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**Astronomy
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Erratum

A spectro-astrometric study of southern pre-main sequence stars

Binaries, outflows, and disc structure down to AU scales

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Key words. accretion, accretion disks – line: formation – line: profiles – stars: activity – stars: evolution – stars: pre-main sequence – Errata, addenda

A computational error was made during the calculations of line profiles in Fig. 7. As a result, the figure was made under the assumption that the positional displacement always lies in the direction of the secondary. We compare our results in Table 2 with previous imaging (Reipurth & Zinnecker 1993) and low-resolution spectrophotometry (Brandner & Zinnecker 1997), and conclude that this assumption is not valid in all cases.

We since made extensive calculations for the case where the positional displacement lies in the direction of the primary. The results for both cases are shown in the revised Fig. 7 (see below). The feature at the secondary has a negative intensity in

some cases, implying that the assumed flux ratio is wrong, and such results are excluded from the figure. In the figure, the line-to-continuum ratio at the secondary is not always higher than the primary anymore. This does not affect the main conclusions of our paper.

References

Brandner, W., & Zinnecker, H. 1997, A&A, 321, 220

Reipurth, B., & Zinnecker, H. 1993, A&A, 278, 81

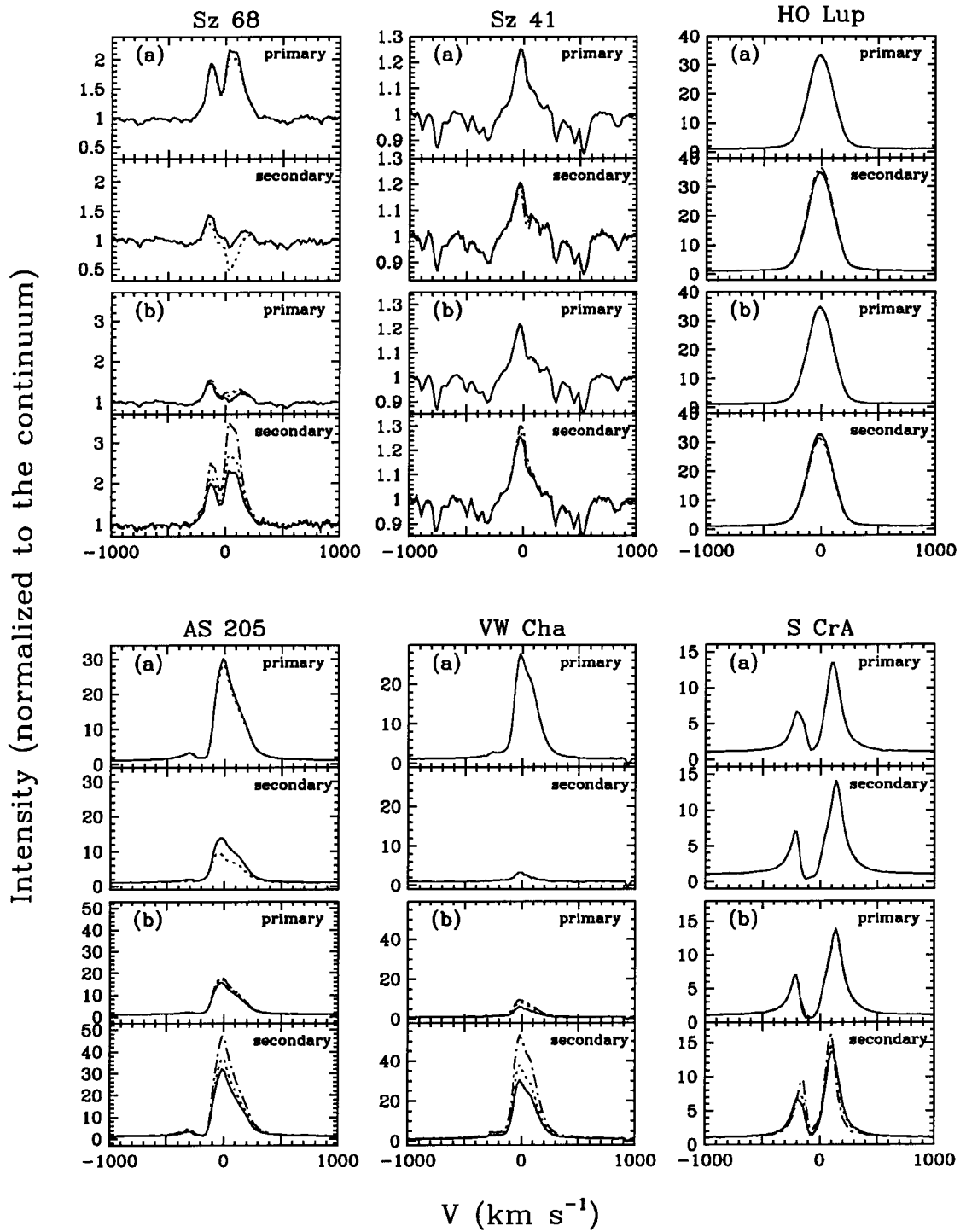


Fig. 7. Profiles of the H α emission for each binary component of known binaries in our sample. For a) and b), we assume that the displacement primarily lies in the direction of primary and secondary, respectively. Dot-dashed, dot, and solid profiles were obtained by assuming the flux ratios of $f = 0.2, 0.4,$ and $0.8,$ respectively, at the continuum level. The intensity scale for each profile is normalized to the continuum. The “primary” and “secondary” components for this figure are defined based on the continuum flux at 6500–6800 Å, different from Fig. 2 and Table 2.