

Models of AM Canum Venaticorum star outbursts

I. Kotko¹, J.-P. Lasota^{1,2}, G. Dubus^{3,2}, and J.-M. Hameury⁴¹ Astronomical Observatory, Jagiellonian University, ul. Orla 171, 30-244 Kraków, Poland
e-mail: I.Kotko@oa.uj.edu.pl² Institut d'Astrophysique de Paris, UMR 7095 CNRS, UPMC Univ Paris 06, 98bis Bd Arago, 75014 Paris, France³ UJF-Grenoble1/ CNRS-INSU, Institut de Planétologie et d'Astrophysique de Grenoble (IPAG) UMR 5274, 38041 Grenoble, France⁴ Observatoire de Strasbourg, CNRS/Université de Strasbourg, 11 rue de l'Université, 67000 Strasbourg, France

Received 2 March 2012 / Accepted 24 May 2012

ABSTRACT

Context. Outbursting AM CVn stars exhibit outbursts similar to those observed in different types of dwarf novae. Their light-curves combine the characteristic features of SU UMa, ER UMa, Z Cam, and WZ Sge-type systems but also show a variety of properties never observed in dwarf novae. The compactness of AM CVn orbits and their unusual chemical composition make these systems valuable testbeds for outburst models.

Aims. We aim for a better understanding of the role of helium in the accretion disc instability mechanism, testing the model for dwarf novae outbursts in the case of AM CVn stars, and aim to explain the outburst light-curves of these ultra-compact binaries.

Methods. We calculated the properties of the hydrogen-free AM CVn stars using our previously developed numerical code adapted to the different chemical composition of these systems and supplemented with formulae accounting for mass transfer rate variations, additional sources of the disc heating, and the primary's magnetic field.

Results. We discovered how helium-dominated discs react to the thermal-viscous instability and were able to reproduce various features of the outburst cycles in the light-curves of AM CVn stars.

Conclusions. The AM CVn outbursts can be explained by the suitably adapted dwarf-nova disc instability model but, as in the case of its application to hydrogen-dominated cataclysmic variables, one has to resort to additional mechanisms to account for the observed superoutbursts, dips, cycling states, and standstills. We show that the enhanced mass-transfer rate, due presumably to variable irradiation of the secondary, must not only be taken into account but is a determining factor that shapes AM CVn star outbursts. The cause of the variable secondary's irradiation has yet to be understood; the best candidate is the precession of a tilted/warped disc.

Key words. accretion, accretion disks – instabilities – binaries: close – stars: dwarf novae

1. Introduction

AM CVn stars (AM CVns) are binary systems with very short orbital period (10–65 min) in which a He or C/O white dwarf primary accretes matter lost by a Roche-lobe filling secondary star, which is believed to be also a helium white dwarf¹ (Nelemans 2005). The observed optical spectra show that accretion discs in these systems are helium-dominated and hydrogen-free. This spectral feature is indeed the criterion for including an ultra-compact binary to the AM CVn class. AM CVns are observed in three distinct luminosity states (see e.g. Bildsten et al. 2006). Two of these are the persistent-luminous (high) and the outbursting states, which are similar to those seen in cataclysmic variable stars (CVs) in which the secondary is hydrogen-rich and usually non-degenerate. The third state, during which the system is persistent and faint, does not have an equivalent among CVs. Smak (1983) remarked that the disc instability model (DIM; see Lasota 2001, for a review and references) which describes dwarf novae (DN) also applies to AM CVn stars and can explain the properties of AM CVn itself (high state) and of GP Com (faint state). Because the instability in question is triggered by opacity changes, the differences between AM CVns and CVs are supposed to be explained mainly by the higher ionization potential

of helium compared to that of hydrogen (see Cannizzo 1984; Tsugawa & Osaki 1997; Lasota et al. 2008). However, as we show in the present paper, other effects related to the helium atomic structure as well as the presence of metals may also influence the outburst cycle properties.

In this context it is also worth stressing that angular momentum transport in discs depends on the microphysics and may be significantly affected by the chemical composition (e.g. the chemical composition influences the Prandtl number, which affects the development of the magneto-rotational instability). One of the methods to test the disc angular momentum transport mechanisms is outburst modelling.

Although still poor compared to that of CVs, the AM CVns observation database has been enriched by recent observations and we know more about these binaries than we did in 1997. Then, in their pioneering work, Tsugawa & Osaki (1997) used a highly simplified description of the DIM (Osaki 1989) and not the full DIM model (such as e.g. Ichikawa & Osaki 1992).

These facts as well as the desire to apply our version of the DIM for helium-rich accretion discs (Lasota et al. 2008; Kotko et al. 2010) to AM CVn outbursts has been the main motivation for writing the present article. In Sect. 2 we discuss the general properties of AM CVn outbursts, concentrating on those we are trying to explain in the following parts of the article. Section 3 describes the DIM in the context of helium-dominated accretion discs. Section 4 briefly recalls what is known about the chemical

¹ The nature of the two shortest orbital period, (X-ray pulsating) systems HM Cnc and V407 Vul remains a mystery.

composition of AM CVn discs and how they comply with the stability criteria set by the DIM. In Sect. 5 we investigate the properties of helium-rich disc models in general and in Sects. 6 and 7 we apply them to various types of AM CVn star outbursts. We discuss our results and summarize the conclusions in Sect. 8.

2. Outburst properties of AM CVn stars

From among 27 systems classified as AM CVn binaries, 11 have been observed to produce outbursts (Levitan et al. 2011a; Ramsay et al. 2012). Because of their relative faintness (~ 14 – 20 mag) the light-curves of these systems are not as well covered and sampled as those of the longer period, hydrogen-dominated CVs. Based on a very incomplete (but soon to be increased by observational campaigns such as the Palomar Transient Factory, Levitan et al. 2011b) AM CVn database one can conclude that most (all ?) outbursting helium-rich systems show superoutbursts or, to be more precise, outbursts that resemble superoutbursts observed in hydrogen-rich dwarf novae such as SU UMa stars (Warner 2003). The most important similarity is the presence of superhumps – photometric humps with periods slightly different from that of the orbital modulation. The amplitudes of these outbursts range from 3.5 to 6 mag; the recurrence times from ~ 45 to ~ 450 days, but this latter parameter has been determined only for five systems with orbital periods between 24.52 and 28.32 min. AM CVns with longer orbital periods are believed (e.g. Levitan et al. 2011a) to have recurrence times longer than three months but this still requires confirmation. One should also stress that in two cases a substantial variation of the recurrence time has been observed: in CR Boo from 46.3 to 14.7 days (Kato et al. 2001b) and in KL Dra from ~ 65 to ~ 44 days (Ramsay et al. 2010, 2012). The superoutbursts of AM CVn stars typically last about 20 days but can be as short as 9 days. On the other hand, in some cases (see e.g. Patterson et al. 1997, 2000) superoutbursts are prolonged by a series of low-amplitude, frequent outbursts: the so-called *cycling state* that may last as long as the “smooth” superoutburst itself. Often the superoutburst is in fact not so smooth because it contains a fairly pronounced dip (Ramsay et al. 2010, 2012). The relation, if any, of these dips to the cycling state and of these two features to the “dipping state” observed in the light-curves of some of WZ Sge-type stars has not been satisfactorily elucidated yet.

Much less is known about the so-called “normal outbursts” that are common in hydrogen-rich dwarf novae, but rarely observed in outbursting AM CVn stars.

From the observational point of view, “normal” outbursts are not very well defined but they are supposed to correspond to the quite narrow eruptions observed in U Gem-type dwarf novae. They are defined in opposition to the superoutbursts whose most characteristic feature are superhumps. Normal outbursts could therefore be defined by the lack of superhumps (but in some cases normal outbursts following a superoutburst do show superhumps).

In the case of CVs, no short period (with orbital periods below the 2–3 h “period gap”) dwarf nova shows a pure normal outburst cycle. The SU UMa stars show a supercycle in which normal outbursts appear between superoutbursts, and binaries of the WZ Sge-type show superoutbursts only. The shortness of period seems to be the decisive factor determining the presence of superoutbursts² which is confirmed by the fact that all AM CVns exhibit these types of outbursts.

² How short is the required period is subject to debate since U Gem, at about 4 h orbital period, showed a superoutburst (Mason et al. 1988; Smak & Waagen 2004; Smak 2006; Schreiber 2007).

Noticing that in CR Boo and V803 Cen the outbursts seen during the cycling state satisfy the empirical Kukarkin-Parenago relation, connecting the amplitude A_n with the recurrence time T_n (see Kotko & Lasota 2012, and Sect. 6.1), Patterson et al. (1997, 2000) concluded that they can be qualified as normal outbursts. On the other hand, Kato et al. (2001b) considered this state as equivalent to the standstill in Z Cam type dwarf novae. Until now the only clear case for relatively well sampled normal outbursts has been provided by observations of PTF1 J0719+4858 (Levitan et al. 2011a). In this case the amplitudes are ~ 2.5 mag, the duration ~ 1 day and the recurrence time about 10 days. In this system normal outbursts appear between superoutbursts as in SU UMa stars. Levitan et al. (2011a) suggested that AM CVn with orbital periods shorter than ~ 27 min are equivalent to SU UMa stars while helium-rich systems with longer periods would correspond to long recurrence time WZ Sge stars showing only superoutbursts. Although not implausible, this hypothesis has to await a more complete observational coverage of AM CVn binaries. Let us just mention that a complete analogy between dwarf novae and their helium-rich cousins might be difficult to establish since e.g. the short-period KL Dra seems to exhibit only (short-recurrence period) superoutbursts (see, however, Ramsay et al. 2012).

2.1. The DIM for AM CVn outbursts: preliminary considerations

A model accounting for the properties of the three categories of AM CVn stars – permanently bright, permanently faint, and outbursting has to take into account the two main differences between these systems and the hydrogen-dominated CVs: the helium-dominated accretion discs and shorter orbital periods, that is, shorter (less extended) accretion discs. As noticed by Smak (1983) the obvious candidate here is the DIM that is used to explain outbursts of hydrogen-rich dwarf novae. It might seem that the adaptation of this model to the case of AM CVn is pretty straightforward because it should consist only in rescaling to helium-rich discs the critical values defining the instability strip. This simplistic view assumes, however, that the DIM is fully successful in explaining most of the dwarf nova outburst properties. Unfortunately this is not the case.

First, the standard DIM does not account for superoutbursts, which is particularly embarrassing for AM CVn stars in which this type of outburst is observed most frequently. Two modifications of the DIM have been proposed to account for superoutbursts of dwarf novae. Osaki (1989) (see also Osaki 1996, for a review) proposed that superoutbursts in SU UMa stars are produced by a tidal-thermal instability (TTI) triggered by a large enhancement of the tidal torque acting on the disc. Despite successes in describing some of the superoutburst properties, the TTI model has encountered some insurmountable difficulties when trying to agree with observations (see e.g. Buat-Ménard & Hameury 2002; Schreiber et al. 2004; Hameury & Lasota 2005, and references therein). Recently, in a remarkable series of articles Smak (2009a,b,c,d) has undermined the basis of the TTI model by casting doubt on the presence of an eccentric disc in the SU UMa stars and providing strong observational arguments in favour of a different model: the enhanced mass-transfer (EMT) model, originally proposed by Osaki (1985) and developed by Lasota et al. (1995), Buat-Ménard & Hameury (2002), Schreiber et al. (2004) and Hameury & Lasota (2005). According to the EMT model: *superoutbursts are due and begin with a major enhancement in the mass transfer rate. During the “flat-top” part of the superoutburst the mass transfer rate*

decreases slowly, causing the observed luminosity to decline. The superoutburst ends when the mass transfer rate decreases below its critical value, resulting in a transition to the quiescent state of the dwarf nova cycle (Smak 2008). In this article we will be testing the EMT superoutburst model for several reasons.

First, there is (mostly indirect) evidence of large variations of the mass-transfer rate at least in some outbursting AM CVn (see e.g. Patterson et al. 2000). Second, as shown by Schreiber et al. (2004), the differences between standard SU UMa stars and ER UMa systems are a natural outcome of the EMT, but not of the TTI model therefore it is unlikely that the latter will be able to describe a system behaving like a WZ Sge, SU UMa, ER UMa, and Z Cam star. Third, since (unchallenged) Smak (2009a) has undermined the very basis of the TTI model for the SU UMa stars it is normal to assume it is also not relevant to AM CVn star outbursts.

Although normal outbursts seem to be infrequent in AM CVn stars, they are the phenomena the DIM is supposed to describe and explain. Since the standard DIM can be tested only on such outbursts, there is an obvious difficulty in applying this model to AM CVns. The recent observation of AM CVn star outbursts that can be clearly identified as normal (Levitan et al. 2011a) promises a more straightforward comparison with the DIM predictions. In particular, it might help solving one important problem with this model: the scaling and re-scaling of the viscosity parameter α . As first noted by Smak (1984a), to obtain dwarf nova outbursts with observed amplitude the viscosity parameter during outburst α_h must be at least four times larger than the quiescent α_c . When implemented, this ad hoc assumption increases the ratio of critical surface densities which in turn allows (because of the S-shaped equilibrium curves) adequately large outburst amplitudes.

As showed by Smak (1999) for dwarf novae, $\alpha_h \approx 0.2$. He arrived at this conclusion in two ways. First, by comparing the observed relation between the normal-outburst decay time and the orbital period ($t_{\text{dec}} - P_{\text{orb}}$) with the model-calculated decay time $t_{\text{dec}} = R_D/v_R$, where R_D is the outer disc radius and v_R the viscous speed in the hot disc. Second, by using the normal-outburst width vs. orbital period ($W - P_{\text{orb}}$) relation to compare with the width of the normal outbursts obtained from the model. The viscous speed is

$$v_R \sim \frac{\nu}{R} \sim \alpha_h c_s^2 v_K^{-1} \sim \alpha_h \frac{\gamma k T_c}{\mu m_H} R^{1/2}, \quad (1)$$

where ν is the kinematic viscosity coefficient, c_s the sound speed, v_K the Keplerian velocity, μ the mean molecular weight, γ the adiabatic index, m_H the hydrogen atomic mass, and T_c is the mid-plane temperature. Therefore the viscous speed in AM CVn discs is half that of hydrogen-dominated CVs (the critical temperatures, for which ionization/recombination of the dominant element in the disc takes place, are twice as high but the molecular weight is four times higher). Since AM CVn discs are shorter (by a factor ~ 2) than those of CVs, decay times in the former should be comparable to those in the latter. The available limited sample seems to confirm that. The correlation found by Smak (1999) between the observed outburst widths $W(d)$ (in days) and the orbital period P_{orb} (in hours) is $W(d) = (2.01 \pm 0.29)P_{\text{orb}}(\text{h})^{(0.78 \pm 0.11)}$. Therefore for $P_{\text{orb}} = 25$ min the normal outburst duration should be ~ 1 , day as observed in PTF 1J0719+4858 by Levitan et al. (2011a). Based on this very limited evidence, one can conclude that the value of the viscosity parameter in hot AM CV discs is close to that of hydrogen-rich dwarf novae (Kotko & Lasota 2012). For convenience we will use the two fiducial values $\alpha_h = 0.1$ and $\alpha_h = 0.2$, hoping that

observations will soon bring more information. The ratio α_h/α_c required to reproduce observed properties of AM CVn outbursts has to be determined through the DIM calculations. This problem will be addressed below. First, we have to establish how the change of chemical composition affects the DIM.

3. The disc instability model for AM CVn stars: application

The model assumes that the disc is geometrically thin, which allows decoupling the radial and vertical structure equations. The radial structure of the disc is described by the equations of mass, angular momentum and energy conservation (see Eqs. (1), (2), (4) in Hameury et al. 1998, hereafter H98) and its vertical structure is assumed to be in hydrostatic equilibrium, which is described by Eqs. (14)–(16) and (23) from H98.

By solving the equations of the disc local vertical structure one obtains the values of effective temperature T_{eff} (or central temperature T_c) and surface density Σ , for which the thermal equilibrium, at a given disc radius R , is calculated. At each R these solutions form on the $T - \Sigma$ plane the well-known S-shaped curve. The lower branch of this curve, terminating at a maximum value of the surface density Σ_{crit}^- , represents stable cold disc equilibria, while the upper branch starting at a critical minimum value of the surface density Σ_{crit}^+ , corresponds to stable, hot equilibrium disc solutions. Solutions on the middle branch are thermally unstable. Their existence is the cause of the dwarf nova outbursts.

For a pure helium disc ($Y = 1$) the critical surface densities (Σ_{crit}^\pm) and the critical effective temperatures (T_{eff}^\pm) can be fitted by the formulae found by Lasota et al. (2008):

$$\Sigma_{\text{crit}}^+ = 528 \alpha_{0.1}^{-0.81} R_{10}^{1.07} M_1^{-0.36} \text{ g cm}^{-2} \quad (2)$$

$$T_{\text{eff}}^+ = 13000 \alpha_{0.1}^{-0.01} R_{10}^{-0.08} M_1^{0.03} \text{ K} \quad (3)$$

$$\Sigma_{\text{crit}}^- = 1620 \alpha_{0.1}^{-0.84} R_{10}^{1.19} M_1^{-0.40} \text{ g cm}^{-2} \quad (4)$$

$$T_{\text{eff}}^- = 9700 \alpha_{0.1}^{-0.01} R_{10}^{-0.09} M_1^{0.03} \text{ K}, \quad (5)$$

whereas for a helium disc with solar metal abundance ($Z = 0.02$, $Y = 0.98$) one has

$$\Sigma_{\text{crit}}^+ = 380 \alpha_{0.1}^{-0.78} R_{10}^{1.06} M_1^{-0.35} \text{ g cm}^{-2} \quad (6)$$

$$T_{\text{eff}}^+ = 11500 \alpha_{0.1}^{-0.01} R_{10}^{-0.08} M_1^{0.03} \text{ K} \quad (7)$$

$$\Sigma_{\text{crit}}^- = 612 \alpha_{0.1}^{-0.82} R_{10}^{1.10} M_1^{-0.37} \text{ g cm}^{-2} \quad (8)$$

$$T_{\text{eff}}^- = 8690 \alpha_{0.1}^{-0.00} R_{10}^{-0.09} M_1^{0.03} \text{ K}, \quad (9)$$

where $\alpha_{0.1}$ is the viscosity parameter in units of 0.1, R_{10} is the disc radius in 10^{10} cm and M_1 the white dwarf primary's mass in solar units (see Appendix A for the complete list of critical values, also for $Z = 0.04$).

The critical effective temperatures are independent of the viscosity parameter α (the very weak dependence in some cases results from the imprecision of the fit) because in thermal equilibrium the flux (hence T_{eff}) is independent of the viscosity mechanism (see e.g. Frank et al. 2002). This allows a fairly simple understanding of the S-curve properties.

From mass and angular-momentum conservation equations one obtains

$$\nu \Sigma \sim \sigma T_{\text{eff}}^4. \quad (10)$$

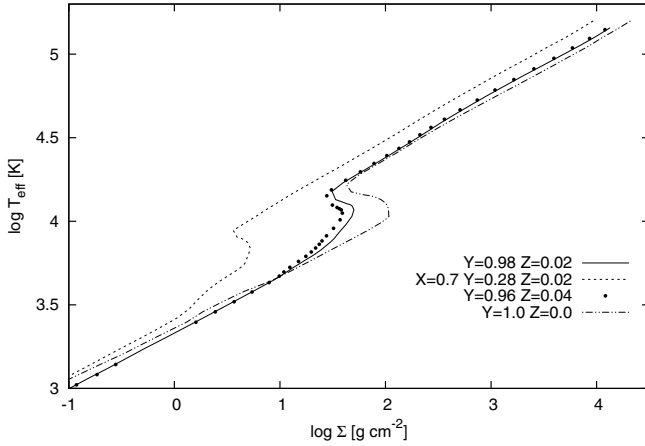


Fig. 1. S-curves for $M_1 = 1.0$, $R = 10^9$ cm, $\alpha = 0.1$, and four chemical compositions.

The kinematic viscosity coefficient is $\nu = 2/3 \alpha c_s H$, where c_s is the speed of sound and H the disc semi-scale-height. Therefore (for a gas-pressure-dominated disc) one has

$$\alpha T_c \Sigma \sim \sigma T_{\text{eff}}^4, \quad (11)$$

which implies that for $T_c \propto \alpha^n$ and $\Sigma \propto \alpha^m$, $n + m = -1$. Equation (11) and the energy transfer equation linking T_c and T_{eff} , determine the slopes of the S-curve's branches $T_{\text{eff}}(\Sigma)$.

In the case of radiative cooling the effective and central temperatures are related through

$$T_c = \left(\frac{3\tau_{\text{tot}}}{8} \right)^{1/4} T_{\text{eff}} \quad (12)$$

where τ_{tot} is the total disc optical depth (see [Dubus et al. 1999](#)). For a fully ionized disc the Rosseland opacity coefficient can be written as $\kappa \sim \Sigma H^{-1} T_c^{-7/2}$ cm²/g from which follows the well-known relation $T_{\text{eff}} \propto \Sigma^{5/14}$ found in the classical [Shakura & Sunyaev \(1973\)](#) solution. On the other hand, a very steep decrease of opacity with temperature induces a change of slope of the $T_{\text{eff}}(\Sigma)$ relation. This is what happens during recombination.

Figure 1 shows S-curves of discs with four chemical compositions: $Y = 1.0$, $Y = 0.98$, $Y = 0.96$, and solar abundance ($X = 0.7$, $Y = 0.28$, $Z = 0.02$). As expected, for all cases the upper branches are parallel (and very close). Since the change of slope at Σ_{crit}^+ is caused by a recombination-induced change in the opacity dependence on temperature, it follows that the upper S-curve bend will be the highest for the element with the highest ionization potential. Hence $\Sigma_{\text{crit}}^+(Y = 1) > \Sigma_{\text{crit}}^+(Y = 0.98) > \Sigma_{\text{crit}}^+(Y = 0.96) > \Sigma_{\text{crit}}^+(\text{solar})$. Cooling by convection, which flattens the temperature profile, affects the position of the upper bend of the S-curve, but not the relative order of the critical surface-densities values.

For the lower bend of the S-curve, the situation is slightly more complicated. According to [Tsugawa & Osaki \(1997\)](#) it appears where the disc becomes optically thin ($\tau \sim 1$) and they explain the higher value of Σ_{crit}^- for helium by its low opacity that must accordingly be compensated for by higher surface density to give $\kappa \Sigma \sim 1$. However, as seen from Table 1, this explanation cannot be correct. As already noticed and explained by [Cannizzo & Wheeler \(1984\)](#); see also [Smak 1999](#) the optical depth at Σ_{crit}^- is *always* (for the standard assumption $\alpha \leq 1$) larger than 1 and the turnover at Σ_{crit}^- is related to the optical depth in a less straightforward way.

Table 1. S-curve critical points for four abundances.

	$Y = 1.0$	$Y = 0.98$	$Y = 0.96$	Solar
Σ_{crit}^+ [g/cm ²]	42.1	30.8	27.8	3.5
Σ_{crit}^- [g/cm ²]	108.0	50.2	39.0	5.2
T_{eff}^- [K]	10 700	11 700	11 200	6960
T_{eff}^+ [K]	15 600	14 250	14 200	8210
$\kappa(\Sigma, T_c^-)$ [cm ² /g]	4.2×10^{-1}	4.5×10^2	5.4×10^2	1.3×10^3
$\tau_{\text{tot}}(\Sigma, T_c^-)$	45.3	2.24×10^4	2.11×10^4	7.03×10^3

Notes. $R = 1.0 \times 10^9$ cm, $\alpha = 0.1$.

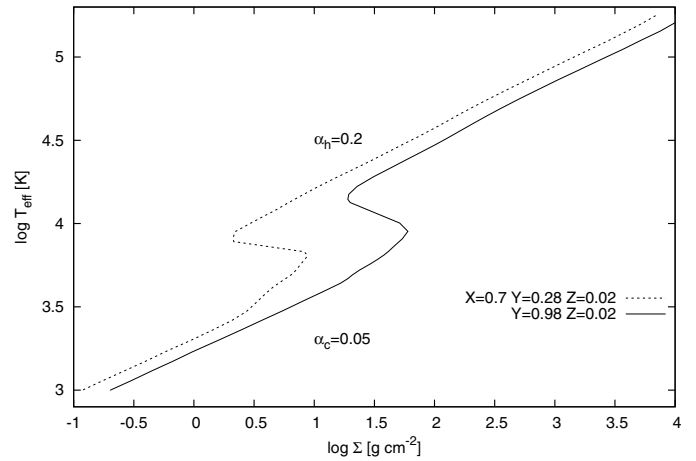


Fig. 2. “Effective” S-curves for combined cold and hot branches corresponding to two values of α : $\alpha_c = 0.05$ and $\alpha_h = 0.2$. The curve on the left corresponds to solar composition, while that on the right corresponds to $Y = 0.98$. $M_1 = 1.0$, $R = 1.0 \times 10^9$ cm.

From Table 1 one can see that at critical points the relation Eq. (12) is approximately satisfied only for $Y = 1$. For the other two cases the temperature gradient is much flatter than predicted by Eq. (12), reflecting the importance of convective energy transport, consistent with the high optical depths of these configurations. Therefore for $Y = 1$ (low opacity) the turnover at Σ_{crit}^- results from the change in opacity temperature-dependence, whereas for the high optical depth cases of $Z = 0.02$ and $Z = 0.04$ the change of slope results from the (opacity-related) change in importance of convective energy transport.

Both (maximum and minimum) critical surface-density values are higher in helium dominated discs than in discs where hydrogen dominates. However, as has been realized by [Smak \(1984a\)](#), the S-curves used in the DIM must be modified if the model is to reproduce the outburst properties of dwarf novae. The viscosity parameter α cannot be constant during outburst, it must be larger in hot discs than in cold, quiescent discs. The DIM uses “effective” S-curves by combining configurations with two different values of α . Observations of dwarf nova normal outbursts imply that the ratio between the two is between 4 and 10. Examples of effective S-curves are shown in Fig. 2 for an α ratio equal 4.0. This ratio has been assumed to be the same for a hydrogen- and helium-dominated disc. Since the α variation used is purely arbitrary, gauged through observed dwarf nova light curves, this assumption is not necessarily valid. We will discuss this problem in more detail in Sect. 5.2.

4. Mass transfer rates and stability

The first to consider the stability of helium accretion discs in AM CVn systems was [Cannizzo \(1984\)](#). Here we use the latest Opal opacity tables and a larger statistics of AM CVn stars to test the stability criteria versus the observed properties of these systems. As discussed above, the critical parameters defining the disc's stability depend on the chemical composition.

4.1. Chemical composition of accretion discs in AM CVn stars

In general, observations clearly show that discs in AM CVn stars are not made of pure helium. We briefly describe the observed AM CVn spectra to show what their disc metallicity might be.

The two persistently faint systems – GP Com and V396 Hya – are very special. Unlike the other (outbursting or persistently bright) AM CVn stars, they show a significant overabundance of NV, in addition to a strong HeI and a weaker HeII line, but not the Si line, that is usually seen in other binaries of this type. This suggests metal poor secondaries. Other binaries in a low state, such as SDSS J0902 or SDSS J1552 also show weak SiII and FeII lines.

In the permanent high-state systems AM CVn or HP Lib the spectrum is dominated by the absorption lines, but sometimes emission in HeII is also detected. In the optical part of the spectrum the asymmetrical, broad absorption line of HeI dominates. The UV spectrum shows broad absorption lines of HeII, NV, NIV, SiIV, CIV ([Wade et al. 2007](#)). The spectra of the erupting systems, CR Boo and V803 Cen have similar characteristics during outbursts.

For GP Com [Strohmayer \(2004\)](#) estimated $X_{\text{He}} = 0.99$ and $Z = 0.01$. Because GP Com and V396 Hya are the two metal-poor systems, we can consider $Z = 0.01$ as lower limit for an AM CVn disc metallicity. However, the exact metallicities of particular systems have not been well determined yet ([Nelemans et al. 2010](#), and priv. comm.).

Throughout, our fiducial model of outbursting AM CVn stars assumes $Y = 0.98$, $Z = 0.02$, but when discussing the general properties of the systems and the model we will also consider other possibilities.

4.2. Stability of observed systems

In the standard version of the DIM the mass transfer rate \dot{M}_{tr} from the secondary is assumed to be constant. Its value determines the stability of the disc.

For pure helium discs the critical values of the accretion rate are ([Lasota et al. 2008](#))

$$\dot{M}_{\text{cr}}^+ = 1.01 \times 10^{17} \alpha_{0.1}^{-0.05} R_{10}^{2.68} M_1^{-0.89} \text{ g s}^{-1} \quad (13)$$

and

$$\dot{M}_{\text{cr}}^- = 3.17 \times 10^{16} \alpha_{0.1}^{-0.02} R_{10}^{2.66} M_1^{-0.89} \text{ g s}^{-1}. \quad (14)$$

To be in a hot (cold) stable equilibrium the accretion rate in a disc must be higher (lower) *everywhere* than the corresponding critical \dot{M} . Therefore the stability conditions are

- a stationary ($\dot{M}(R) = \text{const.}$) disc is hot and stable when $\dot{M}_{\text{tr}} > \dot{M}_{\text{cr}}^+(R_{\text{D}})$;
- a stationary ($\dot{M}(R) = \text{const.}$) disc is cold and stable when $\dot{M}_{\text{tr}} < \dot{M}_{\text{cr}}^-(R_{\text{in}})$.

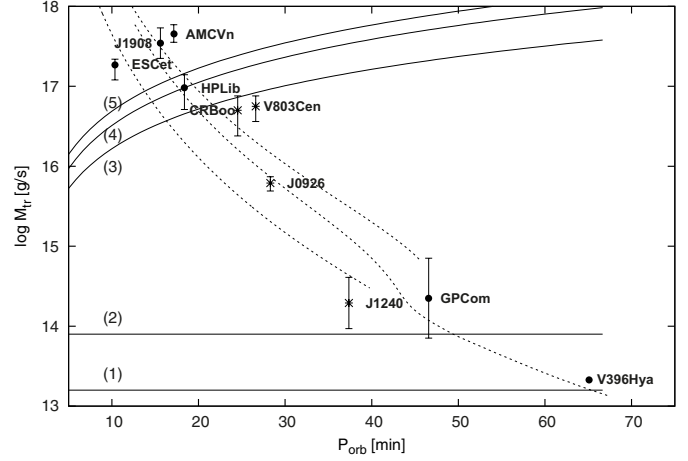


Fig. 3. The $\log \dot{M}_{\text{tr}} - P_{\text{orb}}$ plane. *Dots* – persistent AM CVn systems with known distances; *asterisks* – outbursting AM CVn systems with known distances; *solid lines* – the upper and lower critical limits \dot{M}_{tr} : (1) \dot{M}_{cr}^- for $M_1 = 1.0$, $Y = 1.0$; (2) \dot{M}_{cr}^- for $M_1 = 0.6$, $Y = 1.0$; (3) \dot{M}_{cr}^+ for $M_1 = 1.0$, $Y = 0.96$, $Z = 0.04$; (4) \dot{M}_{cr}^+ for $M_1 = 1.0$, $Y = 1.0$; (5) \dot{M}_{cr}^+ for $M_1 = 0.6$, $Y = 1.0$. (M_1 – in solar units). *Dotted lines*: evolution models for AM CVns through the WD channel (kindly provided by Chris Deloye).

The critical rate \dot{M}_{cr}^- for helium is 12 times higher than for hydrogen-dominated discs consequently the existence of cold stable hydrogen-deficient discs does not require ridiculously low mass-transfer rates as in the case of hydrogen-rich CVs, where it has to be lower than $8 \times 10^{12} \text{ g/s}$. The outer disc radius, R_{D} , is determined from the equation

$$\frac{R_{\text{D}}}{a} = \frac{0.6}{1+q}, \quad (15)$$

([Warner 2003](#)), where a is the binary separation and $q = M_2/M_1$ the mass ratio. Since in AM CVn stars $0.0125 \leq q \leq 0.18$ the dependence of the maximal disc radius on mass ratio is rather weak and can be safely neglected.

The inner radius, R_{in} , is taken to be equal to the radius of the central white dwarf, $R_{\text{in}} = R_{\text{WD}}$, which is determined by the white dwarf mass M_{WD} through the $M - R$ relation ([Nauenberg 1972](#)). The dependence on α in Eqs. (13) and (14) is clearly negligible. The plot of the estimated \dot{M}_{tr} versus P_{orb} for observed AM CVns (see Table 2) compared with the model predictions is shown in Fig. 3. In agreement with the model, the permanently high-state systems ES Cet, SDSS J1908+3940 and AM CVn lie well above the critical mass transfer limit. Another bright and steady system is HP Lib. Although it is very close to the critical line of \dot{M}_{cr}^+ , there still exists a certain set of \dot{M}_{cr}^+ – lines above which HP Lib lies with reasonable parameters. The mass-transfer rates of CR Boo and V803 Cen are very close to the upper critical line, which might explain why these two outbursting systems have been alternately and confusedly classified as analogues of ER UMA, Z Cam, and (not consistent with their mass-transfer rate) of WZ Sge stars. Below we test the hypothesis that they are “superoutbursting” Z Cam-type stars. GP Com and V396 Hya have shown no changes in their brightness till now, they are considered to be low-state systems. From Eq. (14) one concludes that their primary masses M_1 should lie in the range $0.6 - 1.0 M_{\odot}$ (the mass-transfer rate for GP Com is an upper limit, see Table 2).

Table 2. Properties of AM CVns with known orbital periods and distances.

System	P_{orb} (min)	$M_1(M_{\odot})$	$M_2(M_{\odot})$	$\log \dot{M}_{\text{tr}}$ [g/s]	Ref.
ES Cet*	10.35	0.44–0.69	0.062–0.26	$17.27^{+0.07}_{-0.19}$	1, 5
SDSS J1908+3940*	$\sim 15.6^{\text{b}}$?	?	17.35–17.73	7
AM CVn*	17.15	0.71 ± 0.07	0.13 ± 0.01	$17.65^{+0.12}_{-0.1}$	2
HP Lib	18.38	0.80–0.4	0.048–0.088	$16.98^{+0.17}_{-0.27}$	2
CR Boo	24.52	1.10–0.67	0.048–0.088	$16.7^{+0.18}_{-0.32}$	2
V803 Cen	26.6	1.17–0.78	0.059–0.109	$16.75^{+0.13}_{-0.19}$	2
SDSS J0926+3624	28.3	0.85 ± 0.04	0.035 ± 0.003	$15.79^{+0.08}_{-0.1}$	4, 6
SDSS J1240-01	37.36	?	?	$14.29^{+0.32}_{-0.32}$	3
GP Com*	46.57	0.50–0.68	0.009–0.012	$\leq 14.35^{+0.5^{\dagger}}_{-0.5}$	2
V396 Hya*	65.1	?	?	13.33	3

Notes. (*) Persistent systems; (\dagger) upper limit; ($^{\text{b}}$) superhump period (?).

References. (1) [Espaillet et al. \(2005\)](#); (2) [Roelofs et al. \(2007a\)](#); (3) [Bildsten et al. \(2006\)](#); (4) [Deloye et al. \(2007\)](#); (5) [Copperwheat et al. \(2011b\)](#); (6) [Copperwheat et al. \(2011a\)](#); (7) [Fontaine et al. \(2011\)](#). See also [Nelemans \(2010\)](#).

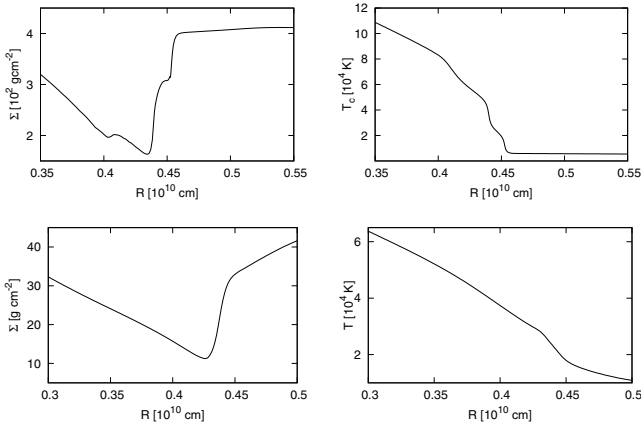


Fig. 4. Cooling front structure: *top, left*: $Y = 1$, Σ profile; *top, right*: $Y = 1$, T_c profile; *bottom, left*: solar, Σ profile; *bottom, right*: solar, T_c profile. Parameters used for calculation are $\alpha_c = 0.05$, $\alpha_h = 0.2$, $M_1 = 1.0$, $\dot{M}_{\text{tr}} = 1.0 \times 10^{15}$ g/s, $\langle R_D \rangle = 1.2 \times 10^{10}$ cm.

5. Modelling outbursts of helium-dominated accretion discs

5.1. The front structure

The structure of heating and cooling fronts is affected by the atomic structure of the dominant elements. Helium has two ionization energy levels: $E_{\text{HeII}} = 24.6$ eV which corresponds to temperature $T_{\text{HeII}} \sim 28\,500$ K for the ionization of the first electron, and $E_{\text{HeIII}} = 54.4$ eV, i.e. $T_{\text{HeIII}} \sim 63\,000$ K, for the ionization of the second electron. The existence of two electrons in the helium atom also opens a wide range of states where one electron is excited and the second one is ionized, or both are excited to various energy levels.

For solar composition discs, a detailed analysis of the cooling and heating fronts has been performed by [Menou et al. \(1999\)](#). They found that the form of the fronts does not depend on viscosity and front location in the disc. In hydrogen-dominated discs the temperature gradients of both heating and cooling front are quite smooth (the bottom row of Fig. 4 shows surface density (left) and central temperature (right) profiles in solar-composition disc), showing no clear breaks. In

comparison, fronts in helium discs exhibit additional features (top row of Fig. 4). The breaks in the slope of the temperature gradient appear at roughly the same temperatures for the heating and cooling fronts. One can identify these temperatures with atomic transitions in helium. The first break appears for $T = 80\,000$ – $92\,000$ K (69–80 eV) and can be assigned to double-photoionization and different negative-ion resonances. The break at $T \sim 50\,000$ K (~ 42 eV) corresponds to the state in which the atom is singly ionized and the second electron is in excited state $n = 2$. In the region of temperatures around $T \sim 28\,000$ K (~ 24 eV) the neutral helium becomes singly ionized. Below this is a range of temperatures at which one electron remains in the ground state while second one is in the excited state. Other points marked on T and Σ profiles are hard to classify unambiguously.

Yet more changes with the addition of metals to helium. We consider here discs with $Y = 0.98$, $Z = 0.02$ and $Y = 0.96$, $Z = 0.04$. At the temperatures available in these discs, elements such as oxygen or carbon become highly ionized. Despite low number fractions, their contribution to the population of free electrons, and in turn their impact on the opacities in the disc, is significant. For instance, the fourth ionization energy of O is $E_{\text{OV}} = 77.7$ eV ($T \sim 82\,000$ K), the fourth ionization energy of C is $E_{\text{CV}} = 64.7$ eV ($T \sim 68\,000$ K) and that of N (thought to be enriched in some AM CVn stars) is $E_{\text{NV}} = 77.47$ eV ($T \sim 81\,760$ K). Because the main opacity sources in the disc are free-free and bound-free transitions (electron scattering, however, is negligible), the additional electrons change the efficiency of the cooling and heating mechanisms. This influences the maximum and minimum temperatures in the disc (see Table 1). N enrichment does not contribute more than other metals to the opacity growth. For solar-composition discs we do not see the features from atomic transitions of helium and metals, because fronts there start to propagate at lower temperature and much lower densities (i.e. hydrogen ionization temperature $\sim 3.5 \times 10^4$ K and $\Sigma \sim 30$ g cm $^{-2}$) at which these transition do not take place or are negligible because of the small contribution of metals to the overall disc composition (such as single ionization of carbon at $T \sim 12\,000$ K).

The influence of the chemical composition on the details of the front structure does not manifest itself in the observational outburst cycle but helps understanding the DIM physics.

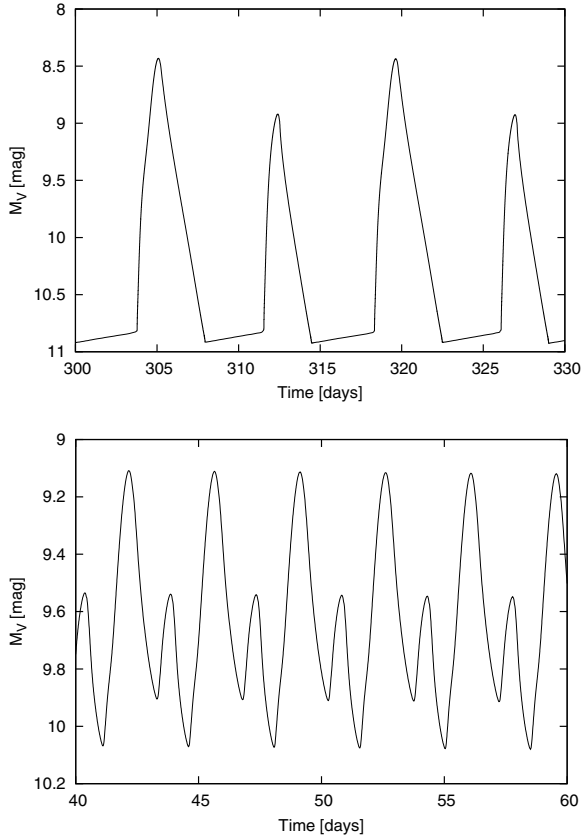


Fig. 5. Light curves calculated within the standard DIM with the same α in the hot and cold state for the discs with two different chemical compositions. *Top:* $Y = 1.0$. *Bottom:* $Y = 0.98$, $Z = 0.02$. Parameters in both cases: $\alpha_h = \alpha_c = 0.1$, $\dot{M}_{tr} = 10^{16}$ g/s, $M_1 = 1.0$ and the mean outer radius is $\langle R_D \rangle = 1.2 \times 10^{10}$ cm.

5.2. Playing with α' s

The critical surface densities depend on α to the power ~ 0.85 so that an α 's ratio of 4, say, corresponds to Σ 's ratio of ~ 3.2 . From Table 1 we can see that this is just the ratio of critical densities for a pure helium disc with $\alpha_h = \alpha_c$. Indeed, for such a disc the DIM produces outbursts with amplitudes up to ~ 2.5 mag with α kept constant (Fig. 5). However, this cannot correspond to real AM CVn outbursts since, as mentioned above, observations show that their discs also contain certain amount of metals. We tested our model for discs with metallicity $Z = 0.0003$, 0.02 , and 0.04 . For $Z \geq 2\%$ one cannot obtain amplitudes larger than 1 mag without changing α . Adding just 2% of metals lowers the amplitude to less than 1 mag (Fig. 5). Nevertheless, as we will see below, in AM CVn the required α jump can be lower than in dwarf novae (see e.g. Fig. 7). One can notice also that changing the chemical composition modifies not only the amplitude but also the shape of the outburst light-curve. This is to be expected because a diminished $\Sigma_{crit}^-/\Sigma_{crit}^+$ ratio favours the appearance of so-called “reflares” (see e.g. Menou et al. 2000; Dubus et al. 2001; Lasota 2001).

The accretion rate at outburst maximum is $\dot{M}_{peak} \sim (\alpha_h/\alpha_c)^{8/7}$ (see e.g. Dubus et al. 2001), hence the necessity of increasing the ratio of the viscosity parameters. There is not much to guide us in performing the change of α . Numerical simulations of the magneto-rotational instability (MRI) which is believed to drive turbulence in Keplerian accretion discs (see Balbus 2003, 2005, and references therein), provide no guide in this case. Indeed, the value of α they produce is an order of

magnitude lower than required by observations (e.g. King et al. 2007). Transients appearing in some MRI simulations produce a viscosity coefficient of the correct order of magnitude but they are too short-lived to be the solution of this fundamental problem (Sorathia et al. 2011; Reynolds, priv. comm.).

The choice of the formula matching the two α s is not without consequences because it can affect the properties of the outburst light-curves (see e.g. Cannizzo et al. 2010). In our code we use the following interpolation formula:

$$\log \alpha(T_c) = \log \alpha_c + [\log \alpha_h - \log \alpha_c] \times \left[1 + \left(\frac{T_0}{T_c} \right)^8 \right]^{-1}. \quad (16)$$

The motivation behind this expression is the wish to keep the values of $\Sigma^-(\alpha)$ and $\Sigma^+(\alpha)$ in the effective S curve equal to the “original” Σ_{crit}^- and Σ_{crit}^+ (H98). Much depends, however, on the choice of T_0 . H98 chose $T_0 \approx T_c(\Sigma_{crit})$ while Dubus et al. (2001) used $T_0 = 0.5[T_c(\Sigma_{crit}^-) + T_c(\Sigma_{crit}^+)]$, which is more convenient for the irradiated disc they were considering (irradiation changes the critical values on the hot branch). In the present paper we are using the H98 prescription.

5.3. Additional sources of disc heating

Additional sources of disc heating are known to be important in H-rich CVs, and we need to test their influence on the modelled light-curves for helium discs.

5.3.1. Outer disc heating

As the mass is transferred from the secondary, its stream hits the outer ring of the accretion disc, forming the so-called hot spot. In the impact region the disc annulus of width ΔR_{hs} is heated at a rate Q_i with efficiency η_i (Buat-Ménard et al. 2001a):

$$Q_i(R) = \eta_i \frac{GM_1 \dot{M}_{tr}}{2R_D} \frac{1}{2\pi R_D \Delta R_{hs}} \exp\left(-\frac{R_D - R}{\Delta R_{hs}}\right), \quad (17)$$

where M_1 is the mass of the primary white dwarf, \dot{M}_{tr} is the mass transfer rate from the secondary and R_D is the outer radius of the disc.

When the system is in the high state most of its luminosity comes from the hot accretion disc, especially from the inner parts through dissipation in the boundary layer, but during the low state as much as half of the visible light can come from the hot spot (Smak 2010). The magnitude of this contribution depends on the inclination of the disc: lower inclination means higher luminosity from the impact region.

When the term $Q_i(R)$ is taken into account in energy conservation equation, the heating from the hot spot reduces the critical values of Σ , which facilitates triggering the outside-in outbursts for mass transfer rates lower than in absence of heatings. It also decreases the lower limit of the \dot{M}_{tr} range for which the disc should be unstable, and lifts the quiescence level. In several cases including this term in the equations improves the description of dwarf nova outbursts considerably (see e.g. Buat-Ménard et al. 2001a,b). On the other hand, Smak (2002) expressed doubts concerning the relevance of hot-spot heating, noticing that most of the impact energy is radiated away at the hot spot location.

Another presumed source of the outer disc heating is the action of tidal stresses (see e.g. Buat-Ménard et al. 2001a). In this case Smak (2002) argues that effect is important only during outburst and limited strictly to the outer disc edge.

The two outer disc heating contributions are of the same order of magnitude. In the following, when testing the DIM’s properties, we will use only hot-spot heating (Eq. (17)) because it represents the presumed effect on the outer disc structure.

5.3.2. Irradiation by the hot white dwarf

As shown by Hameury et al. (1999), in dwarf novae irradiation by the hot white dwarf can seriously affect outburst structure and evolution. In addition to stabilizing the innermost disc regions by maintaining their temperature above the ionization temperature, this irradiation has a *destabilizing* effect on the adjacent parts of the flow. There, irradiation reduces the vertical disc temperature gradient which decreases the critical surface density Σ_{crit}^- . In hydrogen-dominated discs $T_c^- \sim 8000$ K while the accreting white dwarf temperature reaches values from 15 000 to 50 000 K (Sion & Godon 2007). Effective temperatures of accreting white dwarfs in AM CVns are expected to be in the same range (Bildsten et al. 2006), and apparently they are (Sion et al. 2011). However, the critical midplane temperatures of helium-dominated discs are in the range $T_c^- \sim 18\,000\text{--}25\,000$ K which, in general, reduces white dwarf irradiation to an unimportant effect that will be neglected below.

5.4. Magnetic field of the primary

Observations of SDSS J080449.49+161624.8, whose optical spectra have properties similar to those of the hydrogen-rich intermediate polars, suggest that the primary’s magnetic field also plays a role in AM CVn stars (Roelofs et al. 2009). If this is the case, the white dwarf magnetic moment will have a non-negligible influence on the inner part of the disc. The magnetic pressure increases steeply with decreasing radius ($P_{\text{mag}} \sim R^{-6}$) and in quiescence it might exceed the gas and ram pressures of the infalling matter up to radius R_M , disrupting the matter flow. During the outburst the mass accretion rate sharply rises and the ram pressure of matter dominates over the magnetic pressure so that the inner edge of the disc approaches the surface of the white dwarf. This mechanism is summarized in the formula for the inner disc radius (Frank et al. 2002):

$$R_{\text{in}} = R_M = 9.8 \times 10^8 \dot{M}_{15}^{-2/7} M_1^{-1/7} \mu_{30}^{4/7} \text{ cm}, \quad (18)$$

where μ_{30} is the magnetic moment in units of 10^{30} G cm³, M_1 is the mass of the primary in solar masses, and \dot{M}_{15} is the mass accretion rate in units of 10^{15} g s⁻¹.

Including formula (18) in the model prolongs the quiescence time and outburst duration. The explanation is the same for small M_1 : when R_{in} is larger, more mass has to be accumulated to cross $\Sigma_{\text{crit}}^-(R)$, therefore, the quiescence time lengthens. But because more mass is now in the disc, its accretion and the decay from the outburst will take longer. The magnetic field does not influence the outburst amplitude.

6. Normal outbursts of AM CVn stars

From the theoretical point of view, normal outbursts of accretion discs are best defined as outbursts in which only the thermal-viscous instability is at play with no significant mass-transfer variations and no decisive role of the tidal interactions (see Lasota 2001, for a detailed description of normal outbursts).

As mentioned before, there exists a clear difficulty with testing the (helium) DIM on AM CVn outbursts since, in these systems, most of the time it is not clear when and if an observed outburst is normal. Patterson et al. (1997, 2000) proposed using the

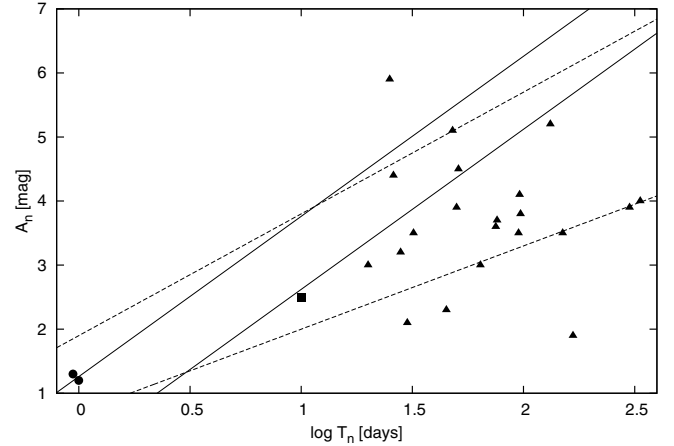


Fig. 6. Amplitudes and recurrence times of normal dwarf nova outbursts and the Kukarkin-Parenago relation. *Triangles*: U Gem stars; *square*: PTF1 J0719+4858 (normal outburst); *circles*: V803 Cen and CR Boo (cycling state rebrightenings). Data from Warner (2003) and Ritter & Kolb (2003). The dotted lines represent the upper and lower uncertainty of the observational KP relation fitted to the systems marked on the plot while the two solid lines represent the theoretical KP relation for helium discs – the lower line has been calculated for $\alpha_n = 0.2$, $R_D = 0.7 \times 10^{10}$ cm, $M_1 = 1.0$ and the upper line for $\alpha_n = 0.2$, $R_D = 1.0 \times 10^{10}$ cm, $M_1 = 1.2$.

empirical relation between the observed outburst amplitude and the recurrence time (the “Kukarkin-Parenago relation”; hereafter “KP relation”) as a test of “normalcy” of outbursts. To evaluate the relevance of this suggestion, we will now briefly discuss this empirical amplitude vs recurrence-time relation and its connection to the DIM (see Kotko & Lasota 2012, for more details).

6.1. The Kukarkin-Parenago relation for AM CVn stars

Kukarkin & Parenago (1934) suggested the existence of a relation between the amplitudes of cataclysmic-variable star outbursts and their recurrence times. As clarified by van Paradijs (1985) this correlation may indeed to be considered to exist if the sample considered is reduced to normal dwarf nova outbursts. Because of the large scatter of the observational data the form of the correlation is sample-dependent and uncertainties are rather large. For a selected subset of U Gem type dwarf novae with well-measured amplitudes and recurrence times (see Table B.1) the correlation has the form³

$$A_n = (1.3 \pm 0.6) + (1.6 \pm 0.3) \log T_n, \quad (19)$$

where T_n is measured in days and A_n in magnitudes.

The only AM CVn system with unambiguous normal outbursts is PTF 1J0719+4858. Their amplitude is $A_n \sim 2.5$ and recurrence time $T_n \sim 10$ d. Figure 6 shows the KP relation and the corresponding amplitudes and recurrence times of selected U Gem stars (marked by triangles). The dotted lines represent the upper and lower uncertainty of the KP relation (Eq. (19)). Also plotted are the amplitudes and recurrence times for PTF 1J0719+4858 and the cycling states of the AM CVn stars V803 Cen and CR Boo. PTF 1J0719+4858 fits the KP relation well. The same is true, as remarked by Patterson et al. (1997, 2000), of the parameters of the cycling states of the two other AM CVns: they are compatible with this relation. However, does

³ For the sample in Warner (2003) the relation is $A_n = (0.7 \pm 0.43) + (1.9 \pm 0.22) \log T_n$.

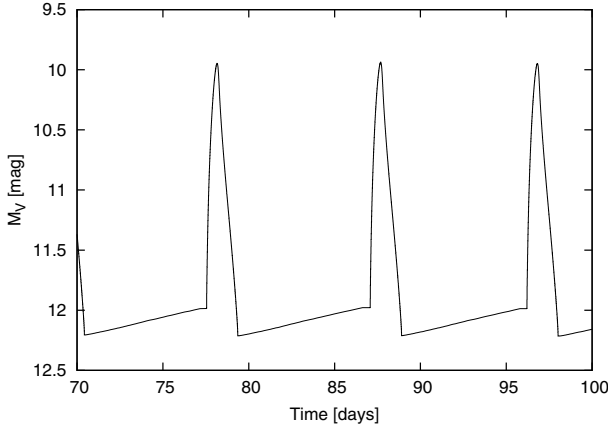


Fig. 7. Normal outburst cycle for $Y = 0.98$, $Z = 0.02$ accretion disc. $\alpha_c = 0.05$, $\alpha_h = 0.1$, $\dot{M}_{16} = 0.1$, $M_1 = 1.0$, $\langle R_D \rangle = 1.1 \times 10^{10}$ cm. The disc extends down to the white dwarf surface. No irradiation of disc or companion is taken into account.

thos imply that they can be classified as normal outbursts? This depends on what the KP relation is supposed to represent.

Using the standard version of the DIM, Kotko & Lasota (2012) derived a theoretical $A_n(T_n)$ relation in the form of

$$A_n \approx C_1 + 2.5 \log T_n, \quad (20)$$

with

$$C_1 \approx 1.5 - 2.5 \log t_{\text{dec}} + BC_{\text{max}} - BC_{\text{min}}, \quad (21)$$

where t_{dec} is (as before) the outburst decay time calculated using the disc size and the viscous speed (see Eq. (1)) and BC_{max} and BC_{min} are the disc bolometric corrections in the high and low states, respectively. The viscosity parameter α_h must be in the range 0.1–0.2 to correspond to observations.

The relation Eqs. (20), (21) is based on the assumption that the whole mass accumulated during the time $t_{\text{quiesc}} + t_{\text{rise}}$ is accreted onto the white dwarf during outburst (see Kotko & Lasota 2012, for details). Normal outbursts are supposed to satisfy this assumption therefore eruptions following the theoretical $A_n(T_n)$ relation can be considered to belong to this category. In Fig. 6 we plotted two theoretical $A_n(T_n)$ relations corresponding to helium discs with parameters $R_D = 0.7 \times 10^{10}$ cm, $M_1 = 1.0$ (PTF 1J0719+4858⁴) and $R_D = 1.0 \times 10^{10}$ cm, $M_1 = 1.2$ (V803 Cen and CR Boo), respectively. In both cases $\alpha_h = 0.2$. With these parameters the amplitudes and recurrence times of helium-dominated discs lie on the theoretical $A_n(T_n)$ relation.

6.2. Model light-curves of normal outbursts

In SU UMa stars the outbursts appearing between the superoutbursts are believed to be of normal type. The case of PTF 1J0719+4858, where a similar sequence is present, suggests that the same is probably true for AM CVn stars.

Figure 7 shows a synthetic light curve produced with the DIM applied to the helium-dominated discs. Although one does not expect to observe the light curve analogous to Fig. 7 among compact helium binaries, the basic DIM should describe this part of the outburst cycle where normal outbursts are present.

We briefly recall how the model’s free parameters α_h , α_c and mass transfer rate \dot{M}_{tr} influence the recurrence times T_n and amplitudes A_n of normal outbursts. The primary mass M_1 , fixing

⁴ The orbital parameters of this system are not yet known, however.

the inner disc radius R_{in} , is often difficult to estimate from observations and then it can be considered a free model parameter as well. The disc radius is constrained by observations, so there is not much freedom left in setting R_D in the model, as long as one intends to refer to the real systems.

The time elapsing between the onsets of two consecutive outbursts extends with increasing ratio β of viscosity parameters: $\beta = \alpha_h/\alpha_c$. Since the fiducial value of α_h is 0.1–0.2, it is α_c which decides about T_n . However, β also controls the value of A_n . Accordingly, the increase of β (by lowering α_c) not only results in longer recurrence times, but also in higher amplitudes. The outbursts will be more frequent in the disc fed with matter at high rates or/and in a system with heavier primary.

High \dot{M}_{tr} is also responsible for the appearance of wide normal outburst, which should not be confused with superoutbursts (see Lasota 2001).

Adding metals to helium increases Σ_{crit}^- , due to the additional non-negligible contribution from highly ionized metals to the opacity. Furthermore, the consequence of augmented Z is the reduction of the total amount of helium and Σ_{crit}^+ drop. The conclusion is that adding elements heavier than He to the disc matter is equivalent to diminishing the Σ ratio (decreasing the β ratio). This results in shorter recurrence times and lower outburst amplitudes.

It is worth emphasizing that metallicity affects the value of α_c , which needs to be “adjusted” to obtain the observed outbursts amplitudes and recurrence times. The difference between α_h and α_c to be explained by MRI may not be as large in a helium disc as usually thought for a solar-composition disc. However, a difference is still required, as argued in Sect. 5.2.

7. AM CVn outburst supercycles

Light curves of AM CVn can be divided into three categories, which we refer to as: KL Dra – like, PTF 1J0719+4858 – like, and CR Boo/V803 Cen – like. The first category exhibits superoutbursts only and might be the equivalent of WZ Sge-type dwarf nova stars⁵. PTF 1J0719+4858 almost certainly corresponds to the SU UMa type, well-known in hydrogen-dominated dwarf novae. The most mysterious are CR Boo and V803 Cen, which have been variedly described as equivalent to WZ Sge, SU UMa, ER UMa, and Z Cam stars. It seems that they belong to all these classes partaking freely of such a wide range of dwarf nova behaviour, to quote Patterson et al. (2000).

As mentioned in Sect. 2.1, in the present context we have chosen to opt for the EMT model as the model explaining the mechanism of superoutbursts appearance.

Although there is ample evidence for both the secondary’s irradiation and the enhanced (during outburst) mass-transfer rate (Smak 2011), the connection between the two is far from being clear. Viallet & Hameury (2007, 2008) have shown that matter heated up by irradiation of the companion’s disc-unshaded surface has cooled down too much by the time it arrives in the vicinity of the L_1 point to affect mass transfer towards the disc. Although the treatment of the cooling is simplified (Smak 2009d), it is unlikely that the correct treatment would modify the result significantly. A solution for this difficulty would be to have direct irradiation of the L_1 region allowed by a tilted/warped disc (Smak 2009e).

⁵ KL Dra might not be really of this type since small amplitude normal could have been missed, but an AM CVn analog of the WZ Sge-type presumably exists.

Below, inspired by previous work, we will attempt to mimic the effect of irradiation on the mass-transfer rate by suitable parametrizations.

7.1. KL Dra and PTF 1J0719+4858 – like outbursts

The general properties of the relatively simple light curves of KL Dra and PTF 1J0719+4858 can be reproduced by using the simple prescription for accretion-irradiation mass-transfer rate increase originally proposed by Hameury et al. (1997):

$$\dot{M}_{\text{tr}} = \max(\dot{M}_{0,\text{tr}}, \gamma \dot{M}_{\text{acc}}), \quad (22)$$

where $0 \leq \gamma \leq 1$. $\dot{M}_{0,\text{tr}}$ corresponds to the “secular” (non-enhanced) mass-transfer rate while \dot{M}_{acc} is the accretion rate onto the primary.

Dealing with better-sampled multi-wavelength light curves such as that of VY Hyi, Schreiber et al. (2004) used a more “refined” prescription using a suitably averaged over time \dot{M}_{acc} but using it here would be exaggerated.

As for hydrogen-dominated dwarf novae (Hameury et al. 2000) the light curve properties depend on the mass-transfer rate, the mass of the primary, and the viscosity parameters, as well as on the assumed white dwarf’s magnetic moment and the value of the parameter γ . In AM CVns one has additionally to fix the chemical composition, i.e. the metallicity which we will assume to be $Z = 0.02$ ($Y = 0.98$) unless stated otherwise. We assumed $\langle R_D \rangle = 1.0 \times 10^{10}$ cm for the average size of the accretion disc.

7.1.1. Superoutbursts only

To obtain a light curve similar to that of KL Dra, we used parameters close to those suggested by Ramsay et al. (2010) for this binary: $\dot{M}_{\text{tr}} = 2 \times 10^{16} \text{ g s}^{-1}$, $M_1 = 0.6 M_{\odot}$. With the choice of the viscosity parameters $\alpha_c = 0.035$, $\alpha_h = 0.2$, choosing $\gamma = 0.8$ and using $\mu_{30} = 1$ (corresponding to a magnetic field $B \sim 1.5 \times 10^3$ G) we obtained the supercycle shown in Fig. 8. The calculated amplitude $A_s \approx 3$ mag, the (super)outburst duration $T_{\text{dur}} \approx 15$ d and the recurrence time $T_{\text{recc}} \approx 57$ d correspond quite well to the observed $A_s \approx 3.5$ mag, $T_{\text{dur}} \approx 14$ d and $T_{\text{recc}} \approx 63$ d.

Including the magnetic field of the primary and the heating by the hot spot was essential for obtaining the light curve in Fig. 8. These two effects stabilize the disc. The magnetic field truncates the inner parts of the disc, increasing $\Sigma_{\text{crit}}^-(R_{\text{in}})$. The disc then has to accumulate more mass to trigger the inside-out outburst. The hot spot heating, in turn, lowers $\Sigma_{\text{crit}}^+(R_D)$ so that more mass has to be accreted before the cooling front can form. This favours triggering of a superoutburst instead of a normal outburst. Both effects contribute to the suppression of normal outbursts. The resulting light curve compares well with the observations (Fig. 8).

In reality, as mentioned above, KL Dra could have normal outbursts that have been missed during observational campaigns. If true, this would mean that one should drop from the disc evolution equations e.g. the hot-spot heating term to obtain normal outbursts. This is what we have done in the next subsection when calculating models of the AM CVn equivalents of SU UMa stars.

7.1.2. Helium “SU UMa stars”

Except for its orbital period (≈ 27.7 min), not much is known about the parameters of the first observed helium SU UMa-type

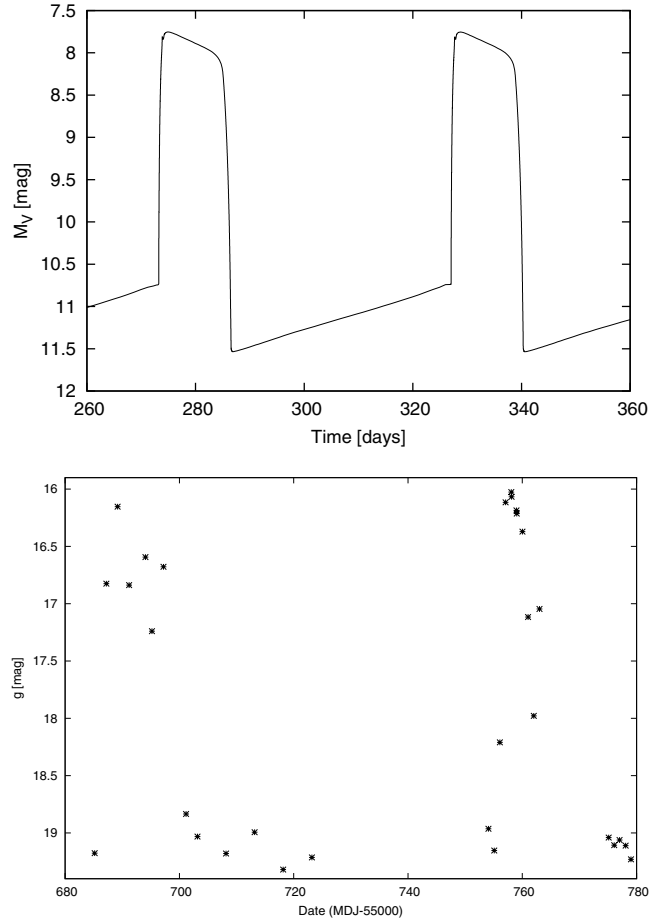


Fig. 8. Top: KL Dra model: $\alpha_c = 0.035$, $\alpha_h = 0.2$, $\dot{M}_{\text{tr}} = 2 \times 10^{16} \text{ g/s}$, $M_1 = 0.6$, $\gamma = 0.8$, $\langle R_D \rangle = 1.2 \times 10^{10}$ cm, $\mu_{30} = 1$. Bottom: the part of KL Dra light curve showing two subsequent superoutbursts. The time between their maxima is 60 days, their amplitude is 3 mag and their duration is ~ 13 days. Data provided by G. Ramsay.

system PTF 1J0719+4858. We assumed that this type of system transfers mass close to its secular mean taking for example $\dot{M}_{\text{tr}} = 6 \times 10^{16} \text{ g s}^{-1}$. Assuming $M_1 = 1.0 M_{\odot}$, $\alpha_c = 0.02$, $\alpha_h = 0.1$, $\gamma = 0.6$ and a disc extending down to the white dwarf’s surface, we obtain the light curve presented in Fig. 9.

Between superoutbursts one notices a slight gradual increase of the minimum and maximum luminosity of the four consecutive normal outbursts. This happens because during each normal outburst \dot{M}_{tr} is enhanced due to the γ -prescription – the disc gains more mass than it loses due to accretion. In consequence, during the sequence of the normal outbursts, the mass is gradually accumulated in the disc and Σ rises everywhere in the disc. This effect is one of the typical features of the EMT model (it is also present for slightly different reasons in the TTI version, see Schreiber et al. 2004; Tsugawa & Osaki 1997). While the general luminosity rise in quiescence is not observed in real systems and is one of the well-known weaknesses of the DIM (Smak 2000), the increase of the normal outburst peak luminosity might have been observed in some systems.

The superoutburst precursor is also typical of the EMT and TTI models (Schreiber et al. 2004). Owing to mass accumulation, the disc arrives to a state where after the rise of a normal outburst the cooling front is no longer able to propagate – the disc becomes stuck in the hot state and a superoutburst begins. This last normal outburst leading to a superoutburst appears as its precursor in the light curve.

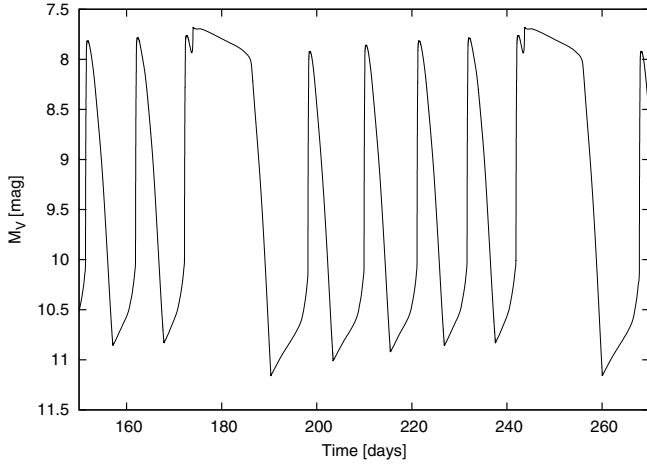


Fig. 9. SU UMa type model for a helium-dominated disc (an “imitation” of PTF 1J0719+4858): $\alpha_c = 0.02$, $\alpha_h = 0.1$, $\dot{M}_{\text{tr}} = 6 \times 10^{16}$ g/s, $M_1 = 1.0$, $\gamma = 0.6$, $\langle R_D \rangle = 1.0 \times 10^{10}$ cm, $\mu_{30} = 0$.

When one compares the model light curve to that of PTF 1J0719+4858 one notes that in our model the amplitude of normal outbursts is larger by 0.5 mag and their duration is four times longer. The normal outburst duration is extended by the \dot{M}_{tr} enhancement and the high amplitude is the result of the assumed α ratio – a lower ratio would result in a lower amplitude. However, lowering this ratio would also lower the superoutburst amplitude and shorten the recurrence times. Tuning the parameters to obtain a better agreement does not make much sense in view of the arbitrariness of the mass-transfer prescription and of the uncertainties of the DIM itself (for a discussion of the weaknesses of the DIM see e.g. Lasota 2001).

7.2. Dips, “cycling”, and standstills

Most (maybe all) non-stationary AM CVn stars exhibit at least one “dip” during the decay from the superoutburst maximum. In some systems, such as CR Boo and V803 Cen, the system becomes stuck in a state in which series of ~ 1.0 mag amplitude, ~ 0.8 – 1.0 d recurrence-time outbursts are present. There is also some uncertainty about the nature of these intermissions in the superoutburst cycle: on the one hand, Patterson et al. (1997) considered them to be a “cycling state” of normal dwarf nova outbursts, on the other, Kato et al. (2001a) interpreted the same part of the (CR Boo) light curve as a standstill (analogous to that of Z Cam dwarf novae), arguing that no superoutbursts are present during the high-luminosity state maintained by the system.

To answer which of these two interpretation is correct more observations are needed and therefore we are not trying here to decide this question. Let us just note that the mass-transfer rates attributed to both CR Boo and V803 Cen (see Table 2) are very close to the critical ones and a phenomenon analogous to Z Cam standstills could indeed be expected for this systems. The main difference is that Z Cam stars have no superoutbursts. We will therefore adapt the Z Cam model (Buat-Ménard et al. 2001b) by combining its mass-transfer modulations with the EMT.

Dips might or might not be related to the cycling states. They are very similar to those observed during the decay from superoutbursts in WZ Sge – type stars. If common to both systems, the dip origins are not connected to the value of the mass-transfer rate but would instead result from the very compact size of their orbit. We tested a simple hypothesis based on this assumption.

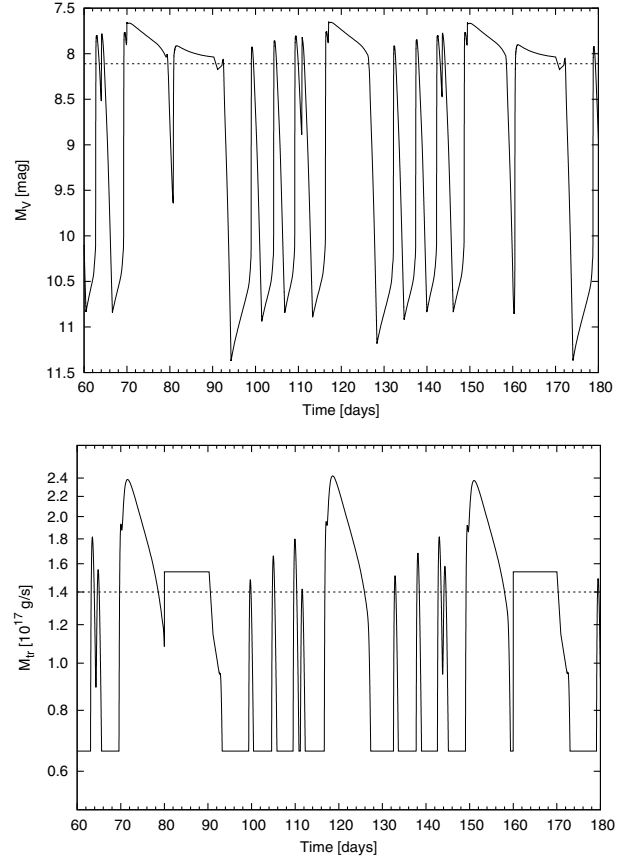


Fig. 10. Light-curve of a superoutbursting helium “Z Cam” star. The parameters are: $\alpha_c = 0.04$, $\alpha_h = 0.2$, $M_1 = 1.0$, $\dot{M}_{0,\text{tr}} = 1.1 \times 10^{17}$ g/s, $\Delta\dot{M}_{\text{tr}} = 0.4\dot{M}_{0,\text{tr}}$, $T_{\text{high}} = 10$ d, $T_{\text{low}} = 70$ d, $\gamma = 0.7$, $\langle R_D \rangle = 1.0 \times 10^{10}$ cm. The dotted lines correspond to the hot stability limit. *Top*: the model lightcurve. *Bottom*: \dot{M}_{tr} modulations.

7.2.1. “Z Cam – type” modulations and (γ) superoutbursts

Adapting the Z Cam models of Buat-Ménard et al. (2001b) to AM CVn stars, we assumed the mass-transfer to be modulated as $\Delta\dot{M}/\langle\dot{M}\rangle = 40\%$ around an average rate close to the critical ($\dot{M}_{0,\text{tr}} = 1.1 \times 10^{17}$ g s $^{-1}$) with the modulations occurring at $t = 10 + k \cdot 70$ days. To take into account the presence of irradiation-induced super-outbursts, we combined this modulation with the mass-transfer modulations given by Eq. (22) with the actual value of the mass-transfer rate instead of $\dot{M}_{0,\text{tr}}$.

Choosing as typical parameters $\alpha_c = 0.04$, $\alpha_h = 0.2$, $M_1 = 1.0 M_{\odot}$, $\langle R_D \rangle = 1.0 \times 10^{10}$ cm and $\gamma = 0.7$ we obtained the light-curve shown in Fig. 10. This light curve shows superoutbursts and standstills, as expected, but features resulting from the superposition of both types of modulations are also present in the form of “dips”. The hypercycle (the cycle between two consecutive superoutbursts followed by a standstill) starts with a major \dot{M}_{tr} enhancement due to γ (about day 70 in Fig. 10). During the decay from the superoutburst maximum, the fall of \dot{M}_{tr} is quenched by a rising modulation through $\dot{M}_{0,\text{tr}} + \Delta\dot{M}_{\text{tr}}$. The mass-transfer rate then rises again (around day 80) and the decay from maximum is reversed by a resulting outside-in heating front that catches up with the propagating cooling front before the latter reaches the inner disc edge – hence a dip-like feature. During the following superoutburst a slight difference in phase of the two mass-transfer modulations allows the cooling front to propagate almost to the disc’s inner end. A similar mechanism produces the dip in narrow outburst preceding the superoutburst.

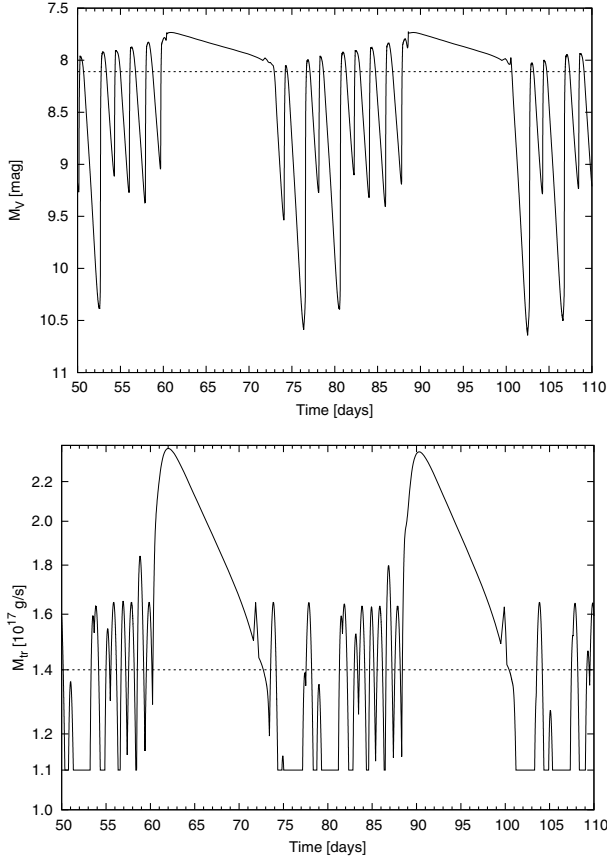


Fig. 11. Light-curve corresponding to sinusoidal + γ modulations of the mass-transfer rate. The parameters are $\alpha_h = 0.2$, $\alpha_c = 0.04$, $\dot{M}_{0,\text{tr}} = 1.1 \times 10^{17}$ g/s, $M_1 = 1.0$, $\langle R_D \rangle = 1.0 \times 10^{10}$ cm, $\gamma = 0.8$, $A = 0.5$, $C = 2$, $\tau = 1$ d. The dotted lines correspond to the hot stability limit. *Top:* the model light-curve. *Bottom:* \dot{M}_{tr} modulations.

This suggest that dips and cycling-state features might result from mass-transfer modulations triggering cooling/heating front “catching” and not from reflections. On the other hand, the indentation at the end of the standstill is the result of cooling front reflection (as described e.g. in [Dubus et al. 1999](#)).

In any case, the presumed Z Cam-effect mechanism applied to outbursting AM CVns produced the required type of light-curve combining superoutbursts with standstills (and normal outbursts).

7.2.2. Sinusoidal plus γ modulation

The prescription for mass-transfer modulations used in the previous section has been chosen not for its realism but because it has been used in modelling the Z Cam star outburst cycle by [Buat-Ménard et al. \(2001b\)](#). Since in AM CVn light curves other peculiar features are present in addition to standstills, we tried other simple forms of mass-transfer rate modulations. The simplest is to modify Eq. (22) as

$$\dot{M}_{\text{tr}} = \max\left(\dot{M}_{0,\text{tr}}\left(1 + A \sin\left(C + \pi t/\tau\right)\right), \left(\gamma \dot{M}_{\text{acc}}\right)\right), \quad (23)$$

where A , C (dimensionless) and τ (time) are adjustable constants; t is the time coordinate.

With $M_1 = 1.0 M_{\odot}$, $\dot{M}_{\text{tr}} = 1.1 \times 10^{17}$ g/s and $\langle R_D \rangle = 1.0 \times 10^{10}$ cm, $\alpha_c = 0.04$, $\alpha_h = 0.2$, choosing $A = 0.5$, $C = 2$, $\tau = 1$ d and $\gamma = 0.8$, one obtains the light-curve shown in [Fig. 11](#). One remarks superoutbursts with amplitude $A_s \sim 3$ mag,

duration $T_s \sim 15$ d and recurrence time of $T_{\text{recc},s} \sim 27$ d. These are “standard” EMT, γ enhancement triggered outbursts.

In addition, short outbursts with a repetition time of ~ 1.8 d and an amplitude ~ 1.2 – 2.4 mag appear during the decay from superoutburst and persist until the following one. One is of course tempted to identify them with the cycling states of CR Boo or V803 Cen. The origin of these short outbursts is quite simple to understand. At the end of the decay of \dot{M}_{tr} from the superoutburst value, the sinusoidal enhancement of \dot{M}_{tr} dominates over the γ -induced \dot{M}_{tr} fluctuations. This short enhancement of \dot{M}_{tr} results in a low-amplitude, short outburst (starting at day ~ 73 in [Fig. 11](#)) because the cooling front that switches off the superoutburst is quenched by a freshly triggered outside-in heating front. This happens because the post cooling-front Σ is still close to the critical value due to \dot{M}_{tr} enhancement through the sinusoidal modulation. However, since the disc’s mass is decreasing, after one day a new cooling front forms immediately and this time it succeeds in propagating till the inner disc’s rim but, because of mass accumulating, a new outburst starts immediately. All heating fronts are triggered by mass-transfer enhancement and therefore are of the outside-in type. Their maximum brightness is slightly increasing because of mass accumulation. Their amplitude depends on the combined strengths of the sinusoidal plus γ enhancement. If the enhancement brings \dot{M}_{tr} above \dot{M}_{cr}^+ , the resulting short outburst will have a small amplitude as the outside-in heating front forms easily and promptly catches up with the cooling front that quenches it. But when the disc’s mass is still quite low and the $\gamma + \text{sine}$ enhancement is not powerful enough to bring \dot{M}_{tr} above \dot{M}_{cr}^+ , the cooling front will propagate unhindered till the low-luminosity level (see e.g. day 81). It seems therefore that suitable modulations of the mass-transfer rate can produce various observed features of outbursting AM CV light curves even with the simplest assumptions about their shapes. However, the mechanisms that are able to produce mass-transfer modulations are still to be firmly identified. There is no doubt that mass-transfer rates in close binaries are highly variable, in some cases with huge amplitudes. Although irradiation of the secondary is the usual suspect, in many cases the observed variability (e.g. of AM Her or VY Scl stars, see [Warner 2003](#)) is not caused by irradiation. A periodic modulation of mass transfer rate could be also achieved by a warped/tilted disc ([Smak 2011](#)). The sinusoidal variation of Eq. (23) could be considered an attempt to represent the effect of the variable irradiation if the modulation time $\tau = 1$ d could find a convincing interpretation in this context.

8. Discussion and conclusion

We have presented a systematic study of the dwarf nova disc instability model applied to outbursts of AM CVn stars. What distinguishes these hydrogen-free compact binaries from their cataclysmic variable cousins is (obviously) the chemical composition and the shorter orbital periods. Both can affect the outburst properties of AM CVn stars. However, because of the still unknown mechanism driving the, in this context necessary, variations of viscosity (or the angular-momentum and disc-heating mechanisms), it is difficult to assess the importance of helium richness in the outburst processes. They are potentially important, especially for very low metallicities, but their real impact must await a better understanding of the basic mechanisms driving accretion in discs in general.

The main conclusion of our study is that outbursts of AM CVn stars can be explained by the suitably adapted dwarf-nova disc instability model. As in the case of the application

of this model to hydrogen-dominated cataclysmic variables, one has to resort to additional effects to account for the observed superoutbursts, dips, cycling states, and standstills. The basic feature that must be included in the model is the variable mass-transfer rate. In particular, we have shown that the enhanced mass-transfer rate, due presumably to variable irradiation of the secondary, must not only be taken into account but is a factor that determines the shape of most AM CVn star outburst light-curves. The cause of variable secondary's irradiation has yet to be understood; the best candidate is the precession of a tilted/warped disc. We postulate the existence of superoutbursting Z Cam type AM CVn stars that have been tentatively identified in observations (Kato et al. 2001b).

The very short periods and low mass-ratios of AM CVns imply the absence of a pure normal-outburst cycle – the test bench of the model. The compactness of AM CVn stars is clearly important because it accentuates the interactions between the binary components. Their mass ratios $0.0125 \leq q \leq 0.18$ suggest the importance of resonances and tidal effects in general. Irradiation is important as well. Unfortunately, as in the case of cataclysmic variables, the implications of these effects on the outbursts is uncertain and subject to controversy. At best, they can be parameterized only in a fairly rough way. The only effect that, presumably important in CVs, can be neglected in AM CVn stars is the irradiation of the inner disc regions by the hot accreting white dwarfs.

Another drawback that might soon be overcome is the usually unsatisfactory quality of observations of the AM CV outbursts. When better and richer sets of data will become available the comparison of the DIM with observations will become a precious source of knowledge about accretion disc physics.

Acknowledgements. We are grateful to Chris Deloye for providing us with evolutionary tracks for AM CVn stars and to Joe Patterson for the CR Boo and V803 Cen light curves and for helpful comments, to Gavin Ramsay for the light curve of KL Dra and to Eran Ofek and David Levitan for very helpful discussions about the AM CVn stars observed by the Palomar Transient Factory. We thank the anonymous referee for his report that helped improving the article. This work has been supported in part by the Polish MNiSW grants PSP/K/PBP/000392 and N N203 380336, the French Space Agency CNES, and the European Commission via contract ERC-StG-200911.

Appendix A: Critical parameters for various chemical compositions of helium-dominated discs

In this Appendix we summarize the critical parameter formulae obtained for various chemical compositions of helium-dominated accretion discs. These formulae have been obtained through fits to numerically calculated S-curves. One should note that whereas the fit of the critical T_{eff} is accurate within few percent, the accuracy of the fit to the critical central temperature is no better than 20%.

1. $Y = 1$

$$\begin{aligned}
 \Sigma^+ &= 528 \alpha_{0.1}^{-0.81} R_{10}^{1.07} M_1^{-0.36} \text{ g cm}^{-2} \\
 \Sigma^- &= 1620 \alpha_{0.1}^{-0.84} R_{10}^{1.19} M_1^{-0.40} \text{ g cm}^{-2} \\
 T_c^+ &= 77000 \alpha_{0.1}^{-0.20} R_{10}^{0.08} M_1^{-0.03} \text{ K} \\
 T_c^- &= 17800 \alpha_{0.1}^{-0.13} R_{10}^{-0.03} M_1^{0.01} \text{ K} \\
 T_{\text{eff}}^+ &= 13000 \alpha_{0.1}^{-0.01} R_{10}^{-0.08} M_1^{0.03} \text{ K} \\
 T_{\text{eff}}^- &= 9700 \alpha_{0.1}^{-0.01} R_{10}^{-0.09} M_1^{0.03} \text{ K} \\
 \dot{M}^+ &= 1.01 \times 10^{17} \alpha_{0.1}^{-0.05} R_{10}^{2.68} M_1^{-0.89} \text{ g s}^{-1} \\
 \dot{M}^- &= 3.17 \times 10^{16} \alpha_{0.1}^{-0.02} R_{10}^{2.66} M_1^{-0.89} \text{ g s}^{-1}
 \end{aligned} \tag{A.1}$$

2. $Y = 0.98, Z = 0.02$

$$\begin{aligned}
 \Sigma^+ &= 380 \alpha_{0.1}^{-0.78} R_{10}^{1.06} M_1^{-0.35} \text{ g cm}^{-2} \\
 \Sigma^- &= 612 \alpha_{0.1}^{-0.82} R_{10}^{1.10} M_1^{-0.37} \text{ g cm}^{-2} \\
 T_c^+ &= 71400 \alpha_{0.1}^{-0.21} R_{10}^{0.08} M_1^{-0.03} \text{ K} \\
 T_c^- &= 23600 \alpha_{0.1}^{-0.14} R_{10}^{-0.00} M_1^{0.00} \text{ K} \\
 T_{\text{eff}}^+ &= 11500 \alpha_{0.1}^{-0.01} R_{10}^{-0.08} M_1^{0.03} \text{ K} \\
 T_{\text{eff}}^- &= 8690 \alpha_{0.1}^{-0.00} R_{10}^{-0.09} M_1^{0.03} \text{ K} \\
 \dot{M}^+ &= 6.22 \times 10^{16} \alpha_{0.1}^{-0.05} R_{10}^{2.67} M_1^{-0.89} \text{ g s}^{-1} \\
 \dot{M}^- &= 2.04 \times 10^{16} \alpha_{0.1}^{-0.02} R_{10}^{2.62} M_1^{-0.87} \text{ g s}^{-1}
 \end{aligned} \tag{A.2}$$

3. $Y = 0.96, Z = 0.04$

$$\begin{aligned}
 \Sigma^+ &= 322 \alpha_{0.1}^{-0.78} R_{10}^{1.04} M_1^{-0.35} \text{ g cm}^{-2} \\
 \Sigma^- &= 459 \alpha_{0.1}^{-0.81} R_{10}^{1.08} M_1^{-0.36} \text{ g cm}^{-2} \\
 T_c^+ &= 66800 \alpha_{0.1}^{-0.22} R_{10}^{0.07} M_1^{-0.02} \text{ K} \\
 T_c^- &= 25100 \alpha_{0.1}^{-0.14} R_{10}^{0.00} M_1^{-0.00} \text{ K} \\
 T_{\text{eff}}^+ &= 10700 \alpha_{0.1}^{-0.01} R_{10}^{-0.09} M_1^{0.03} \text{ K} \\
 T_{\text{eff}}^- &= 8350 \alpha_{0.1}^{-0.00} R_{10}^{-0.10} M_1^{0.03} \text{ K} \\
 \dot{M}^+ &= 4.76 \times 10^{16} \alpha_{0.1}^{-0.06} R_{10}^{2.65} M_1^{-0.88} \text{ g s}^{-1} \\
 \dot{M}^- &= 1.74 \times 10^{16} \alpha_{0.1}^{-0.02} R_{10}^{2.61} M_1^{-0.87} \text{ g s}^{-1}
 \end{aligned} \tag{A.3}$$

Appendix B: Amplitudes and recurrence times for a selection of U Gem-type dwarf novae and outbursting AM CVn stars

Table B.1. List of U Gem-type and AM CVn systems (three last rows) plotted in Fig. 6.

	P_{orb} [h]	A_n [mag]	T_n [d]
IP Peg	3.79	3.5	95
AR And	3.91	5.9	25
UU Aql	3.92	5.1	48
CW Mon	4.24	3.5	150
U Gem	4.25	5.2	132
TW Vir	4.38	3.2	28
SS Aur	4.39	4.5	51
EX Dra	5.04	3.0	20
CZ Ori	5.25	4.4	26
BV Pup	6.36	2.1	30
BF Eri	6.50	3.0	64
SS Cyg	6.60	3.9	50
AF Cam	7.78	3.6	75
MU Cen	8.21	2.3	45
CH UMa	8.24	4.0	335
RU Peg	8.99	3.7	76
AT Ara	9.01	3.8	97
DX And	10.57	3.9	300
EY Cyg	11.02	4.1	96
V442 Cen	11.04	3.5	32
BV Cen	14.67	1.9	167
CR Boo	0.41	1.2	1
V803 Cen	0.44	1.3	0.95
PTF1J0719	0.45	2.5	10

Notes. The columns are: P_{orb} is the orbital period, A_n is the normal outburst amplitude and T_n is the normal outburst recurrence time. Data from Warner (2003) and Ritter & Kolb (2003).

References

- Balbus, S. A. 2003, *ARA&A*, 41, 555
- Balbus, S. A. 2005, in *The Astrophysics of Cataclysmic Variables and Related Objects*, eds. J.-M. Hameury, & J.-P. Lasota, *Proc. ASP Conf.*, 330, 185
- Bildsten, L., Townsley, D. M., Deloye, C. J., & Nelemans, G. 2006, *ApJ*, 640, 466
- Buat-Ménard, V., & Hameury, J.-M. 2002, *A&A*, 386, 891
- Buat-Ménard, V., Hameury, J.-M., & Lasota, J.-P. 2001a, *A&A*, 366, 612
- Buat-Ménard, V., Hameury, J.-M., & Lasota, J.-P. 2001b, *A&A*, 369, 925
- Cannizzo, J. K. 1984, *Nature*, 311, 443
- Cannizzo, J. K., & Wheeler, J. C. 1984, *ApJS*, 55, 367
- Cannizzo, J. K., Still, M. D., Howell, S. B., Wood, M. A., & Smale, A. P. 2010, *ApJ*, 725, 1393
- Copperwheat, C. M., Marsh, T. R., Littlefair, S. P., et al. 2011a, *MNRAS*, 410, 1113
- Copperwheat, C. M., Marsh, T. R., Dhillon, V. S., et al. 2011b, *MNRAS*, 413, 3068
- Deloye, C. J., Taam, R. E., Winisdoerffer, C., & Chabrier, G. 2007, *MNRAS*, 381, 525
- Dubus, G., Lasota, J.-P., Hameury, J.-M., & Charles, P. 1999, *MNRAS*, 303, 139
- Dubus, G., Hameury, J.-M., & Lasota, J.-P. 2001, *A&A*, 373, 251
- Espaillet, C., Patterson, J., Warner, B., & Woudt, P. 2005, *PASP*, 117, 189
- El-Khoury, W., & Li, J. 1999, *Annapolis Workshop on Magnetic Cataclysmic Variables*, 157, 393
- Fontaine, G., Brassard, P., Green, E. M., et al. 2011, *ApJ*, 726, 92
- Frank, J., King, A., & Raine, D. J. 2002, *Accretion Power in Astrophysics*, eds. J. Frank, A. King, & D. Raine (Cambridge, UK: Cambridge University Press), 398
- Hameury, J.-M., & Lasota, J.-P. 2005, *A&A*, 443, 283
- Hameury, J.-M., Lasota, J.-P., & Hure, J.-M. 1997, *MNRAS*, 287, 937
- Hameury, J.-M., Menou, K., Dubus, G., Lasota, J.-P., & Huré, J.-M. 1998, *MNRAS*, 298, 1048
- Hameury, J.-M., Lasota, J.-P., & Dubus, G. 1999, *MNRAS*, 303, 39
- Hameury, J.-M., Lasota, J.-P., & Warner, B. 2000, *A&A*, 353, 244
- Ichikawa, S., & Osaki, Y. 1992, *PASJ*, 44, 15
- Idan, I., Lasota, J. P., Hameury, J.-M., & Shaviv, G. 1999, *Phys. Rep.*, 311, 213
- Idan, I., Lasota, J.-P., Hameury, J.-M., & Shaviv, G. 2010, *A&A*, 519, A117
- Kato, T., Nogami, D., Baba, H., Hanson, G., & Poyner, G. 2000, *MNRAS*, 315, 140
- Kato, T., Reszelski, M., Poyner, G., et al. 2001a, *Information Bulletin on Variable Stars*, 5120, 1
- Kato, T., Stubbings, R., Monard, B., Pearce, A., & Nelson, P. 2001b, *Information Bulletin on Variable Stars*, 5091, 1
- Kato, T., Stubbings, R., Monard, B., et al. 2004, *PASJ*, 56, 89
- King, A. R., Pringle, J. E., & Livio, M. 2007, *MNRAS*, 376, 1740
- Kotko, I., & Lasota, J.-P. 2012, *A&A*, submitted
- Kotko, I., Lasota, J.-P., & Dubus, G. 2010, *Astron. Nachr.*, 331, 231
- Kukarkin, B. W., & Parenago, P. P. 1934, *Var. Star. Bull.*, 4, 44
- Lasota, J.-P. 2001, *New A Rev.*, 45, 449
- Lasota, J. P., Hameury, J. M., & Hure, J. M. 1995, *A&A*, 302, L29
- Lasota, J.-P., Dubus, G., & Kruk, K. 2008, *A&A*, 486, 523
- Levitán, D., Fulton, B. J., Groot, P. J., et al. 2011a, *ApJ*, 739, 68
- Levitán, D., Kulkarni, S., Prince, T., Nissanke, S., & Vallisneri, M. 2011b, *APS Meeting Abstracts*, 11008
- Mason, K. O., Cordova, F. A., Watson, M. G., & King, A. R. 1988, *MNRAS*, 232, 779
- Menou, K., Hameury, J.-M., & Stehle, R. 1999, *MNRAS*, 305, 79
- Menou, K., Hameury, J.-M., Lasota, J.-P., & Narayan, R. 2000, *MNRAS*, 314, 498
- Nauenberg, M. 1972, *ApJ*, 175, 417
- Nelemans, G. 2005, *The Astrophysics of Cataclysmic Variables and Related Objects*, eds. J.-M. Hameury, & J.-P. Lasota, *Proc. ASP Conf.*, 330, 27
- Nelemans, G. 2010, http://www.astro.ru.nl/~nelemans/dokuwiki/doku.php?id=verification_binaries:am_cvn_stars
- Nelemans, G., Yungelson, L. R., van der Sluys, M. V., & Tout, C. A. 2010, *MNRAS*, 401, 1347
- Osaki, Y. 1985, *A&A*, 144, 369
- Osaki, Y. 1989, *PASJ*, 41, 1005
- Osaki, Y. 1996, *PASP*, 108, 39
- Patterson, J., Kemp, J., Shambrook, A., et al. 1997, *PASP*, 109, 1100
- Patterson, J., Walker, S., Kemp, J., et al. 2000, *PASP*, 112, 625
- Payne-Gaposchkin, C. 1957, in *The Galactic Novae* (North-Holland Publishing Company and Interscience Publishers), Chap. 8
- Payne-Gaposchkin, C. 1977, *Novae and Related Stars*, ed. M. Friedjung (Dordrecht: D. Reidel Publishing Co), *Astrophys. Space Sci. Lib.*, 65, 3
- Ramsay, G., Groot, P. J., Marsh, T., et al. 2006, *A&A*, 457, 623
- Ramsay, G., Kotko, I., Barclay, T., et al. 2010, *MNRAS*, 407, 1819
- Ramsay, G., Barclay, T., Steeghs, D., et al. 2012, *MNRAS*, 419, 2836
- Ritter, H., & Kolb, U. 2003, *A&A*, 404, 301
- Roelofs, G. H. A., Groot, P. J., Benedict, G. F., et al. 2007a, *ApJ*, 666, 1174
- Roelofs, G. H. A., Nelemans, G., & Groot, P. J. 2007b, *MNRAS*, 382, 685
- Roelofs, G. H. A., Groot, P. J., Steeghs, D., et al. 2009, *MNRAS*, 394, 367
- Schreiber, M. R. 2007, *A&A*, 466, 1025
- Schreiber, M. R., Hameury, J.-M., & Lasota, J.-P. 2004, *A&A*, 427, 621
- Shakura, N. I., & Sunyaev, R. A. 1973, *A&A*, 24, 337
- Shears, J., Brady, S., Koff, R., Goff, W., & Boyd, D. 2011, *JBAA*, in press [arXiv:1104.0107]
- Sion, E. M., & Godon, P. 2007, *15th European Workshop on White Dwarfs*, 372, 563
- Sion, E. M., Linnell, A. P., Godon, P., & Ballouz, R.-L. 2011, *ApJ*, 741, 63
- Skillman, D. R., Patterson, J., Kemp, J., et al. 1999, *PASP*, 111, 1281
- Smak, J. 1983, *Acta Astron.*, 33, 333
- Smak, J. 1984a, *Acta Astron.*, 34, 161
- Smak, J. 1984b, *PASP*, 96, 5
- Smak, J. 1989, *Acta Astron.*, 39, 201
- Smak, J. 1999, *Acta Astron.*, 49, 391
- Smak, J. 2000, *New A Rev.*, 44, 171
- Smak, J. 2002, *Acta Astron.*, 52, 263
- Smak, J. 2006, *Acta Astron.*, 56, 365
- Smak, J. 2008, *Acta Astron.*, 58, 55
- Smak, J. 2009a, *Acta Astron.*, 59, 89
- Smak, J. 2009b, *Acta Astron.*, 59, 103
- Smak, J. 2009c, *Acta Astron.*, 59, 109
- Smak, J. 2009d, *Acta Astron.*, 59, 121
- Smak, J. 2009e, *Acta Astron.*, 59, 419
- Smak, J. 2011, *Acta Astron.*, 61, 59
- Smak, J. 2010, *Acta Astron.*, 60, 83
- Smak, J., & Waagen, E. O. 2004, *Acta Astron.*, 54, 433
- Sorathia, K. A., Reynolds, C. S., Stone, J. M., & Beckwith, K. 2011 [arXiv:1106.4019]
- Strohmer, T. E. 2004, *ApJ*, 614, 358
- Tsugawa, M., & Osaki, Y. 1997, *PASJ*, 49, 75
- van Paradijs, J. 1983, *A&A*, 125, L16
- van Paradijs, J. 1985, *A&A*, 144, 199
- Viallet, M., & Hameury, J.-M. 2007, *A&A*, 475, 597
- Viallet, M., & Hameury, J.-M. 2008, *A&A*, 489, 699
- Wade, R. A., Eracleous, M., & Flohic, H. M. L. G. 2007, *AJ*, 134, 1740
- Warner, B. 1987, *MNRAS*, 227, 23
- Warner, B. 2003, *Cataclysmic Variable Stars*, Cambridge Astrophysics Series, 28
- Wood, M. A., Casey, M. J., Garnavich, P. M., & Haag, B. 2002, *MNRAS*, 334, 87