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STEPS TOWARD DETERMINATION OF THE SIZE AND STRUCTURE OF THE BROAD-LINE REGION IN ACTIVE GALACTIC NUCLEI. IX. ULTRAVIOLET OBSERVATIONS OF FAIRALL


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Reinhold, P.; Pascual, P.M.; Allion, D.; Clavel, J.; Crenshaw, D. Michael; Horne, K.; Kriss, Gerard A.; Krolik, Julian Henry; Malkan, Matthew A.; Netzer, Hagai; O'Brien, Paul T.; Peterson, Bradley M.; Reichert, G.A.; Wamsteker, W.; Alexander, T.; Barr, P.; Blandford, R.D.; Bergman, J.N.; Carone, T.E.; Clements, S.; Courvoisier, T.-J.; De Robertis, M.M.; Dietrich, M.; Dottori, H.; Edelson, R.A.; Filippenko, A. V.; Gaskell, C. Martin; Huchra, J. P.; Hutchings, J.B.; Kollatschny, W.; Koratkar, A.P.; Korista, Kirk T.; Laor, Ari; MacAlpine, G.M.; Martin, P. G.; Maoz, D.; McCollum, B.; Morris, S.L.; Perola, G.C.; Pogge, Richard William; Ptak, R.L.; Recondo-Gonzalez, M.C.; Rodriguez-Espinoza, J.M.; Rokaki, E.L.; Santos-Lleo, M.; Sekiguchi, K.; Schull, J.M.; Snuders, M.A.J.; Sparke, L.S.; Stirpe, G.M.; Stoner, R.E.; Sun, W.-H.; Wagner, S.J.; Wanders, I.; Wilkes, Belinda J.; Winge, C.; and Zheng, W., "STEPS TOWARD DETERMINATION OF THE SIZE AND STRUCTURE OF THE BROAD-LINE REGION IN ACTIVE GALACTIC NUCLEI. IX. ULTRAVIOLET OBSERVATIONS OF FAIRALL" (1997). *Martin Gaskell Publications*. 8.

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STEPS TOWARD DETERMINATION OF THE SIZE AND STRUCTURE OF THE
BROAD-LINE REGION IN ACTIVE GALACTIC NUCLEI. IX. ULTRAVIOLET
OBSERVATIONS OF FAIRALL 9

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Received 1996 October 24; accepted 1996 December 16

ABSTRACT

An 8 month monitoring campaign on the Seyfert 1 galaxy Fairall 9 has been conducted with the *International Ultraviolet Explorer* in an attempt to obtain reliable estimates of continuum-continuum and continuum–emission-line delays for a high-luminosity active galactic nucleus (AGN). While the results of this campaign are more ambiguous than those of previous monitoring campaigns on lower luminosity sources, we find general agreement with the earlier results: (1) there is no measurable lag between ultraviolet continuum bands, and (2) the measured emission-line time lags are very short. It is especially notable that the Ly α + N v emission-line lag is about 1 order of magnitude smaller than determined from a previous campaign by Clavel, Wamsteker, & Glass (1989) when Fairall 9 was in a more luminous state. In other well-monitored sources, specifically NGC 5548 and NGC 3783, the highest ionization lines are found to respond to continuum variations more rapidly than the lower ionization lines, which suggests a radially ionization-stratified broad-line region. In this case, the results are less certain, since none of the emission-line lags are very well determined. The best-determined emission line lag is Ly α + N v, for which we find that the centroid of the continuum–emission-line cross-correlation function is $\tau_{\text{cent}} \approx 14\text{--}20$ days. We measure a lag $\tau_{\text{cent}} \lesssim 4$ days for He II $\lambda 1640$; this result is consistent with the ionization-stratification pattern seen in lower luminosity sources, but the relatively large uncertainties in the emission-line lags measured here cannot rule out similar lags for Ly α + N v and He II $\lambda 1640$ at a high level of significance. We are unable to determine a reliable lag for C IV $\lambda 1550$, but we note that the profiles of the variable parts of Ly α and C IV $\lambda 1550$ are not the same, which does not support the hypothesis that the strongest variations in these two lines arise in the same region.

Subject headings: galaxies: active — galaxies: individual (Fairall 9) — galaxies: Seyfert — ultraviolet: galaxies

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1. INTRODUCTION

With the discovery of coordinated continuum and emission-line variability in active galactic nuclei (AGNs), it was realized by many authors that variability might provide a very valuable tool for unveiling the structure of the innermost regions of these extremely luminous objects. The rapid response of the emission lines to continuum variations strongly argues that the line emission is driven by photoionization by the AGN central source, which is generally supposed to be an accretion disk surrounding a supermassive black hole. Thus, by measuring the time-delayed response of the emission lines to continuum variations, it is in principle possible to infer the distribution and kinematics of the line-emitting gas through a process known as “reverberation mapping” (e.g., Blandford & McKee 1982). Multiwavelength measurements of continuum and emission-line variability also provide important observational constraints on the central source itself by providing a probe (1) of the unobservable ionizing continuum, through the response of the emission lines to changes in the observable continuum, and (2) of the structure of the accretion disk itself, since in any model with a radial temperature gradient (such as a thin accretion disk), different continuum wavelengths arise principally at different disk radii.

On account of the rapid and irregular nature of the continuum variations, it was also realized early that enormous observational efforts were needed to exploit the potential of variability studies. The techniques to obtain the necessary information (reviewed in Gondhalekar, Horne, & Peterson 1994) require long series of high-quality spectroscopic data, preferably taken at nearly regular intervals. The *International Ultraviolet Explorer (IUE)* is extremely well suited to this task, which motivated the establishment of international consortia in the late 1980s to conduct a series of major observing campaigns using both *IUE* and ground-based facilities to obtain the necessary spectroscopic observations. The largest of these groups, the International AGN Watch (Alloin et al. 1994), has now completed multiwavelength programs on NGC 5548 (Clavel et al. 1991, [Paper I in this series]; Peterson et al. 1991 [Paper II], 1992 [Paper III], 1994 [Paper VII]; Dietrich et al. 1993 [Paper IV]; Maoz et

al. 1993; Romanishin et al. 1995; Korista et al. 1995 [Paper VIII]), NGC 3783 (Reichert et al. 1994 [Paper V]; Stirpe et al. 1994 [Paper VI]; Alloin et al. 1995), and NGC 4151 (Crenshaw et al. 1996; Kaspi et al. 1996; Warwick et al. 1996; Edelson et al. 1996). The general results of these AGN monitoring campaigns can be summarized as follows:

1. The UV and optical continua vary simultaneously within the accuracy of the data and the sampling frequency (typically ± 2 days or less).

2. The emission-line lags (or mean response times to continuum variations) are found to be quite small, typically ranging between a few days to several weeks.

3. The high-ionization lines (e.g., N v $\lambda 1240$, He II $\lambda 1640$, and C IV $\lambda 1550$) respond faster to continuum changes than the low-ionization lines (e.g., Mg II $\lambda 2800$, H β).

The luminosity range for which these results have been obtained [$10^{39} \lesssim L_{\lambda}(1450 \text{ \AA}) \lesssim 10^{40} \text{ ergs s}^{-1} \text{ \AA}^{-1}$] is still too small to test critically the expected $R \propto L^{1/2}$ relationship between the broad-line region (BLR) size R and continuum luminosity L predicted by simple photoionization equilibrium arguments (e.g., Peterson 1993), and extension of variability studies to higher luminosity is therefore necessary. In this regard, the luminous Seyfert galaxy Fairall 9 [$L_{\lambda}(1450 \text{ \AA}) > 10^{41} \text{ ergs s}^{-1} \text{ \AA}^{-1}$; $z = 0.047$] is of interest, as it is one of the few high-luminosity sources that have been monitored with *IUE*. *IUE* observations with an average sampling interval of about 90 days suggested emission-line lags of about 115 ± 70 days for Ly α and about 200 ± 80 days for C IV $\lambda 1550$ (Clavel, Wamsteker, & Glass 1989; see also Koratkar & Gaskell 1989; Lub & de Ruiter 1992; Shields, Ferland, & Peterson 1995; Recondo-González et al. 1997) relative to the UV continuum during a time when Fairall 9 underwent a dramatic change in luminosity. Between 1978 and 1985, the UV flux of Fairall 9 decreased nearly monotonically by a factor of about 30.

In view of its well-documented UV variability history and its importance in the R - L relationship, we decided to undertake a more intensive monitoring program on Fairall 9. Our goal was to obtain observations with the same sampling frequency (about one observation every 4 days) and dura-

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TABLE 1
LOG OF OBSERVATIONS

SWP Image Number (1)	UT Date (2)	UT Start (hh:mm:ss) (3)	JD (-2,449,000) (4)	Exposure Time (minutes) (5)	Number of Segments (6)
50641 ^a	1994 Apr 28	04:26:11	470.684	140	7
50671	1994 May 01	03:37:45	473.652	110	5
50700	1994 May 05	04:26:31	477.684	100	6
50731	1994 May 10	04:07:40	482.672	116	6
50762	1994 May 14	04:15:34	486.676	125	5
50807	1994 May 18	04:15:26	490.676	130	5
50860	1994 May 22	04:40:16	494.695	105	4
50898	1994 May 25	04:28:58	497.688	115	4
50936	1994 May 29	04:37:18	501.691	103	5
50994	1994 Jun 03	10:01:08	506.918	90	5
51034	1994 Jun 07	09:52:11	510.910	120	6
51068	1994 Jun 11	10:33:10	514.941	100	5
51101	1994 Jun 15	05:56:38	518.746	90	5
51118	1994 Jun 18	05:55:23	521.746	80	4
51119	1994 Jun 18	08:02:33	521.836	80	4
51120	1994 Jun 18	10:17:51	521.930	70	3
51129	1994 Jun 19	06:02:21	522.750	80	4
51142	1994 Jun 20	09:04:24	523.879	80	3
51143	1994 Jun 20	11:05:39	523.961	80	3
51163	1994 Jun 22	22:22:29	526.434	130	5
51236	1994 Jun 26	22:33:21	530.441	115	5
51271	1994 Jun 30	02:08:09	533.590	92	4
51317	1994 Jul 04	00:30:41	537.520	85	4
51375	1994 Jul 09	00:24:16	542.516	110	5
51418	1994 Jul 13	01:04:46	546.543	87	3
51485	1994 Jul 16	23:56:50	550.496	115	4
51546 ^a	1994 Jul 21	00:11:04	554.508	93	4
51596 ^a	1994 Jul 25	05:44:51	558.738	85	5
51636	1994 Jul 28	05:52:08	561.746	85	4
51637	1994 Jul 28	08:11:03	561.840	85	4
51677	1994 Aug 01	05:37:08	565.734	85	4
51678	1994 Aug 01	07:47:03	565.824	60	3
51717 ^{a, b}	1994 Aug 05	05:08:57	569.715	90	1
51741 ^b	1994 Aug 08	02:08:09	572.590	90	1
51742 ^b	1994 Aug 08	03:59:47	572.664	50	1
51792 ^{a, b}	1994 Aug 13	01:46:10	577.574	90	1
51793 ^b	1994 Aug 13	03:45:50	577.656	75	1
51832 ^b	1994 Aug 17	01:45:48	581.574	86	1
51833 ^b	1994 Aug 17	03:36:05	581.648	70	1
51861 ^b	1994 Aug 21	01:49:12	585.574	85	1
51862 ^b	1994 Aug 21	03:48:34	585.660	75	1
51911	1994 Aug 24	22:07:08	589.422	90	4
51973	1994 Aug 28	21:50:07	593.410	130	7
51998	1994 Sep 01	20:12:50	597.340	117	5
52039	1994 Sep 05	20:03:05	601.336	113	5
52072	1994 Sep 09	20:29:44	605.352	115	4
52106	1994 Sep 13	20:48:25	609.367	100	4
52142	1994 Sep 17	16:10:44	613.172	130	5
52174	1994 Sep 21	20:46:26	617.367	95	5
52194 ^a	1994 Sep 24	23:55:24	620.497	35	2
52195	1994 Sep 25	01:03:41	620.543	90	4
52211	1994 Sep 25	20:46:46	621.367	100	4
52235	1994 Sep 28	03:45:56	623.656	90	4
52236	1994 Sep 28	06:13:53	623.758	35	1
52279 ^a	1994 Oct 02	02:18:04	627.596	90	4
52343 ^a	1994 Oct 06	21:53:13	632.412	95	4
52344	1994 Oct 07	00:14:33	632.508	60	2
52370 ^a	1994 Oct 10	21:56:02	636.414	80	4
52371 ^a	1994 Oct 10	23:57:39	636.498	50	2
52406	1994 Oct 15	02:16:39	640.594	80	5
52516	1994 Oct 19	22:05:33	645.422	100	5
52624	1994 Oct 23	21:50:52	649.410	82	4
52625	1994 Oct 23	23:54:31	649.496	55	2
52675	1994 Oct 27	21:54:02	653.414	70	3
52676	1994 Oct 27	23:47:07	653.492	55	2
52702	1994 Oct 31	21:29:18	657.395	70	3
52703	1994 Oct 31	23:28:26	657.477	45	2
52724	1994 Nov 04	14:50:23	661.117	90	4
52762	1994 Nov 08	16:37:19	665.191	110	5
52790	1994 Nov 12	16:36:56	669.191	110	5
52820	1994 Nov 17	16:24:58	674.184	95	5

TABLE 1—Continued

SWP Image Number (1)	UT Date (2)	UT Start (hh:mm:ss) (3)	JD (−2,449,000) (4)	Exposure Time (minutes) (5)	Number of Segments (6)
52845	1994 Nov 20	16:59:31	677.207	70	3
52895	1994 Nov 24	16:40:42	681.195	100	4
52963	1994 Dec 02	13:58:18	689.082	130	6
52997	1994 Dec 06	14:41:52	693.113	100	4
53034	1994 Dec 10	14:24:25	697.102	100	5
53053	1994 Dec 14	06:36:57	700.776	70	3
53106 ^a	1994 Dec 17	14:35:48	704.108	90	4
53160 ^a	1994 Dec 22	15:18:53	709.138	20	1
53191	1994 Dec 26	14:38:53	713.109	75	4

^a Image dropped because the target was not properly centered within the aperture.

^b Exposure guided on offset star in the FES field.

tion (about 8 months) used in the original NGC 5548 and NGC 3783 *IUE* campaigns in order to facilitate a more direct comparison with these lower luminosity Seyfert galaxies. The observations taken during this campaign are described in § 2, the data are presented in § 3, and the results are discussed in § 4. Finally, the main conclusions are summarized in § 5.

2. OBSERVATIONS

From 1994 April 28 to December 26 ($\Delta T = 242$ days), Fairall 9 was observed with *IUE* at nearly regular intervals of about 4 days. All observations were made through the large aperture ($10'' \times 20''$) of the short wavelength spectrograph (1150–1950 Å) in the low-dispersion mode (~ 6 Å). During the 63 different observing epochs of 4 hr (occasionally 8 hr) each, a total of 80 spectra were recorded with the short wavelength prime (SWP) camera (Harris & Sonneborn 1987).

The fine error sensor (FES) anomaly, which is attributed to scattered light entering the *IUE* telescope tube, has not been reported to contaminate any SWP image, since it was found for the first time in 1991 (Weinstein & Carini 1992). However, the solar-like spectrum of this scattered light affects the FES, a photomultiplier with a bandpass from 3900 Å to 9000 Å (S-20 cathode), with a sensitivity strongly peaked at 4600 Å (Holm & Rice 1981). The FES is used for target acquisition and fine guiding during exposures; hence, these operations are significantly influenced by the presence of the scattered light. For most *IUE* observations of Fairall 9, there was no star within the FES field of view suitable for tracking. In these cases, the total integration time was achieved by adding up a number of shorter exposures, typically 15–30 minutes long. After each of these segments, the *IUE* telescope was moved to point to a nearby bright star to check the pointing stability.

As a consequence of the scattered light, the photometric capabilities of the FES could not be used to obtain a simultaneous optical light curve for Fairall 9. However, an optical light curve has been obtained from the nearly simultaneous ground-based optical campaign, and these data will be presented separately (Santos-Lleó et al. 1997).

The log of the *IUE* observations (Table 1) includes, for each SWP image number (col. [1]), the UT date (col. [2]), time (col. [3]), and Julian Date (col. [4]) of the start of the exposure, the total integration time (col. [5]), and the number of segments in which this integration was achieved (col. [6]).

3. DATA ANALYSIS AND RESULTS

All the images have been processed with the new *IUE*

Final Archive processing software, NEWSIPS (Nichols et al. 1993). The major advantage of this software is the method of the raw science data registration, which significantly reduces the fixed pattern noise in the *IUE* images and improves the photometric correction. NEWSIPS also includes a weighted slit extraction method (Kinney, Bohlin, & Neill 1991) and rederived absolute flux calibrations (González-Riestra, Cassatella, & de la Fuente 1992). All of these improvements result in higher photometric accuracy and higher signal-to-noise ratio spectra, when compared to the data processed with the old software (IUESIPS; Turnrose & Thompson 1984).

Because of the acquisition and tracking problems mentioned in § 2, we have checked every image to identify those in which Fairall 9 may have gone out of the aperture for some time during the total exposure. A first test was made by checking the FES recentering errors after those exposures that were not guided; a benchmark test made with an *IUE* standard star has shown that up to ± 25 units in the *X*-axis and ± 7 units in the *Y*-axis can be accepted as tolerance limits to ensure that the target did remain in the aperture (Rodríguez-Pascual 1997). Another check is based on the position within the *IUE* SWP camera at which the spectrum was recorded. In the spatially resolved spectrum (SILO file of NEWSIPS processing software), we have measured the shift in the spatial direction with respect to the center of the aperture. The wavelength of the peak of the strongest emission lines (*Ly* α and C iv $\lambda 1550$) has been used as reference for the shift in the spectral direction. The 12 spectra recorded more than 4 pixels away from the expected position in either spatial or spectral directions have been rejected and are not considered in this paper.

3.1. Continuum and Emission-Line Measurements

As shown in Figure 1, the continuum flux density has been measured in four bands that are free of obvious emission or absorption features in the mean spectrum: 1380–1400 Å, 1500–1520 Å, 1780–1810 Å, and 1860–1900 Å (observed wavelengths). The corresponding flux densities with their associated uncertainties are listed in Table 2 and plotted in Figure 2. Note that the integration times have been adjusted to reach an optimal exposure in the peaks of the strongest emission lines (*Ly* α , C iv $\lambda 1550$) at the cost of underexposing the continuum, where the flux density is 5 times weaker; this results in relatively large uncertainties in the continuum flux-density measurements.

With regard to the emission lines, no attempt has been made to deconvolve the observed profiles into components. Consequently, the *Ly* α flux includes the flux from the N v $\lambda 1240$ line as well, and we will refer to the measured flux as

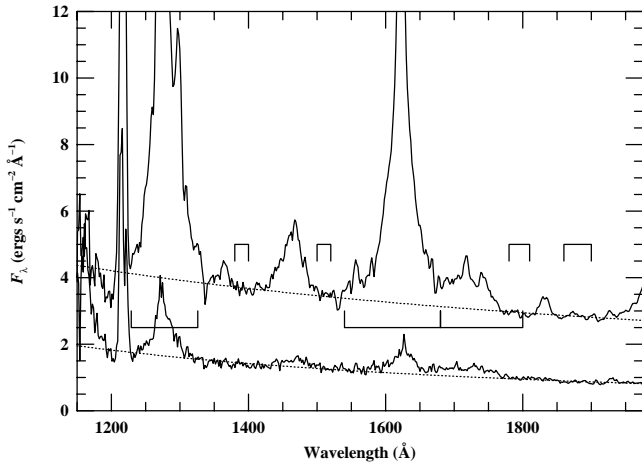


FIG. 1.—Mean (top) and rms (bottom) SWP spectra of Fairall 9. The bands in which the continuum has been measured are marked above the spectrum, and the dotted lines show power-law fits to the fluxes in these bands in both the mean and rms spectra. The Ly α + N v, C iv λ 1550, and He II λ 1640 integration ranges are marked below the mean spectrum.

Ly α + N v. The measurements of C iv λ 1550 and He II λ 1640 are to some extent (<2%) affected by the overlap of the red wing of C iv λ 1550 with the blue wing of He II λ 1640. The observed wavelength ranges used for integration of the Ly α + N v, Si iv λ 1400, C iv λ 1550, and He II λ 1640 line fluxes are, respectively, 1229–1326 Å, 1400–1500 Å, 1540–1680 Å, and 1680–1800 Å, as marked in Figure 1.

The continuum underneath the emission lines has been estimated by linear interpolation between the four continuum bands described above. The absorption lines have not

TABLE 2

CONTINUUM FLUXES

Julian Date (-2,449,000) (1)	$F_{\lambda}(1390 \text{ \AA})^a$ (2)	$F_{\lambda}(1510 \text{ \AA})^a$ (3)	$F_{\lambda}(1795 \text{ \AA})^a$ (4)	$F_{\lambda}(1880 \text{ \AA})^a$ (5)
473.652.....	2.45 ± 0.16	2.06 ± 0.13	1.80 ± 0.10	2.06 ± 0.09
477.684.....	2.06 ± 0.22	2.49 ± 0.12	2.00 ± 0.16	2.00 ± 0.15
482.672.....	2.59 ± 0.16	2.30 ± 0.09	1.91 ± 0.10	1.85 ± 0.15
486.676.....	2.62 ± 0.08	2.60 ± 0.11	1.99 ± 0.10	2.26 ± 0.09
490.676.....	2.56 ± 0.20	2.39 ± 0.11	2.16 ± 0.15	2.13 ± 0.09
494.695.....	2.48 ± 0.14	2.32 ± 0.12	1.97 ± 0.13	1.88 ± 0.09
497.688.....	2.96 ± 0.13	2.70 ± 0.14	2.02 ± 0.11	2.21 ± 0.09
501.691.....	3.11 ± 0.13	2.87 ± 0.10	2.63 ± 0.20	2.55 ± 0.11
506.918.....	2.89 ± 0.29	2.70 ± 0.22	2.49 ± 0.20	2.27 ± 0.09
510.910.....	3.40 ± 0.30	3.61 ± 0.13	2.75 ± 0.21	2.64 ± 0.12
514.941.....	2.88 ± 0.25	3.17 ± 0.16	2.43 ± 0.19	2.44 ± 0.10
518.746.....	3.46 ± 0.33	3.47 ± 0.15	2.78 ± 0.48	2.97 ± 0.15
521.746.....	3.91 ± 0.22	3.46 ± 0.16	2.66 ± 0.23	2.99 ± 0.18
521.836.....	3.44 ± 0.19	3.54 ± 0.15	2.69 ± 0.23	2.78 ± 0.11
521.930.....	3.55 ± 0.17	3.54 ± 0.14	2.55 ± 0.27	2.56 ± 0.13
522.750.....	3.56 ± 0.26	3.46 ± 0.24	2.87 ± 0.15	2.67 ± 0.24
523.879.....	3.77 ± 0.19	3.89 ± 0.15	3.11 ± 0.10	2.56 ± 0.09
523.961.....	5.47 ± 0.75	3.66 ± 0.16	2.63 ± 0.15	2.82 ± 0.16
526.434.....	3.42 ± 0.33	3.36 ± 0.14	2.66 ± 0.15	2.64 ± 0.09
530.441.....	3.36 ± 0.30	3.30 ± 0.18	2.88 ± 0.15	2.69 ± 0.09
533.590.....	3.22 ± 0.18	3.16 ± 0.12	2.36 ± 0.18	2.34 ± 0.09
537.520.....	2.96 ± 0.30	2.69 ± 0.13	2.48 ± 0.20	2.18 ± 0.13
542.516.....	3.00 ± 0.16	2.56 ± 0.12	2.30 ± 0.32	2.20 ± 0.11
546.543.....	3.01 ± 0.14	2.58 ± 0.08	2.04 ± 0.15	1.99 ± 0.10
550.496.....	2.17 ± 0.17	2.24 ± 0.14	2.06 ± 0.12	2.19 ± 0.10
561.746.....	2.76 ± 0.33	2.41 ± 0.16	2.00 ± 0.15	2.31 ± 0.13
561.840.....	2.86 ± 0.32	2.61 ± 0.14	2.47 ± 0.21	2.36 ± 0.13
565.734.....	3.07 ± 0.39	2.73 ± 0.14	2.39 ± 0.26	2.42 ± 0.13
565.824.....	3.33 ± 0.46	3.13 ± 0.22	2.76 ± 0.18	2.31 ± 0.12
572.590.....	2.62 ± 0.18	2.93 ± 0.10	2.28 ± 0.09	2.51 ± 0.11
572.664.....	3.04 ± 0.37	2.74 ± 0.23	2.62 ± 0.33	2.28 ± 0.15
577.656.....	3.22 ± 0.25	3.29 ± 0.18	2.90 ± 0.20	2.35 ± 0.14
581.574.....	2.99 ± 0.24	2.60 ± 0.14	2.45 ± 0.15	2.53 ± 0.21
581.648.....	2.91 ± 0.37	2.90 ± 0.20	2.46 ± 0.21	2.80 ± 0.19
585.574.....	3.55 ± 0.13	3.07 ± 0.12	2.90 ± 0.20	2.42 ± 0.18
585.660.....	3.70 ± 0.17	3.17 ± 0.17	3.05 ± 0.14	2.37 ± 0.16
589.422.....	3.41 ± 0.21	3.25 ± 0.18	2.78 ± 0.13	2.82 ± 0.14
593.410.....	3.72 ± 0.19	3.41 ± 0.16	2.70 ± 0.21	2.67 ± 0.11
597.340.....	4.15 ± 0.23	3.53 ± 0.10	3.03 ± 0.14	3.01 ± 0.11
601.336.....	4.10 ± 0.28	3.81 ± 0.13	2.73 ± 0.22	2.81 ± 0.15
605.352.....	3.87 ± 0.23	3.76 ± 0.18	2.67 ± 0.17	2.80 ± 0.13
609.367.....	3.27 ± 0.28	3.60 ± 0.23	2.97 ± 0.13	2.70 ± 0.15
613.172.....	3.12 ± 0.21	3.11 ± 0.13	2.69 ± 0.17	2.50 ± 0.10
617.367.....	3.08 ± 0.18	3.21 ± 0.15	2.61 ± 0.09	2.57 ± 0.10
620.543.....	3.41 ± 0.27	3.42 ± 0.22	2.92 ± 0.14	2.67 ± 0.13
621.367.....	3.62 ± 0.23	3.66 ± 0.14	2.93 ± 0.15	2.90 ± 0.11
623.656.....	3.76 ± 0.29	3.25 ± 0.12	2.68 ± 0.17	2.60 ± 0.10
623.758.....	4.06 ± 0.52	2.83 ± 0.26	2.49 ± 0.21	2.79 ± 0.20
632.508.....	3.20 ± 0.38	2.78 ± 0.18	2.73 ± 0.19	2.60 ± 0.13
640.594.....	5.38 ± 0.19	5.24 ± 0.15	4.17 ± 0.17	4.07 ± 0.13
645.422.....	5.75 ± 0.24	5.23 ± 0.07	4.15 ± 0.15	4.19 ± 0.19
649.410.....	6.22 ± 0.22	5.68 ± 0.10	4.22 ± 0.19	3.97 ± 0.16
649.496.....	5.78 ± 0.30	5.28 ± 0.22	4.57 ± 0.26	4.09 ± 0.21
653.414.....	6.39 ± 0.31	5.97 ± 0.28	5.00 ± 0.20	4.29 ± 0.21
653.492.....	5.92 ± 0.27	5.84 ± 0.24	4.48 ± 0.20	4.43 ± 0.21
657.395.....	6.52 ± 0.39	5.69 ± 0.28	4.79 ± 0.25	4.49 ± 0.15
657.477.....	6.11 ± 0.61	5.54 ± 0.27	4.83 ± 0.27	4.45 ± 0.15
661.117.....	6.99 ± 0.34	6.48 ± 0.14	4.70 ± 0.24	4.77 ± 0.13
665.191.....	6.68 ± 0.26	6.27 ± 0.13	5.10 ± 0.20	4.65 ± 0.21
669.191.....	6.91 ± 0.25	6.29 ± 0.20	5.10 ± 0.18	4.58 ± 0.13
674.184.....	7.65 ± 0.24	6.66 ± 0.15	4.87 ± 0.20	4.76 ± 0.23
677.207.....	7.31 ± 0.30	6.93 ± 0.26	5.46 ± 0.31	4.88 ± 0.18
681.195.....	7.31 ± 0.31	6.61 ± 0.17	5.45 ± 0.20	4.96 ± 0.19
689.082.....	7.28 ± 0.23	5.83 ± 0.18	5.19 ± 0.27	4.63 ± 0.10
693.113.....	5.80 ± 0.18	5.63 ± 0.10	4.62 ± 0.26	4.16 ± 0.19
697.102.....	5.92 ± 0.35	4.91 ± 0.19	4.11 ± 0.24	3.86 ± 0.19
700.776.....	4.61 ± 0.23	4.33 ± 0.16	4.09 ± 0.22	3.71 ± 0.12
713.109.....	5.89 ± 0.47	5.67 ± 0.20	4.31 ± 0.18	4.74 ± 0.19

^a Units are 10^{-14} ergs cm^{-2} s^{-1} \AA^{-1} .

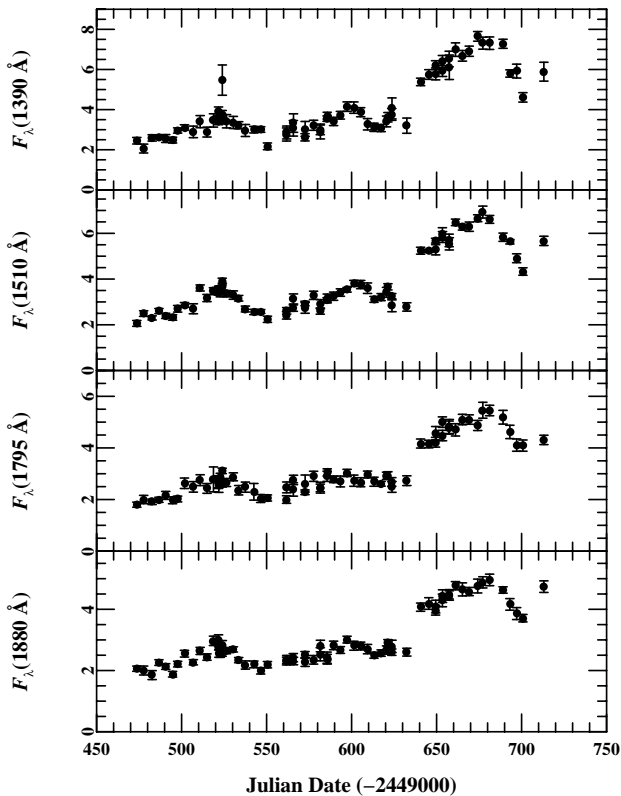


FIG. 2.—Light curves of the continuum fluxes at 1390 Å, 1510 Å, 1795 Å, and 1880 Å (observed wavelengths).

been removed from the emission-line fluxes, since the strongest absorption features (C II, Si II, C IV) seen in the average spectrum are difficult to identify even in the best-exposed single spectrum. The fluxes and their associated uncertainties are listed in Table 3, and the emission-line light curves are shown in Figure 3.

Initially, the error estimates for the line and continuum measurements are based on the errors assigned by the NEWSIPS extraction software to the individual pixels in the relevant integration range for each feature. However, it became clear from the resulting light curves that the uncertainties estimated in this fashion are too large. We will make the assumption that the *relative* magnitudes of the errors obtained in the initial estimate are correct, but that all of the uncertainties should be scaled down by a constant multiplicative factor.

The uncertainties in the line and continuum fluxes can be estimated by measurements closely spaced in time, where differences should be attributable primarily to random errors rather than intrinsic variations. We consider the ratio of two measurements of the same flux f_i and f_j that are separated by some short time interval $\Delta t_{ij} = t_i - t_j$. For Δt_{ij} sufficiently small, f_i and f_j are essentially independent measurements of the same quantity, and thus the uncertainties in their ratio ought to be $\sigma_{ij}^2 \approx \sigma_i^2 + \sigma_j^2$, where σ_n is the uncertainty associated with measurement f_n . Averaged over many pairs of observations, the variance of the distribution of f_i/f_j will be $\sigma_{ij}^2/\langle f^2 \rangle$, where $\langle f^2 \rangle$ is the mean square flux, provided that there is no intrinsic variability on timescales shorter than the largest value of Δt_{ij} . Thus, for closely spaced points (i.e., those with Δt_{ij} less than some small

TABLE 3

EMISSION-LINE FLUXES

Julian Date (-2,449,000) (1)	Ly α λ 1215 ^a (2)	C IV λ 1550 ^a (3)	He II λ 1640 ^a (4)
473.652.....	555 \pm 15	450 \pm 11	49.8 \pm 10.1
477.684.....	561 \pm 23	415 \pm 18	75.9 \pm 16.2
482.672.....	546 \pm 17	383 \pm 13	53.7 \pm 11.9
486.676.....	570 \pm 12	403 \pm 9	38.1 \pm 8.3
490.676.....	548 \pm 18	370 \pm 14	48.2 \pm 12.6
494.695.....	580 \pm 16	410 \pm 12	63.3 \pm 11.1
497.688.....	555 \pm 15	416 \pm 11	46.7 \pm 10.2
501.691.....	568 \pm 18	427 \pm 14	62.8 \pm 13.1
506.918.....	618 \pm 20	478 \pm 15	...
510.910.....	570 \pm 24	434 \pm 18	57.9 \pm 16.7
514.941.....	619 \pm 22	459 \pm 17	74.4 \pm 15.2
518.746.....	559 \pm 35	358 \pm 27	79.7 \pm 23.8
521.746.....	611 \pm 28	430 \pm 22	46.2 \pm 19.4
521.836.....	626 \pm 22	457 \pm 17	82.1 \pm 15.4
521.930.....	626 \pm 23	480 \pm 18	67.2 \pm 16.2
522.750.....	607 \pm 25	429 \pm 20	42.4 \pm 17.5
523.879.....	592 \pm 16	453 \pm 12	68.2 \pm 10.8
523.961.....	636 \pm 26	445 \pm 20	77.1 \pm 17.8
526.434.....	621 \pm 19	421 \pm 15	52.6 \pm 13.6
530.441.....	587 \pm 19	398 \pm 14	51.8 \pm 13.2
533.590.....	604 \pm 17	441 \pm 14	74.4 \pm 12.4
537.520.....	601 \pm 26	442 \pm 20	58.1 \pm 18.0
542.516.....	606 \pm 22	450 \pm 17	77.6 \pm 15.3
546.543.....	580 \pm 17	461 \pm 14	55.4 \pm 12.1
550.496.....	609 \pm 16	425 \pm 13	39.3 \pm 11.8
561.746.....	574 \pm 22	426 \pm 17	61.2 \pm 15.6
561.840.....	528 \pm 27	398 \pm 20	42.1 \pm 18.5
565.734.....	542 \pm 29	445 \pm 22	83.4 \pm 20.2
565.824.....	553 \pm 25	427 \pm 19	75.5 \pm 17.2
572.590.....	603 \pm 15	452 \pm 12	80.9 \pm 10.3
572.664.....	587 \pm 33	458 \pm 25	82.7 \pm 22.5
577.656.....	557 \pm 25	448 \pm 20	...
581.574.....	606 \pm 24	429 \pm 19	68.8 \pm 17.0
581.648.....	591 \pm 32	430 \pm 24	67.3 \pm 22.0
585.574.....	585 \pm 21	411 \pm 16	59.6 \pm 14.8
585.660.....	548 \pm 20	390 \pm 16	60.5 \pm 14.4
589.422.....	575 \pm 19	423 \pm 15	76.6 \pm 13.7
593.410.....	614 \pm 22	404 \pm 17	58.3 \pm 15.1
597.340.....	603 \pm 19	404 \pm 15	73.4 \pm 13.6
601.336.....	644 \pm 28	454 \pm 21	83.6 \pm 19.1
605.352.....	647 \pm 23	441 \pm 18	73.0 \pm 15.8
609.367.....	601 \pm 22	427 \pm 17	64.3 \pm 15.1
613.172.....	656 \pm 19	447 \pm 15	90.4 \pm 13.6
617.367.....	625 \pm 15	459 \pm 12	77.1 \pm 10.3
620.543.....	550 \pm 21	422 \pm 16	68.0 \pm 15.1
621.367.....	598 \pm 20	411 \pm 16	61.1 \pm 14.1
623.656.....	603 \pm 21	453 \pm 16	89.3 \pm 14.8
623.758.....	598 \pm 34	434 \pm 27	65.2 \pm 23.5
632.508.....	564 \pm 26	428 \pm 20	69.8 \pm 18.0
640.594.....	707 \pm 22	477 \pm 17	...
645.422.....	640 \pm 24	398 \pm 18	76.0 \pm 16.6
649.410.....	705 \pm 25	421 \pm 19	78.1 \pm 17.4
649.496.....	693 \pm 34	434 \pm 26	110.6 \pm 23.4
653.414.....	659 \pm 31	446 \pm 24	91.1 \pm 21.2
653.492.....	777 \pm 30	469 \pm 23	123.7 \pm 20.7
657.395.....	658 \pm 31	408 \pm 24	111.6 \pm 21.6
657.477.....	685 \pm 34	453 \pm 26	102.1 \pm 23.4
661.117.....	650 \pm 28	421 \pm 22	77.9 \pm 19.6
665.191.....	774 \pm 30	451 \pm 23	90.4 \pm 20.6
669.191.....	743 \pm 24	461 \pm 18	115.8 \pm 16.8
674.184.....	712 \pm 29	500 \pm 22	100.4 \pm 19.8
677.207.....	764 \pm 35	458 \pm 27	110.0 \pm 23.7
681.195.....	762 \pm 31	454 \pm 23	107.7 \pm 21.1
689.082.....	869 \pm 23	440 \pm 18	107.7 \pm 15.9
693.113.....	723 \pm 27	504 \pm 21	111.6 \pm 18.8
697.102.....	763 \pm 34	468 \pm 26	81.5 \pm 23.1
700.776.....	726 \pm 24	481 \pm 18	105.0 \pm 16.7
713.109.....	785 \pm 30	476 \pm 23	78.4 \pm 20.7

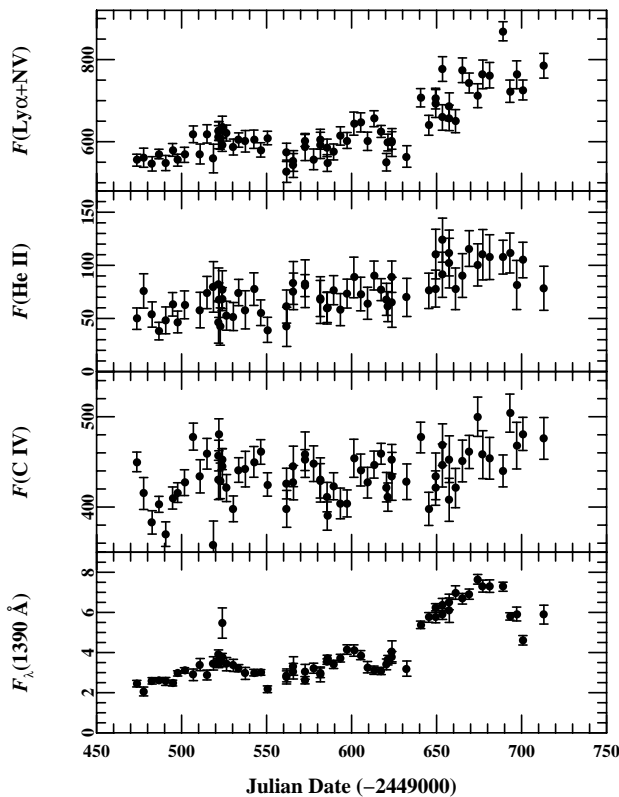
^a Units are 10^{-14} ergs cm^{-2} s^{-1} .

FIG. 3.—Light curves of the UV emission lines in which significant variability was detected. The bottom panel shows the 1390 Å continuum light curve for easy comparison.

value), we can compute the variance of the distribution in f_i/f_j from the actual flux measurements and compare this directly with the mean value of $\sigma_{ij}^2/\langle f^2 \rangle$ computed from the error estimates. We find that the root mean square (rms) of the f_i/f_j distribution is in all cases smaller than expected from the quoted error estimates and thus reduce all of the error estimates by a constant multiplicative factor (which ranges from 1.03 in the case of Ly α + N v to 3.4 for the 1510 Å continuum) to bring the two error estimates into agreement. This is a conservative approach: if there is intrinsic variability on short timescales, then the rms of the f_i/f_j distribution ought to be larger than the mean value of $\sigma_{ij}^2/\langle f^2 \rangle$, contrary to what is found. The true uncertainties may be even smaller than the rescaled error estimates, since any low-level variability that occurs on short timescales is attributed to random error in this estimate. Operationally, the factors we use for rescaling the NEWSIPS uncertainties are determined by considering pairs of observations separated by less than 1 day. We also find that if we increase the maximum time interval to 4 days, which greatly increases the number of pairs of observations in the calculation, then the rescaling factors increase (as expected), but only slightly, indicating that there is little intrinsic variability on such short timescales.

The uncertainties quoted in Tables 2 and 3 and illustrated in the figures are the rescaled values computed as described.

3.2. Variability

The variability parameters of the continuum and emission lines are given in Table 4. The entries in Table 4 are the unweighted mean flux $\langle f \rangle$ (col. [2]) and rms flux σ (col. [3]) for the entire campaign. These are defined in the usual way, i.e., for N individual observations f_i the mean is

$$\langle f \rangle = \frac{1}{N} \sum_{i=1}^N f_i, \quad (1)$$

and σ is the square root of the variance

$$\sigma^2 = \frac{1}{N-1} \sum_{i=1}^N (f_i - \langle f \rangle)^2. \quad (2)$$

Also in Table 4 are two common measures of the variability F_{var} (col. [4]) and R_{max} (col. [5]). The quantity R_{max} is simply the ratio of maximum to minimum flux. The quantity F_{var} is an estimate of the rms of the intrinsic variability relative to the mean flux, i.e., corrected for the effect of

random errors, i.e.,

$$F_{\text{var}} = \frac{(\sigma^2 - \Delta^2)^{1/2}}{\langle f \rangle}, \quad (3)$$

where the quantity Δ^2 is the mean square value of the uncertainties Δ_i associated with the fluxes f_i , i.e.,

$$\Delta^2 = \frac{1}{N} \sum_{i=1}^N \Delta_i^2. \quad (4)$$

As found in previous studies of other AGNs (Clavel et al. 1991; Reichert et al. 1994; Korista et al. 1995; Crenshaw et al. 1996), the amplitude of the continuum variations in Fairall 9 decreases at longer UV wavelengths. Whether this is due to an intrinsic hardening of the featureless continuum as the source becomes more luminous or to the presence of an intrinsically less variable feature, whose contribution increases with wavelength, cannot be addressed with the limited wavelength range covered by the SWP camera. However, based on the analysis of *IUE* archival data obtained between 1979 to 1991 and published optical data, Recondo-González et al. (1997) conclude that the Balmer continuum and the pseudocontinuum formed by blended emission lines of Fe II contribute significantly to the continuum flux in the UV range covered by the SWP camera. For a continuum flux similar to that observed during the 1994 campaign, their estimates of the combined contribution of Balmer continuum and Fe II lines to the bands centered at 1795 Å and 1880 Å are 0.56 and 0.51×10^{-14} ergs cm $^{-2}$ s $^{-1}$ Å $^{-1}$, respectively. Subtraction of these values from the mean flux densities in Table 4 results in all the continuum bands varying with the same amplitude ($F_{\text{var}} = 0.34 \pm 0.03$).

Another test that can be done to determine whether or not the AGN continuum hardens as it becomes brighter is to compare on a linear plot simultaneously measured fluxes in two well-separated wave bands. In the case of NGC 5548, for example, a plot of the optical flux (at 5100 Å) versus the short-wavelength UV flux (at 1350 Å) shows (1) that the best fit to these requires some curvature, and (2) that the optical axis intercept is nonzero (Fig. 3 of Peterson 1991). The sense of the curvature in this relationship means that the amplitude of the UV variations is greater than the amplitude of the optical variations, i.e., the variable part of the spectrum becomes harder as the continuum gets brighter. The nonzero intercept of the optical axis gives a measure of the constant starlight contribution to the optical fluxes, which in this case is in good agreement with the estimate obtained by image decomposition (Romanishin et al. 1995).

In Figure 4, we plot $F_{\lambda}(1880 \text{ Å})$ versus $F_{\lambda}(1390 \text{ Å})$ to test for curvature that would indicate a nonlinear relationship between these two continuum bands. A linear fit to these data (as shown in Fig. 4) provides a good fit ($\chi_v^2 = 3.08$); indeed, a quadratic gives a slightly worse fit ($\chi_v^2 = 3.11$) and the curvature coefficient is statistically insignificant. We conclude that, unlike the case of NGC 5548, and as previously concluded on the basis of existing data on this object, in the case of Fairall 9 there is no evidence that the AGN continuum shape changes as the luminosity changes. To test more directly the hypothesis that the relationship between the two continuum bands is linear, we fit the data to the equation

$$\log F_{\lambda}(1880 \text{ Å}) = C + \gamma \log F_{\lambda}(1390 \text{ Å}), \quad (5)$$

TABLE 4
VARIABILITY PARAMETERS

Feature (1)	Mean Flux ^a (2)	Standard Deviation ^a (3)	F_{var} (4)	R_{max} (5)
$F_{\lambda}(1390 \text{ Å})$	4.11	1.54	0.367	3.72 ± 0.42
$F_{\lambda}(1510 \text{ Å})$	3.82	1.35	0.350	3.36 ± 0.24
$F_{\lambda}(1795 \text{ Å})$	3.14	1.04	0.325	3.03 ± 0.24
$F_{\lambda}(1880 \text{ Å})$	3.02	0.92	0.300	2.68 ± 0.24
$F(\text{Ly}\alpha)$	628	73	0.110	1.65 ± 0.10
$F(\text{C IV } \lambda 1550)$	436	29	0.049	1.41 ± 0.12
$F(\text{He II } \lambda 1640)$	75	21	0.155	3.25 ± 0.89

^a Units are 10^{-14} ergs cm $^{-2}$ s $^{-1}$ Å $^{-1}$ for continuum fluxes and 10^{-14} ergs cm $^{-2}$ s $^{-1}$ for emission lines.

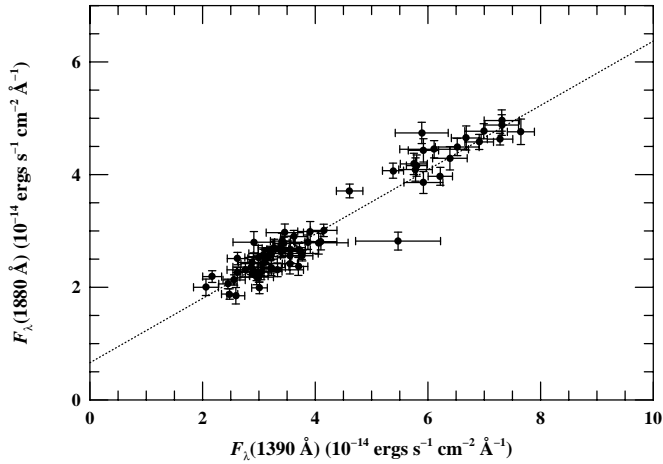


FIG. 4.—Simultaneous measurements of the continuum flux at 1390 Å and 1880 Å, from the data in Table 2. The dotted line shows the best linear fit to these data. There is no evidence for curvature in this relationship, which would indicate that the spectral index of the continuum depends on luminosity.

where C is a constant, which we assume to be the y -intercept of Figure 4 (i.e., $C = 6.6 \times 10^{-15} \text{ ergs s}^{-1} \text{ cm}^{-2} \text{ Å}^{-1}$). This yields $\gamma = 0.992 \pm 0.031$, consistent with a linear relationship. This will be tested more critically over a wider wavelength range using the combined UV and optical data on this object (Santos-Lleó et al. 1997).

Among the emission lines, the largest variations occur in Ly α + N v and He II λ 1640, and C IV λ 1550 varies only weakly. The light curves we measure for Si IV λ 1400 show no statistically significant variability, and therefore we do not include these measurements in the tables or figures. We must point out, however, that the Si IV λ 1400 feature appears, albeit weakly, in the rms spectrum shown in Figure 1. Weak variations did in fact occur, but the measured light curve is dominated by random noise.

The continuum light curves show three “events” of similar duration (~ 70 days), but with noticeably different amplitudes (Fig. 2). The first two events show variations of up to about 50%, while during the last event the continuum varied by more than a factor of 2. The Ly α + N v light curve is qualitatively similar to that of the continuum, with two low-amplitude events ($\sim 20\%$) followed by a stronger ($\sim 50\%$) third one. The most remarkable difference is that the fading of the continuum by the end of the campaign is not clearly seen in the Ly α + N v light curve. Although still smaller than that of the continuum, the amplitude of the He II λ 1640 variations is the largest among the emission lines. The detection of the first two events in the He II λ 1640 and C IV λ 1550 light curves is quite marginal, but the onset of the third event is clear even in these noisy light curves.

3.3. Cross-Correlation Analysis

The emission-line time delays, or lags, relative to the continuum variations can be quantified by cross-correlation of the continuum and emission-line light curves. Before computing the cross-correlation functions (CCFs), multiple measurements obtained during single observing shifts were averaged, weighted in each case by the inverse square of their assigned uncertainties. The CCFs have been computed by using three different algorithms: (1) the interpolation cross-correlation function (ICCF) of Gaskell & Sparke

(1986), (2) the discrete correlation function (DCF) method of Edelson & Krolik (1988), and (3) the z -transformed discrete correlation function (ZDCF) method of Alexander (1996). The ICCF and DCF algorithms are as described by White & Peterson (1994). All three methods give similar, although not identical, results. The CCFs between the continuum at 1390 Å and the other continuum bands (Fig. 5) show that they are all highly correlated, with no measurable delay among the different bands. In Table 5 we summarize the characteristics of the CCFs, where τ_{peak} is the location of the ICCF peak, which has value r_{max} ; τ_{cent} is the centroid for ICCF values greater than $0.8r_{\text{max}}$; and FWHM is the full width of the ICCF at $0.5r_{\text{max}}$. Also in Table 5 are the ZDCF results; τ_{ML} is the maximum likelihood estimate of the location of the true peak in the CCF, and the errors associated with it represent a 68% confidence interval (see Alexander 1996 and Edelson et al. 1996). The maximum likelihood uncertainties quoted in Table 5 are more conservative error estimates than we obtain using either the analytic formula of Gaskell & Peterson (1987), which yields estimated errors $\Delta\tau \approx 4\text{--}7$ days for all ICCFs computed in this paper, or error estimates based on Monte Carlo simulations (e.g., White & Peterson 1994), which yield $\Delta\tau \lesssim 2$ days for all ICCFs presented here.

The results of cross-correlating the continuum at 1390 Å and the three emission-line light curves of Table 3 and Figure 3 are shown in Figure 6 and summarized in Table 5. Only Ly α + N v and He II λ 1640 are highly correlated with the continuum ($r_{\text{max}} > 0.8$). For C IV λ 1550, $r_{\text{max}} = 0.63$ is statistically significant (for $N = 55$ epochs) at greater than

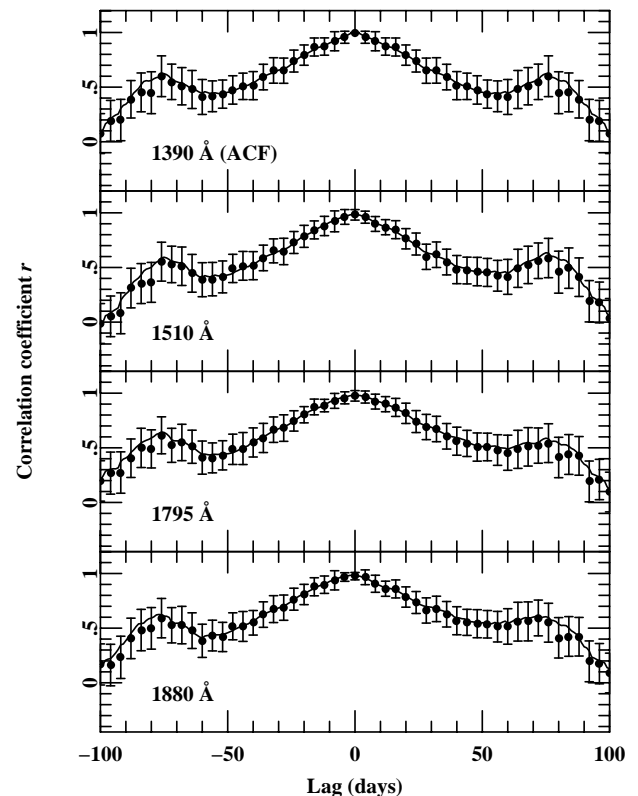


FIG. 5.—Continuum cross-correlation functions (CCFs). The panels show the CCFs (ICCF as solid lines; DCF as individual points with associated error bars) for the 1390 Å continuum and (a) itself (i.e., the ACF), (b) the 1510 Å continuum, (c) the 1795 Å continuum, and (d) the 1880 Å continuum.

TABLE 5
CROSS-CORRELATION RESULTS

FEATURE	ICCF/DCF				ZDCF		
	r_{\max}	FWHM (days)	τ_{peak} (days)	τ_{cent} (days)	r_{\max}	τ_{ML} (days)	τ_{cent} (days)
1390 Å (ACF).....	1.00	88	0	0.0
1510 Å.....	0.98	87	-1	-0.5	0.94	$3.7^{+1.9}_{-7.0}$	-0.4
1795 Å.....	0.98	104	1	0.0	0.93	$3.7^{+3.2}_{-7.6}$	1.6
1880 Å.....	0.98	131	-2	-1.0	0.94	$-3.9^{+7.2}_{-2.3}$	7.8
Ly α λ 1215.....	0.91	106	14	14.6	0.86	$3.6^{+12.0}_{-2.2}$	19.5
C IV λ 1550.....	0.63	75	34::	31.5::	0.51	$43.8^{+25.4}_{-24.5}$	29.7
He II λ 1640.....	0.85	145	10	3.7	0.63	$-3.9^{+17.4}_{-9.7}$	4.2

99.9% confidence. However, the CCF is very flat-topped and the width of the CCF is somewhat smaller than that of the continuum autocorrelation function (ACF). We therefore attach no significance to the measured C IV λ 1550 lag.

3.4. Line Profile Variations

The low amplitude of the line variations limits the depth of the analysis of the profile variability. However, inspection of Figures 2 and 3 shows that, during this observing campaign, Fairall 9 went through two different levels of brightness, and a simple comparison can be made of the profiles in the high and low states. For the high state, we have averaged all the spectra from JD 2,449,660 to the end of the campaign, and for the low state, the spectra taken between JD 2,449,530 and JD 2,449,580 have been averaged. The Ly α , C IV λ 1550, and He II λ 1640 line profiles for the two

states are shown in Figure 7, after continuum subtraction. Although the error bars are rather large, the difference profile of C IV λ 1550 appears to be narrower than the Ly α and He II λ 1640 difference profiles. In an attempt to quantify roughly the changes in the line profiles, each difference profile has been fitted with a Gaussian (shown as a thick line in the bottom panels of Fig. 7). We find that the width of the C IV λ 1550 difference profile ($4180 \pm 340 \text{ km s}^{-1}$) is less than half the width of the Ly α ($10600 \pm 460 \text{ km s}^{-1}$) and He II λ 1640 ($13890 \pm 900 \text{ km s}^{-1}$) difference profiles. Also, the center of the Gaussian fitted to C IV λ 1550 is more redshifted relative to the galaxy redshift ($c \Delta z = 1570 \pm 170 \text{ km s}^{-1}$) than are either Ly α ($880 \pm 130 \text{ km s}^{-1}$) or He II λ 1640 ($270 \pm 400 \text{ km s}^{-1}$). However, given the low level of emission-line variability detected in this experiment, we are reluctant to draw strong conclusions based on these limited data. We draw attention to these difference because they are of potential importance should they be substantiated by better data.

4. DISCUSSION

The selection of Fairall 9 for an intensive monitoring effort with *IUE* was intended to extend the very limited sample of intensively monitored AGNs to higher luminosity and thus test critically the predicted BLR radius-luminosity relationship $R \propto L^{1/2}$. A previous effort by Clavel et al. (1989) (see also Koratkar & Gaskell 1989; Lub & de Ruiter 1992; Shields, Ferland, & Peterson 1995; Recondo-González et al. 1997) yielded lags of about 115 ± 70 days for Ly α and about 200 ± 80 days for C IV λ 1550. These lags were measured during a factor of 30 decrease in luminosity and are based on data that are not well sampled on short timescales. Indeed, cross-correlation lags are strongly influenced by inflection points in light curves, and the previously published Fairall 9 light curves are not well sampled at the time the continuum and emission lines passed through their respective minima. Moreover, a plot of the observed C IV λ 1550 time lag versus UV continuum for reasonably well observed AGNs (e.g., Fig. 3 of Peterson 1994) suggests that the published emission-line time lags for Fairall 9 are somewhat larger than would be expected by extrapolation from the lower luminosity Seyfert galaxies, assuming *either* the predicted $R \propto L^{1/2}$ *or* the possibly shallower luminosity dependence weakly suggested by the other data. For example, if we scale the NGC 5548 results⁴⁴ (Clavel et al.

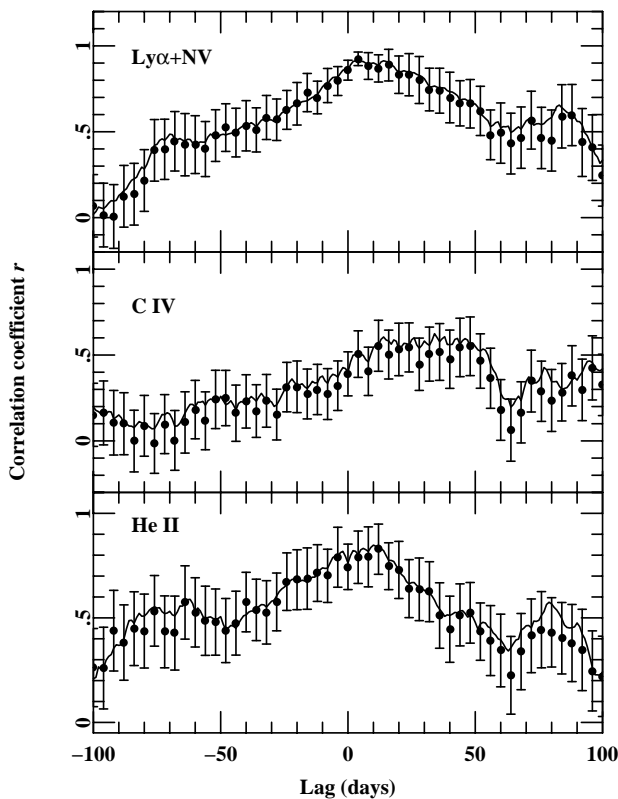


FIG. 6.—Continuum-emission-line CCFs. The panels show the CCFs (ICCF as solid lines; DCF as individual points with associated error bars) for the 1390 Å continuum and (top) Ly α + N v, (middle) C IV λ 1550, and (bottom) He II λ 1640.

⁴⁴ The values of τ_{cent} in NGC 5548 are about 11 days for both lines, using the interpolation CCF as described by White & Peterson (1994).

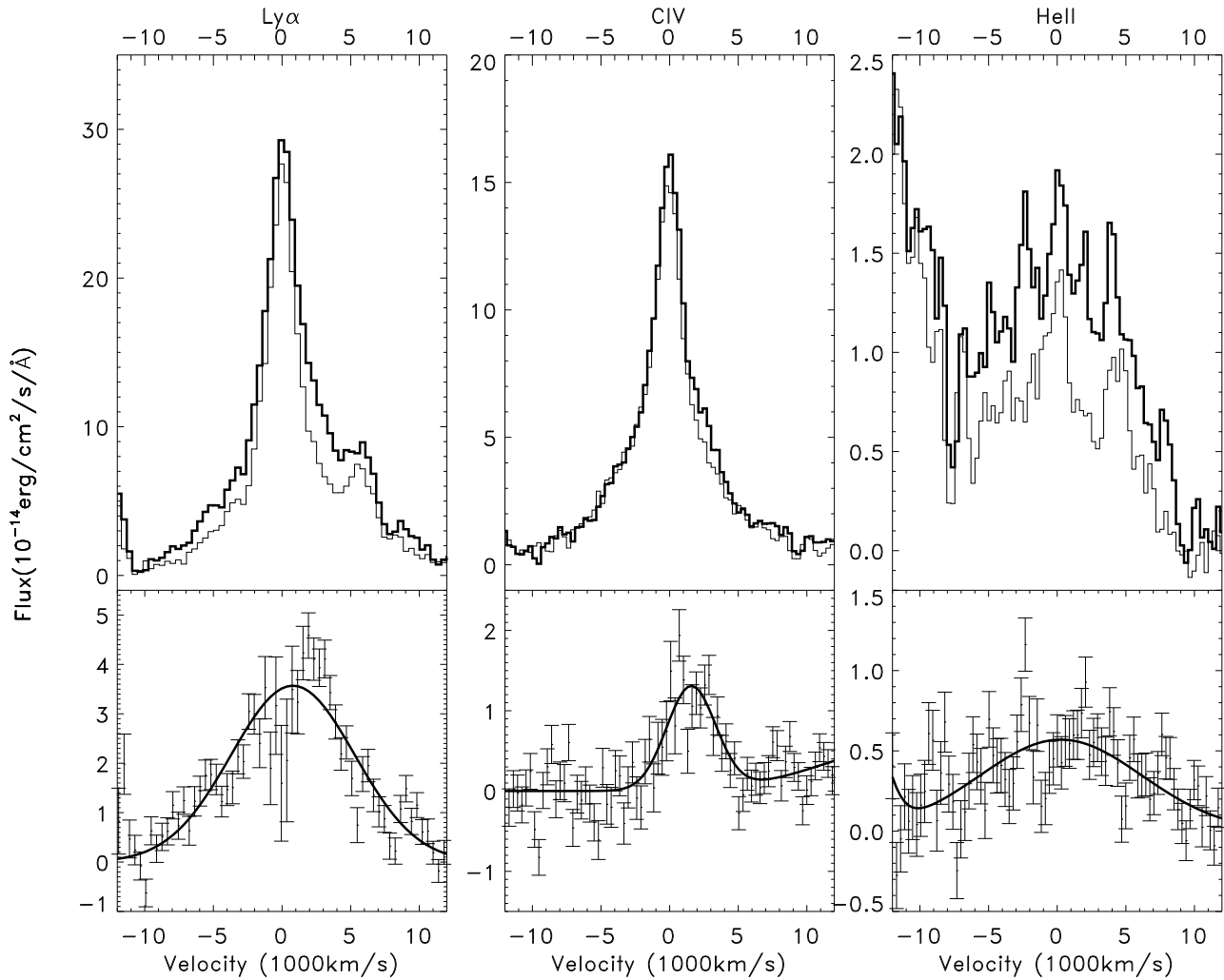


FIG. 7.—The average profiles at high (*thick lines*) and low (*thin lines*) states (*top*) and their differences (*bottom*) are shown for $\text{Ly}\alpha$ (*left*), $\text{C IV } \lambda 1550$ (*middle*), and $\text{He II } \lambda 1640$ (*right*). Note that the overlap in the profiles of $\text{C IV } \lambda 1550$ and $\text{He II } \lambda 1640$ is not present in their difference profiles. The thick lines in the bottom panels show the Gaussian fits to the difference profiles.

1991) assuming the $L^{1/2}$ luminosity dependence, then we would expect an $\text{Ly}\alpha$ $\text{C IV } \lambda 1550$ lag of between 13 and 73 days for Fairall 9 for its low and high historical states, respectively. These considerations led us specifically to look for more rapid BLR response than could have been detected in the original Clavel et al. (1989) experiment.

For $\text{Ly}\alpha + \text{N V}$ we find $\tau_{\text{cent}} \approx 14\text{--}20$ days for this experiment, which is not much smaller than the value of 28 days we would expect by scaling the NGC 5548 results to the mean luminosity of Fairall 9 during this campaign. Unfortunately, relatively weak variability during the first several months of the campaign precluded an accurate measurement of the $\text{C IV } \lambda 1550$ response time. Formally, we obtain $\tau_{\text{cent}} \approx 30$ days for $\text{C IV } \lambda 1550$, although we place no confidence in this measurement on account of the low signal-to-noise ratio of the $\text{C IV } \lambda 1550$ light curve.

The reason for the dramatic difference between the Clavel et al. (1989) $\text{Ly}\alpha + \text{N V}$ lag and that reported here is not obvious. Just as the previous data were not sensitive to short-timescale emission-line response, the present campaign was not sensitive to *long*-timescale ($R/c \gtrsim 100$ days) emission-line lags. Among likely possibilities are the following:

1. The BLR did not change between the two experiments. On account of their very different sampling characteristics, the two monitoring programs sampled completely different scales in a very extended BLR. The lack of an obvious turnover in the $\text{Ly}\alpha + \text{N V}$ light curve at the end of the campaign may be consistent with this interpretation.

2. As a consequence of the significant change in mean luminosity between the two campaigns (a factor of $\sim 3\text{--}6$, depending on how it is measured), there has been a change in the *optical depth* in BLR clouds in the sense that the more recent experiment samples clouds that are currently optically thick (and thus produced highly variable emission lines), but which were optically thin (and producing less variable emission lines) in the previous campaign. Such dramatic changes are possible on account of the short recombination time for high-density BLR clouds, $\tau_{\text{rec}} \approx 10^9/n_e(\text{cm}^{-3})$ hr. This could occur if much of the line response we see in this campaign originates in clouds that have column densities only somewhat deeper than the ionized column; if the continuum flux increased by a significant factor, then these clouds would be completely ionized and thus show no recombination-line variability with further increases in continuum flux. Similar clouds at larger

distances would remain partially neutral and responsive to continuum changes.

3. There has been a significant rearrangement of the BLR gas between the two experiments. The crossing time for an emission-line cloud in the BLR is $\tau_{\text{dyn}} \approx c\tau_{\text{cent}}/\Delta V$, where ΔV is the emission-line width. Taking $\Delta V = 10^4 \text{ km s}^{-1}$ and $\tau_{\text{cent}} \approx 150$ days yields $\tau_{\text{dyn}} \approx 10$ yr, which is approximately the interval of time between the observations reported here and those made as Fairall 9 passed through the minimum state recorded by Clavel et al. (1989); it is certainly plausible that a significant rearrangement of the BLR gas has occurred between the two experiments.

Detailed modeling beyond the scope of this paper will be required to support further or to eliminate these and other possibilities. Unfortunately, a longer campaign that might have discriminated among the various possibilities is simply not possible on account of the termination of *IUE* operations.

The ICCF and DCF indicate that He II $\lambda 1640$, with $\tau_{\text{cent}} \lesssim 4$ days, responds more rapidly than does Ly α + N v, which is consistent with what has been found in other sources and is consistent with a radially ionization-stratified BLR. However, the uncertainties in all of these measurements are sufficiently large that we cannot rule out the possibility that Ly α + N v and He II $\lambda 1640$ actually respond with the same lag.

The C IV $\lambda 1550$ results, however, remain puzzling. In other well-studied AGNs, the response times and line profiles for Ly α and C IV $\lambda 1550$ are similar. While these data do not provide a strong constraint on the C IV $\lambda 1550$ response time, the variable part of the C IV $\lambda 1550$ line profile appears to be significantly narrower and redshifted relative to the Ly α and He II $\lambda 1640$ lines. It appears to be rather doubtful that the peak response of C IV $\lambda 1550$ and Ly α occurs in the same gas, as might be more easily inferred for NGC 5548, for example.

Given (1) the possible similarity of the Ly α + N v and He II $\lambda 1640$ responses, (2) our expectation from other experiments that Ly α and C IV $\lambda 1550$ ought to respond on the same timescale, and (3) some, albeit unreliable, indications that the C IV $\lambda 1550$ response is longer than Ly α + N v in Fairall 9, we might well ask if the response we are seeing in Ly α is in fact not Ly α at all, but N v $\lambda 1240$, which, again based on other sources, we expect to respond like He II $\lambda 1640$. The contribution of the N v $\lambda 1240$ emission line to the Ly α + N v light curve may cause the relatively short lag found if the variations in this high-ionization line are sufficiently large. We have considered this explanation and reject it for the following reasons. First, there is no structure in the difference profile shown in Figure 7 that may suggest that N v $\lambda 1240$ significantly contributes to the variations. Second, N v $\lambda 1240$ appears only weakly in the rms spectrum (Fig. 1), again arguing that the amplitude of N v $\lambda 1240$

is insufficient to contribute in a significant way to the Ly α + N v variability. Third, as a final check, we remeasured the Ly α flux as before, but we extended the measurement only to 1228 Å (in the rest frame) to minimize contamination by N v $\lambda 1240$; this yielded Ly α light curves and CCFs that did not differ in any substantive way from the Ly α + N v results reported here.

5. SUMMARY

The results of the *IUE* monitoring campaign on a bright Seyfert 1 nucleus (Fairall 9) qualitatively support the general conclusions drawn from previous campaigns on the less luminous AGNs NGC 3783, NGC 4151, and NGC 5548. In particular, the various continuum bands respond simultaneously, within our ability to measure them, and the emission lines respond to continuum variations on timescales that are very short (i.e., days), although the relatively noisy emission-line light curves measured in this experiment result in large uncertainties in the emission-line time lags. The principal conclusions of this study are as follows:

1. Large-amplitude (factor of 2) UV continuum variations occur in Fairall 9 on a timescale (~ 70 days) similar to the fastest variations detected (at a similar amplitude) in NGC 5548, which is about 6 times less luminous.

2. No significant lags have been found among the four UV continuum bands measured, although the amplitude of the variations decreases with increasing wavelength. Unlike NGC 5548, we do not find evidence that the AGN continuum changes shape (i.e., spectral index) as the luminosity changes. The wavelength baseline over which this has been tested will be extended to the optical region in a forthcoming paper (Santos-Lleó et al. 1997).

3. Among the emission lines, He II $\lambda 1640$ shows the largest fractional variations. The variations in Ly α + N v and C IV $\lambda 1550$ are much smaller.

4. The lags (τ_{cent}) between the emission lines and the UV continuum are $\lesssim 4$ days for He II $\lambda 1640$ and about 14–20 days for Ly α + N v. The uncertainties of these determinations are large (~ 5 –10 days), so it cannot be concluded with any confidence that the lags for these two lines are not the same, although this is much clearer in other sources. The small variations in C IV $\lambda 1550$ make it difficult to estimate a lag, but it seems to be greater than for Ly α + N v and He II $\lambda 1640$.

5. The line profile variations in Ly α and He II $\lambda 1640$ differ from those in C IV $\lambda 1550$. The variable part of the C IV $\lambda 1550$ line appears to be narrower and more redshifted than the variable parts of Ly α and He II $\lambda 1640$.

This work has been supported in part by NASA through grant NAG 5-3233 to Ohio State University.

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