

Contents lists available at [ScienceDirect](http://ScienceDirect)

# Physics Letters B

[www.elsevier.com/locate/physletb](http://www.elsevier.com/locate/physletb)


## Planck scale physics and topology change through an exactly solvable model


 Francisco S.N. Lobo<sup>a</sup>, Jesus Martinez-Asencio<sup>b</sup>, Gonzalo J. Olmo<sup>c,d</sup>, D. Rubiera-Garcia<sup>d,\*</sup>
<sup>a</sup> Centro de Astronomia e Astrofísica da Universidade de Lisboa, Campo Grande, Ed. C8, 1749-016 Lisboa, Portugal

<sup>b</sup> Departamento de Física Aplicada, Facultad de Ciencias, Fase II, Universidad de Alicante, Alicante E-03690, Spain

<sup>c</sup> Departamento de Física Teórica and IFIC, Centro Mixto Universidad de Valencia – CSIC, Universidad de Valencia, Burjassot, 46100, Valencia, Spain

<sup>d</sup> Departamento de Física, Universidade Federal da Paraíba, 58051-900 João Pessoa, Paraíba, Brazil

### ARTICLE INFO

#### Article history:

Received 2 January 2014

Received in revised form 19 February 2014

Accepted 19 February 2014

Available online 25 February 2014

Editor: M. Trodden

#### Keywords:

Modified gravity

Palatini formalism

Planck-scale physics

Dynamical Vaidya solutions

Topology change

### ABSTRACT

We consider the collapse of a charged radiation fluid in a Planck-suppressed quadratic extension of General Relativity (GR) formulated à la Palatini. We obtain exact analytical solutions that extend the charged Vaidya-type solution of GR, which allows to explore in detail new physics at the Planck scale. Starting from Minkowski space, we find that the collapsing fluid generates wormholes supported by the electric field. We discuss the relevance of our findings in relation to the quantum foam structure of space–time and the meaning of curvature divergences in this theory.

© 2014 Published by Elsevier B.V. This is an open access article under the CC BY license (<http://creativecommons.org/licenses/by/3.0/>). Funded by SCOAP<sup>3</sup>.

## 1. Introduction

An outstanding question in quantum gravitational physics is whether large metric fluctuations at the Planck scale may induce a change in topology. In fact, as suggested by Wheeler, at scales below the Planck length, the highly nonlinear and strongly interacting metric fluctuations may endow space–time with a foam-like structure [1]. Thus, this hints that the geometry, and the topology, may be constantly fluctuating, so that space–time may take on all manners of nontrivial topological structures, such as wormholes [2]. However, one does encounter a certain amount of criticism to Wheeler's notion of space–time foam, for instance, in that stability considerations may place constraints on the nature or even on the existence of Planck-scale foam-like structures [3]. Nevertheless, the notion of space–time foam is generally accepted, in that this picture leads to topology-changing quantum amplitudes and to interference effects between different space–time topologies [4], although these possibilities have met with some disagreement [5].

Due to the multiply-connected nature of wormholes, their respective creation/generation inevitably involves the problematic issue of topology change [6,7]. The possibility that inflation might provide a natural mechanism for the enlargement of Planck-size wormholes to macroscopic size has been explored [8]. In fact, the construction of general relativistic traversable wormholes, with the idealization of impulsive phantom radiation, has been considered extensively in the literature [9–11]. Another problematic aspect in wormhole physics is that these geometries violate the pointwise energy conditions [12]. However, this issue may be avoided in modified gravity, where the normal matter threading the wormholes may in principle be imposed to satisfy the energy conditions, and it is the higher order curvature terms that sustain these geometries [13]. In fact, the general approach to wormhole physics is to run the gravitational field equations in the reverse direction, namely, consider first an interesting space–time metric and then through the field equations, the distribution of the stress-energy tensor components is deduced. However, one may rightly argue that this approach in solving the field equations lacks physical justification, and that a more physical motivation would be to consider a plausible distribution of matter–energy.

In this work we follow the latter route and find that in a quadratic extension of GR formulated in the Palatini formalism wormholes can be generated dynamically. This result follows by

\* Corresponding author.

E-mail addresses: [flobo@cii.fc.ul.pt](mailto:flobo@cii.fc.ul.pt) (F.S.N. Lobo), [jesusmartinez@ua.es](mailto:jesusmartinez@ua.es) (J. Martinez-Asencio), [gonzalo.olmo@csic.es](mailto:gonzalo.olmo@csic.es) (G.J. Olmo), [drubiera@fisica.ufpb.br](mailto:drubiera@fisica.ufpb.br) (D. Rubiera-Garcia).

probing Minkowski space-time with a charged null fluid, thus producing a Vaidya-type configuration. Now, switching off the flux, the metric settles down into a static configuration where the existence of a wormhole geometry becomes manifest. The topologically nontrivial character of the wormhole allows us to define the electric charge in terms of lines of electric field trapped in the topology. The size of the wormhole changes in such a way that the density of lines of force at the throat is given by a universal quantity. These facts allow to consistently interpret these solutions as geons in Wheeler's sense [1] and have important consequences for the issue of the foam-like structure of space-time [2].

## 2. Dynamical charged fluid in Ricci-squared Palatini theories

Consider a theory defined by the action [14]

$$S[g, \Gamma, \psi_m] = \frac{1}{2\kappa^2} \int d^4x \sqrt{-g} f(R, Q) + S_m[g, \psi_m], \quad (1)$$

where  $\mathcal{L}_G = f(R, Q)/(2\kappa^2)$  represents the gravity Lagrangian,  $\kappa^2$  is a constant with suitable dimensions (in GR,  $\kappa^2 \equiv 8\pi G/c^3$ ),  $g_{\mu\nu}$  is the space-time metric,  $R = g^{\mu\nu} R_{\mu\nu}$ ,  $Q = g^{\mu\alpha} g^{\nu\beta} R_{\mu\nu} R_{\alpha\beta}$ ,  $R_{\mu\nu} = R^\rho{}_{\mu\rho\nu}$ ,  $R^\alpha{}_{\beta\mu\nu} = \partial_\mu \Gamma_{\nu\beta}^\alpha - \partial_\nu \Gamma_{\mu\beta}^\alpha + \Gamma_{\mu\lambda}^\alpha \Gamma_{\nu\beta}^\lambda - \Gamma_{\nu\lambda}^\alpha \Gamma_{\mu\beta}^\lambda$  is the Riemann tensor of the connection  $\Gamma_{\mu\nu}^\lambda$ , and  $S_m[g, \psi_m]$  represents the matter action (with  $\psi_m$  the matter fields). We work in the Palatini formalism, where  $g_{\mu\nu}$  and  $\Gamma_{\mu\nu}^\lambda$  are regarded as independent fields. Setting the torsion to zero for simplicity, the field equations imply that  $\Gamma_{\mu\nu}^\lambda$  can be written as the Levi-Civita connection of a metric  $h_{\mu\nu}$  defined as [15]

$$h^{\mu\nu} = \frac{g^{\mu\alpha} \Sigma_\alpha{}^\nu}{\sqrt{\det \hat{\Sigma}}}, \quad h_{\mu\nu} = (\sqrt{\det \hat{\Sigma}}) \Sigma^{-1}{}_\mu{}^\alpha g_{\alpha\nu}. \quad (2)$$

We have defined the matrices  $\hat{\Sigma}$  and  $\hat{P}$ , whose components are  $\Sigma_\alpha{}^\nu \equiv (f_R \delta_\alpha^\nu + 2f_Q P_\alpha{}^\nu)$  and  $P_\mu{}^\nu \equiv R_{\mu\alpha} g^{\alpha\nu}$ , with  $f_X \equiv df/dX$ . From the metric variation, one finds

$$2f_Q \hat{P}^2 + f_R \hat{P} - \frac{f}{2} \hat{I} = \kappa^2 \hat{T}, \quad (3)$$

which represents a quadratic algebraic equation for  $P_\mu{}^\nu$  as a function of  $[\hat{T}]_\mu{}^\nu \equiv T_{\mu\alpha} g^{\alpha\nu}$ . This implies that  $R = [\hat{P}]_\mu{}^\mu$ ,  $Q = [\hat{P}^2]_\mu{}^\mu$ , and  $\Sigma_\alpha{}^\nu$  are just functions of the matter sources. With elementary algebraic manipulations, Eq. (3) can be cast as

$$R_\mu{}^\nu(h) = \frac{\kappa^2}{\sqrt{\det \hat{\Sigma}}} (\mathcal{L}_G \delta_\mu^\nu + T_\mu{}^\nu). \quad (4)$$

This representation of the metric field equations implies that  $h_{\mu\nu}$  satisfies a set of GR-like second-order field equations, with the right-hand side being determined by  $\hat{T}$ . Since  $h_{\mu\nu}$  and  $g_{\mu\nu}$  are algebraically related, it follows that  $g_{\mu\nu}$  also verifies second-order equations. We note that whenever  $T_\mu{}^\nu = 0$ , Eq. (4) recovers the vacuum Einstein equations [15,16] (with possibly a cosmological constant, depending on the form of  $\mathcal{L}_G$ ) which, similarly as in other Palatini theories [18,17,19], guarantees the absence of ghost-like instabilities.

Let us consider the matter sector as described by a spherically symmetric flux of ingoing charged matter with a stress-energy tensor given by

$$T_{\mu\nu}^{\text{flux}} = \rho_{\text{in}} k_\mu k_\nu, \quad (5)$$

where  $k_\mu$  is a null vector, satisfying  $k_\mu k^\mu = 0$ , and  $\rho_{\text{in}}$  is the energy density of the flux. The electric field generated by this

flux contributes to the total stress-energy by means of  $T_{\mu\nu}^{\text{em}} = \frac{1}{4\pi} [F_{\mu\alpha} F_\nu{}^\alpha - \frac{1}{4} F_{\alpha\beta} F^{\alpha\beta} g_{\mu\nu}]$ . If we consider a line element of the form

$$ds^2 = -A(x, v) e^{2\psi(x, v)} dv^2 + 2e^{\psi(x, v)} dv dx + r^2(v, x) d\Omega^2, \quad (6)$$

then Maxwell's equations,  $\nabla_\mu F^{\mu\nu} = 4\pi J^\nu$ , where  $J^\nu \equiv \Omega(v) k^\nu$  is the current of the ingoing flux, lead to  $r^2 e^{\psi(x, v)} F^{xv} = q(v)$ , where  $q(v)$  is an integration function. These equations also imply that  $\Omega(v) \equiv q_v/4\pi r^2$ .

To proceed further, we focus on a simple quadratic extension of GR,

$$f(R, Q) = R + l_p^2 (aR^2 + Q), \quad (7)$$

where  $l_p \equiv \sqrt{\hbar G/c^3}$  is the Planck length and  $a$  is a free parameter. We note that the renormalizability of quantum fields in curved space-times requires [20] a high-energy completion of the Einstein–Hilbert Lagrangian including quadratic curvature terms such as those appearing in (7). Moreover, these higher-order curvature corrections typically appear in approaches to quantum gravity such as those based on string theory [21], and also when GR is regarded as an effective theory of quantum gravity [22]. The Palatini formulation has been particularly successful in the construction of an effective Lagrangian [23] that captures the Hamiltonian dynamics of loop quantum cosmology [24]. This theory cures the big bang singularity producing a cosmic bounce at the Planck density. The effective Palatini Lagrangian admits a power series expansion with quadratic and higher-order curvature corrections. The Lagrangian (7) is also able to replace the big bang singularity by a cosmic bounce in both isotropic and anisotropic cosmologies [25]. As a working hypothesis, we assume that the action (7) is able to capture the main effects of the quantum gravitational degrees of freedom in the form of an effective geometry, i.e., we assume that the geometrical nature of gravitation is not spoiled by the quantum gravitational effects. The model (7) is thus regarded as an effective model useful to describe dynamical processes and obtain some hints on the topology change issue, but in a purely classical continuum scenario with the quantum effects encoded in  $l_p^2$ .

Having specified the matter sources and the gravity Lagrangian, one finds  $R = 0$ ,  $Q = \kappa^2 q^4/4\pi r^8$ , and

$$\Sigma_\mu{}^\nu = \begin{pmatrix} \sigma_- & \sigma_{\text{in}} & 0 & 0 \\ 0 & \sigma_- & 0 & 0 \\ 0 & 0 & \sigma_+ & 0 \\ 0 & 0 & 0 & \sigma_+ \end{pmatrix}, \quad (8)$$

where  $\sigma_\pm = 1 \pm \frac{\kappa^2 l_p^2 q^2(v)}{4\pi r^4}$  and  $\sigma_{\text{in}} = \frac{2\kappa^2 l_p^2 \rho_{\text{in}}}{1 - 2\kappa^2 l_p^2 q^2(v)/4\pi r^4}$ . The field equations (4) can thus be written as

$$R_\mu{}^\nu(h) = \begin{pmatrix} -\frac{\kappa^2 q^2(v)}{8\pi r^4 \sigma_+} & \frac{e^{-\psi} \kappa^2 \rho_{\text{in}}}{\sigma_+ \sigma_-} & 0 & 0 \\ 0 & -\frac{\kappa^2 q^2(v)}{8\pi r^4 \sigma_+} & 0 & 0 \\ 0 & 0 & \frac{\kappa^2 q^2(v)}{8\pi r^4 \sigma_-} & 0 \\ 0 & 0 & 0 & \frac{\kappa^2 q^2(v)}{8\pi r^4 \sigma_-} \end{pmatrix}. \quad (9)$$

The strategy now is to solve for  $h_{\mu\nu}$  first and then use (2) and (8) to obtain  $g_{\mu\nu}$ . For this purpose, we define a line element for  $h_{\mu\nu}$  of the form

$$d\tilde{s}^2 = -F(v, x) e^{2\xi(v, x)} dv^2 + 2e^{\xi(v, x)} dv dx + \tilde{r}^2(v, x) d\Omega^2. \quad (10)$$

The field equations then imply that we can take  $\tilde{r}(v, x) = x$ ,  $e^{\xi(v,x)} = 1$ , and  $F(x, v) = 1 - 2M(x, v)/x$ , with  $M(x, v)$  satisfying

$$M_x = \frac{\kappa^2 q^2}{16\pi r^2}, \quad M_v = \frac{\kappa^2 \rho_{in} r^2}{2}, \quad (11)$$

where the relation (2) establishes that

$$r^2(x, v) = \frac{x^2 + \sqrt{x^4 + l_p^2 \kappa^2 q^2(v)/\pi}}{2}. \quad (12)$$

By direct integration of  $M_x$ , we find

$$M(x, v) = M_0 + \gamma(v) + \frac{\tilde{\kappa}^2 q^2(v)}{4} \int \frac{dx}{r^2} \Big|_{v=const}, \quad (13)$$

where  $M_0$  is a constant and  $\gamma(v)$  an integration function. Computing  $M_v$  and comparing with Eq. (11), one finds

$$\rho_{in} r^2 = \frac{2}{\kappa^2} \left[ \gamma_v + \frac{\kappa^2 q q_v}{8\pi} \int \frac{dx}{r^2 \sigma_+} \right], \quad (14)$$

which is fully consistent with  $\nabla_\mu (T_{flux}^{\mu\nu} + T_{em}^{\mu\nu}) = 0$ .

We have thus found a complete solution to the Palatini theory (7) coupled to a stream of charged null radiation. In this solution,  $q_v \equiv \partial_v q(v)$  and  $\gamma_v \equiv \partial_v \gamma(v)$  are free functions. Given the structure of the mass function in Eq. (13) and to make contact with previous results on static configurations [14] of the same model (7), we find it useful to write it as

$$M(x, v) = M_0 + \gamma(v) + \frac{r_q(v)^2}{4r_c(v)} \left( \int dz G_z \right) \Big|_{z=r/r_c}, \quad (15)$$

where  $r_q^2(v) \equiv \kappa^2 q^2/4\pi$ ,  $r_c(v) \equiv \sqrt{r_q(v)l_p}$ ,  $z(x, v) \equiv r(x, v)/r_c(v)$  and

$$G_z \equiv \frac{dG}{dz} = \frac{z^4 + 1}{z^4 \sqrt{z^4 - 1}}, \quad (16)$$

where we have used the relation  $dr/dx = \sigma_-^{1/2}/\sigma_+$  (at constant  $v$ ). From Eq. (16) the function  $G(z)$  can be written as an infinite power series and its form was given in [14]. Using the compact notation

$$M(x, v) = M(v) [1 + \delta_1(v) G(z)], \quad (17)$$

where  $M(v) = M_0 + \gamma(v) \equiv r_S(v)/2$  and  $\delta_1(v) = \frac{1}{2r_S(v)} \sqrt{\frac{r_q(v)^3}{l_p}}$ , the line element (6) becomes

$$ds^2 = - \left[ \frac{1}{\sigma_+} \left( 1 - \frac{1 + \delta_1(v) G(z)}{\delta_2(v) z \sigma_-^{1/2}} \right) - \frac{2l_p^2 \kappa^2 \rho_{in}}{\sigma_- (1 - \frac{2r_c^4}{r^4})} \right] dv^2 + \frac{2}{\sigma_+} dv dx + r^2(x, v) d\Omega^2, \quad (18)$$

where  $\delta_2(v) \equiv \frac{r_c(v)}{r_S(v)}$ . Eq. (18) constitutes the main result of this Letter.

### 3. Dynamic generation of wormholes

Consider now that the initial state is flat Minkowski space and assume that a charged perturbation of compact support propagates within the interval  $[v_i, v_f]$ . Given the relation (12), which can be

written as  $r^2(x, v) = \frac{x^2 + \sqrt{x^4 + 4r_c^4(v)}}{2}$ , it follows that for  $v < v_i$  we have  $r^2(x, v) = x^2$ , which extends from zero to infinity. Entering the  $v \geq v_i$  region, this radial function, which measures the area of the 2-spheres of constant  $x$  and  $v$ , never becomes smaller than

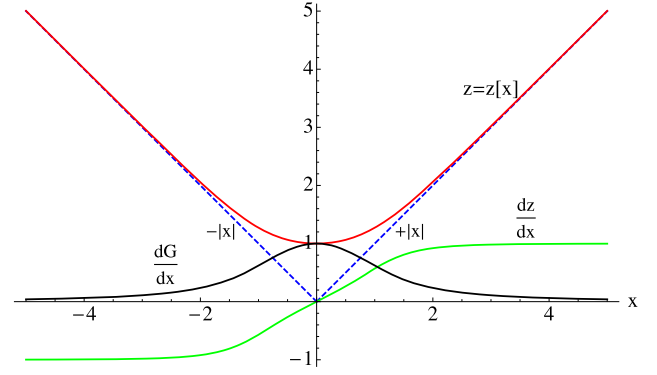


Fig. 1. The minimum of  $z(x)$  implies the existence of a wormhole extension of the geometry, with  $x$  covering the whole real axis  $-\infty < x < +\infty$ . Note the smoothness of  $dG/dx$  at  $x=0$ . In this plot,  $r_c = 1$ .

$r_c^2(v)$ , where this minimum is reached at  $x=0$ . If we now consider the region  $v > v_f$ , in which  $\rho_{in}$  is again zero, the result is a static geometry. One can verify [14] that in this static geometry curvature scalars generically diverge at  $x=0$  except if the charge-to-mass ratio  $\delta_1^f \equiv \delta_1(v \geq v_f)$  takes the value  $\delta_1^f = \delta_1^* \simeq 0.572$ . This  $\delta_1^*$  is a constant that appears in the series expansion of

$$G(z) \approx -1/\delta_1^* + 2\sqrt{z-1} + \dots \quad (19)$$

as  $z \rightarrow 1$ . The smoothness of the geometry when  $\delta_1^f = \delta_1^*$ , together with the fact that  $r(x)$  reaches a minimum at  $x=0$ , allow to naturally extend the coordinate  $x$  to the negative real axis, thus showing that the radial function  $r^2(x)$  bounces off to infinity as  $x \rightarrow -\infty$  (see Fig. 1). This puts forward the existence of a wormhole structure with its throat located at  $x=0$ .

It is important to note that the flux  $\Phi = \int_S *F$ , where  $*F$  is the 2-form dual to the Faraday tensor, through any closed 2-surface  $S$  enclosing  $x=0$  is non-zero and can be used to define the charge inside  $S$ ,  $\Phi = 4\pi q(v)$ , on a purely topological basis [2], i.e., without the need to introduce a metric or an affine connection. One thus finds that the density of lines of force at  $x=0$ , given by  $\Phi/4\pi r_c^2(v) = \sqrt{c^7/2\hbar G^2}$ , is constant and independent of the particular values of the charge and mass of the solution. This shows that the (topological) wormhole structure exists even when  $\delta_1(v) \neq \delta_1^*$  and, therefore, is insensitive to the existence of local curvature divergences.

We also note that for  $x \gg r_c(v)$ , the line element (18) quickly recovers the expected GR prediction in both the dynamic and the static case. Event horizons are thus expected in general, though there exist nonsingular naked configurations with  $\delta_1^f = \delta_1^*$  [16,14].

### 4. Discussion and summary

The generation of wormholes outlined above shows that a dynamical change in the space-time topology is possible when charged fluids interact with a manifold governed by a quadratic Palatini theory. This process differs radically in nature relative to the construction of general relativistic traversable wormholes, with the idealization of impulsive phantom radiation considered in the literature [9–11]. More specifically, in the latter it was shown that the adequate synchronization of the energies and the emission timing of two opposing streams of phantom radiation may support a static traversable wormhole [10]. The theory presented here generates static wormholes by means of a finite pulse of electrically charged radiation, without the need to keep two exotic-energy streams active continuously or to synchronize them across the wormhole.

Regarding the size of the wormholes, we note that if instead of using  $l_p^2 = \hbar G/c^3$  to characterize the curvature corrections one considers a different length scale, say  $l_\epsilon^2$ , then their area would be given by  $A_{WH} = (\frac{l_\epsilon}{l_p}) \frac{2N_q}{N_q} A_p$ , where  $A_p = 4\pi l_p^2$ ,  $N_q = |q/e|$  is the number of charges, and  $N_q^c \approx 16.55$ . Though this could allow to reach sizes orders of magnitude larger than the Planck scale, it does not seem very likely that macroscopic wormholes could arise from any viable theory of this form, though the effects of other matter/energy sources might be nontrivial.

The meaning and implications of classical singularities have been a subject of intense debate in the literature for years. Their existence in GR is generally interpreted as a signal of the limits of the theory, where the quantum effects of gravity should become relevant and an improved theory would be necessary. This is, in fact, the reason that motivates our study of quadratic corrections beyond GR.

It is thus remarkable that for a certain value of the charge-to-mass ratio all curvature scalars of the theory become finite. The absence of divergences in this particular case has allowed us to unveil a nontrivial topological structure, which turns out to be common to all the solutions of the theory, even to those with  $\delta_1 \neq \delta_1^*$ . This conclusion follows from the fact that the electric charge can be defined purely on topological grounds as the flux of a two-form through a closed 2-surface, i.e., without the need to introduce neither a metric nor a connection for this purpose. The independence of this flux on the particular value of the mass of the solutions, and the existence of a conserved density of lines of force at the surface  $x = 0$  ( $r = r_c$ ) confirms that all the solutions possess similar topological properties. The existence of geometric singularities when  $\delta_1 \neq \delta_1^*$  could thus be seen similarly as the situation that arises, for instance, when a sphere and a cube are compared: both have the same topology, but the geometry is ill-defined along the edges and at the vertices of the cube.

To the light of this example, the physical significance of curvature singularities loses strength. In fact, as pointed out above and shown in detail in [14], the curvature divergences for the static wormhole solutions arising in quadratic Palatini gravity with electrovacuum fields are much weaker than their counterparts in GR (from  $\sim 1/r^8$  in GR to  $\sim (\delta_1 - \delta_1^*)^2/(r - r_c)^3$  in our model), which signals a smoothing out of the geometry as compared to GR as the minimal surface  $r = r_c$  is approached. As a result of the existence of this minimal surface, the function  $r^2$  never drops below the scale  $r_c^2$  and implies that the total energy stored in the electric field is finite (see [19,16] for details), which clearly contrasts with the infinite result that GR yields. Therefore, even though curvature scalars may diverge, physical magnitudes such as total mass, energy, and electric charge are insensitive to those divergences, which puts forward the importance of the topology to address physical questions associated to curvature divergences.

In the dynamical case discussed here, transient curvature divergences may also arise due to the  $\rho_{in}$  term in the  $g_{\nu\nu}$  component of the metric. A glance at (18) indicates that there might be divergences at  $r = r_c(v)$ , where  $\sigma_-$  vanishes, and at  $r^4 = 2r_c^4(v)$ . However, if one considers the case  $\delta_1(v) = \delta_1^*$ , which avoids the divergences in the static case, a short algebra confirms that the geometry is smooth at  $r = r_c(v)$  also in the dynamical case, with  $\rho_{in} \rightarrow 0$  there. Therefore, no large metric fluctuations occur at the wormhole throat for this choice of  $\delta_1(v)$ . At  $r = 2^{1/4}r_c(v)$ , a divergence occurs, though it disappears when the infalling flux ceases. The details of this transient will be discussed elsewhere [26].

We note that since in our theory the field equations outside the matter sources recover those of vacuum GR, Birkhoff's theorem must hold in those regions. This means that for  $v < v_i$  we

have Minkowski space, whereas for  $v > v_f$  we have a Reissner–Nordström-like geometry. The departure from Reissner–Nordström is due to the Planck scale corrections of the Lagrangian, which are excited by the presence of an electric field, and only affect the microscopic structure, which is of order  $\sim r_c(v)$ . Due to the spherical symmetry and the second-order character of the field equations, Birkhoff's theorem guarantees the staticity of the solutions for  $v > v_f$ .

In concluding, we have found an exact analytical solution for the dynamical process of collapse of a null fluid carrying energy and electric charge in a quadratic extension of GR formulated à la Palatini. This extends the well-known Vaidya–Bonnor solution of GR [27] to a new scenario that allows to explore in detail new physics at the Planck scale. Focusing on the initial and final static configurations, we have shown that wormholes can be formed out of Minkowski space by means of a pulse of charged radiation, which contrasts with previous approaches in the literature requiring artificial configurations and synchronization of two opposing streams of phantom energy.

Though we have considered an idealized pressureless and ultra-relativistic charged fluid in a highly symmetric scenario, we believe that our results support the view that space–time could have a foam-like microstructure with wormholes generated by fluctuations of the quantum vacuum involving the spontaneous creation/annihilation of very energetic charged particles. In fact, all spherically symmetric charged solutions of our theory able to extend their electric field down to scales of order  $r_c \sim l_p$  possess a wormhole structure. We understand that pressure and other dispersion effects should act so as to prevent the effective concentration of charge and energy on scales of order  $\sim l_p$ , thus suppressing wormhole production in low energy scenarios.

Let us finally stress that though the wormhole structures found here may develop curvature divergences, they are characterized by well-defined and finite electric charge and total energy [19,16]. The physical role that such divergences may play, if any, is thus uncertain and requires a more in-depth analysis. To fully understand these issues, our model should be improved to include several important aspects, such as the presence of gauge fields and/or fermionic degrees of freedom, or to consider the dynamics of counter-streaming effects due to the presence of simultaneous ingoing and outgoing fluxes. The impact that higher-order curvature corrections in the action could have on the wormhole solutions found here is also a fundamental test to assess the robustness of our conclusions. In this sense, it has been recently found [19] that a Born–Infeld-like gravity theory [17,18], which contains higher than quadratic contributions in the Ricci tensor, also leads to wormhole solutions of the type found here when formulated à la Palatini. This supports the view that dynamical topology change might be a robust feature of metric-affine (Palatini) theories.

## Acknowledgements

F.S.N.L. acknowledges financial support of the Fundação para a Ciência e Tecnologia through an Investigador FCT Research contract, with reference IF/00859/2012, funded by FCT/MCTES (Portugal), and grants CERN/FP/123615/2011 and CERN/FP/123618/2011. G.J.O. is supported by the Spanish grant FIS2011-29813-C02-02, the Consolider Program CPANPHY-1205388, and the JAE-doc program of the Spanish Research Council (CSIC). D.R.G. is supported by CNPq (Brazilian agency) through project No. 561069/2010-7 and thanks the hospitality and partial support of the Theoretical Physics Department of the University of Valencia. This work has also been supported by CNPq through project No. 301137/2014-5.

## References

- [1] J.A. Wheeler, *Phys. Rev.* 97 (1955) 511.
- [2] C.W. Misner, J.A. Wheeler, *Ann. Phys.* 2 (1957) 525.
- [3] I.H. Redmount, W.-M. Suen, *Phys. Rev. D* 47 (1993) 2163; I.H. Redmount, W.-M. Suen, *Phys. Rev. D* 49 (1994) 5199.
- [4] M. Visser, *Lorentzian Wormholes: From Einstein to Hawking*, AIP, New York, 1995.
- [5] A. Anderson, B.S. DeWitt, *Found. Phys.* 16 (1986) 91.
- [6] R. Geroch, *J. Math. Phys.* 8 (1967) 782; S.W. Hawking, *Phys. Rev. D* 46 (1992) 603.
- [7] M. Visser, *Phys. Rev. D* 41 (1990) 1116; M. Visser, *Phys. Rev. D* 43 (1991) 402.
- [8] T.A. Roman, *Phys. Rev. D* 47 (1993) 1370.
- [9] L.A. Gergely, *Phys. Rev. D* 65 (2002) 127503; L.A. Gergely, *Phys. Rev. D* 58 (1998) 084030.
- [10] S.A. Hayward, *Phys. Rev. D* 65 (2002) 124016.
- [11] S.A. Hayward, H. Koyama, *Phys. Rev. D* 70 (2004) 101502.
- [12] M. Morris, K.S. Thorne, *Am. J. Phys.* 56 (1988) 395.
- [13] T. Harko, F.S.N. Lobo, M.K. Mak, S.V. Sushkov, *Phys. Rev. D* 87 (2013) 067504.
- [14] G.J. Olmo, D. Rubiera-Garcia, *Phys. Rev. D* 86 (2012); G.J. Olmo, D. Rubiera-Garcia, *Int. J. Mod. Phys. D* 21 (2012) 1250067; G.J. Olmo, D. Rubiera-Garcia, *Eur. Phys. J. C* 72 (2012) 2098.
- [15] G.J. Olmo, D. Rubiera-Garcia, *Phys. Rev. D* 88 (2013) 084030.
- [16] F.S.N. Lobo, G.J. Olmo, D. Rubiera-Garcia, *J. Cosmol. Astropart. Phys.* 07 (2013) 011.
- [17] S. Deser, G.W. Gibbons, *Class. Quantum Gravity* 15 (1998) L35.
- [18] M. Bañados, P.G. Ferreira, *Phys. Rev. Lett.* 105 (2010) 011101.
- [19] G.J. Olmo, D. Rubiera-Garcia, H. Sanchis-Alepuz, arXiv:1311.0815 [hep-th].
- [20] L. Parker, D.J. Toms, *Quantum Field Theory in Curved Space-Time: Quantized Fields and Gravity*, Cambridge University Press, UK, 2009; N.D. Birrel, P.C.W. Davies, *Quantum Fields in Curved Space*, Cambridge University Press, UK, 1982.
- [21] M. Green, J. Schwarz, E. Witten, *Superstring Theory*, Cambridge University Press, UK, 1987; T. Ortin, *Gravity and Strings*, Cambridge University Press, UK, 2004.
- [22] J.A.R. Cembranos, *Phys. Rev. Lett.* 102 (2009) 141301.
- [23] G.J. Olmo, P. Singh, *J. Cosmol. Astropart. Phys.* 0901 (2009) 030.
- [24] M. Bojowald, *Living Rev. Relativ.* 11 (2008) 4; G.A. Mena-Marugan, *AIP Conf. Proc.* 1130 (2009) 89; G.A. Mena-Marugan, *J. Phys. Conf. Ser.* 314 (2011) 012012; A. Ashtekar, P. Singh, *Class. Quantum Gravity* 28 (2011) 213001.
- [25] C. Barragan, G.J. Olmo, H. Sanchis-Alepuz, *Phys. Rev. D* 80 (2009) 024016; C. Barragan, G.J. Olmo, *Phys. Rev. D* 82 (2010) 084015.
- [26] F.S.N. Lobo, J. Martinez-Asencio, G.J. Olmo, D. Rubiera-Garcia, in preparation.
- [27] W.B. Bonnor, P.C. Vaidya, *Gen. Relativ. Gravit.* 1 (1970) 127.