¹, AND Y. G. TSAMIS³

¹ Instituto de Astrofísica, Facultad de Física, Pontificia Universidad Católica de Chile, Av. Vicuña Mackenna 4860, 782-0436 Macul, Santiago, Chile; amesad@astro.puc.cl

² Instituto de Radioastronomía y Astrofísica, Universidad Nacional Autónoma de México, Campus Morelia, Apartado Postal 3-72, 58090 Morelia, Michoacán, México

³ Department of Physics and Astronomy, University College London, London WC1E 6BT, UK

Received 2016 April 24; revised 2016 May 24; accepted 2016 May 25; published 2016 June 30

ABSTRACT

We report the first direct imaging of protoplanetary disks in the star-forming region of Carina, the most distant massive cluster in which disks have been imaged. Using the Atacama Large Millimeter/sub-millimeter Array (ALMA), the disks are observed around two young stellar objects (YSOs) that are embedded inside evaporating gaseous globules and exhibit jet activity. The disks have an average radius of 60 au and total masses of 30 and 50 M_{Jup} . Given the measured masses, the minimum timescale required for planet formation ($\sim 1\text{--}2$ Myr) and the average age of the Carina population ($\sim 1\text{--}4$ Myr), it is plausible that young planets are present or their formation is currently ongoing in these disks. The non-detection of millimeter emission above the 4σ threshold ($\sim 7M_{\text{Jup}}$) in the core of the massive cluster Trumpler 14, an area containing previously identified proplyd candidates, suggests evidence for rapid photo-evaporative disk destruction in the cluster's harsh radiation field. This would prevent the formation of giant gas planets in disks located in the cores of Carina's dense subclusters, whereas the majority of YSO disks in the wider Carina region remain unaffected by external photoevaporation.

Key words: protoplanetary disks – stars: formation – stars: protostars – techniques: high angular resolution – techniques: interferometric

1. INTRODUCTION

A significant fraction of stars are born in groups (Blaauw 1964; Lada & Lada 2003) where massive and low-mass stars co-exist and evolve. In these regions, the far-ultraviolet (FUV) and extreme ultraviolet (EUV) radiation from the most massive stars can curtail disk lifetimes of the low-mass protostars via external photoevaporation, caused by the heating and dissociating effects of ionizing radiation (e.g., Johnstone et al. 1998; Adams et al. 2004; Clarke 2007; Gorti et al. 2016). Depending on the intensity and the duration of the exposure, photoevaporation by external irradiation can potentially lead to disk dispersal on timescales shorter than the minimum 1–2 Myr required for planet formation (Hubickyj et al. 2005; Lissauer et al. 2009; Najita & Kenyon 2014; ALMA Partnership et al. 2015). Although internal photoevaporation due to FUV and X-ray radiation from the protostar itself may be important, this effect generally acts on longer timescales of about 3–5 Myr (e.g., Owen et al. 2012; Gorti et al. 2016). Understanding the impact of external photoevaporation on disks within massive clusters is key to assessing their chances for survival and their planet formation capacity.

Richert et al. (2015) found no evidence for significant disk dispersal by analyzing young stellar objects (YSOs) in a large sample of stellar associations. Due to stellar crowding and/or PAH contamination, their analysis is inconclusive for the cores of the massive clusters in Orion, M 17, and Carina, which is exactly where external photoevaporation could be important. A clear example is Orion's rich population of proplyds (O'Dell et al. 1993), where millimeter observations show that disks have been eroded down to $\sim 3M_{\text{Jup}}$ within 0.03 pc from the Trapezium cluster (e.g., Eisner et al. 2008; Mann & Williams 2010; Mann et al. 2014).

Whereas the Orion nebula is dominated by a single O7V-type star (θ^1 Ori C; Simón-Díaz et al. 2006), the Carina nebula,

at a distance of ~ 2.3 kpc (Smith & Brooks 2008, hereafter SB08), hosts nearly 100 O-type stars and tens of thousands of lower-mass young stars (Povich et al. 2011a, 2011b; Preibisch et al. 2011a, 2011b). The clusters Trumpler (Tr) 14 and 16 have ultraviolet (UV) luminosities about 20 and 60 times higher than θ^1 Ori C and they are home of the most massive and luminous stars in the Galaxy (see the review by SB08). As a prototype of massive star-forming regions in our Galaxy, Carina is ideally suited to an investigation of disk evaporation: its older age (1–4 Myr; SB08) could allow enough time for a deeper action of the external photoevaporation process, especially in the core of the massive clusters.

In this Letter, we present novel observations with the Atacama Large Millimeter/sub-millimeter Array (ALMA) of YSO disks in Carina and investigate the implications for disk survival in the face of external photoevaporation.

2. OBSERVATIONS AND DATA REDUCTION

We observed four regions covering a range of separations from Tr 16 and Tr 14, centered on sources identified by Smith et al. (2003) as potential protoplanetary disks (see Figure 1). We focused on the evaporating gas globules (EGGs) named 104–593, 105–600 (frEGG-1; Sahai et al. 2012), and 104–598, as well as a section of the Tr 14 core that contains the best candidates for bona fide proplyds in Carina (Smith et al. 2010).

The observations were performed on 2015 October 27 as part of the Cycle 3 program 2015.1.01323.S. We used 45 of the 12 m antennas, yielding baselines with projected lengths from 21 to 16, 196 m. The positions of the phase centers of each pointing are given in Table 1. Continuum images were obtained in Band 6, with integration times of 160 s and a frequency coverage of 220–238 GHz within a primary beam with a full width at half maximum of about $27''$. The digital

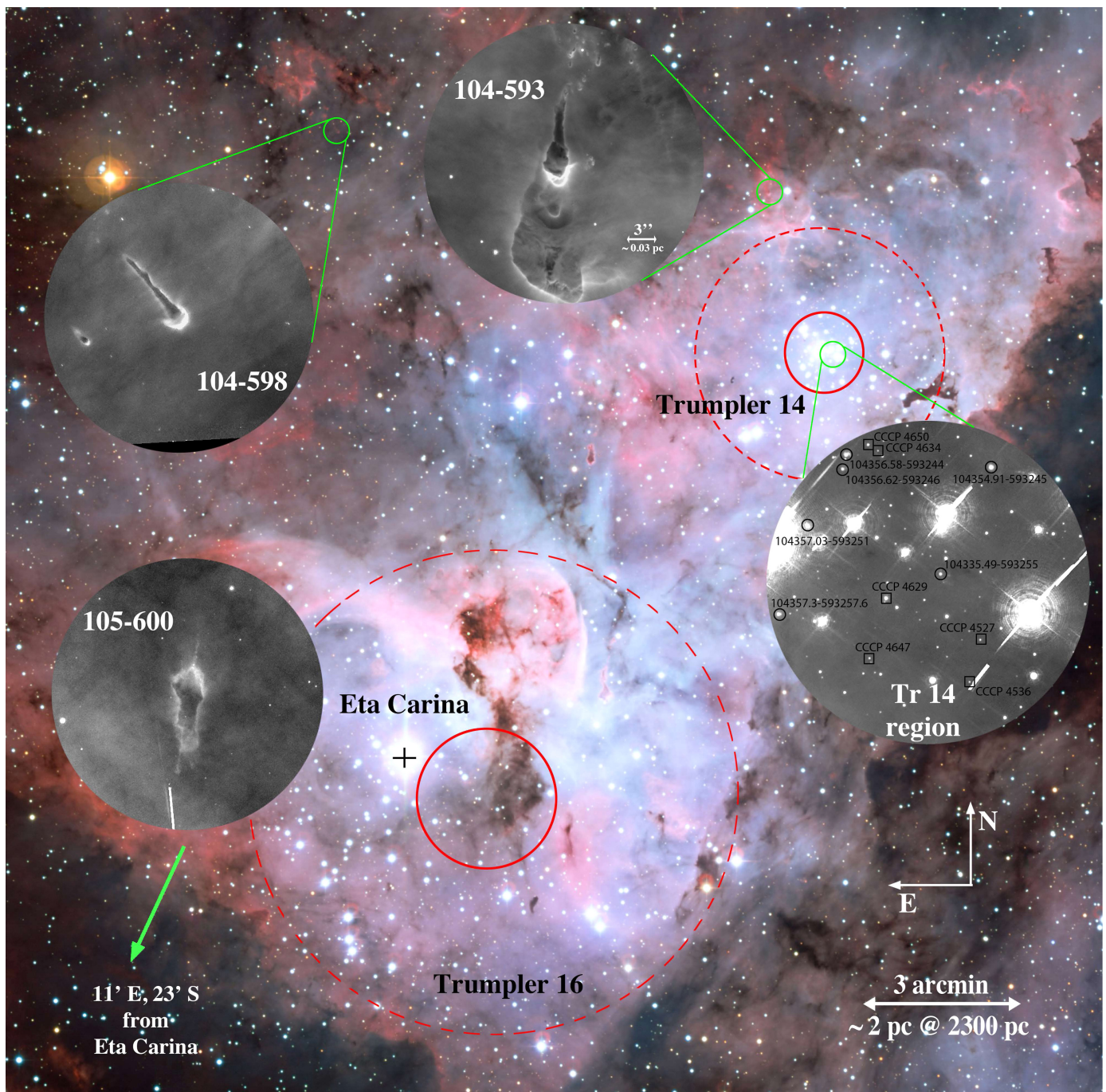


Figure 1. ALMA fields and the clusters Tr 16 and Tr 14 shown on a wide-field optical image of the Carina nebula (ESO). Red circles mark upper and lower limits to the size of irradiation zones around each cluster (see Section 4). The ALMA fields are represented by the close-up images observed with the *Hubble Space Telescope* and its Advanced Camera for Surveys instrument in the F658N narrow-band filter. Within the Tr 14 region, the circles contain protoplanetary candidates, whereas the squares mark sources with near-infrared excess.

correlator was configured to four wide spectral windows in dual polarization mode centered at 223.99 GHz, 234.07 GHz, 220.87 GHz, and 237.02 GHz (the average frequency of 228.5 GHz, corresponding to $\lambda \approx 1.3$ mm). The effective bandwidth for the 223 and 234 GHz spectral windows was 2000 MHz, and 938 MHz for the 220 and 237 GHz windows. The weather conditions were very good and stable with an average precipitable water vapor of 1.06 mm and system temperature of 60–200 K. The ALMA calibration included simultaneous observations of the 183 GHz water line with

water vapor radiometers, used to reduce the atmospheric phase noise. Quasars J1107–4449 and J1049–6003 were used to calibrate the bandpass and the gain fluctuations, respectively.

Data reduction was performed using the Common Astronomy Software Applications package (CASA 4.2.2; McMullin et al. 2007). Imaging of the calibrated visibilities was done using the task CLEAN. The resulting root-mean-square noise for the continuum emission (σ) was about 60 μ Jy at an angular resolution of $0''.03 \times 0''.02$ with a position angle (P.A.) of $-46^\circ.34$. We thus resolve structures with physical sizes larger

Table 1
Physical Parameters of the Protoplanetary Disks Detected in Carina

Source	Position ^a		Flux Density ^b	Spectral ^c	Deconvolved Size ^d			Disk Mass ^e	
	α_{2000}	δ_{2000}	229 GHz	Index	Major	Minor	P.A.	40 K	20 K
	(^h ^m ^s)	([°] ['] ^{''})	(mJy)		(milliarcsecond)		([°])	(M_{Jup})	
104–593	10 44 05.360	–59 29 40.77	1.5 ± 0.2	2.9	45 ± 7	25 ± 5	17 ± 4	45	105
105–600	10 46 32.905	–60 03 53.67	0.90 ± 0.09	2.5	50 ± 5	30 ± 6	110 ± 12	27	62
104–598	10 45 14.200	–59 28 25.00	$\leq 0.24^f$	≤ 7	≤ 16
Tr 14 region	10 43 55.490	–59 32 55.00	$\leq 0.24^f$	≤ 7	≤ 16

Notes.

^a Positional errors are $0''.0075$ for R.A. and $0''.005$ for decl.

^b Obtained from averaging all the four spectral windows.

^c From the density fluxes between the 220 and 237 GHz spectral windows.

^d From the GAUSSFIT task of the CASA package.

^e See Section 3.

^f 4σ upper limit.

than $70 \text{ au} \times 45 \text{ au}$, or roughly the size of the solar system. In deriving the integrated fluxes, we used a Briggs robustness parameter of 0.5 in the CLEAN task to obtain a better signal-to-noise ratio at the cost of losing some spatial resolution. We averaged all line-free channels in the four spectral windows for the construction of the 1.3 mm continuum images.

3. DISKS WITHIN CARINA'S EGGS

In 104–593 and 105–600 we detect disks embedded within the core of the optically opaque EGGs (Figure 2). Table 1 gives the sizes and integrated fluxes of the disks determined with the GAUSSFIT task of CASA from a natural-weighted image and within a $0''.2 \times 0''.2$ box centered on the detection. The disks have peak signals about 5σ and 10σ , and both are marginally resolved with an average deconvolved size of $0''.05 \times 0''.03 \approx 120 \text{ au} \times 70 \text{ au}$. Assuming a geometrically thin, circular disk, a radius R_{disk} of $\sim 60 \text{ au}$ is implied in both cases, inclined at $\sim 35^\circ$ to the line of sight.

The disk nature of these ALMA detections is further supported by their association with mid-infrared sources (Povich et al. 2011a). Our positions match the YSOs PCYC 429 for 104–593 and PCYC 1173 for 105–600 within the errors; these are Herbig AeBe stars with masses of $2.8 \pm 1.6 M_\odot$ and $1.8 \pm 0.9 M_\odot$, respectively. The mid-infrared spectra of both targets imply that they are Class 0/I sources, indicating that they still retain a massive envelope (Povich et al. 2011a; Reiter & Smith 2013). Evidence for ongoing accretion onto the protostars is provided by the collimated outflows emerging from both systems in the form of Herbig-Haro objects (Smith et al. 2010; Reiter & Smith 2013). The spatial distribution of the bipolar outflows also corroborates the disk nature of the sources since the disks are located at the origin of each outflow, and their major axes are aligned almost perpendicular to them (Figure 2). In the case of 105–600, the presence of an hourglass-shaped near-infrared reflection nebula is additional evidence for a rotationally flattened envelope or disk (Sahai et al. 2012).

By comparing the flux densities between the spectral windows at 220 and 237 GHz, spectral indices higher than 2 are determined for the two sources, implying that the observed emission is most likely produced by dust grains (see Table 1). Assuming that the dust is optically thin and isothermal (Hildebrand 1983), these indices suggest that the grains have

grown to larger sizes than are seen in interstellar medium dust (Ossenkopf & Henning 1994), which is consistent with expectations for grain evolution in the mid-plane of evolved, settled disks (Acke et al. 2004).

The total disk mass (M_{disk}) can be derived under the same assumptions, being proportional to the mm flux density S_ν :

$$M_{\text{disk}} = \frac{S_\nu D^2}{B_\nu(T_{\text{dust}}) \kappa_\nu}, \quad (1)$$

where D is the distance to Carina, $B_\nu(T_{\text{dust}})$ is the Planck function at the dust temperature T_{dust} , and κ_ν is the absorption coefficient per unit of total (gas + dust) mass density. The estimated values for the total masses associated with the dust continuum emission are presented in Table 1. We have used a dust mass opacity $\kappa_{1300 \mu\text{m}} = 0.015 \text{ cm}^2 \text{ g}^{-1}$ (Ossenkopf & Henning 1994, assuming a 100:1 gas-to-dust mass ratio) and $T_{\text{dust}} = 40 \text{ K}$, as appropriate for disks around intermediate mass stars (e.g., Natta et al. 2000, p. 559). For the brighter of the two sources, 104–593, the peak observed flux density of approximately 0.6 mJy/beam is roughly equal to the 230 GHz Planck function at 40 K, implying that the central parts of the disk may be optically thick, leading to the true masses being underestimated. The mass estimates are uncertain within a factor of a few due to potential deviations of the dust opacity from our adopted value, depending on the degree of grain growth within the disk, plus uncertainties in the dust temperature (see Table 1).

We derive masses of about $50 M_{\text{Jup}}$ and $30 M_{\text{Jup}}$ for 104–593 and 105–600, respectively. These masses are on the upper end of the typical M_{disk} distribution found in Class I sources in less hostile environments such as Taurus, Ophiuchus, and Orion (e.g., Eisner & Carpenter 2006; Andrews & Williams 2007; Andrews et al. 2013; Williams et al. 2013; Carpenter et al. 2014; Mann et al. 2014; Cieza et al. 2015). The disks can be considered protoplanetary because the measured masses are well above the minimum M_{disk} of about $10 M_{\text{Jup}}$ required for a pre-solar nebula to develop a planetary system (Weidenschilling 1977). Additionally, the average age of the Carina population ($\sim 1\text{--}4 \text{ Myr}$; SB08) perfectly fits the required minimum timescale to form planets ($\sim 1\text{--}2 \text{ Myr}$; see Section 1). Planet formation may also occur on shorter scales in Herbig

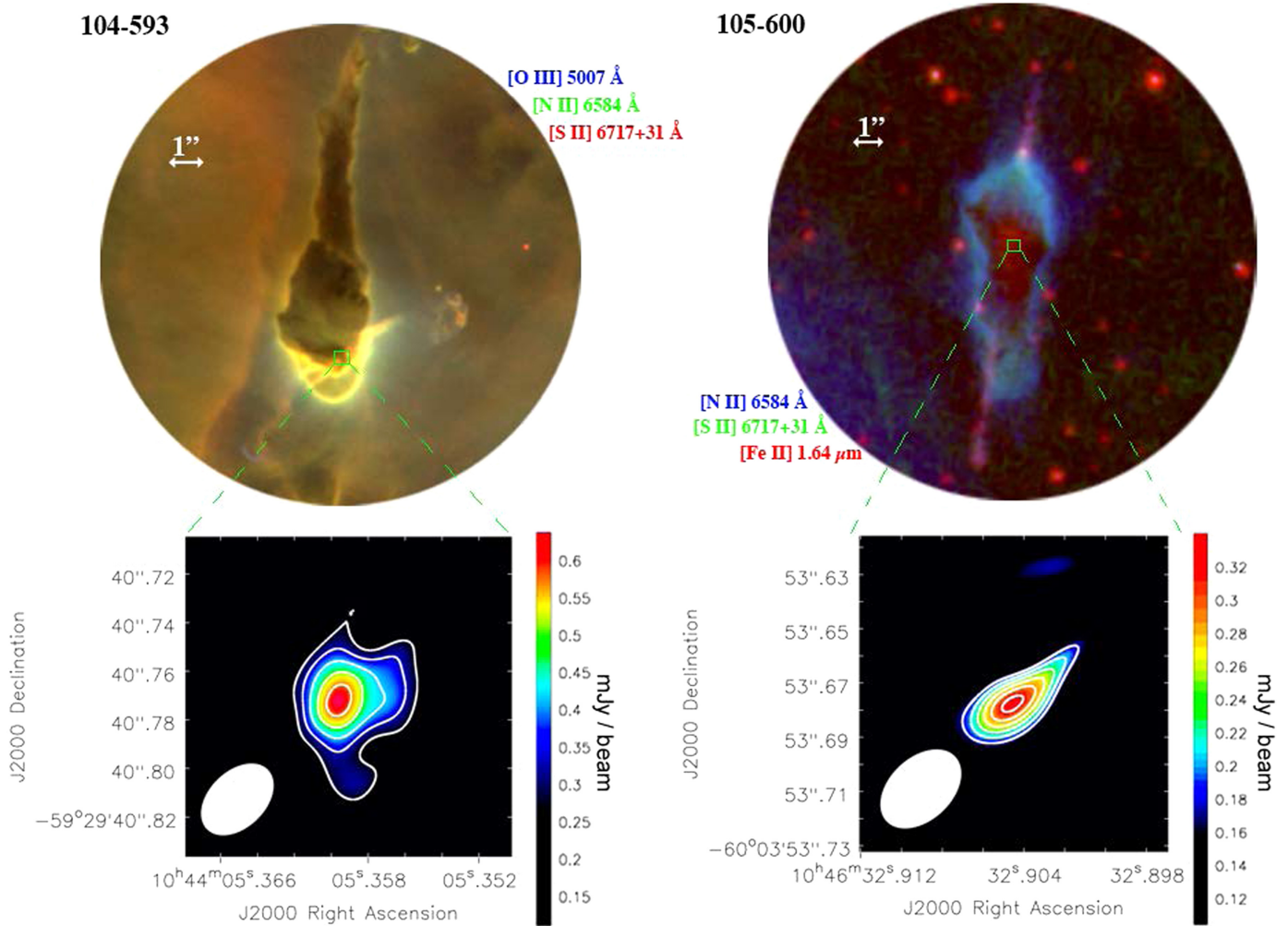


Figure 2. Top: RGB images of the EGGs 104–593 (left-hand side) and 105–600 (right-hand side) constructed from narrow-band *Hubble Space Telescope* images from Wide Field Camera 3 (WFC3), Advanced Camera for Surveys (ACS), and Wide Field/Planetary Camera 2 (WFPC2). For 104–593, WFC3 visual band filters in lines of [O III] (blue), [N II] (green), and [S II] (red). For 105–600, ACS visual band [N II] filter (blue), WFPC2 visual band [S II] filter (green), and WFC3 infrared [Fe II] filter (red). The images also show the Herbig-Haro objects HH1066 in 104–593 and HH1006 in 105–600. Bottom: ALMA detections at 1.3 mm continuum emission. The white contours represent values at 3.8σ , 5.3σ , 6.8σ , 8.3σ , and 9.8σ mJy beam^{-1} for 104–593, and 2.9σ , 3.4σ , 3.9σ , 4.4σ , 5.0σ , and 5.5σ mJy beam^{-1} for 105–600. The white ellipses show the beam size of $0''.03 \times 0''.02$.

systems (Mendigutía et al. 2012), so it is plausible that young planets are forming or already present within these EGGs.

No millimeter emission was detected above the 4σ threshold within either the 104–598 or the Tr 14 fields, yielding an upper limit of $\sim 7 M_{\text{Jup}}$ to the mass of any disk that might be present (see Table 1). In the case of 104–598, no infrared source is detected (Povich et al. 2011a; Preibisch et al. 2011a), implying that the EGG does not (yet) contain a YSO. Similar starless EGGs are seen in the Cygnus OB2 (Wright et al. 2012) and are especially common in M 16 (Hester et al. 1996; McCaughrean & Andersen 2002).

In the Tr 14 region, a total of 12 YSOs are found among candidate proplyds (Ascenso et al. 2007; Smith et al. 2010) and sources with near-infrared excess (Preibisch et al. 2011a). Unfortunately, half of them are located near the field edge where the sensitivity is lower (Figure 1). These YSO disks must be much less massive than those detected in 104–593 and 105–600. If the disks were warmer than 40 K (which may well be the case due to external heating from the nearby massive stars), then the 4σ limit would be even lower than $7 M_{\text{Jup}}$, being

roughly similar to the median mass of $\sim 5\text{--}8 M_{\text{Jup}}$ for Class II disks (e.g., Andrews & Williams 2005; Eisner 2012). According to the M_{disk} distribution in low-mass star-forming regions (e.g., Williams & Cieza 2011), we expected to detect about half of the YSOs above the threshold. The high flux of ionizing and dissociating radiation produced by the massive members of Tr 14 might explain why we have not detected massive disks in this field.

4. PHOTOEVAPORATION OF DISKS AND EGGS

Theoretical values of the disk mass photoevaporation rate \dot{M} can depend on the strength of the illuminating radiation field; the disk size and YSO mass; and whether it is non-ionizing FUV or ionizing EUV radiation that controls the evaporation (e.g., Johnstone et al. 1998). In the simple case of EUV radiation ($\lambda < 912 \text{ \AA}$), the incident photon flux F_{EUV} balances the recombination rate per unit area in the evaporation flow plus the hydrogen ion flux through the ionization front,

allowing us to estimate

$$\dot{M}_{\text{EUV}} \simeq 4 \times 10^{-9} \left(\frac{F_{\text{EUV}}}{10^{10} \text{ s}^{-1} \text{ cm}^2} \right)^{1/2} \times \left(\frac{R_{\text{disk}}}{100 \text{ au}} \right)^{3/2} M_{\odot} \text{ yr}^{-1}. \quad (2)$$

A sufficient flux F_{FUV} of FUV radiation (from 912 Å to 2000 Å) can increase the global \dot{M} via a warm atomic disk wind, which inflates the ionization front to several times R_{disk} (e.g., Johnstone et al. 1998). The physics of FUV evaporation is less well understood and can depend on the rate of viscous spreading of the disk (Anderson et al. 2013). It is most efficient in the supercritical case (see Adams et al. 2004) when the sound speed in the neutral wind exceeds the gravitational escape speed from the disk, requiring $T \simeq 1000$ K or $F_{\text{FUV}} > 10^4 G_0$ if $R_{\text{disk}} = 60$ au, where $G_0 = 1.6 \times 10^{-3} \text{ erg s}^{-1} \text{ cm}^{-2}$ is the mean FUV field in the solar neighborhood.

For the Tr 14 region, where we do not detect any disks, the conditions are certainly extreme. The cluster has an UV luminosity (Q_H) of $2 \times 10^{50} \text{ photons}^{-1}$ (SB08), implying $F_{\text{EUV}} = 10^{14} \text{ photons}^{-1} \text{ cm}^{-2}$ and $F_{\text{FUV}} \approx 10^6 G_0$. The \dot{M}_{EUV} alone then gives $2 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$ for $R_{\text{disk}} = 60$ au, and FUV radiation would likely increase this by factors of 2–4. The corresponding evaporation timescale M_{disk}/\dot{M} is then as short as 10^5 years for an initial disk mass of $50 M_{\text{Jup}}$. Photoevaporation becomes inefficient once the disk shrinks to a radius $R_{\text{disk}} \approx 3(M_*/M_{\odot})$ au, where the escape velocity is of the order of the ionized sound speed. If $M_{\text{disk}} \propto R_{\text{disk}}$ (corresponding to a surface density profile $\sim R^{-1}$), then the disk mass in this phase will be $M_{\text{disk}} < 0.003 M_*$. Thus, centrally condensed disks may persist in the core of Tr 14 with compact sizes (< 10 au) and low masses ($< 3 M_{\text{Jup}}$).

Similar timescales are found by simply scaling results from the Orion proplyds, where empirical estimates of \dot{M} have a median value $\dot{M} = 10^{-7} M_{\odot}/\text{year}$ (Henney & O’Dell 1999; Tsamis et al. 2013). From Equation (2), the Q_H for Tr 14 and Tr 16 would predict $\dot{M}_{\text{EUV}} > 10^{-7} M_{\odot} \text{ yr}^{-1}$ for a $R_{\text{disk}} = 100$ au within radii of $45'' \approx 0.5$ pc and $80'' \approx 0.9$ pc, respectively. These EUV-irradiation zones are shown in Figure 1 (solid red circles), although for Tr 16 the picture is more complex than indicated because the high-mass stars are spread over $10'$ (Povich et al. 2011b). Adding the FUV contribution, the irradiation zones become larger (the dashed red circles in Figure 1) from the condition $F_{\text{FUV}} > 10^4 G_0$. Any disk remaining in these areas for 1 Myr will have lost approximately $0.1 M_{\odot} \approx 100 M_{\text{Jup}}$. For disks with initial masses and sizes similar to our detected sources, this would be more than sufficient to shrink them below our detection limit.

Analogous physics operates in the evaporation of EGGs. The EUV controls \dot{M} in this case, as the FUV-heated layer is thin compared with the EGG radius. Assuming a roughly spherical shape, 105–600 has a radius of $R_{\text{EGG}} \approx 1840$ au and is located at a projected distance of 17 pc from Tr 16, which has a Q_H of $6 \times 10^{50} \text{ photons}^{-1}$ (SB08). If the true distance is not much greater than the projected one, then the incident flux is $F_{\text{EUV}} = 6.5 \times 10^9 \text{ photons}^{-1} \text{ cm}^{-2}$, implying $\dot{M} = 2.5 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$. Within a factor of two, the same result is derived from the $H\alpha$ surface brightness of the globule’s bright rim (measured on the images from the Advanced Camera for Surveys; Smith et al. 2010), which gives a model-independent estimate of F_{EUV} , albeit subject to

uncertainties in the foreground dust extinction. The EGG mass is estimated as $M_{\text{EGG}} = 0.1\text{--}0.3 M_{\odot}$ (Sahai et al. 2012), leading to a timescale M_{EGG}/\dot{M} of approximately 1 Myr. For 104–593, the EGG has a similar size, but is found at a projected distance of 3.3 pc from Tr 14, resulting in a larger ionizing flux $F_{\text{EUV}} \approx 2 \times 10^{11} \text{ photons}^{-1} \text{ cm}^{-2}$. No mass estimate is available for this EGG, but if it were similar to 105–600, then it would be evaporated in about 2×10^5 years.

Once the EGGs have been dispersed, then the disks themselves will be subject to external photoevaporation. Taking our observed parameters of $M_{\text{disk}} = 50 M_{\text{Jup}} \approx 0.05 M_{\odot}$ and $R_{\text{disk}} = 60$ au, we find \dot{M}_{EUV} of 10^{-9} to $10^{-8} M_{\odot} \text{ yr}^{-1}$, corresponding to timescales of 5 to 50 Myr for the two disks. The FUV fluxes at the detected disk positions are weak, $F_{\text{FUV}} \approx 4500 G_0$ for 104–593 and $F_{\text{FUV}} \approx 300 G_0$ for 105–600, resulting in a relatively cool neutral wind ($T < 300$ K) that would have difficulty escaping the stellar gravitational potential. Thus, the EUV-derived timescales are accurate if the systems do not move significantly from their current position.

5. CONCLUSIONS

High spatial resolution observations at millimeter wavelengths with ALMA allow us to overcome the problem of stellar crowding in the densely populated center of Tr 14. Therefore, given the number of YSO sources within our observed field (Figure 1), the non-detection of any disks in the exposed core of the cluster is evidence for accelerated disk dispersal. This is consistent with the disk fraction as measured by the near-infrared excess emission among X-ray selected YSOs in Carina being significantly smaller than in lower-mass clusters of a similar age (Preibisch et al. 2011a).

Roughly 50% of the YSO population in Carina is found in sparse small groups that are widely distributed over an area of several square parsecs (Preibisch et al. 2011b; Zeidler et al. 2016). The newly detected Carina disks belong to this distributed population and are both located at larger distances from the ionizing sources where external evaporation will not threaten their survival over the next 10 million years. Disk formation and evolution in most of these YSOs will probably proceed similarly to those in less massive clusters, and planetary formation should occur without being severely affected by the cluster environment (Richert et al. 2015). Only the more clustered YSOs will experience significant external photoevaporation. The lack of disks more massive than about $10 M_{\text{Jup}}$ in this subpopulation will potentially affect the formation of planetary systems similar to the solar system, mostly preventing the building-up of gas giant planets.

We thank the constructive suggestions to the anonymous referee. This paper makes use of the ALMA data ADS/JAO. ALMA#2015.1.01323.S. ALMA is a partnership of ESO (representing its member states), NSF (USA) and NINS (Japan), together with NRC (Canada), NSC and ASIAA (Taiwan), and KASI (Republic of Korea), in cooperation with the Republic of Chile. The Joint ALMA Observatory is operated by ESO, AUI/NRAO, and NAOJ. Our color compositions are based on the programs 12050 and 13391 made with the NASA/ESA *HST*, and obtained from the HLA, a collaboration between the STScI/NASA, the ST-ECF/ESA, and the CADK/NRC/CSA. A.M.D. acknowledges support from the FONDECYT project 3140383 and thanks H. Cánovas the stimulating discussions. L.Z. and W.J.H. acknowledge

support from CONACyT and DGAPA-UNAM, Mexico, through grants CB152913 and PAPIIT IN111215. T.H.P. acknowledges support from the FONDECYT Regular grant no. 1121005.

Facility: ALMA

REFERENCES

- Acke, B., van den Ancker, M. E., Dullemond, C. P., van Boekel, R., & Waters, L. B. F. M. 2004, *A&A*, **422**, 621
- Adams, F. C., Hollenbach, D., Laughlin, G., & Gorti, U. 2004, *ApJ*, **611**, 360
- ALMA, Partnership, Brogan, C. L., Pérez, L. M., et al. 2015, *ApJL*, **808**, L3
- Anderson, K. R., Adams, F. C., & Calvet, N. 2013, *ApJ*, **774**, 9
- Andrews, S. M., Rosenfeld, K. A., Kraus, A. L., & Wilner, D. J. 2013, *ApJ*, **771**, 129
- Andrews, S. M., & Williams, J. P. 2005, *ApJ*, **631**, 1134
- Andrews, S. M., & Williams, J. P. 2007, *ApJ*, **671**, 1800
- Ascenso, J., Alves, J., Vicente, S., & Lago, M. T. V. T. 2007, *A&A*, **476**, 199
- Blaauw, A. 1964, *ARA&A*, **2**, 213
- Carpenter, J. M., Ricci, L., & Isella, A. 2014, *ApJ*, **787**, 42
- Cieza, L., Williams, J., Kourkchi, E., et al. 2015, *MNRAS*, **454**, 1909
- Clarke, C. J. 2007, *MNRAS*, **376**, 1350
- Eisner, J. A. 2012, *ApJ*, **755**, 23
- Eisner, J. A., & Carpenter, J. M. 2006, *ApJ*, **641**, 1162
- Eisner, J. A., Plambeck, R. L., Carpenter, J. M., et al. 2008, *ApJ*, **683**, 304
- Gorti, U., Liseau, R., Sandor, Z., & Clarke, C. 2016, *SSRv*, doi:10.1007/s11214-015-0228-x
- Henney, W. J., & O'Dell, C. R. 1999, *AJ*, **118**, 2350
- Hester, J. J., Scowen, P. A., Sankrit, R., et al. 1996, *AJ*, **111**, 2349
- Hildebrand, R. H. 1983, *QJRAS*, **24**, 267
- Hubickyj, O., Bodenheimer, P., & Lissauer, J. J. 2005, *Icar*, **179**, 415
- Johnstone, D., Hollenbach, D., & Bally, J. 1998, *ApJ*, **499**, 758
- Lada, C. J., & Lada, E. A. 2003, *ARA&A*, **41**, 57
- Lissauer, J. J., Hubickyj, O., D'Angelo, G., & Bodenheimer, P. 2009, *Icar*, **199**, 338
- Mann, R. K., Di Francesco, J., Johnstone, J., et al. 2014, *ApJ*, **784**, 82
- Mann, R. K., & Williams, J. P. 2010, *ApJ*, **725**, 430
- McCaughrean, M. J., & Andersen, M. 2002, *A&A*, **389**, 513
- McMullin, J. P., Waters, B., Schiebel, D., Young, W., & Golap, K. 2007, in *ASP Conf. Ser. 376, Astronomical Data Analysis Software and Systems XVI*, ed. R. A. Shaw, F. Hill, & D. J. Bell (San Francisco, CA: ASP), 127
- Mendigutía, I., Mora, A., Montesinos, B., et al. 2012, *A&A*, **543**, A59
- Najita, J. R., & Kenyon, S. J. 2014, *MNRAS*, **445**, 3315
- Natta, A., Grinin, V., & Mannings, V. 2000, in *Protostars and Planets IV*, ed. V. Mannings, A. P. Boss, & S. S. Russell (Tucson, AZ: Univ. Arizona Press), 559
- O'Dell, C. R., Wen, Z., & Hu, X. 1993, *ApJ*, **410**, 696
- Ossenkopf, V., & Henning, T. 1994, *A&A*, **291**, 943
- Owen, J. E., Clarke, C. J., & Ercolano, B. 2012, *MNRAS*, **422**, 1880
- Povich, M. S., Smith, N., Majewski, S. R., et al. 2011a, *ApJS*, **194**, 14
- Povich, M. S., Townsley, L. K., Broos, P. S., et al. 2011b, *ApJS*, **194**, 6
- Preibisch, T., Hodgkin, S., Irwin, M., et al. 2011a, *ApJS*, **194**, 10
- Preibisch, T., Ratzka, T., Kuderna, B., et al. 2011b, *A&A*, **530**, A34
- Reiter, M., & Smith, N. 2013, *MNRAS*, **433**, 2226
- Richert, A. J. W., Feigelson, E. D., Getman, K. V., & Kuhn, M. A. 2015, *ApJ*, **811**, 10
- Sahai, R., Güsten, R., & Morris, M. R. 2012, *ApJL*, **761**, L21
- Simón-Díaz, S., Herrero, A., Esteban, C., & Najarro, F. 2006, *A&A*, **448**, 351
- Smith, N., Bally, J., & Morse, J. A. 2003, *ApJL*, **587**, L105
- Smith, N., Bally, J., & Walborn, N. R. 2010, *MNRAS*, **405**, 1153
- Smith, N., & Brooks, K. J. 2008, in *Handbook of Star Forming Regions*, Vol. II, ed. B. Reipurth (San Francisco, CA: ASP), 138
- Tsamis, Y. G., Flores-Fajardo, N., Henney, W. J., Walsh, J. R., & Mesa-Delgado, A. 2013, *MNRAS*, **430**, 3406
- Weidenschilling, S. J. 1977, *Ap&SS*, **51**, 153
- Williams, J. P., & Cieza, L. A. 2011, *ARA&A*, **49**, 67
- Williams, J. P., Cieza, L. A., Andrews, S. M., et al. 2013, *MNRAS*, **435**, 1671
- Wright, N. J., Drake, J. J., Drew, J. E., et al. 2012, *ApJL*, **746**, L21
- Zeidler, P., Preibisch, T., Ratzka, T., Roccatagliata, V., & Petr-Gotzens, M. G. 2016, *A&A*, **585**, A49