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The CoRoT B-type binary HD50230: a prototypical hybrid pulsator with g-mode period and p-mode frequency spacings[★]

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ABSTRACT

Context. B-type stars are promising targets for asteroseismic modelling, since their frequency spectrum is relatively simple.

Aims. We deduce and summarise observational constraints for the hybrid pulsator, HD 50230, earlier reported to have deviations from a uniform period spacing of its gravity modes. The combination of spectra and a high-quality light curve measured by the CoRoT satellite allow a combined approach to fix the position of HD 50230 in the HR diagram.

Methods. To describe the observed pulsations, classical Fourier analysis was combined with short-time Fourier transformations and frequency spacing analysis techniques. Visual spectra were used to constrain the projected rotation rate of the star and the fundamental parameters of the target. In a first approximation, the combined information was used to interpret multiplets and spacings to infer the true surface rotation rate and a rough estimate of the inclination angle.

Results. We identify HD 50230 as a spectroscopic binary and characterise the two components. We detect the simultaneous presence of high-order g modes and low-order p and g-modes in the CoRoT light curve, but were unable to link them to line profile variations in the spectroscopic time series. We extract the relevant information from the frequency spectrum, which can be used for seismic modelling, and explore possible interpretations of the pressure mode spectrum.

Key words. Stars: oscillations; Stars: variables: early-type; Stars; fundamental parameters – Stars: individual: HD 50230 – Techniques: photometric

1. Introduction

Prior to the launch of the CoRoT space mission, HD 50230 (spectral type B3, V mag = 8.95) was a poorly known early type star. It is located in an overlapping region in the pulsational Hertzsprung-Russel diagram, where gravity modes and pressure modes are expected to occur. The target was therefore chosen to

be observed as part of the study of massive stars within CoRoT's asteroseismology programme (Baglin et al. 2006; Michel et al. 2006), although ground-based all-sky monitoring programmes reveal no variability (e.g., ASAS, Pojmanski 2002). The CoRoT light curve, as shown in Fig. 1 in Degroote et al. (2010), immediately shows that detection limitations lie at the basis of this non-detection. Indeed, these authors revealed a gravity mode spacing in the low-frequency region below 2 d^{-1} that was linked with the chemical gradient left by the receding core and suggests extra-mixing around the convective core.

In addition to the gravity modes, β Cep-like pressure modes (above 4 d^{-1}) are clearly visible in the Fourier spectrum (Suppl. Fig. 3 in Degroote et al. (2010)), but up to now they were not studied in detail. This simultaneous occurrence of different types of modes yields the potential of constraining the whole interior of the star, from the core (probed by the gravity modes), to the envelope (probed by the pressure modes).

The simultaneous presence of pressure and gravity modes was already observed in hot B-stars from ground-based data, but the detected gravity mode frequencies in these stars are not always predicted to be excited by the theoretical models. This concerns the stars ν Eri (Handler et al. 2004; De Ridder et al.

* The CoRoT space mission was developed and is operated by the French space agency CNES, with participation of ESA's RSSD and Science Programmes, Austria, Belgium, Brazil, Germany, and Spain. Based on observations made with the ESO telescopes at La Silla Observatory under the ESO Large Programme LP182.D-0356, and on observations made with the Mercator Telescope, operated on the island of La Palma by the Flemish Community, at the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias, and on observations obtained with the HERMES spectrograph, which is supported by the Fund for Scientific Research of Flanders (FWO), Belgium, the Research Council of K.U.Leuven, Belgium, the Fonds National Recherches Scientific (FNRS), Belgium, the Royal Observatory of Belgium, the Observatoire de Genève, Switzerland and the Thüringer Landessternwarte Tautenburg, Germany.

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2004), 12 Lac (Handler et al. 2006; Desmet et al. 2009) and γ Peg (Handler et al. 2009). In all these examples, however, the pressure modes are dominant over the gravity modes. Space-based data from Kepler (Balona et al. 2011) uncovered seven additional B-type hybrid stars.

In this paper, we aim at preparing an in-depth seismic study of HD 50230, motivated by the detection of period spacings. To this end, we perform an analysis of the pressure modes visible in the CoRoT spectrum. Subsequently, we determine the fundamental parameters of the star via a set of spectra obtained with three different instruments within a time span of three years. The observed spectroscopic variability is then finally linked back to the results obtained from the CoRoT light curve.

2. The CoRoT light curve

2.1. Frequency analysis

The CoRoT observations of HD 50230 were made during the first long run in the direction of the Galactic anticentre (LRa01), and lasted ~ 137 d. The observations started at CoRoT Julian Day 2846.95139 (or Julian Day 2454391.95139). The raw light curve contains nearly 400 000 observations of which ~ 40 000 flagged data points are removed because of various instrumental and environmental causes (Auvergne et al. 2009). The power in the frequency spectrum is concentrated in the low-frequency region between 0 and 15 d^{-1} . At higher frequencies ($15\text{--}100 \text{ d}^{-1}$), the only power present is due to harmonics of the satellite’s orbital frequency ($f_{\text{sat}} \approx 13.97 \text{ d}^{-1}$), which shows slight variations both in frequency and amplitude during the course of observations. Beyond 120 d^{-1} , no significant power is present.

The long-term linear trend in the light curve is linked to the decline of the CCD’s gain (Auvergne et al. 2009). We treated it as a multiplicative instrumental effect and divided the light curve with a fitted linear polynomial. This way, we neglect a small deviation from the linear trend between day 70 and day 100, which is also visible in the light curves of the simultaneously observed A5 star HD 49862 and A2 star HD 49294, and is therefore probably of instrumental origin. This introduces low amplitude low frequencies (e.g. f_{013} in Table A.2). For a description of the noise properties and the window function, we refer to Degroote et al. (2009a), where a detailed analysis for LRc01 is presented.

We used an iterative prewhitening procedure, each time with a nonlinear update of the last 20 frequencies (see Degroote et al. 2010) to extract the frequency content from the CoRoT light curve. The results are shown in Fig. 1 and listed in Table A.2. This table contains a list of all frequency identification numbers, the amplitudes, frequency values and phases (including errors), as well as the signal-to-noise ratio (S/N) and the BIC value. All amplitudes range between 5 and 1800 ppm, or, equivalently, between $6 \mu\text{mag}$ and 1.8 mmag . The general trend is that at higher frequencies, the amplitudes become lower, with a few exceptions (e.g. f_{034} and f_{050}). If we limit the analysis to the first 20 frequencies, no signs of nonlinear mode behaviour are detected in the frequency spectrum of HD 50230, neither in the form of harmonics nor higher-order combination frequencies (e.g. Fig. 2). If we include the frequencies beyond the first 20, we do identify harmonics and even combination frequencies. However, since the chance of finding combinations increases with the number of detected frequencies, these may be coincidental, and, therefore, only harmonics were taken into consideration. The frequencies for which harmonics have been detected are identified in Table A.2 in the last column, together with the relative strength present in the higher harmonics, compared to the main

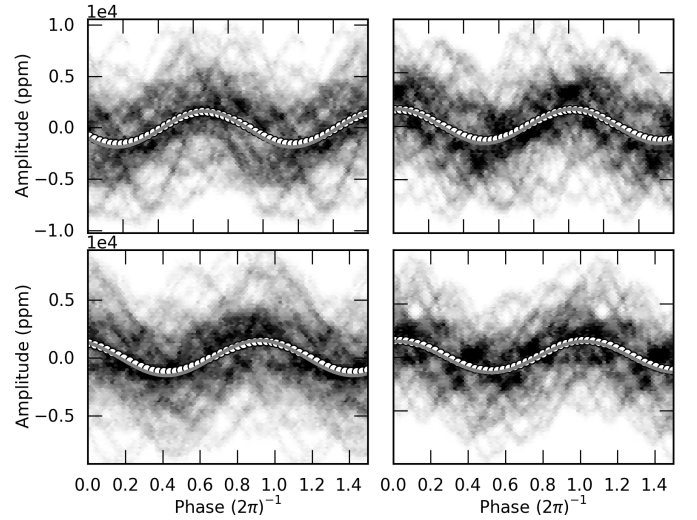


Fig. 2. Phase-folded light curves of the highest amplitude frequencies show that only one sine component is needed to fit the shape (*from left to right, top to bottom*: f_{001} , f_{002} , f_{004} , and f_{005}). The black dots are observations, the white circles are phase-binned averages, the thin grey line is a sine-fit (the exponent of the y-axis scale is given at the top of the plot).

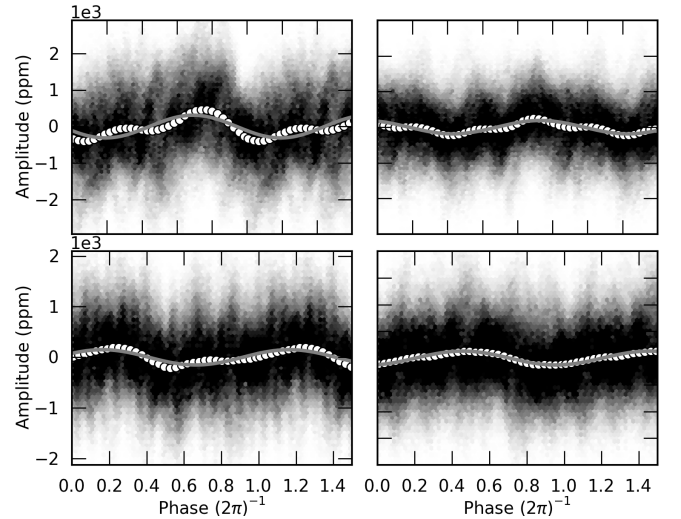


Fig. 3. Phase-folded light curves, using the frequencies (*from left to right, top to bottom*: f_{020} , f_{033} , f_{038} , and f_{047}). A clear deviation from the single-harmonic case in the form of a distorted phase shape is visible. For an explanation of the symbols, see Fig. 2.

frequency. As an example, for f_{012} , we list $A_r(4) = 0.38$, which means that four harmonics of f_{012} are detected, and that the sum of the amplitudes of all these harmonics amounts to 38% of the amplitude of f_{012} itself. Frequencies with a high A_r have a highly distorted phase shape (e.g. Fig 3).

To test the validity of the sine model, the variance reduction was computed for an increasing number of single-sinusoids. Because our method of iteratively prewhitening sinusoids necessarily implies an increasing variance reduction, we computed the χ^2 (with constant error terms) and Bayesian information criterion (BIC) as well (see also Degroote et al. 2009b). The latter criterion is only meaningful in a relative way, and weighs the introduction of extra parameters against the gain in variance reduction. The model with the lowest BIC value has the ideal num-

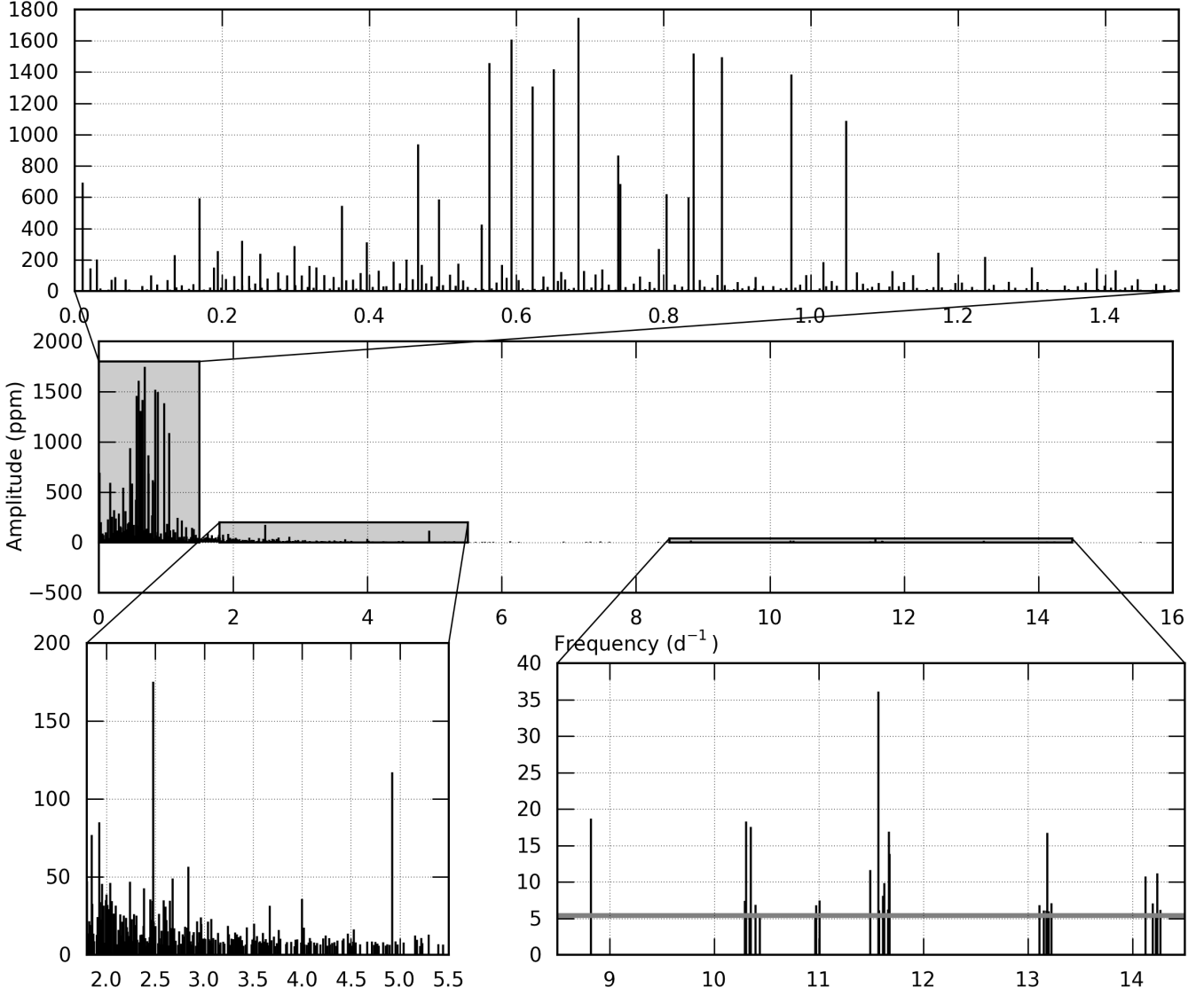


Fig. 1. Frequency diagram (units of d^{-1} on the x -axis, and ppm on the y -axis) containing all frequencies from the iterative prewhitening analysis. The centre panel shows the whole range of detected frequencies. The top panel is a zoom-in on the gravity mode region. The bottom right panel is a zoom-in on the pressure mode region (grey line is the detection threshold). The bottom left panel shows the intermediate region with two modes clearly standing out amongst the other frequencies..

ber of parameters of all tested models. In general, the BIC and χ^2 show a decreasing trend when more single-sinusoids are considered, which means that sinusoids are considered to be appropriate models to explain the observed variability (Fig. 4). When the BIC for each model is explicitly compared to the previous model, we conclude that in some cases adding three parameters to the model does *not* outweigh the resulting variance reduction. Because the BIC puts more weight on the number of parameters used, it puts the threshold for accepting sinusoids from the complete frequency analysis at ~ 200 frequencies (see column 9 in Table A.2), while the χ^2 allows the use of 300 frequencies. Below these thresholds, the differences between the consecutive BIC and χ^2 are no longer monotonically decreasing.

For the last nine frequencies of the 566 found from the linear frequency analysis, the BIC starts to rise, implying that the additional number of parameters are not worth the gain in variance reduction. This coincides with the use of a p -value of $p = 0.1\%$ as a stop criterion under the assumption of an exponential distribution of the Fourier periodogram, and a Bonferroni correction

for multiple trial frequencies. The stepwise nonlinearly updated version of the frequency analysis finds nine frequencies less than the linear analysis (using the p value as stop criterion), while it still explains slightly more of the variance.

The high-amplitude gravity mode spectrum is discussed in detail by Degroote et al. (2010). For completeness, we list the frequencies connected to the period spacing in Table 1. They are also marked in italics in Table A.2. In the following, we focus on the (low-amplitude) features in the frequency spectrum above $1.5 d^{-1}$.

2.2. Interpretation of the p -mode frequency spectrum

The majority of the frequencies above $1.5 d^{-1}$ have amplitudes below 50 ppm, making it difficult to separate them. The origin of these closely packed frequencies (between ~ 1.5 and $\sim 5 d^{-1}$, see, e.g., the lower left panel in Fig. 1) is not obvious. They are not instrumental, since for quiet stars, the noise level of the CoRoT photometry is much lower. We do not expect them to be an ar-

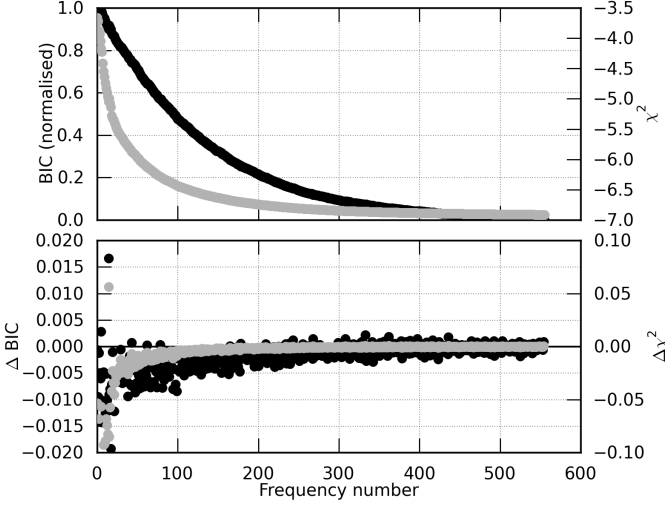


Fig. 4. Goodness-of-fit statistics of the sine-model. Both the BIC (*black*) and χ^2 (*grey*) show a general decreasing trend (*top panel*), suggesting that sines are a good way of modelling the light curve. Despite the general decrease, however, not all frequencies are adequately modelled by sines: some frequencies in particular induce a worse BIC ($\Delta\text{BIC} > 0$) (*lower panel*).

Table 1. Periods and associated period spacings detected in the gravity mode spectrum.

Frequency ID	Period (s)	ΔP (s)
f_{052}	69854 ± 7	9241.379 ± 21
f_{001}	79096 ± 20	9640.178 ± 20
f_{016}	88736 ± 5	9522.432 ± 9
f_{011}	98258 ± 7	9186.452 ± 10
f_{005}	107445 ± 8	9562.853 ± 12
f_{006}	117008 ± 10	9233.798 ± 15
f_{028}	126241 ± 11	9430.037 ± 33
f_{101}	126241 ± 32	-

tifact from the prewhitening procedure, partly because we used nonlinear optimisation, but mostly because the frequencies are located far away ($\gg 1/T$ with T the total time span of the observations) from the dominant frequencies. They could originate from many closely spaced time-variable modes, or background stellar noise. A short-time Fourier transformation (STFT) shows hints for the time-variability of the modes (Fig. 5).

In addition, some time-independent frequencies clearly stand out, and are therefore good candidate pulsation modes (Fig. 1). This concerns 17 frequencies, marked in boldface in Table A.2. Some of these candidate pressure modes are isolated (Fig. 6), while others have a clear multiplet structure (Fig. 7, lower panel). In the STFT (Fig. 8), the latter show the typical interference patterns connected to beatings: when the spacing between two frequencies is x/T , we expect to see x times of constructive or destructive interference.

The autocorrelation of the periodogram between 10 and 15 d^{-1} (Fig. 7) reveals a well-resolved splitting $\Delta f_{\text{obs}} = 0.044 \pm 0.007 \text{ d}^{-1}$, possibly related to rotational splitting. For slow rotators, the rotational splitting of pressure modes can be approximated with (Ledoux 1951)

$$\Delta f_{\text{rot}} = m\beta_{n\ell}\Omega,$$

where $\beta_{n\ell}$ is dependent on the stellar structure, number of nodes n and degree ℓ , but typically close to unity. Ω is the rotation

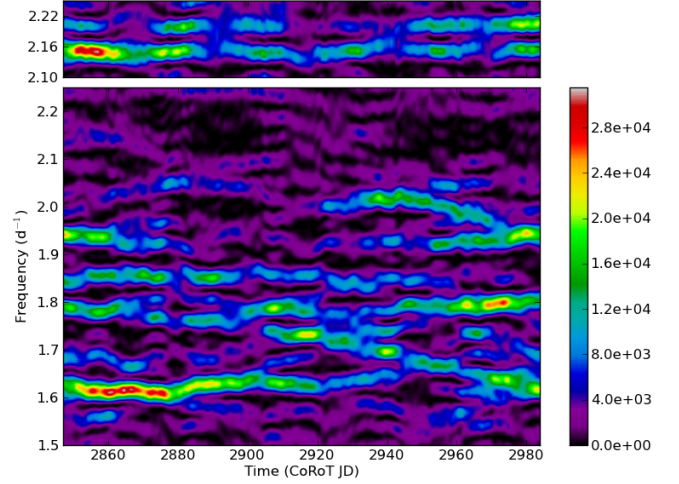


Fig. 5. Short-time Fourier transformation of the frequency region between 1.5 and 2.3 d^{-1} with a rectangular window of width 30 d^{-1} . All signal outside of these regions is removed. Clear signal is present (colours denote squared amplitude in ppm^2), and the frequencies and amplitudes show hints of variability. As a reference, the top panel shows a copy of the upper part of plot, where two artificial constant sine curves are introduced.

frequency of the star, and m the azimuthal order of the mode. The value of Δf_{obs} is too low to be a large frequency separation between pressure modes of equal degree and consecutive radial order in the asymptotic approximation. Also, it is unlikely that this value is the low-frequency separation between modes of different degree, since the exact same pattern emerges around 10.4 d^{-1} and 11.6 d^{-1} , which would imply that the large spacing is exactly the same for all degrees and that all modes of a given radial order but different degree cluster around the same value. The first-order interpretation of Δf_{obs} as the rotation frequency Ω is compatible (but delivers no additional constraints) with stellar models between 5 and $10 M_{\odot}$ and stellar radii between 2 and $9 R_{\odot}$ and a projected rotational equatorial velocity $v_{\text{eq}} \approx 6.9 \text{ km s}^{-1}$, with a restraint of the inclination angle i to be $i > 20^{\circ}$.

A hint for a larger spacing around 0.8 d^{-1} is also discernible in the two panels of Fig. 7. The peak in the autocorrelation at 0.8 d^{-1} in the top panel of Fig. 7 is linked to the multiplets around 13.2 d^{-1} and 14 d^{-1} and the multiplets around 10.3 d^{-1} and 11 d^{-1} . The peak at 1.25 d^{-1} is tied to the multiplets around 10.3 d^{-1} and 11.56 d^{-1} . Detailed seismic modelling is needed to interpret this as a large frequency separation between equal-degree modes, though the radial order of these modes is probably too low to be in the asymptotic regime.

3. Fundamental parameters

3.1. Multicolour photometry: an early B-type main sequence star

An estimate of the fundamental parameters of HD 50230 is obtained by fitting model atmospheres to observed multicolour photometry. We chose this approach over the application of existing calibration schemes (e.g., Cramer 1984), because fitting model atmospheres directly allows for an easy insertion of up-to-date atmosphere models, and allows for the combined use of different photometric systems. For early-type stars, the shape of the spectral energy distribution (SED) in the blue part, before

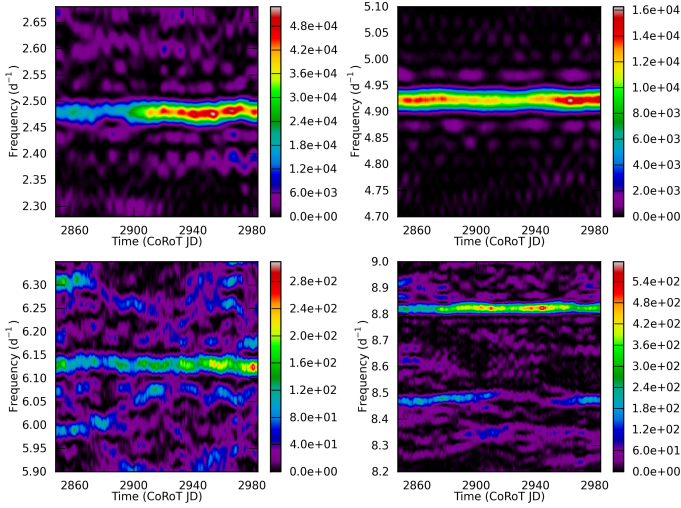


Fig. 6. Short-time Fourier transformations of four isolated frequencies in the high-frequency region (rectangular window of width 30 d^{-1} , colours denote squared amplitude (ppm^2)). In contrast to Fig. 5, the frequencies appear throughout the whole light curve.

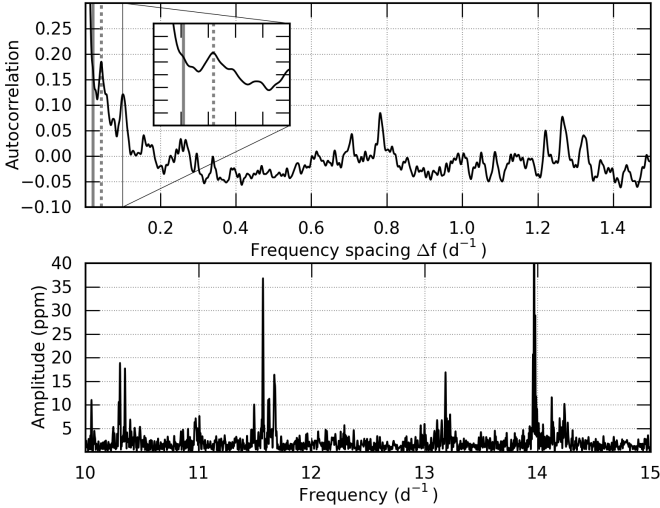


Fig. 7. Autocorrelation (upper panel) of the frequency spectrum (lower panel) between 10 and 15 d^{-1} (with the frequencies outside this region removed) uncovers a multiplet spacing of $\Delta_f = 0.044 \text{ d}^{-1}$ that was measured between the peaks near $\sim 10.4 \text{ d}^{-1}$ and $\sim 11.6 \text{ d}^{-1}$. The peak around $\sim 0.8 \text{ d}^{-1}$ originates from the spacing between the structures around 13.2 d^{-1} and 14 d^{-1} (bottom panel), and those at 10.3 d^{-1} and 11 d^{-1} . The thick grey line is plotted at a separation of $3/T$ for comparison with the frequency resolution.

(and including) the Balmer jump, is most sensitive to changes in the fundamental parameters. In general, SEDs are not very sensitive to binarity, because the luminosity scales with the square of the radius of the star and fourth power of effective temperature. For main sequence binaries, the brightest star is expected to be the hottest and the largest. To cover a wide spectral range and to emphasise the UV region, we collected TD1 (Thompson et al. 1978), ANS (Wesselius et al. 1982), Geneva (Mermilliod et al. 1997), and 2MASS (Cutri et al. 2003) measurements. Although there are also Johnson (Johnson 1965) measurements available, we decided not to include them here. The Johnson filter sys-

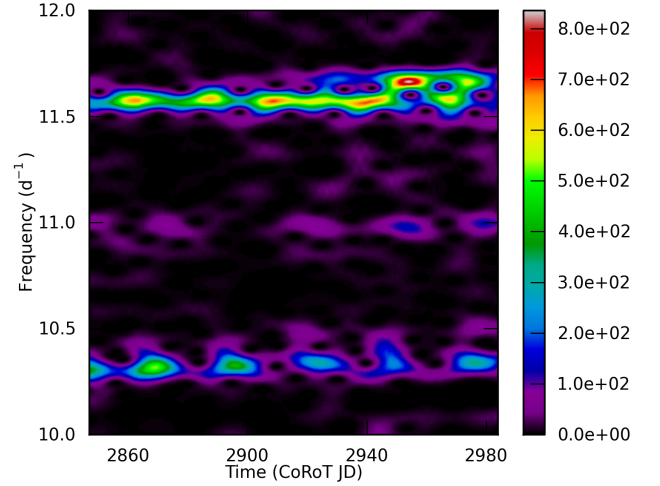


Fig. 8. The short-time Fourier transformation of the frequency region between 10 and 12 d^{-1} (with a Hamming window to improve the dynamical range, window width of 20 d^{-1}). The typical succession of constructive and destructive interference are typical signs of a barely resolved multiplet structure. The separation between the frequencies is of the order of $6/T$.

tem is often used to denote a wide range of slightly different response curves, and including them does not increase the wavelength coverage.

To convert the Vega-system magnitudes to fluxes, we used the calibration of Rufener & Nicolet (1988), Maíz Apellániz (2006), and Wesselius et al. (1980) in combination with the Vega model of Bohlin & Gilliland (2004). We chose to fit the TLUSTY model atmospheres of Lanz & Hubeny (2007), because they include non-LTE effects, which can influence the shape of the SED of early-type stars.

The comparison between model atmospheres and observations for the determination of fundamental parameters was made via a grid-based approach as described in Degroote et al. (2011). The goodness-of-fit and confidence intervals (CI) of the parameters were determined using a χ^2 statistic with five degrees of freedom (the effective temperature T_{eff} , surface gravity $\log g$, metallicity Z , interstellar reddening $E(B-V)$, and angular diameter θ). Because of the grid-based approach, correlation effects on the parameter uncertainties are naturally taken into account.

After fixing the fundamental parameters of HD 50230, we searched for near-infrared photometry obtained by the WISE satellite (Wright et al. 2010), which we calibrated following Cutri et al. (2011). We observed no significant infrared excess. The results are listed in Table 2 and the final SED is shown in Fig 9.

3.2. Spectroscopy: a suspected binary

Between 11 October 2008 and 10 January 2012, we have gathered 60 high-resolution spectra from three different instruments: 17 spectra from the CORALIE spectrograph ($R \approx 50\,000$, $3870\text{--}6900\text{\AA}$, Baranne et al. 1996) mounted on the Euler telescope at La Silla (Chile), 40 spectra from the HERMES spectrograph ($R \approx 85\,000$, $3770\text{--}9000\text{\AA}$, Raskin et al. 2011) on the Mercator telescope on La Palma (Spain), and four spectra from the HARPS spectrograph in the high-efficiency mode (EGGS, $R \approx 80\,000$, $3800\text{--}6800\text{\AA}$, Mayor et al. 2003) on the ESO 3.6m telescope at La Silla (Chile). A full log of the observations can

Table 2. Fundamental parameters of HD 50230 and their 95% CI.

Parameter	Photometric			Spectroscopic	Adopted ^c
	CI _{low}	CI _{50%}	CI _{high}		
$T_{1,\text{eff}}$ (K) ^a	15800	18000	21500	18500 ± 1000	18000 ± 1500
$T_{2,\text{eff}}$ (K)	-	-	-	≤ 16000	≤ 16000
$\log[g_1]$ (cm/s ²) (dex)	3.00 ^d	-	4.50	3.8 ± 0.3	3.8 ± 0.3
$\log[g_2]$ (cm/s ²) (dex)	-	-	-	-	~ 4
$\log Z/Z_\odot$ ^b (dex)	-1.0	0.3	0.3	-	-
$E(B - V)$ (mag)	0.12	0.19	0.27	-	0.19 ± 0.08
θ (mas)	0.037	0.040	0.045	-	0.040 ± 0.005

Notes. ^(a) Subscript ‘1’ and ‘2’ denote primary and secondary, respectively. The secondary component is not detectable from multicolour photometry. ^(b) $Z_\odot = 0.02$ dex. ^(c) The adopted values are compromises between photometric and spectroscopic determinations where both are available. ^(d) The $\log g$ is photometrically not constrained.

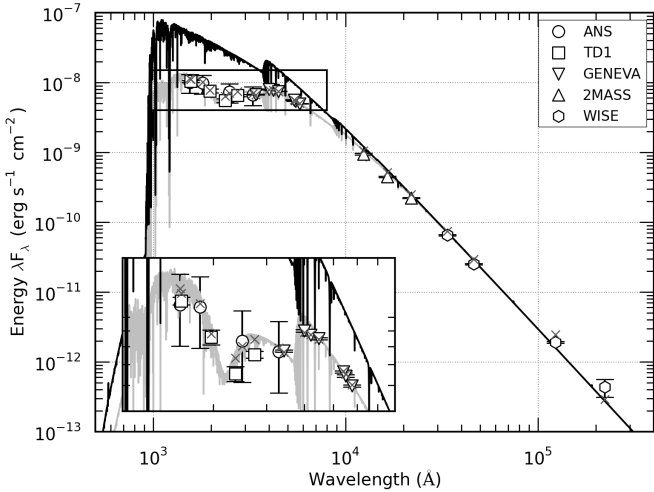


Fig. 9. Spectral energy distribution of HD 50230. Observations and their associated errors are shown with different open markers (see legend). Synthetic photometry from the best-fitting model (grey solid line, the unreddened model is shown as the black solid line) are marked with crosses. The inset is a zoom on the ultraviolet region.

be found in Table A.1. The first column shows the UT of mid-exposure, the second column contains the name of the instrument, and the last two columns give the exposure time and S/N of the spectra around 4500 Å.

A first look at the spectra shows many narrow and shallow lines, suggesting a low projected rotational velocity. The high resolution of the HARPS spectra allows the Fourier decomposition method (see, e.g., Simón-Díaz & Herrero 2007) to measure $v_{\text{eq}} \sin i$ values as low as $\sim 2 \text{ km s}^{-1}$ irrespective of any additional pulsational broadening effects. To obtain the best possible value, we averaged the two highest signal-to-noise ratio spectra and applied the Fourier decomposition on a set of 78 lines (from ions AlII, AlIII, ArII, CII, FeII, FeIII, NII, NeI, OI, OII, SII, SIII, SiII, and SiIII), with depth and isolation as selection criteria. We derive $v_{\text{eq}} \sin i = 6.9 \pm 1.5 \text{ km s}^{-1}$, assuming a linear limb-darkening law with $\epsilon = 0.35$. This low value is consistent with the observation that the two peaks of the narrow MgII 4481 doublet are visibly separated. HD 50230 is therefore a slow rotator.

To improve the estimation of the fundamental parameters based on the spectral features, we constructed an average spectrum. The individual spectra were normalised via spline fits to the continuum, and were corrected for radial velocity shifts.

The radial velocity of each spectrum was computed via the least-squares deconvolution algorithm (LSD) of Donati et al. (1997), with varying assumptions on the template spectra and selected lines. The measurements show a clear continuously rising trend during the three years of monitoring (Fig. 10). To check whether the variability is caused by pulsations or binarity, we used the generalised least-square method of Zechmeister & Kürster (2009) to fit sines and Keplerian orbits to the data. This method has the advantage over other period-finding algorithms such as the Scargle periodogram (Scargle 1982) that it also fits a constant term to the data, avoiding the problem of determining the average radial velocity in case of poor sampling. At the same time, the method still has the benefits of fitting functions directly to the data, in contrast to phase-binning methods such as the phase dispersion minimisation technique (Stellingwerf 1978). No significant periods are found. We conclude that the rising trend is most likely caused by a companion in a wide binary system. Instrumental zeropoint offsets, which can be of the order of 1.5 km s^{-1} (Uytterhoeven et al. 2008), are ruled out because the peak-to-peak difference is much larger and the trend is also visible in the spectra from HARPS or HERMES only. This is confirmed in the time-average, radial-velocity-corrected spectrum, where broad absorption features are discernible. These are most apparent around the He and Si lines (Fig. 11), but also in the red wing of the $H\alpha$ line, where a small asymmetry is observed (Fig. A.4, bottom right panel). It is highly unlikely that the broad features are due to pulsationally induced line profile variations, since they are not symmetric in the averaged spectra, and the deviation is only apparent close to the continuum.

3.3. Spectroscopic fundamental parameters

The fundamental parameters of HD 50230 as determined from photometry are listed in Table 2. The surface gravity is not constrained except for a lower limit of $\log g \approx 3$ dex. Alternatively, the fundamental parameters can be determined with the help of spectra in two ways: by fitting the line profiles directly with the help of synthesised spectra (e.g., Lehmann et al. 2011), or by comparing observed and predicted equivalent widths (EWs) (e.g., Morel et al. 2006). The former method has the advantage of a more concise way of handling blends, while the latter benefits from a reduction in the number of free parameters (i.e., $v_{\text{eq}} \sin i$ and the radial velocity). We combined the two approaches by using the EW method (for a description of the grid, see Morel et al. 2006) to estimate the effective temperature of the primary component based on carefully selected lines, and turned to the spectral fitting method to incorporate the secondary component.

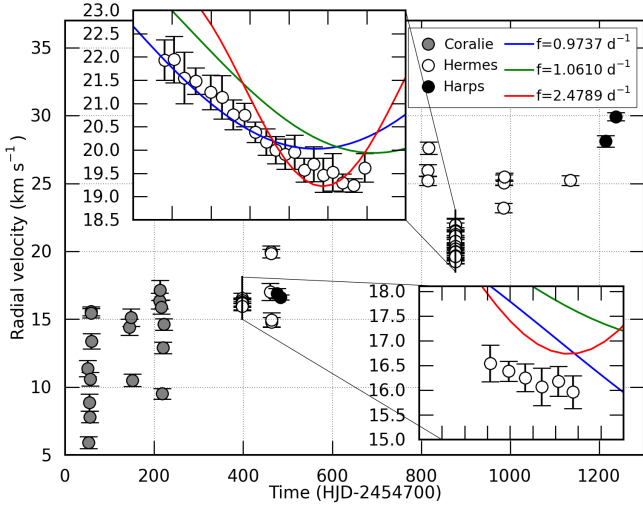


Fig. 10. Radial velocity derived from each individual spectrum via the LSD method. The inset in the upper left corner shows a zoom on an eight hour continuous monitoring of HD 50230, the inset in the lower right corner shows a zoom on a two-hour continuous monitoring. In colours, three fits with different frequencies are shown, vertically displaced to remove the day-to-day radial velocity shift.

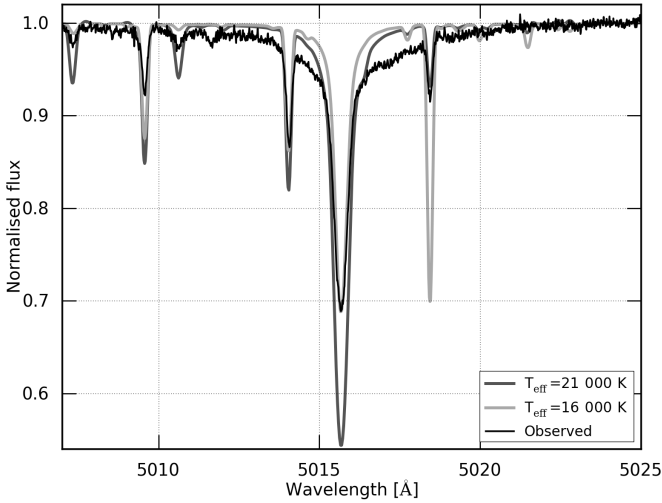


Fig. 11. HeI profile (black) of the averaged spectra. In grey, two single-star model HeI profiles are shown with different effective temperatures. No single adopted temperature can explain the wings of the profile, suggesting the presence of a broad-lined companion.

The EW-ratio of ions belonging to different ionisation stages of the same element are highly sensitive to changes in T_{eff} , but are fairly insensitive to binarity, as long as the continuum flux ratio of both components is similar in the wavelength range under study, or trivially when the secondary component is not visible in the spectrum. In the temperature range between 16 000 and 22 000 K, useful ions are SiII and SiIII, and SII and SIII. For the Si lines, this includes the ratios of SiII6371 to SiIII4567, SiIII4574, SiIII5739 and SiIII4552. The SiII5041, SiII5056, SiII4128, and SiII4130 lines were discarded because of the uncertainties in the atomic modelling of these high-energy transitions (Simón-Díaz 2010). For the sulfur atom, we selected SII5032, SII4815, SII4230, SII4162, SII4716, SII4991, SII4294,

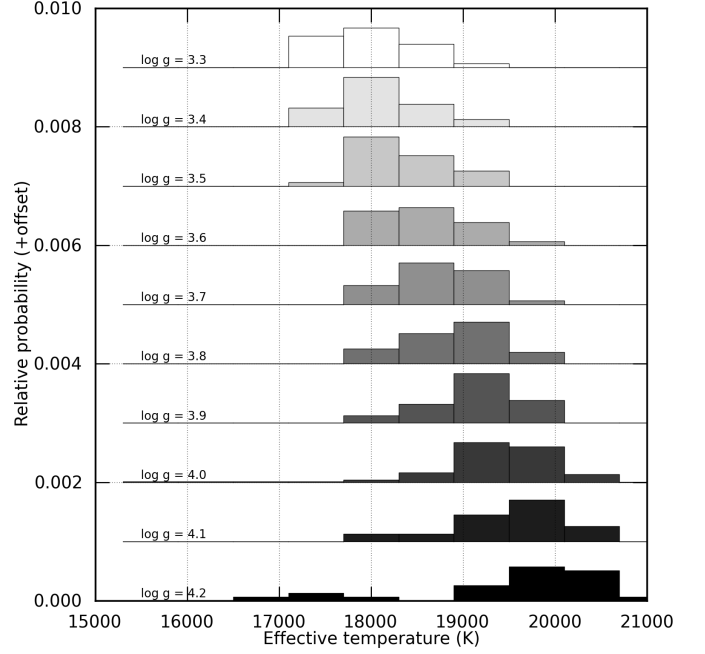


Fig. 12. Histograms of effective temperatures of the primary component derived from line ratios of SiII/SiIII and SiII/SiIII, assuming different values for the surface gravity $\log g$.

SII4463, SII4824, SII4656, SII4792, SII4278, and SII4217 from the first ionisation stage, and SIII4284 and SIII4253 from the second ionisation stage. The observed EWs of the selected unblended features were calculated by numerical integration of the time-averaged normalised spectra (weighted with the S/N), and compared to the synthetic EWs. We find $17500 < T_{\text{eff}} < 19500$, and notice that lower surface gravity values imply lower effective temperatures (Fig. 12). We attempted to derive a value for the surface gravity and individual abundances by calculating the abundance of each element predicted by each individual line, adopting each $\log g$ parameter in the grid, and then minimising the scatter. These efforts failed, largely because of the unknown scaling factor that has to be included to compensate for the presence of the companion.

Subsequently, we use the χ^2 fitting method to match the observed HARPS spectra with a grid of TLUSTY synthetic spectra (Lanz & Hubeny 2007) when $T_{\text{eff}} \geq 15000$, and ATLAS synthetic spectra (Palacios et al. 2010) when $T_{\text{eff}} < 15000$. Because the secondary component is not visible in all selected lines, we first performed a grid search using a template spectrum of the HeI6678 line and varying the radial velocity of the component v'_r and projected rotational velocity of the component $v'_{\text{eq}} \sin i'$. From this, we fixed the values to $v'_r = 40 \text{ km s}^{-1}$ and $v'_{\text{eq}} \sin i' = 117 \text{ km s}^{-1}$, though the uncertainties are several tens of km s^{-1} on both quantities. We assumed equal metallicities for both components, and demanded that the companion is less evolved than the primary ($\log g' > \log g$). The χ^2 minimisation method was performed with effective temperatures of the primary between 16 000 K and 22 000 K and between 15 000 K and 22 000 K for the secondary, surface gravities between 3.5 and 4.5 for the primary and secondary. The metallicity was chosen to be 0.1, 0.2, 0.5, or 1.0 times the solar value. For all 23 lines, the synthetic spectrum was calculated by taking the average of the synthetic profiles from the primary and secondary, weighted with the synthetic Eddington flux. Each i th profile was

fitted to the data together with a linear continuum via

$$\chi_i^2 = \frac{1}{N} \sum_{\lambda=\lambda_0}^{\lambda_n} \left(\frac{F_{i,\text{obs}}(\lambda)}{a_i\lambda + b_i} - F_{i,\text{syn}}(\lambda) \right)^2,$$

and the total χ^2 value was calculated as the squared sum of the residuals of all line profiles,

$$\chi_1^2 = \frac{1}{N} \sum_{\lambda \in P} \left(\frac{F_{i,\text{obs}}(\lambda)}{a_i\lambda + b_i} - F_{i,\text{syn}}(\lambda) \right)^2,$$

where the index λ runs over all wavelengths in all profiles P . As an alternative test statistic, the χ^2 was also calculated as the average of the χ_i^2 values of all N_p profiles,

$$\chi_2^2 = \frac{1}{N_p} \sum_i \chi_i^2.$$

In the former statistic, the broad Balmer and helium profiles dominate because they cover a broader wavelength range. The latter statistic is designed to give equal weight to all profiles, regardless of the intrinsic width.

The fit procedure does not result in an accurate determination of the two components of the binary system, because the fits to several lines are not satisfactory (e.g., Figs A.1-A.6). This could possibly be explained by shortcomings in atomic physics (e.g., in Si), or by metal mixtures significantly different from the solar mixture that is adopted in the models. A histogram weighted with the χ^2 -values sets the effective temperature and surface gravity of the primary at $T_{\text{eff}} = 17\,000\text{ K}$ and $\log g = 3.5$ (cgs), while the minimum χ^2 is attained at $T_{\text{eff}} = 19\,000\text{ K}$ and $\log g = 4.0$ (cgs). This range is compatible with the temperature determination from the SED fit and the EW method, which are less influenced by the fainter companion. The effective temperature of the companion is $T_{\text{eff}} \leq 15\,000\text{ K}$ according to the χ^2 fitting (the minimum is obtained at the edges of the grid), while the weighted histogram sets it to $T_{\text{eff}} = 16\,000\text{ K}$. The surface gravity is found to be between $\log g = 4.0 - 4.5$ (cgs). The most likely value for the metallicity is half the solar value, but cannot be accurately determined from the used fitting procedure. A fit to all lines with $T_{\text{eff}} = 19\,000\text{ K}$ and $\log g = 4.0$ (cgs) for the primary, and $T_{\text{eff}} = 15\,000\text{ K}$ and $\log g = 4.0$ (cgs), is shown in Figs A.1 to A.6. While the binary fit procedure fails at characterising the secondary, the binary line profile fits result in a better explanation for the asymmetry or broad line components in some of the lines. We conclude that $T_{\text{eff}} = 18\,500 \pm 1000\text{ K}$ and $\log g = 3.8 \pm 0.3$ (cgs) is an acceptable range for the fundamental parameters of the primary, and $T_{\text{eff}} \leq 16\,000\text{ K}$ and $\log g \sim 4$ (cgs) for the secondary.

4. Line profile variability

Pulsational variability can cause radial velocity changes or variability in the line profile itself, depending on the degree and azimuthal order of the mode. They can be detected and identified via a pixel-to-pixel frequency analysis of the profiles (e.g., Zima 2006), or via characterisation of the line profile's moments (Aerts et al. 1992). Given the narrow lines, we opted for the latter method.

The combination of the scarce time sampling of the spectroscopic measurements and the continuously rising trend in radial

velocities (Fig. 10) impedes the detection of sinusoidal variability due to pulsations in the line profiles. For that reason, we applied a linear fit to both features simultaneously, using a function of the form

$$M(t_i) = a_0 + a_1 t_i + a_2 \sin[2\pi(ft_i + a_3)]. \quad (1)$$

Here, a_0 and a_1 represent the parameters of the straight-line fit, and a_2 , f , and a_3 are the amplitude, frequency and phase of the sinusoidal component. We fitted the linear parameters a_j to the time series of moments calculated at all observing times t_i , for a grid of test frequencies f that includes all frequencies detected in the CoRoT light curve. To improve the SNR of the individual profiles, we used the LSD method (Donati et al. 1997) to compute the average profile of the deepest metallic spectral lines. From a fit evaluation via the F -statistic we conclude that there are no significant frequencies. A similar procedure excluding a linear trend also yields no significant frequencies.

5. Discussion and conclusions

We classified HD 50230 as a wide double-lined spectroscopic binary consisting of a component with a low projected rotational velocity, and a component with a moderate projected rotation velocity. From the EW ratios, we established an effective temperature around 18500 K for one component, which agrees with multicolour photometry. From this, we conclude that the narrow-lined component is the primary, and the broad-line component the (cooler) secondary. This is confirmed by a spectral line fitting procedure, where an upper limit for the effective temperature of 16000 K is found. We estimate both stars to be main-sequence stars, and rule out the possibility of a subdwarf companion based on luminosity arguments, although their surface gravities could not be determined accurately from the wings of the Balmer lines, because of the blend of the two components.

The identification of HD 50230 as a spectroscopic binary complicates the interpretation of the light curve compared to the single-star hypothesis. However, the period spacing discovered in Degroote et al. (2010) is unlikely to be affected by the binary light curve, partially because of the high amplitudes of the modes considered, but mainly because the chance of observing such a regular pattern as the combination of two independent frequency patterns, are slim. Less obvious is the connection between the gravity modes and the pressure modes as originating from the same star. The companion is less luminous (the flux ratio is below 0.7) and cooler than the primary, placing it far below the β Cep instability strip. It is therefore unlikely to exhibit this family of pressure modes. Moreover, there are strong hints for rotationally split multiplets in the pressure mode regime, of which the narrow spacing is incompatible with fast rotation as occurs in the secondary component. Possibly, the low-frequency region is contaminated by the signal of the companion at low amplitudes because it resides in the SPB strip and may have a flux as high as half that of the primary. Even so, we do not expect that the gravity-mode spectrum, and in particular the gravity mode period spacing, originates from the companion, since cool B-type stars do not typically exhibit a variability spectrum this dense (see, e.g., Balona et al. 2011; Degroote et al. 2011). We finally remark that, although we give a list of 556 frequencies detected in the CoRoT light curve, the second half of that list should be taken with caution, because they do not satisfy the BIC criterion.

The spectroscopic line profiles show obvious signs of variability on short time scales. In particular, an eight-hour continuous monitoring of the star with HERMES shows a continuously

declining trend in the radial velocities, in contrast to the previously described orbital *rising* trend. The data set at hand did not suffice, however, to find clear periodicities. The origin of this problem could lie in the dense gravity mode region, making the resulting complex beating pattern impossible to separate from the noise. It is hard to imagine, however, that we could detect frequencies in the line profile variations which are not present in the CoRoT light curve at photometric micromagnitude precision. This makes $f_{034} \approx 2.48 \text{ d}^{-1}$ and $f_{050} = 4.92 \text{ d}^{-1}$ the most likely candidates for causing LPV. The latter frequency is excluded as the main contributor to the radial velocity dispersion, based on the eight-hour continuous observation run, which should cover at least one cycle. If we focus on the measurements taken in close succession phased to f_{034} (Fig. 10, red curve), we can see that this frequency is indeed likely to be present in the data, but other frequencies are equally successful at fitting the observations (e.g., Fig. 10, blue and green curve).

In conclusion, we have gathered observational information to begin an in-depth seismic study of the hybrid B-type pulsator HD 50230. Despite its slow surface rotation of the order of $v_{\text{eq}} \approx 10 \text{ km s}^{-1}$, the seismic modelling will have to take into account rotational effects when interpreting small deviations from uniform period spacings (Aerts & Dupret 2011). The ultimate future goal is to reconcile the observational information from the period spacings with that of the lower-order mode spectrum, which will allow a description of the whole interior of the star, from the convective core to the radiative envelope.

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Appendix A: Appendix**Table A.1.** Logbook of spectroscopic observations.

Calendar day (UT) ^a	instrument	T_{exp} (s)	SNR ^b
2008-10-11 08:14:16	CORALIE	1801	56
2008-10-13 07:57:30	CORALIE	1801	50
2008-10-15 08:07:28	CORALIE	1801	52
2008-10-16 08:16:20	CORALIE	1801	58
2008-10-17 08:48:56	CORALIE	1801	64
2008-10-18 08:11:23	CORALIE	1801	61
2008-10-18 08:42:42	CORALIE	1801	58
2008-10-19 08:33:39	CORALIE	1500	58
2009-01-13 04:26:21	CORALIE	1801	54
2009-01-16 06:33:07	CORALIE	1801	55
2009-01-19 04:45:44	CORALIE	1801	54
2009-03-22 02:07:09	CORALIE	1801	58
2009-03-23 01:57:27	CORALIE	1801	56
2009-03-25 02:35:33	CORALIE	1801	44
2009-03-27 03:16:36	CORALIE	1801	48
2009-03-29 02:16:51	CORALIE	1801	52
2009-03-31 02:14:55	CORALIE	1801	55
2009-09-23 04:22:44	HERMES	800	63
2009-09-23 04:46:24	HERMES	1200	71
2009-09-23 05:07:17	HERMES	1200	79
2009-09-23 05:28:11	HERMES	1200	74
2009-09-23 05:49:07	HERMES	1200	76
2009-09-23 06:07:56	HERMES	950	70
2009-11-25 04:49:04	HERMES	1500	74
2009-11-27 01:22:10	HERMES	900	67
2009-11-27 01:38:09	HERMES	900	75
2009-11-28 06:11:23	HERMES	1200	58
2009-11-28 06:32:13	HERMES	1200	58
2009-12-10 06:43:32	HARPS	1200	147
2009-12-18 06:36:40	HARPS	1400	148
2010-11-14 04:05:54	HERMES	1800	101
2010-11-14 05:17:00	HERMES	1057	74
2010-11-16 04:51:22	HERMES	2200	103
2011-01-15 21:22:00	HERMES	1000	55
2011-01-15 21:39:35	HERMES	1000	42
2011-01-15 21:58:45	HERMES	1200	56
2011-01-15 22:19:36	HERMES	1200	59
2011-01-15 22:47:28	HERMES	1200	59
2011-01-15 23:08:19	HERMES	1200	59
2011-01-15 23:29:09	HERMES	1200	62
2011-01-15 23:49:59	HERMES	1200	66
2011-01-15 00:09:59	HERMES	1100	65
2011-01-15 00:30:18	HERMES	1000	65
2011-01-15 00:47:48	HERMES	1000	71
2011-01-15 01:05:19	HERMES	1000	64
2011-01-15 01:22:49	HERMES	1000	61
2011-01-15 01:40:19	HERMES	1000	62
2011-01-15 01:58:12	HERMES	1000	69
2011-01-15 02:15:43	HERMES	1000	65
2011-01-15 02:33:13	HERMES	1000	62
2011-01-15 02:51:33	HERMES	1100	65
2011-01-15 03:11:33	HERMES	1200	73
2011-01-15 03:32:24	HERMES	1200	71
2011-05-03 21:19:35	HERMES	2200	66
2011-05-04 20:52:21	HERMES	1800	95
2011-05-04 21:23:11	HERMES	1800	93
2011-05-05 21:14:42	HERMES	1800	79
2011-09-30 05:39:25	HERMES	2400	116
2011-12-18 05:16:13	HARPS	1200	132
2012-01-10 05:00:07	HARPS	1800	148

a Time of mid-exposure

b Measured around 4550 Å

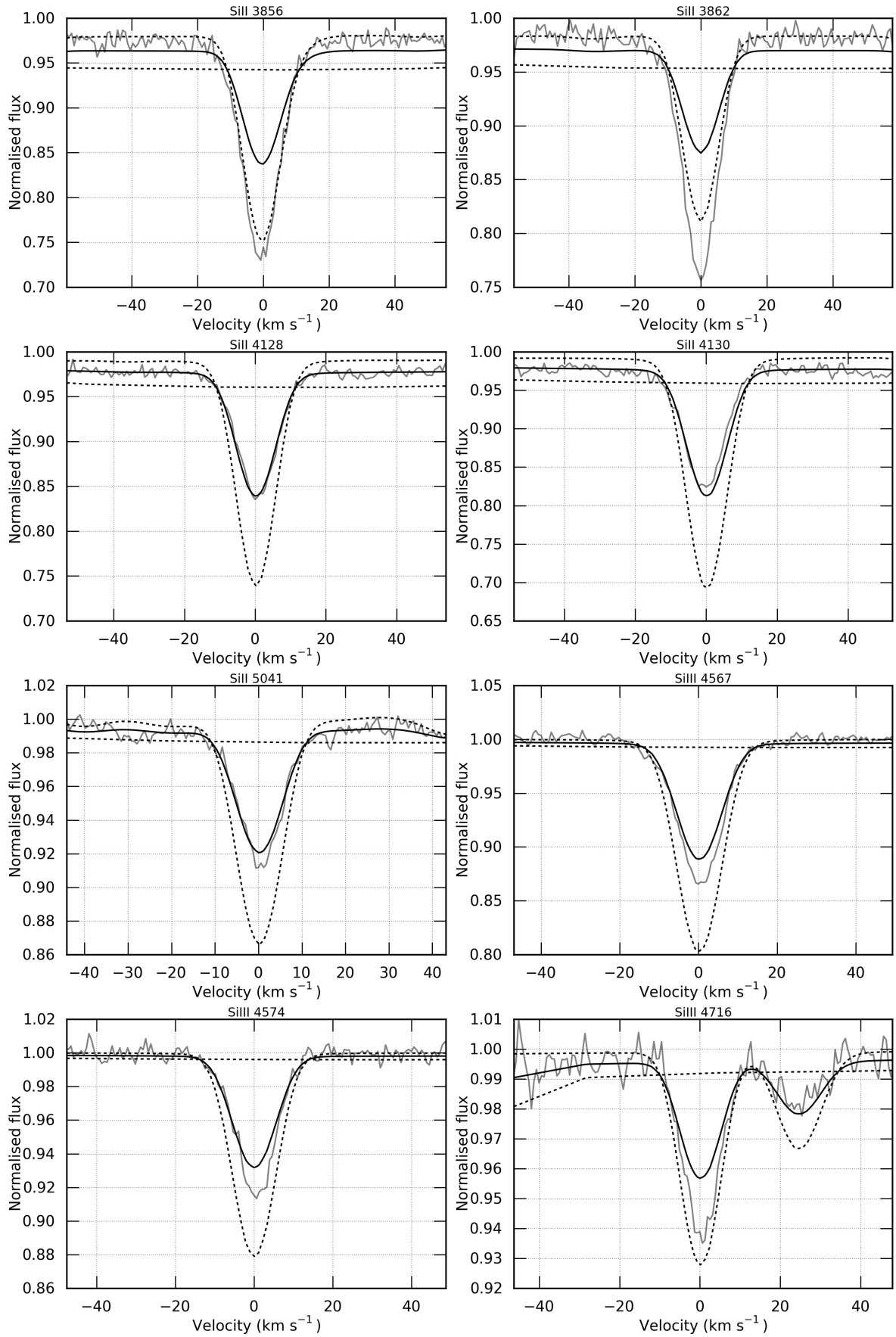


Fig. A.1. Binary fit to a selection of spectral lines. In grey, the average of all observed HARPS spectra of HD 50230 is plotted. Dashed black lines denote the contributions of the separate companions, the solid black line shows the obtained binary fit. Note that due to the simultaneous fitting of the continuum with the lines, the line continuum level can be off if the line does not fit well (i.e., the average quadratic distance between the observed and synthetic line profile is minimised). The fit consists of a primary component with $T_{\text{eff}} = 18\,500$ K and $\log g = 4.0$ (cgs), and $T_{\text{eff}} = 15\,000$ K and $\log g = 4.0$ (cgs) for the secondary.

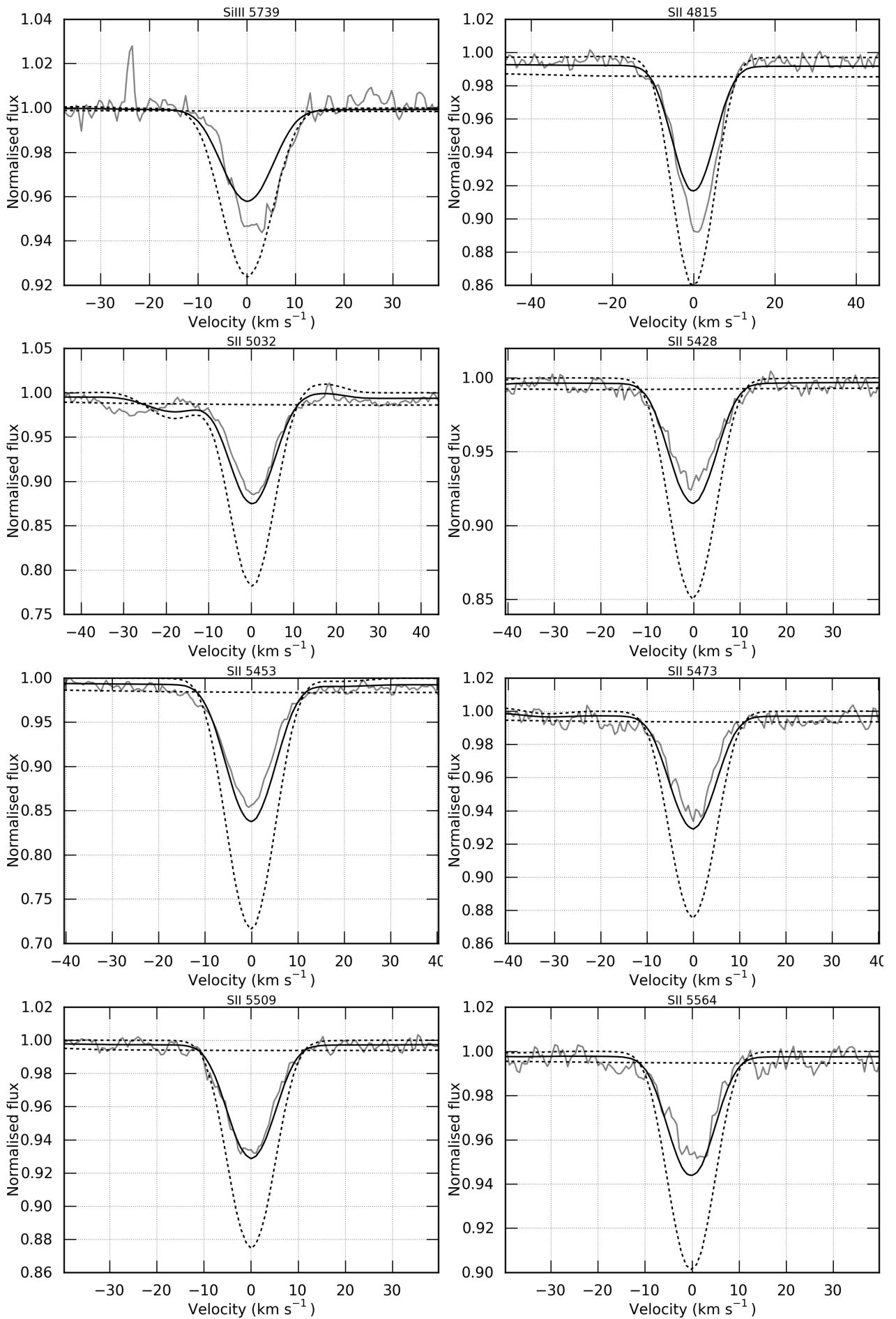


Fig. A.2. Same as Fig. A.1, but for additional spectral lines.

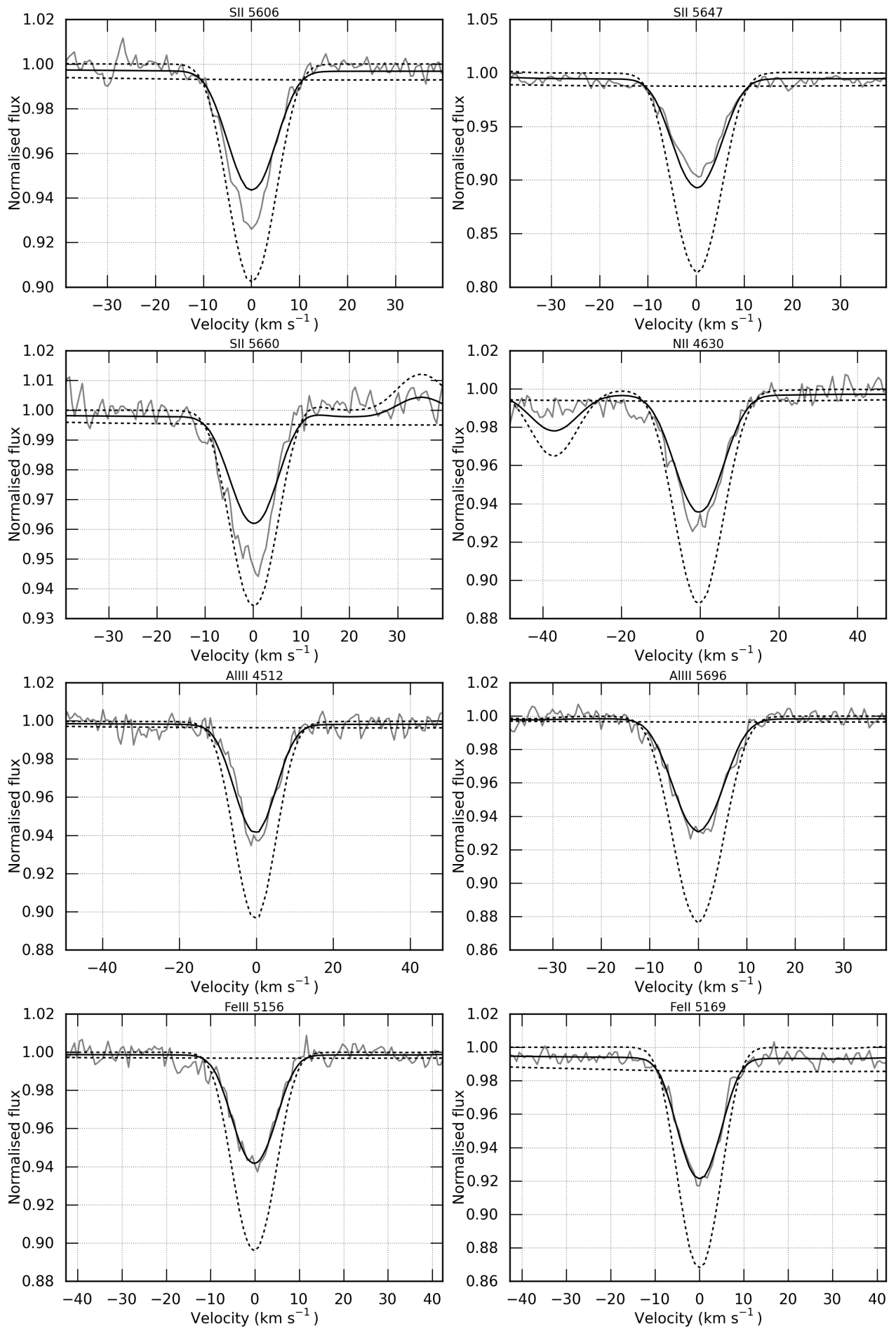


Fig. A.3. Same as Fig. A.1, but for additional spectral lines.

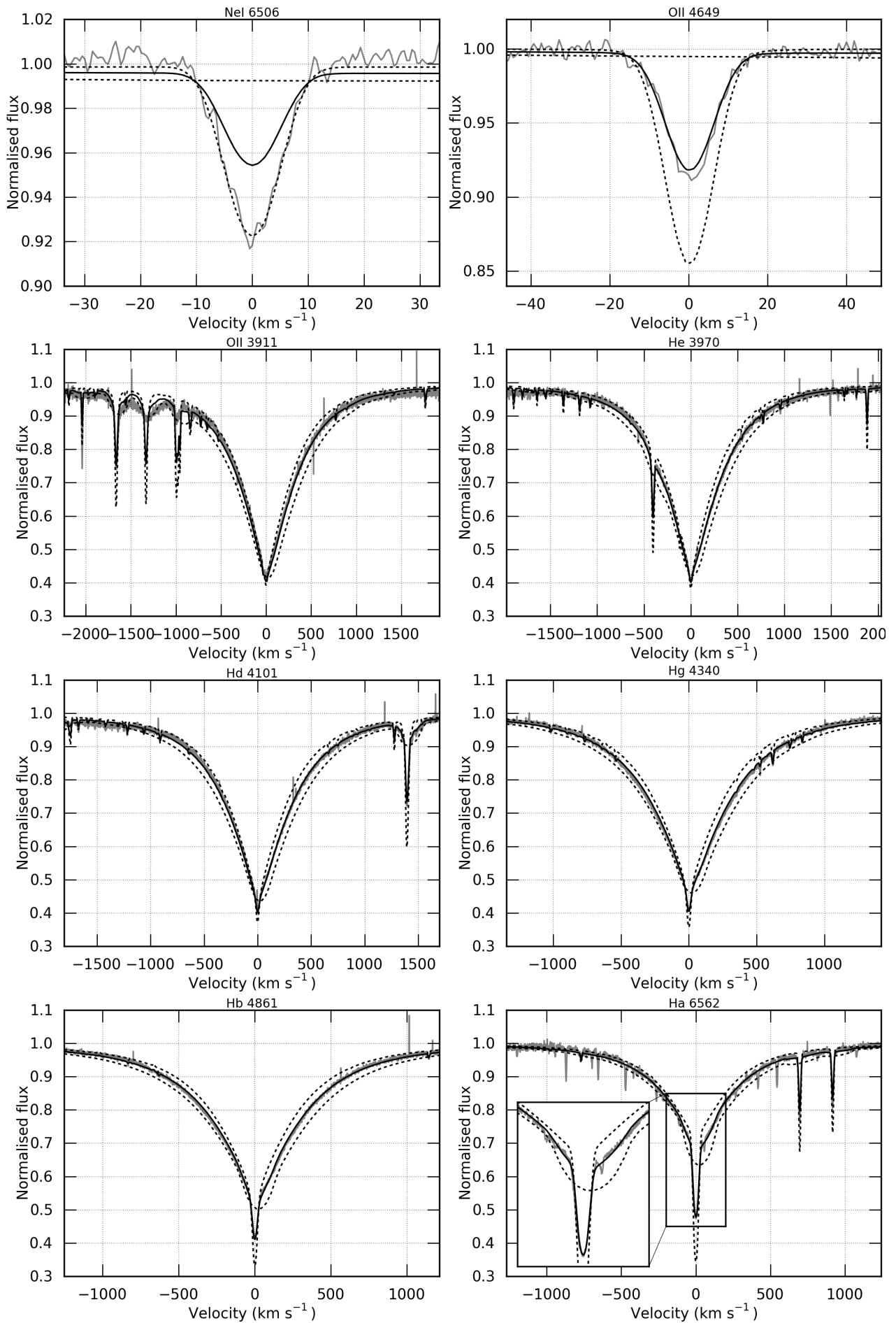


Fig. A.4. Same as Fig. A.1, but for additional spectral lines.

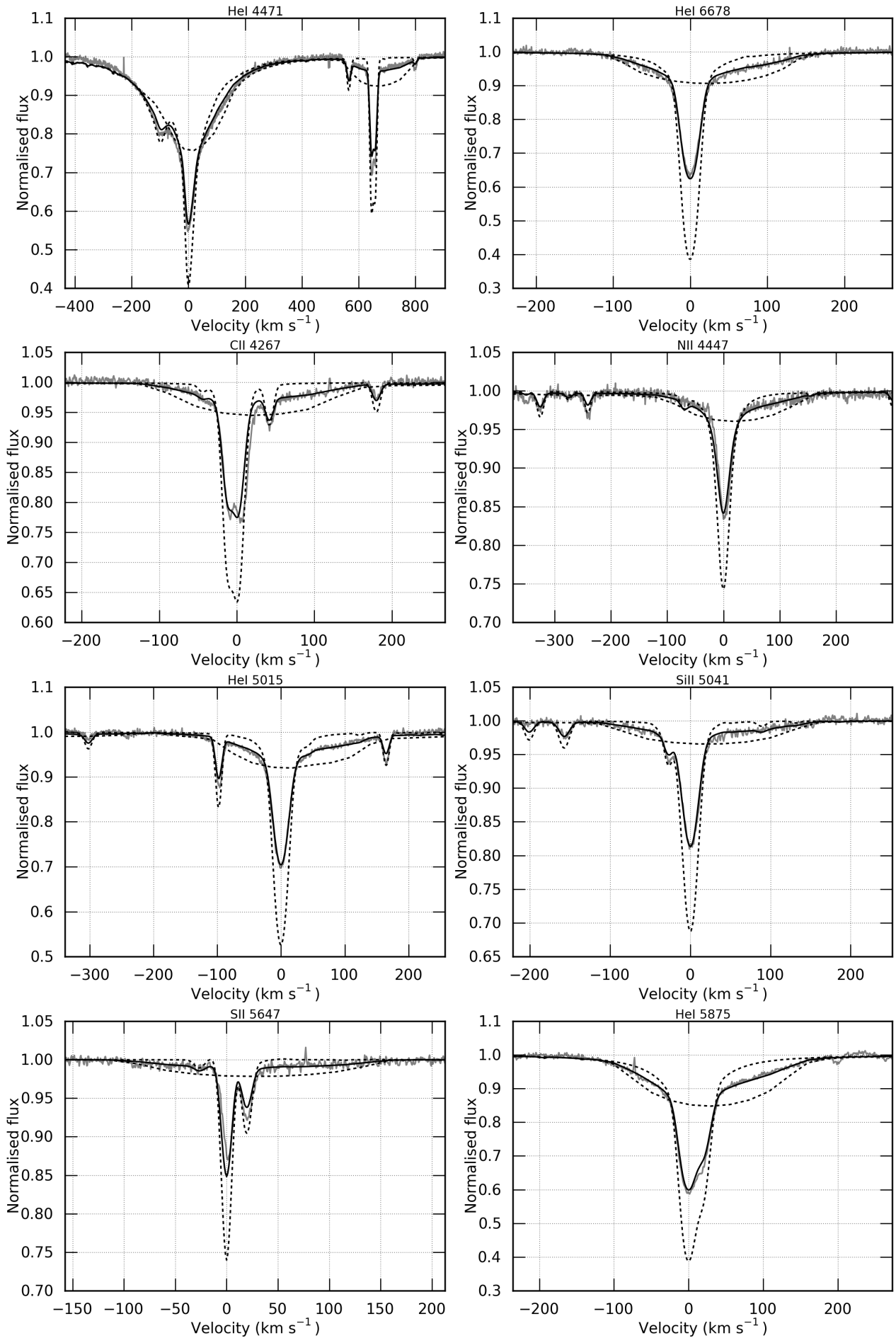


Fig. A.5. Same as Fig. A.1, but for additional spectral lines.

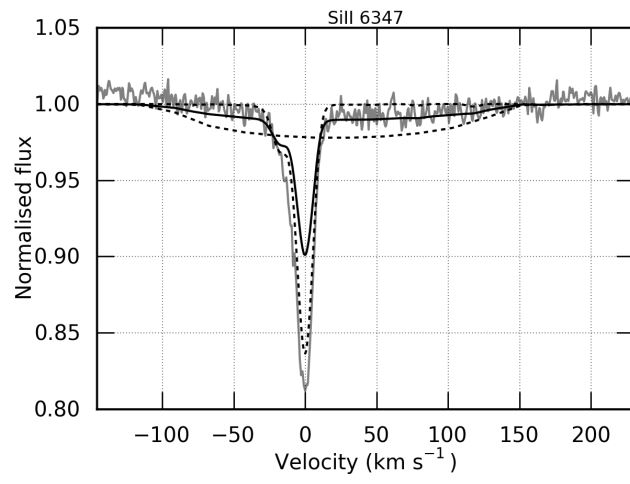


Fig. A.6. Same as Fig. A.1, but for additional spectral lines.

Table A.2. Frequency list and statistics from a Fourier analysis with iterative nonlinear updates of the parameters.

nr	A [ppm]	$\sigma(A)^1$ [ppm]	f^2 [d ⁻¹]	$\sigma(f)^1$ [10 ⁻³ d ⁻¹]	ϕ [2 π]	$\sigma(\phi)^1$ [2 π]	SNR ³	BIC ⁴	Remark ⁵ s
f_{001}	1746.3	8.05	0.68440	0.02	-0.417	0.005	26.9	1.00000	
f_{002}	1607.6	7.57	0.59342	0.02	0.228	0.005	26.5	0.99938	
f_{003}	1518.2	7.13	0.84085	0.02	0.330	0.005	26.0	0.98990	
f_{004}	1457.3	6.66	0.56340	0.02	0.228	0.005	27.5	0.98873	
f_{005}	1495.0	6.14	0.87931	0.02	0.361	0.004	29.2	0.98372	
f_{006}	1384.3	5.64	0.97367	0.02	-0.469	0.004	30.2	0.97772	
f_{007}	1417.3	5.09	0.65088	0.01	-0.124	0.004	32.5	0.98052	
f_{008}	1307.5	4.55	0.62213	0.01	0.485	0.004	32.6	0.96622	
f_{009}	1087.9	4.13	1.04813	0.02	-0.300	0.004	29.1	0.95598	
f_{010}	937.6	3.78	0.46656	0.02	0.384	0.004	27.6	0.95117	
f_{011}	619.1	3.59	0.80413	0.02	0.045	0.005	20.4	0.94036	
f_{012}	593.1	3.42	0.16960	0.02	-0.023	0.006	19.6	0.92786	$A_r(4) = 0.38$
f_{013}	693.1	3.23	0.01081	0.02	0.272	0.005	20.2	0.92268	Instrumental
f_{014}	544.6	3.08	0.36311	0.02	0.150	0.005	19.5	0.91255	
f_{015}	585.4	2.91	0.49494	0.02	-0.348	0.005	20.0	0.90041	
f_{016}	866.2	2.75	0.73841	0.02	-0.475	0.005	19.9	0.91700	
f_{017}	599.3	2.61	0.83393	0.02	0.069	0.005	17.9	0.90789	
f_{018}	424.7	2.50	0.55296	0.03	0.043	0.006	15.7	0.89985	$A_r(3) = 0.35$
f_{019}	684.3	2.41	0.74129	0.03	-0.473	0.007	14.2	0.88055	
f_{020}	311.3	2.33	0.39690	0.03	-0.442	0.007	13.6	0.87340	$A_r(5) = 0.99$
f_{021}	320.9	2.27	0.22743	0.03	0.140	0.008	12.8	0.87238	
f_{022}	287.4	2.22	0.29863	0.03	-0.028	0.008	12.5	0.86679	
f_{023}	255.4	2.17	0.19445	0.03	-0.108	0.008	11.4	0.85458	$A_r(2) = 0.45$
f_{024}	268.9	2.12	0.79364	0.03	0.179	0.008	11.4	0.84810	$f_{024} = 2f_{020} - 0.022/T$
f_{025}	244.7	2.08	1.17330	0.03	0.201	0.009	11.2	0.84306	$A_r(3) = 0.40$
f_{026}	229.4	2.04	0.13589	0.04	0.095	0.009	10.7	0.83706	
f_{027}	238.3	2.00	0.25214	0.04	0.083	0.009	10.7	0.83195	$A_r(3) = 0.32$
f_{028}	218.1	1.96	1.23685	0.04	0.231	0.009	10.8	0.82696	$A_r(2) = 0.16$
f_{029}	201.7	1.93	0.03016	0.04	-0.468	0.010	9.5	0.82026	
f_{030}	160.7	1.90	0.31895	0.04	-0.149	0.010	9.2	0.81623	$A_r(2) = 0.58$
f_{031}	184.9	1.88	1.01719	0.04	0.083	0.010	9.2	0.81571	
f_{032}	200.8	1.85	0.45078	0.04	-0.194	0.010	9.2	0.80995	$A_r(2) = 0.29$
f_{033}	187.7	1.82	0.43327	0.04	0.319	0.011	9.0	0.80301	$A_r(3) = 0.62$
f_{034}	175.2	1.80	2.47889	0.04	-0.388	0.011	9.0	0.79831	
f_{035}	174.8	1.77	0.52121	0.04	-0.089	0.010	9.1	0.79338	$A_r(2) = 0.46$
f_{036}	167.0	1.75	0.58049	0.04	0.362	0.010	9.1	0.78711	
f_{037}	145.2	1.73	1.38875	0.04	-0.275	0.011	8.5	0.78256	
f_{038}	151.0	1.71	0.32835	0.05	0.087	0.012	8.2	0.77902	$A_r(4) = 0.76$
f_{039}	151.6	1.69	1.30040	0.05	-0.168	0.011	8.2	0.77234	
f_{040}	167.2	1.67	0.47159	0.05	-0.120	0.012	8.1	0.76297	$A_r(2) = 0.58$
f_{041}	137.9	1.65	0.71678	0.05	-0.059	0.012	8.1	0.75858	$A_r(3) = 0.26$
f_{042}	144.9	1.63	0.02140	0.05	0.272	0.012	7.8	0.75683	
f_{043}	150.1	1.61	0.18925	0.05	-0.228	0.012	7.8	0.74966	$A_r(3) = 0.60$
f_{044}	133.0	1.59	1.41419	0.05	0.155	0.012	7.8	0.74226	$A_r(3) = 0.89$

Table A.2. continued.

nr	A [ppm]	$\sigma(A)^1$ [ppm]	f^2 [d ⁻¹]	$\sigma(f)^1$ [10 ⁻³ d ⁻¹]	ϕ^4 [2 π]	$\sigma(\phi)^1$ [2 π]	SNR ³	BIC ⁴	Remarks ⁵
f_{045}	127.9	1.58	1.11094	0.05	0.073	0.012	7.7	0.74298	$A_r(2) = 0.11$
f_{046}	121.9	1.56	0.66097	0.05	0.455	0.013	7.6	0.73855	
f_{047}	129.9	1.55	0.41300	0.05	-0.206	0.012	7.7	0.73442	$A_r(4) = 0.97$
f_{048}	128.2	1.53	0.69196	0.05	-0.247	0.013	7.4	0.73181	$A_r(2) = 0.12$
f_{049}	118.6	1.52	0.27648	0.05	0.229	0.013	7.3	0.72311	$f_{049} = 0.5f_{018} + 0.000/T$
f_{050}	117.1	1.50	4.92176	0.05	0.347	0.013	18.6	0.71576	
f_{051}	119.4	1.49	1.06269	0.05	0.044	0.013	7.5	0.70890	$A_r(2) = 0.13$
f_{052}	93.4	1.48	0.63683	0.05	-0.452	0.013	7.6	0.70599	$f_{052} = 2f_{030} - 0.148/T$
f_{053}	101.9	1.46	0.33909	0.05	0.457	0.012	7.7	0.70061	$f_{053} = 2f_{012} - 0.014/T$
f_{054}	102.6	1.45	0.87331	0.05	-0.366	0.013	7.5	0.69268	
f_{055}	97.0	1.43	0.23695	0.05	-0.095	0.013	7.4	0.68498	$f_{055} = 0.5f_{040} + 0.158/T$
f_{056}	100.4	1.42	0.10382	0.05	-0.496	0.013	7.4	0.67717	
f_{057}	114.9	1.41	0.38843	0.05	-0.120	0.013	7.4	0.67354	$f_{057} = 2f_{023} - 0.063/T$
f_{058}	105.8	1.40	0.70762	0.06	0.419	0.014	6.8	0.66919	$f_{058} = 0.5f_{044} + 0.072/T$
f_{059}	99.8	1.39	0.30854	0.06	-0.328	0.014	6.9	0.66234	$A_r(2) = 0.14$
f_{060}	101.8	1.38	0.99383	0.06	0.231	0.014	6.9	0.65480	
f_{061}	89.0	1.37	0.05512	0.06	-0.334	0.014	6.8	0.65309	
f_{062}	89.6	1.36	0.92507	0.06	-0.193	0.015	6.6	0.64865	$A_r(2) = 0.86$
f_{063}	92.9	1.35	0.76802	0.06	-0.474	0.014	6.8	0.64319	$A_r(2) = 0.35$
f_{064}	91.3	1.34	1.62385	0.06	0.080	0.015	6.7	0.64346	
f_{065}	100.2	1.33	0.28863	0.06	0.436	0.014	6.8	0.64007	$A_r(2) = 0.37$
f_{066}	101.5	1.32	1.13907	0.06	-0.274	0.014	6.8	0.63358	
f_{067}	73.5	1.32	0.66625	0.06	0.309	0.015	6.6	0.62966	$A_r(2) = 0.11$
f_{068}	95.4	1.31	0.21685	0.06	0.463	0.015	6.5	0.62193	$f_{068} = 0.5f_{033} + 0.029/T$
f_{069}	93.7	1.30	0.48483	0.06	-0.183	0.015	6.6	0.61849	
f_{070}	104.5	1.29	0.50992	0.06	-0.320	0.014	6.8	0.61112	$f_{070} = 3f_{012} + 0.155/T$
f_{071}	84.5	1.28	0.58681	0.06	0.146	0.014	6.7	0.60721	$f_{071} = 0.5f_{025} + 0.021/T$
f_{072}	89.8	1.27	0.35177	0.06	-0.094	0.015	6.5	0.60056	$A_r(2) = 0.10$
f_{073}	85.0	1.26	1.92772	0.06	-0.025	0.015	6.5	0.59638	
f_{074}	79.6	1.25	0.26204	0.06	0.085	0.016	6.3	0.59465	$f_{074} = 0.5f_{035} + 0.196/T$
f_{075}	76.9	1.24	1.85059	0.06	0.451	0.016	6.3	0.58989	$f_{075} = 2f_{062} + 0.063/T$
f_{076}	75.4	1.24	0.45931	0.06	0.393	0.016	6.2	0.58218	
f_{077}	76.9	1.23	1.78985	0.07	-0.206	0.016	6.1	0.57715	$A_r(2) = 0.15$
f_{078}	75.6	1.22	1.44419	0.07	-0.168	0.017	5.9	0.57326	
f_{079}	72.4	1.22	1.68407	0.07	0.274	0.017	5.9	0.56962	$A_r(2) = 0.12$
f_{080}	77.0	1.21	0.20545	0.07	0.481	0.017	5.7	0.56995	$f_{080} = 0.5f_{047} - 0.143/T$
f_{081}	70.0	1.20	0.60310	0.07	0.478	0.017	5.7	0.56431	$A_r(2) = 0.77$
f_{082}	70.0	1.20	0.12612	0.07	0.125	0.017	5.8	0.56102	$f_{082} = 0.5f_{027} + 0.006/T$
f_{083}	70.7	1.19	0.84926	0.07	0.115	0.017	5.8	0.55322	$A_r(2) = 0.44$
f_{084}	73.5	1.19	0.06925	0.07	0.098	0.018	5.8	0.55258	
f_{085}	67.8	1.18	0.52774	0.07	0.136	0.018	5.6	0.54622	
f_{086}	64.2	1.17	1.02852	0.07	0.135	0.018	5.6	0.54176	$A_r(2) = 0.54$
f_{087}	68.6	1.17	0.36892	0.07	-0.167	0.018	5.7	0.53726	
f_{088}	71.8	1.16	0.05014	0.08	0.448	0.019	5.3	0.53612	
f_{089}	64.7	1.16	0.65648	0.08	0.294	0.020	5.2	0.53081	$f_{089} = 2f_{038} - 0.029/T$
f_{090}	58.4	1.15	0.78122	0.08	0.414	0.020	5.2	0.52715	$A_r(2) = 0.16$
f_{091}	58.8	1.15	1.26909	0.08	0.028	0.020	5.2	0.52142	$A_r(3) = 0.64$
f_{092}	57.9	1.15	1.30827	0.08	-0.281	0.020	5.2	0.51766	$A_r(3) = 0.20$
f_{093}	56.5	1.14	2.83825	0.08	-0.312	0.020	5.2	0.51582	$A_r(2) = 0.17$
f_{094}	57.4	1.14	0.90071	0.08	0.201	0.020	5.1	0.51145	$f_{094} = 2f_{032} - 0.116/T$
f_{095}	57.2	1.13	1.12668	0.08	-0.225	0.021	5.1	0.50750	
f_{096}	72.9	1.13	0.37814	0.08	-0.159	0.021	5.1	0.49946	$f_{096} = 2f_{043} - 0.049/T$
f_{097}	55.6	1.12	1.77753	0.08	-0.058	0.021	5.1	0.49521	$A_r(2) = 0.15$
f_{098}	54.0	1.12	1.60206	0.08	-0.283	0.021	5.1	0.49298	
f_{099}	53.5	1.11	1.37348	0.08	-0.223	0.021	5.1	0.48582	
f_{100}	53.2	1.11	13.96810	0.08	-0.127	0.021	17.3	0.48434	Satellite orbit
f_{101}	51.5	1.11	1.09234	0.09	0.254	0.021	5.0	0.47594	$A_r(2) = 0.17$
f_{102}	54.3	1.10	0.57314	0.09	-0.254	0.022	4.9	0.47366	$A_r(2) = 0.18$
f_{103}	47.2	1.10	1.80208	0.09	0.497	0.022	4.9	0.47249	
f_{104}	48.9	1.10	2.67758	0.09	0.230	0.022	4.8	0.46774	$A_r(2) = 0.14$
f_{105}	53.9	1.09	1.20536	0.09	0.166	0.022	4.8	0.46314	$f_{105} = 2f_{081} - 0.115/T$
f_{106}	48.7	1.09	1.74056	0.09	-0.075	0.022	4.8	0.46072	
f_{107}	49.3	1.09	0.44186	0.09	0.168	0.022	4.8	0.45699	$A_r(3) = 1.01$
f_{108}	45.9	1.08	1.46922	0.09	0.313	0.023	4.8	0.45485	$A_r(2) = 0.18$
f_{109}	47.4	1.08	0.75959	0.09	0.280	0.023	4.8	0.45317	$A_r(2) = 0.35$
f_{110}	45.5	1.08	1.95625	0.09	0.297	0.023	4.8	0.44774	$A_r(2) = 0.47$

Table A.2. continued.

nr	A [ppm]	$\sigma(A)^1$ [ppm]	f^2 [d ⁻¹]	$\sigma(f)^1$ [10 ⁻³ d ⁻¹]	ϕ^4 [2 π]	$\sigma(\phi)^1$ [2 π]	SNR ³	BIC ⁴	Remarks ⁵
f_{111}	46.6	1.07	1.07092	0.09	-0.406	0.023	4.8	0.44539	$A_r(2) = 0.56$
f_{112}	48.9	1.07	0.47748	0.09	0.446	0.023	4.7	0.44218	$A_r(2) = 0.38$
f_{113}	46.8	1.07	2.24188	0.09	-0.251	0.023	4.7	0.44003	$A_r(3) = 0.82$
f_{114}	51.3	1.06	1.19644	0.09	0.053	0.023	4.8	0.43701	
f_{115}	45.7	1.06	1.64129	0.09	0.284	0.023	4.8	0.43335	
f_{116}	48.0	1.06	0.24525	0.09	-0.092	0.023	4.8	0.42863	
f_{117}	47.0	1.06	1.57257	0.10	-0.029	0.024	4.6	0.42394	$A_r(2) = 0.42$
f_{118}	46.1	1.05	2.03992	0.10	-0.188	0.024	4.7	0.41939	$A_r(3) = 0.90$
f_{119}	42.6	1.05	2.38517	0.10	-0.350	0.024	4.7	0.41634	$A_r(2) = 0.17$
f_{120}	43.8	1.05	0.16127	0.10	0.489	0.025	4.5	0.41450	
f_{121}	41.0	1.04	0.81525	0.10	-0.061	0.025	4.5	0.41026	
f_{122}	40.4	1.04	1.71274	0.10	0.166	0.026	4.4	0.40696	$A_r(3) = 0.85$
f_{123}	41.9	1.04	1.65705	0.10	-0.279	0.026	4.4	0.40583	
f_{124}	41.0	1.04	0.72557	0.11	-0.460	0.026	4.4	0.40180	
f_{125}	40.6	1.04	0.98561	0.10	-0.465	0.026	4.4	0.39719	$f_{125} = 3f_{038} + 0.078/T$
f_{126}	42.0	1.03	0.11206	0.10	0.235	0.026	4.4	0.39549	
f_{127}	40.0	1.03	1.50299	0.10	0.051	0.026	4.5	0.39284	
f_{128}	40.8	1.03	1.61370	0.11	0.287	0.026	4.4	0.38854	
f_{129}	39.5	1.03	1.24862	0.11	-0.121	0.027	4.4	0.38595	$A_r(2) = 0.35$
f_{130}	36.1	1.02	11.56967	0.11	0.255	0.027	14.4	0.38210	
f_{131}	38.0	1.02	0.50039	0.11	-0.434	0.027	4.3	0.37680	$A_r(2) = 0.20$
f_{132}	37.8	1.02	0.88320	0.11	0.335	0.027	4.2	0.37628	$f_{132} = 2f_{107} - 0.071/T$
f_{133}	36.8	1.02	0.14552	0.11	-0.134	0.028	4.2	0.37310	$f_{133} = 0.5f_{065} + 0.165/T$
f_{134}	37.3	1.02	0.29985	0.11	-0.160	0.027	4.3	0.37076	
f_{135}	35.5	1.01	2.44967	0.11	-0.388	0.028	4.2	0.36641	$A_r(2) = 0.24$
f_{136}	36.0	1.01	1.43633	0.11	-0.498	0.028	4.2	0.36486	
f_{137}	35.8	1.01	3.99983	0.11	0.071	0.028	5.1	0.36188	Possibly instrumental
f_{138}	34.7	1.01	2.64995	0.11	0.183	0.028	4.2	0.35841	
f_{139}	35.7	1.01	0.22756	0.12	0.200	0.029	4.2	0.35722	
f_{140}	37.0	1.00	1.48038	0.12	-0.474	0.029	4.2	0.35313	
f_{141}	34.4	1.00	2.05624	0.12	-0.450	0.029	4.2	0.35013	$f_{141} = 2f_{086} - 0.110/T$
f_{142}	31.9	1.00	2.01131	0.12	-0.154	0.029	4.2	0.34738	Possibly instrumental
f_{143}	33.8	1.00	1.93591	0.12	-0.401	0.029	4.2	0.34492	$A_r(2) = 0.59$
f_{144}	34.9	1.00	2.58606	0.12	0.292	0.029	4.2	0.34288	$A_r(2) = 0.52$
f_{145}	34.5	0.99	2.47262	0.12	0.110	0.029	4.2	0.33920	$f_{145} = 2f_{028} - 0.148/T$
f_{146}	33.7	0.99	1.03535	0.12	0.345	0.029	4.2	0.33534	$A_r(2) = 1.00$
f_{147}	33.9	0.99	1.39917	0.12	-0.492	0.029	4.2	0.33264	
f_{148}	32.5	0.99	0.93513	0.12	-0.000	0.030	4.1	0.32844	
f_{149}	32.6	0.99	1.85675	0.12	-0.289	0.030	4.1	0.32622	$A_r(2) = 0.36$
f_{150}	32.1	0.99	0.09185	0.12	-0.403	0.031	4.0	0.32545	
f_{151}	31.4	0.98	3.67033	0.12	0.392	0.031	4.6	0.32495	
f_{152}	31.5	0.98	2.09615	0.13	0.146	0.031	4.1	0.32232	
f_{153}	35.4	0.98	1.34512	0.12	0.334	0.031	4.1	0.32065	$A_r(2) = 0.38$
f_{154}	30.7	0.98	1.36284	0.13	0.259	0.031	4.0	0.31779	
f_{155}	33.9	0.98	1.55137	0.13	-0.403	0.031	4.1	0.31699	$A_r(2) = 0.36$
f_{156}	32.2	0.98	1.53504	0.12	-0.331	0.031	4.2	0.31200	$f_{156} = 2f_{063} - 0.138/T$
f_{157}	30.8	0.97	2.60446	0.13	0.488	0.031	4.1	0.31075	$A_r(2) = 0.18$
f_{158}	31.1	0.97	0.42345	0.13	0.207	0.031	4.1	0.30748	$f_{158} = 0.5f_{083} - 0.161/T$
f_{159}	33.5	0.97	0.51761	0.13	0.156	0.031	4.1	0.30559	$f_{159} = 0.5f_{146} - 0.010/T$
f_{160}	35.1	0.97	1.99554	0.13	0.201	0.031	4.1	0.30262	Possibly instrumental
f_{161}	38.6	0.97	2.00191	0.13	0.309	0.032	4.1	0.29921	Possibly instrumental
f_{162}	30.4	0.97	27.94528	0.12	-0.098	0.031	17.1	0.29531	Satellite orbit
f_{163}	31.2	0.97	0.91566	0.13	0.420	0.033	4.0	0.29223	
f_{164}	31.4	0.96	0.94908	0.13	0.346	0.032	4.1	0.29052	
f_{165}	30.7	0.96	1.02092	0.13	0.161	0.033	4.0	0.28531	$f_{165} = 0.5f_{118} + 0.132/T$
f_{166}	31.5	0.96	1.97272	0.13	-0.355	0.033	4.0	0.28380	
f_{167}	28.1	0.96	0.82520	0.13	0.351	0.033	4.0	0.28389	$f_{167} = 2f_{047} - 0.110/T$
f_{168}	29.6	0.96	0.49233	0.13	-0.187	0.033	4.0	0.27961	
f_{169}	30.3	0.96	1.76840	0.14	0.306	0.034	3.9	0.27793	
f_{170}	29.3	0.96	2.02220	0.14	-0.169	0.034	3.9	0.27610	$A_r(2) = 0.57$
f_{171}	31.0	0.95	1.11960	0.14	-0.394	0.035	3.9	0.27349	$f_{171} = 0.5f_{113} - 0.183/T$
f_{172}	28.0	0.95	0.85563	0.14	-0.265	0.035	3.9	0.27158	$f_{172} = 0.5f_{122} - 0.102/T$
f_{173}	26.2	0.95	2.53849	0.14	-0.147	0.035	3.9	0.26847	$f_{173} = 2f_{091} + 0.042/T$
f_{174}	27.0	0.95	1.08335	0.14	-0.056	0.035	3.9	0.26700	
f_{175}	26.0	0.95	2.28315	0.14	-0.040	0.036	3.8	0.26398	
f_{176}	25.9	0.95	2.14239	0.15	0.033	0.036	3.8	0.26393	$f_{176} = 2f_{111} + 0.074/T$

Table A.2. continued.

nr	A [ppm]	$\sigma(A)^1$ [ppm]	f^2 [d ⁻¹]	$\sigma(f)^1$ [10 ⁻³ d ⁻¹]	ϕ^4 [2 π]	$\sigma(\phi)^1$ [2 π]	SNR ³	BIC ⁴	Remarks ⁵
f_{177}	26.5	0.95	0.64255	0.15	0.179	0.036	3.8	0.26017	
f_{178}	30.5	0.95	1.73403	0.15	0.088	0.037	3.8	0.25833	
f_{179}	29.0	0.95	1.10690	0.15	-0.362	0.037	3.8	0.25359	$f_{179} = 2f_{018} + 0.134/T$
f_{180}	26.4	0.94	2.07685	0.15	0.163	0.037	3.8	0.25418	$A_r(2) = 0.29$
f_{181}	26.3	0.94	1.21916	0.15	0.471	0.037	3.8	0.25185	$A_r(2) = 0.65$
f_{182}	25.3	0.94	0.74839	0.15	0.461	0.037	3.8	0.24990	$A_r(4) = 0.94$
f_{183}	26.4	0.94	0.53394	0.15	0.367	0.037	3.8	0.24858	
f_{184}	26.4	0.94	0.40486	0.15	-0.476	0.037	3.8	0.24654	
f_{185}	25.9	0.94	0.31661	0.15	0.279	0.038	3.7	0.24507	
f_{186}	25.1	0.94	2.17863	0.15	0.131	0.038	3.7	0.24300	
f_{187}	25.1	0.94	2.30777	0.15	-0.492	0.038	3.8	0.24279	
f_{188}	23.6	0.94	0.13809	0.15	0.255	0.038	3.7	0.24019	
f_{189}	27.2	0.94	1.72341	0.16	-0.061	0.039	3.7	0.23943	
f_{190}	24.2	0.93	1.91068	0.15	0.087	0.038	3.8	0.23720	
f_{191}	23.7	0.93	2.20193	0.16	0.405	0.039	3.7	0.23468	
f_{192}	21.5	0.93	1.82492	0.16	-0.237	0.039	3.7	0.23316	
f_{193}	23.9	0.93	0.35887	0.16	0.064	0.040	3.7	0.23074	$f_{193} = 0.5f_{041} + 0.066/T$
f_{194}	22.9	0.93	3.07279	0.16	-0.127	0.041	3.7	0.23013	
f_{195}	23.0	0.93	1.79673	0.17	0.355	0.041	3.6	0.22925	
f_{196}	29.7	0.93	0.41948	0.17	0.108	0.042	3.6	0.22651	
f_{197}	23.2	0.93	1.74856	0.17	-0.473	0.042	3.6	0.22276	
f_{198}	22.5	0.93	13.97701	0.17	0.323	0.041	10.2	0.22024	Satellite orbit
f_{199}	22.3	0.93	0.18392	0.17	-0.345	0.042	3.6	0.21930	
f_{200}	22.2	0.93	1.42711	0.17	0.433	0.042	3.6	0.21645	
f_{201}	22.6	0.92	1.94717	0.17	0.275	0.042	3.6	0.21596	
f_{202}	21.5	0.92	0.19901	0.17	0.415	0.042	3.6	0.21226	$f_{202} = 0.5f_{020} + 0.076/T$
f_{203}	22.8	0.92	2.26110	0.17	0.060	0.043	3.6	0.21058	$A_r(2) = 0.37$
f_{204}	22.4	0.92	2.46297	0.17	0.222	0.043	3.6	0.20838	
f_{205}	20.9	0.92	0.86643	0.17	-0.126	0.043	3.6	0.20784	$f_{205} = 2f_{033} - 0.016/T$
f_{206}	22.4	0.92	1.63364	0.17	-0.486	0.043	3.6	0.20533	$A_r(2) = 0.36$
f_{207}	21.6	0.92	1.40788	0.17	0.308	0.043	3.6	0.20499	
f_{208}	21.2	0.92	2.16186	0.17	0.070	0.043	3.6	0.20276	
f_{209}	21.2	0.92	3.03903	0.17	0.034	0.043	3.6	0.20032	
f_{210}	21.1	0.92	1.27779	0.17	0.492	0.043	3.6	0.19980	
f_{211}	22.3	0.92	2.61413	0.17	0.327	0.043	3.6	0.20027	
f_{212}	20.8	0.92	27.93571	0.18	0.384	0.044	13.8	0.19858	Satellite orbit
f_{213}	24.8	0.92	1.16669	0.18	-0.355	0.044	3.6	0.19542	
f_{214}	23.1	0.91	1.17812	0.17	0.344	0.042	3.8	0.19501	
f_{215}	21.3	0.91	2.48424	0.18	-0.201	0.044	3.6	0.19264	$A_r(2) = 0.67$
f_{216}	22.5	0.91	0.70205	0.18	0.387	0.044	3.6	0.19066	
f_{217}	20.4	0.91	0.67811	0.18	0.077	0.044	3.7	0.18875	$f_{217} = 4f_{012} - 0.038/T$
f_{218}	20.9	0.91	1.59447	0.18	-0.464	0.044	3.6	0.18712	
f_{219}	20.8	0.91	0.32461	0.18	-0.343	0.044	3.6	0.18454	
f_{220}	20.7	0.91	1.65142	0.18	0.406	0.044	3.6	0.18337	$f_{220} = 4f_{047} - 0.080/T$
f_{221}	21.4	0.91	2.92779	0.18	-0.183	0.044	3.7	0.18332	$A_r(2) = 0.37$
f_{222}	24.0	0.91	2.96672	0.18	0.284	0.044	3.7	0.17946	
f_{223}	21.1	0.91	0.25417	0.19	0.049	0.046	3.5	0.17881	
f_{224}	20.0	0.91	3.51008	0.19	-0.037	0.046	3.5	0.17848	$A_r(2) = 0.75$
f_{225}	19.1	0.91	1.83725	0.19	-0.332	0.047	3.5	0.17538	
f_{226}	20.0	0.91	1.14423	0.19	0.311	0.047	3.5	0.17520	
f_{227}	19.5	0.91	0.78760	0.19	-0.434	0.046	3.6	0.17517	$f_{227} = 0.5f_{117} + 0.181/T$
f_{228}	19.0	0.91	2.29942	0.19	-0.333	0.047	3.5	0.17162	
f_{229}	18.8	0.90	0.54465	0.19	0.489	0.047	3.5	0.16948	
f_{230}	18.7	0.90	8.82103	0.20	-0.054	0.048	8.8	0.16778	
f_{231}	18.4	0.90	0.44999	0.20	-0.466	0.049	3.4	0.16625	
f_{232}	18.4	0.90	2.50208	0.20	0.012	0.049	3.4	0.16512	
f_{233}	18.5	0.90	3.24274	0.20	-0.021	0.049	3.5	0.16378	$A_r(2) = 0.64$
f_{234}	18.0	0.90	2.80504	0.20	-0.478	0.050	3.4	0.16106	
f_{235}	18.5	0.90	2.37979	0.20	-0.458	0.050	3.4	0.15982	
f_{236}	16.8	0.90	0.95925	0.20	0.353	0.050	3.4	0.16034	$A_r(2) = 0.92$
f_{237}	18.2	0.90	1.29309	0.20	0.168	0.050	3.4	0.15710	$f_{237} = 0.5f_{144} + 0.008/T$
f_{238}	18.3	0.90	10.30467	0.20	-0.308	0.050	8.6	0.15718	
f_{239}	18.7	0.90	0.23764	0.20	-0.150	0.050	3.4	0.15475	$f_{239} = 0.5f_{112} - 0.150/T$
f_{240}	18.0	0.90	1.69262	0.20	0.024	0.050	3.4	0.15533	$A_r(2) = 0.52$
f_{241}	17.6	0.90	3.43621	0.20	-0.346	0.050	3.6	0.15204	
f_{242}	17.7	0.90	2.21007	0.20	0.336	0.051	3.4	0.15202	

Table A.2. continued.

nr	A [ppm]	$\sigma(A)^1$ [ppm]	f^2 [d ⁻¹]	$\sigma(f)^1$ [10 ⁻³ d ⁻¹]	ϕ^4 [2 π]	$\sigma(\phi)^1$ [2 π]	SNR ³	BIC ⁴	Remarks ⁵
f_{243}	17.7	0.90	0.90706	0.21	0.114	0.051	3.4	0.15104	
f_{244}	17.1	0.90	0.60867	0.20	-0.013	0.051	3.4	0.14870	$f_{244} = 0.5f_{181}-0.124/T$
f_{245}	17.5	0.90	10.34957	0.21	0.444	0.051	8.5	0.14994	$A_r(2) = 0.55$
f_{246}	17.4	0.90	4.01892	0.21	0.260	0.052	3.9	0.14785	Possibly instrumental
f_{247}	17.4	0.89	2.82658	0.21	-0.138	0.052	3.4	0.14675	
f_{248}	19.7	0.89	2.97142	0.21	0.217	0.052	3.4	0.14466	
f_{249}	17.1	0.89	0.56659	0.21	0.012	0.053	3.4	0.14298	$f_{249} = 3f_{043}-0.159/T$
f_{250}	16.7	0.89	1.01209	0.21	0.091	0.053	3.4	0.14156	$f_{250} = 0.5f_{170}+0.136/T$
f_{251}	21.6	0.89	0.97908	0.21	-0.267	0.052	3.4	0.14073	$f_{251} = 0.5f_{110}+0.131/T$
f_{252}	19.8	0.89	0.96687	0.22	0.316	0.053	3.4	0.13899	$f_{252} = 0.5f_{143}-0.148/T$
f_{253}	16.7	0.89	13.18488	0.22	0.234	0.054	8.0	0.13625	
f_{254}	16.3	0.89	4.52950	0.22	-0.139	0.054	4.1	0.13575	
f_{255}	17.1	0.89	2.66101	0.22	-0.319	0.054	3.4	0.13387	
f_{256}	16.9	0.89	11.67037	0.22	-0.146	0.054	8.1	0.13239	
f_{257}	17.0	0.89	2.69353	0.22	-0.439	0.054	3.4	0.13290	
f_{258}	16.7	0.89	0.71568	0.22	0.050	0.054	3.4	0.13055	
f_{259}	16.6	0.89	0.03496	0.22	0.263	0.054	3.4	0.13038	
f_{260}	15.5	0.89	1.91969	0.22	-0.304	0.055	3.3	0.12899	$f_{260} = 2f_{236}+0.164/T$
f_{261}	16.7	0.89	1.51783	0.22	-0.324	0.056	3.3	0.12804	$f_{261} = 2f_{109}-0.185/T$
f_{262}	16.4	0.89	0.92371	0.23	0.222	0.056	3.3	0.12674	
f_{263}	16.1	0.89	1.07721	0.23	-0.356	0.056	3.3	0.12679	
f_{264}	15.9	0.89	0.34665	0.23	0.186	0.056	3.3	0.12521	$f_{264} = 0.5f_{048}+0.091/T$
f_{265}	15.3	0.89	0.38252	0.23	-0.202	0.058	3.2	0.12492	
f_{266}	15.9	0.89	3.76236	0.23	-0.275	0.058	3.5	0.12462	
f_{267}	15.4	0.89	2.56437	0.24	0.067	0.059	3.2	0.12363	
f_{268}	15.0	0.89	2.62509	0.24	-0.206	0.059	3.2	0.12159	
f_{269}	14.9	0.88	1.75619	0.24	-0.405	0.059	3.2	0.11930	$f_{269} = 0.5f_{224}+0.158/T$
f_{270}	16.0	0.88	2.12473	0.24	-0.322	0.059	3.2	0.12073	$f_{270} = 2f_{051}-0.089/T$
f_{271}	14.8	0.88	1.55895	0.24	-0.116	0.059	3.2	0.11866	
f_{272}	14.3	0.88	1.24222	0.24	0.322	0.059	3.2	0.11807	$f_{272} = 0.5f_{215}+0.014/T$
f_{273}	15.0	0.88	0.27928	0.24	-0.339	0.059	3.3	0.11729	
f_{274}	14.9	0.88	2.84841	0.24	0.085	0.060	3.2	0.11579	
f_{275}	15.0	0.88	2.74424	0.24	0.203	0.060	3.2	0.11606	
f_{276}	14.8	0.88	2.21672	0.24	-0.254	0.061	3.2	0.11442	
f_{277}	13.6	0.88	1.86757	0.25	0.309	0.061	3.2	0.11278	
f_{278}	13.9	0.88	0.62489	0.25	0.478	0.061	3.2	0.11307	$f_{278} = 0.5f_{129}+0.079/T$
f_{279}	14.0	0.88	3.60054	0.25	-0.296	0.061	3.4	0.11239	
f_{280}	14.3	0.88	0.15533	0.25	0.018	0.062	3.2	0.11240	$f_{280} = 0.5f_{059}+0.146/T$
f_{281}	14.5	0.88	2.89341	0.25	0.463	0.063	3.2	0.11065	$A_r(2) = 0.59$
f_{282}	14.0	0.88	3.14177	0.26	-0.138	0.063	3.2	0.11011	
f_{283}	13.7	0.88	1.80766	0.26	0.046	0.064	3.1	0.10864	
f_{284}	13.6	0.88	4.46867	0.26	0.093	0.065	3.6	0.10788	
f_{285}	14.1	0.88	1.66323	0.26	0.398	0.064	3.1	0.10700	$A_r(3) = 1.11$
f_{286}	14.4	0.88	1.67549	0.26	-0.151	0.064	3.2	0.10624	$A_r(2) = 0.78$
f_{287}	13.5	0.88	0.67124	0.26	0.376	0.065	3.1	0.10567	$f_{287} = 0.5f_{153}-0.181/T$
f_{288}	13.8	0.88	11.67659	0.26	0.431	0.066	6.8	0.10334	
f_{289}	13.1	0.88	3.25252	0.27	-0.024	0.066	3.1	0.10285	
f_{290}	13.1	0.88	2.78090	0.27	-0.417	0.067	3.1	0.10248	$A_r(2) = 0.89$
f_{291}	13.3	0.88	1.60731	0.27	-0.225	0.068	3.0	0.10184	
f_{292}	12.9	0.88	5.29461	0.27	-0.379	0.068	3.9	0.10103	
f_{293}	14.3	0.88	3.31652	0.28	0.266	0.069	3.1	0.10013	
f_{294}	13.1	0.88	3.33408	0.27	0.485	0.068	3.1	0.09869	
f_{295}	12.7	0.88	1.90239	0.28	-0.050	0.069	3.0	0.09817	
f_{296}	12.7	0.88	1.15772	0.28	0.080	0.070	3.0	0.09610	$A_r(2) = 0.76$
f_{297}	12.5	0.88	1.35039	0.28	-0.169	0.069	3.0	0.09539	
f_{298}	12.3	0.88	5.15586	0.29	0.267	0.071	3.3	0.09539	$A_r(2) = 0.54$
f_{299}	12.3	0.87	4.24198	0.29	0.201	0.071	3.3	0.09492	$f_{299} = 3f_{044}-0.081/T$
f_{300}	12.4	0.87	2.15004	0.29	-0.096	0.071	3.0	0.09374	$f_{300} = 3f_{041}-0.042/T$
f_{301}	11.8	0.87	2.08924	0.29	-0.088	0.072	3.0	0.09445	
f_{302}	12.1	0.87	0.22190	0.29	-0.220	0.072	3.0	0.09286	$f_{302} = 0.5f_{107}+0.133/T$
f_{303}	13.0	0.87	2.36163	0.29	0.419	0.072	3.0	0.09185	
f_{304}	12.6	0.87	2.88413	0.29	0.176	0.072	3.0	0.09109	
f_{305}	12.0	0.87	2.23286	0.29	-0.255	0.072	3.0	0.08905	
f_{306}	14.4	0.87	2.94834	0.29	0.372	0.072	3.0	0.08935	
f_{307}	13.7	0.87	0.55490	0.29	-0.325	0.072	3.0	0.08800	$f_{307} = 0.5f_{045}-0.078/T$
f_{308}	13.0	0.87	0.58102	0.27	-0.433	0.068	3.2	0.08754	

Table A.2. continued.

nr	A [ppm]	$\sigma(A)^1$ [ppm]	f^2 [d ⁻¹]	$\sigma(f)^1$ [10 ⁻³ d ⁻¹]	ϕ^4 [2 π]	$\sigma(\phi)^1$ [2 π]	SNR ³	BIC ⁴	Remarks ⁵
f_{309}	12.3	0.87	0.77542	0.29	-0.493	0.072	3.0	0.08779	$f_{309} = 0.5f_{155}-0.035/T$
f_{310}	12.0	0.87	3.53966	0.29	-0.157	0.072	3.2	0.08926	
f_{311}	12.3	0.87	0.47108	0.29	0.016	0.073	3.0	0.08756	
f_{312}	11.6	0.87	3.71277	0.30	-0.107	0.073	3.2	0.08682	$f_{312} = 2f_{149}-0.099/T$
f_{313}	12.1	0.87	3.37120	0.30	0.357	0.073	3.0	0.08601	
f_{314}	11.9	0.87	1.62001	0.30	-0.027	0.073	3.0	0.08648	$f_{314} = 0.5f_{233}-0.186/T$
f_{315}	11.9	0.87	1.32111	0.30	0.275	0.074	3.0	0.08413	
f_{316}	11.7	0.87	0.89558	0.30	-0.250	0.074	3.0	0.08270	$f_{316} = 0.5f_{077}+0.089/T$
f_{317}	11.7	0.87	1.39063	0.30	0.309	0.074	3.0	0.08137	$f_{317} = 0.5f_{290}+0.025/T$
f_{318}	11.7	0.87	2.78957	0.30	-0.465	0.074	3.0	0.08112	
f_{319}	12.7	0.87	2.34791	0.30	-0.499	0.075	3.0	0.07978	$f_{319} = 2f_{025}+0.178/T$
f_{320}	11.6	0.87	11.49174	0.30	-0.094	0.075	5.9	0.08025	
f_{321}	11.9	0.87	1.48878	0.30	-0.271	0.075	3.0	0.08003	
f_{322}	11.2	0.87	0.63471	0.31	-0.366	0.076	3.0	0.07951	$f_{322} = 0.5f_{091}+0.023/T$
f_{323}	11.3	0.87	3.35202	0.31	-0.020	0.076	3.1	0.07827	$f_{323} = 2f_{286}+0.144/T$
f_{324}	11.4	0.87	6.12939	0.31	-0.162	0.076	4.1	0.07812	
f_{325}	10.6	0.87	2.29301	0.31	-0.451	0.077	3.0	0.07780	
f_{326}	11.2	0.87	14.23666	0.31	0.257	0.078	6.3	0.07753	
f_{327}	11.0	0.87	41.91439	0.31	-0.026	0.077	9.5	0.07539	Satellite orbit
f_{328}	10.7	0.87	3.09242	0.32	0.250	0.079	2.9	0.07525	
f_{329}	10.4	0.87	0.09817	0.32	0.352	0.079	2.9	0.07595	
f_{330}	11.0	0.87	1.49688	0.32	0.375	0.080	2.9	0.07369	$f_{330} = 2f_{182}+0.014/T$
f_{331}	11.3	0.87	2.44265	0.32	0.046	0.080	2.9	0.07464	
f_{332}	10.7	0.87	14.12407	0.32	-0.254	0.080	6.2	0.07392	
f_{333}	10.7	0.87	5.21914	0.33	0.091	0.081	3.4	0.07315	
f_{334}	10.7	0.87	4.08033	0.33	-0.286	0.081	3.1	0.07221	$f_{334} = 2f_{118}+0.067/T$
f_{335}	10.6	0.87	0.52596	0.33	-0.078	0.082	2.9	0.07438	
f_{336}	10.7	0.87	2.39777	0.33	0.407	0.082	2.9	0.07164	
f_{337}	10.5	0.87	2.55815	0.33	0.490	0.083	2.9	0.07090	
f_{338}	10.3	0.87	3.27998	0.33	-0.281	0.083	2.9	0.06980	
f_{339}	10.7	0.87	4.27821	0.34	0.497	0.083	3.2	0.06981	
f_{340}	10.5	0.87	4.21662	0.33	-0.327	0.083	3.2	0.06814	$A_r(2) = 0.64$
f_{341}	10.9	0.87	2.95919	0.34	-0.317	0.083	2.9	0.06722	$A_r(2) = 0.60$
f_{342}	10.5	0.86	13.95843	0.34	-0.289	0.084	5.9	0.06772	Satellite orbit
f_{343}	10.3	0.86	3.77847	0.34	0.312	0.084	3.0	0.06863	
f_{344}	10.2	0.86	1.52682	0.34	-0.150	0.086	2.8	0.06576	
f_{345}	10.0	0.86	55.88281	0.35	0.003	0.087	8.4	0.06601	Satellite orbit
f_{346}	9.8	0.86	1.70267	0.35	-0.278	0.088	2.8	0.06419	
f_{347}	10.0	0.86	0.07403	0.35	0.101	0.088	2.8	0.06417	
f_{348}	9.9	0.86	0.28723	0.35	-0.338	0.088	2.8	0.06300	$f_{348} = 0.5f_{102}+0.090/T$
f_{349}	9.8	0.86	1.19017	0.36	-0.266	0.088	2.8	0.06323	$f_{349} = 3f_{020}-0.073/T$
f_{350}	9.7	0.86	0.69606	0.36	-0.327	0.089	2.8	0.06367	
f_{351}	10.0	0.86	3.29494	0.36	-0.279	0.089	2.8	0.06262	
f_{352}	9.6	0.86	2.31598	0.36	-0.069	0.089	2.8	0.06084	$f_{352} = 2f_{296}+0.074/T$
f_{353}	9.8	0.86	11.62631	0.36	-0.381	0.089	5.2	0.06156	
f_{354}	9.5	0.86	0.39076	0.36	0.241	0.090	2.8	0.06043	$f_{354} = 0.5f_{090}+0.020/T$
f_{355}	9.6	0.86	7.47322	0.36	-0.001	0.090	4.4	0.06041	
f_{356}	9.3	0.86	1.05705	0.37	-0.054	0.091	2.7	0.05946	
f_{357}	10.3	0.86	2.99213	0.36	-0.211	0.090	2.7	0.05829	
f_{358}	9.6	0.86	5.17493	0.36	0.248	0.091	3.2	0.05923	$f_{358} = 0.5f_{245}+0.019/T$
f_{359}	9.4	0.86	0.83242	0.36	-0.485	0.090	2.8	0.05833	$f_{359} = 0.5f_{285}+0.109/T$
f_{360}	10.1	0.86	3.04748	0.37	0.396	0.091	2.8	0.05700	
f_{361}	9.4	0.86	1.42052	0.37	-0.100	0.092	2.7	0.05719	$f_{361} = 0.5f_{093}+0.192/T$
f_{362}	9.6	0.86	3.11592	0.37	-0.208	0.092	2.8	0.05545	
f_{363}	10.1	0.86	3.01142	0.37	-0.120	0.091	2.8	0.05384	
f_{364}	9.7	0.86	3.49287	0.37	0.220	0.092	2.8	0.05565	
f_{365}	10.0	0.86	3.62949	0.37	0.361	0.092	2.9	0.05504	
f_{366}	9.2	0.86	0.59391	0.37	0.217	0.093	2.8	0.05505	
f_{367}	9.9	0.86	3.95127	0.37	-0.182	0.093	2.9	0.05537	
f_{368}	9.7	0.86	3.93394	0.37	-0.309	0.091	3.0	0.05421	
f_{369}	9.4	0.86	3.38629	0.38	0.406	0.094	2.8	0.05310	$f_{369} = 2f_{240}+0.143/T$
f_{370}	9.0	0.86	2.01215	0.38	0.284	0.094	2.7	0.05366	Possibly instrumental
f_{371}	9.1	0.86	4.36189	0.38	-0.387	0.095	3.0	0.05337	
f_{372}	9.3	0.86	2.97994	0.38	0.092	0.095	2.7	0.05318	
f_{373}	9.3	0.86	1.78734	0.39	0.492	0.096	3.0	0.05243	
f_{374}	8.9	0.86	1.74109	0.38	-0.018	0.095	2.8	0.05180	

Table A.2. continued.

nr	A [ppm]	$\sigma(A)^1$ [ppm]	f^2 [d ⁻¹]	$\sigma(f)^1$ [10 ⁻³ d ⁻¹]	ϕ^4 [2 π]	$\sigma(\phi)^1$ [2 π]	SNR ³	BIC ⁴	Remarks ⁵
f_{375}	9.1	0.86	3.52887	0.39	-0.385	0.096	2.8	0.05069	$A_r(2) = 0.62$
f_{376}	8.7	0.86	3.57094	0.38	-0.187	0.095	2.9	0.05132	
f_{377}	8.7	0.86	1.31381	0.39	0.224	0.096	2.8	0.05041	$f_{377} = 4f_{038} + 0.057/T$
f_{378}	10.4	0.86	4.42475	0.39	-0.370	0.096	3.0	0.04918	
f_{379}	9.3	0.86	0.17556	0.39	0.177	0.096	2.8	0.05018	$f_{379} = 0.5f_{072} - 0.044/T$
f_{380}	9.3	0.86	3.68772	0.39	0.247	0.096	2.9	0.04795	
f_{381}	9.2	0.86	3.06057	0.39	0.429	0.097	2.8	0.04791	
f_{382}	8.8	0.86	0.04595	0.39	0.335	0.097	2.8	0.04828	
f_{383}	8.8	0.86	8.47783	0.39	0.109	0.097	4.6	0.04697	
f_{384}	9.0	0.86	2.70846	0.39	-0.389	0.097	2.8	0.04634	
f_{385}	8.6	0.86	2.86154	0.39	-0.484	0.098	2.8	0.04720	
f_{386}	8.7	0.86	1.87532	0.40	0.388	0.098	2.8	0.04598	
f_{387}	9.2	0.86	4.41457	0.40	0.256	0.099	3.0	0.04543	
f_{388}	8.6	0.86	4.95584	0.40	0.114	0.100	3.1	0.04520	
f_{389}	8.7	0.86	1.21103	0.40	0.264	0.100	2.7	0.04386	
f_{390}	8.6	0.86	1.44801	0.40	0.078	0.099	2.8	0.04497	$f_{390} = 0.5f_{281} + 0.179/T$
f_{391}	8.6	0.86	2.49354	0.40	0.120	0.100	2.7	0.04376	
f_{392}	8.4	0.86	4.66631	0.40	-0.357	0.100	2.7	0.04244	
f_{393}	8.6	0.86	2.18476	0.40	-0.498	0.100	2.8	0.04181	$f_{393} = 2f_{101} + 0.010/T$
f_{394}	8.4	0.86	0.73486	0.41	-0.425	0.101	2.8	0.04195	$f_{394} = 0.5f_{108} + 0.035/T$
f_{395}	8.7	0.86	2.39013	0.41	-0.081	0.102	2.7	0.04199	
f_{396}	8.7	0.86	2.41797	0.41	-0.343	0.101	2.8	0.04198	
f_{397}	8.4	0.86	1.68414	0.41	-0.325	0.102	2.8	0.04117	
f_{398}	8.4	0.86	27.94975	0.41	0.465	0.102	6.9	0.04006	Satellite orbit
f_{399}	8.4	0.86	4.52239	0.41	-0.168	0.102	2.8	0.04149	$f_{399} = 2f_{203} + 0.026/T$
f_{400}	8.5	0.86	3.61001	0.41	-0.484	0.102	2.9	0.04037	
f_{401}	8.7	0.86	4.09032	0.41	0.402	0.102	2.9	0.04044	
f_{402}	8.3	0.86	3.19477	0.41	0.429	0.103	2.8	0.04102	
f_{403}	8.2	0.86	69.84934	0.41	0.483	0.103	7.4	0.04121	Satellite orbit
f_{404}	7.8	0.86	3.64579	0.41	0.493	0.103	2.8	0.03950	
f_{405}	8.4	0.86	0.61539	0.41	-0.219	0.103	2.8	0.03998	
f_{406}	8.4	0.86	2.68597	0.42	-0.219	0.104	2.8	0.04044	
f_{407}	8.1	0.86	3.84631	0.42	-0.451	0.104	2.8	0.04041	
f_{408}	8.4	0.86	0.46474	0.42	-0.162	0.104	2.8	0.03829	
f_{409}	8.5	0.86	0.88752	0.42	0.081	0.104	2.8	0.03764	$f_{409} = 0.5f_{097} - 0.170/T$
f_{410}	8.4	0.85	0.12954	0.41	0.031	0.103	2.8	0.03695	
f_{411}	8.2	0.85	1.26295	0.42	-0.239	0.104	2.8	0.03775	
f_{412}	8.1	0.85	11.61368	0.42	-0.016	0.104	4.5	0.03745	
f_{413}	8.5	0.85	4.71472	0.42	-0.304	0.105	3.1	0.03708	
f_{414}	8.0	0.85	0.33354	0.42	-0.475	0.105	2.8	0.03658	$f_{414} = 0.5f_{067} + 0.057/T$
f_{415}	8.3	0.85	3.88956	0.42	0.458	0.105	2.9	0.03486	
f_{416}	8.9	0.85	2.27733	0.42	0.456	0.105	2.8	0.03566	
f_{417}	8.0	0.85	2.33818	0.42	0.447	0.105	2.8	0.03484	
f_{418}	8.7	0.85	0.84165	0.43	0.271	0.106	2.8	0.03605	$f_{418} = 0.5f_{079} - 0.052/T$
f_{419}	8.1	0.85	0.81659	0.42	0.434	0.104	2.9	0.03525	$f_{419} = 0.5f_{206} - 0.031/T$
f_{420}	8.4	0.85	1.22566	0.43	0.453	0.107	2.8	0.03476	$f_{420} = 0.5f_{135} + 0.113/T$
f_{421}	8.1	0.85	4.33115	0.43	-0.324	0.107	2.9	0.03396	
f_{422}	8.1	0.85	3.56359	0.43	-0.119	0.108	2.9	0.03349	
f_{423}	8.0	0.85	6.92527	0.44	-0.331	0.108	3.6	0.03299	
f_{424}	7.9	0.85	3.16592	0.44	-0.069	0.108	2.8	0.03298	
f_{425}	7.9	0.85	2.93498	0.43	0.117	0.108	2.8	0.03058	
f_{426}	7.9	0.85	3.36063	0.44	-0.422	0.109	2.8	0.03148	
f_{427}	7.9	0.85	4.54844	0.44	0.202	0.109	2.8	0.03012	
f_{428}	8.2	0.85	4.75052	0.44	-0.050	0.108	3.1	0.02967	
f_{429}	8.0	0.85	4.82933	0.44	-0.436	0.108	3.1	0.02885	
f_{430}	7.9	0.85	5.03889	0.44	-0.000	0.109	2.8	0.02874	
f_{431}	7.9	0.85	1.46473	0.44	0.150	0.109	2.8	0.02866	$f_{431} = 0.5f_{221} + 0.115/T$
f_{432}	7.7	0.85	2.67136	0.44	0.195	0.110	2.8	0.02880	
f_{433}	7.8	0.85	1.36877	0.44	0.155	0.110	2.8	0.02833	
f_{434}	7.7	0.85	1.98046	0.45	-0.275	0.111	2.8	0.02848	Possibly instrumental
f_{435}	7.8	0.85	4.06769	0.45	-0.054	0.111	2.9	0.02823	
f_{436}	7.9	0.85	4.12386	0.45	0.037	0.111	3.0	0.02731	
f_{437}	7.6	0.85	2.43402	0.45	-0.264	0.111	2.8	0.02587	
f_{438}	7.6	0.85	0.27298	0.45	-0.102	0.112	2.8	0.02761	
f_{439}	7.6	0.85	3.58454	0.45	0.071	0.112	2.9	0.02752	
f_{440}	7.6	0.85	2.13089	0.45	-0.028	0.113	2.8	0.02733	

Table A.2. continued.

nr	A	$\sigma(A)^1$	f^2	$\sigma(f)^1$	ϕ^4	$\sigma(\phi)^1$	SNR ³	BIC ⁴	Remarks ⁵
	[ppm]	[ppm]	[d ⁻¹]	[10 ⁻³ d ⁻¹]	[2 π]	[2 π]			
f_{441}	7.5	0.85	4.48234	0.46	0.226	0.114	3.0	0.02683	$f_{441} = 2f_{113}-0.193/T$
f_{442}	7.5	0.85	4.30884	0.46	0.147	0.114	3.0	0.02594	
f_{443}	7.6	0.85	55.88856	0.46	0.495	0.114	6.4	0.02630	Satellite orbit
f_{444}	7.5	0.85	0.99974	0.45	-0.339	0.112	2.8	0.02647	$f_{444} = 2f_{131}-0.143/T$
f_{445}	7.5	0.85	4.15401	0.46	-0.493	0.115	2.9	0.02578	$f_{445} = 2f_{180}+0.043/T$
f_{446}	7.4	0.85	11.00738	0.46	0.468	0.114	4.2	0.02633	
f_{447}	7.4	0.85	2.22466	0.46	-0.133	0.115	2.8	0.02546	
f_{448}	7.4	0.85	1.08933	0.46	-0.308	0.115	2.8	0.02576	
f_{449}	7.4	0.85	10.29305	0.47	0.047	0.115	4.1	0.02542	
f_{450}	7.4	0.85	2.52967	0.47	0.052	0.116	2.8	0.02431	
f_{451}	7.5	0.85	5.22982	0.47	0.054	0.116	3.1	0.02294	
f_{452}	7.4	0.85	7.32390	0.47	0.298	0.116	3.7	0.02171	
f_{453}	7.3	0.85	0.21387	0.47	-0.158	0.117	2.8	0.02131	
f_{454}	7.8	0.85	3.99108	0.47	0.271	0.117	2.9	0.02167	Possibly instrumental
f_{455}	7.4	0.85	3.74334	0.48	-0.105	0.118	2.9	0.02129	
f_{456}	7.0	0.85	14.19349	0.48	0.376	0.118	4.4	0.02139	
f_{457}	7.1	0.85	2.76358	0.48	-0.434	0.119	2.8	0.02147	
f_{458}	7.0	0.85	0.37488	0.48	-0.381	0.120	2.8	0.02073	$f_{458} = 0.5f_{182}+0.093/T$
f_{459}	6.6	0.85	4.02442	0.48	0.487	0.120	2.8	0.02098	$f_{459} = 2f_{370}+0.015/T$
f_{460}	7.3	0.85	3.08213	0.48	0.132	0.120	2.8	0.02090	
f_{461}	7.1	0.85	13.22461	0.49	-0.234	0.121	4.1	0.02068	
f_{462}	6.8	0.85	10.39497	0.49	-0.057	0.121	4.0	0.01999	
f_{463}	6.7	0.85	2.84282	0.49	0.192	0.121	2.8	0.02009	
f_{464}	7.0	0.85	2.26958	0.49	0.399	0.122	2.8	0.01938	
f_{465}	6.8	0.85	1.33883	0.49	-0.079	0.121	2.8	0.01901	$f_{465} = 0.5f_{104}+0.006/T$
f_{466}	6.9	0.85	0.43958	0.49	-0.202	0.122	2.8	0.02003	
f_{467}	7.1	0.85	4.77080	0.50	0.014	0.123	2.9	0.01915	$f_{467} = 2f_{119}+0.063/T$
f_{468}	7.1	0.85	3.20698	0.50	0.461	0.123	2.8	0.01907	
f_{469}	6.7	0.85	4.87159	0.50	-0.062	0.124	2.9	0.01770	
f_{470}	6.7	0.85	4.59364	0.50	0.011	0.125	2.9	0.01806	
f_{471}	6.9	0.85	1.13330	0.50	0.355	0.125	2.8	0.01795	
f_{472}	6.6	0.85	2.03506	0.51	-0.212	0.126	2.8	0.01745	
f_{473}	6.7	0.85	1.37741	0.51	-0.195	0.126	2.8	0.01823	
f_{474}	6.8	0.85	13.10985	0.51	-0.474	0.126	4.0	0.01720	
f_{475}	6.8	0.85	10.97466	0.51	-0.173	0.126	3.9	0.01627	
f_{476}	6.8	0.85	5.39358	0.51	0.459	0.127	3.0	0.01687	
f_{477}	6.7	0.85	2.10838	0.51	-0.427	0.128	2.7	0.01674	$f_{477} = 0.5f_{340}+0.010/T$
f_{478}	6.6	0.85	3.15099	0.51	-0.424	0.128	2.8	0.01644	
f_{479}	6.8	0.85	3.94218	0.52	-0.104	0.128	2.8	0.01546	
f_{480}	6.9	0.85	3.83732	0.52	0.057	0.128	2.8	0.01604	
f_{481}	6.9	0.85	4.00760	0.51	-0.264	0.127	2.8	0.01577	Possibly instrumental
f_{482}	6.7	0.85	4.28592	0.52	0.204	0.129	2.9	0.01455	
f_{483}	6.6	0.85	1.58750	0.52	0.136	0.129	2.8	0.01561	$f_{483} = 4f_{020}-0.014/T$
f_{484}	6.7	0.85	2.54445	0.52	-0.097	0.130	2.8	0.01447	
f_{485}	6.7	0.85	2.57659	0.52	-0.448	0.128	2.8	0.01462	$f_{485} = 0.5f_{298}-0.184/T$
f_{486}	6.6	0.85	4.18834	0.52	0.328	0.130	2.9	0.01388	
f_{487}	6.5	0.85	7.60723	0.52	-0.474	0.130	3.3	0.01416	
f_{488}	7.7	0.85	1.95361	0.53	-0.372	0.131	2.7	0.01474	
f_{489}	7.7	0.85	1.96530	0.48	-0.483	0.120	3.0	0.01368	
f_{490}	6.6	0.85	1.77921	0.52	-0.073	0.130	2.8	0.01219	
f_{491}	6.6	0.85	4.97732	0.53	-0.266	0.131	2.9	0.01190	
f_{492}	6.5	0.85	5.44163	0.53	0.154	0.131	3.0	0.01126	
f_{493}	6.7	0.85	3.48481	0.53	0.198	0.131	2.8	0.01104	
f_{494}	6.5	0.85	1.47918	0.53	-0.456	0.131	2.8	0.01210	$f_{494} = 0.5f_{341}-0.056/T$
f_{495}	6.5	0.85	2.86967	0.53	-0.494	0.132	2.8	0.01032	
f_{496}	6.4	0.85	5.75209	0.54	-0.177	0.134	3.0	0.01034	
f_{497}	6.3	0.85	3.26579	0.54	0.378	0.135	2.8	0.00968	
f_{498}	6.4	0.85	4.04844	0.54	-0.060	0.134	2.9	0.00946	
f_{499}	6.3	0.85	0.75581	0.54	-0.177	0.135	2.9	0.00842	$f_{499} = 3f_{027}-0.084/T$
f_{500}	6.3	0.85	4.49436	0.54	-0.289	0.135	2.9	0.00810	
f_{501}	6.2	0.85	14.26652	0.55	-0.056	0.136	3.9	0.00706	
f_{502}	6.4	0.85	1.09939	0.55	0.080	0.136	2.8	0.00719	
f_{503}	6.3	0.85	2.73405	0.55	-0.395	0.137	2.8	0.00714	
f_{504}	6.3	0.85	4.73636	0.54	0.269	0.135	2.9	0.00596	
f_{505}	6.3	0.85	10.29874	0.55	-0.281	0.136	3.6	0.00702	
f_{506}	6.2	0.85	11.66210	0.55	0.180	0.137	3.7	0.00630	

Table A.2. continued.

nr	A [ppm]	$\sigma(A)^1$ [ppm]	f^2 [d ⁻¹]	$\sigma(f)^1$ [10 ⁻³ d ⁻¹]	ϕ^4 [2 π]	$\sigma(\phi)^1$ [2 π]	SNR ³	BIC ⁴	Remarks ⁵
f_{507}	6.1	0.85	12.99895	0.55	0.261	0.137	3.7	0.00682	Satellite orbit
f_{508}	6.5	0.85	3.42425	0.56	-0.158	0.138	2.8	0.00716	$f_{508} = 2f_{122} - 0.169/T$
f_{509}	6.3	0.85	3.32618	0.55	-0.253	0.136	2.8	0.00725	$f_{509} = 2f_{285} - 0.039/T$
f_{510}	6.2	0.85	0.41076	0.56	0.172	0.139	2.8	0.00714	
f_{511}	6.0	0.85	4.14168	0.57	-0.135	0.141	2.8	0.00677	
f_{512}	5.9	0.85	2.64051	0.57	0.393	0.141	2.7	0.00633	
f_{513}	6.4	0.85	4.89558	0.57	-0.402	0.141	2.9	0.00657	
f_{514}	5.9	0.85	0.65311	0.57	-0.457	0.142	2.7	0.00558	$f_{514} = 0.5f_{092} - 0.141/T$
f_{515}	6.1	0.85	5.87680	0.57	-0.453	0.142	2.9	0.00571	
f_{516}	6.0	0.85	3.66425	0.58	-0.183	0.143	2.7	0.00560	
f_{517}	6.1	0.85	3.49962	0.57	0.086	0.142	2.8	0.00616	
f_{518}	6.1	0.85	13.96591	0.57	-0.227	0.142	3.7	0.00597	Satellite orbit
f_{519}	5.9	0.85	97.79681	0.58	0.001	0.145	5.4	0.00639	Satellite orbit
f_{520}	5.7	0.85	1.76348	0.58	-0.426	0.145	2.7	0.00543	$f_{520} = 0.5f_{375} - 0.131/T$
f_{521}	5.9	0.85	1.02777	0.58	-0.042	0.145	2.7	0.00565	
f_{522}	5.8	0.85	4.51141	0.59	-0.178	0.146	2.8	0.00522	
f_{523}	5.9	0.84	4.85630	0.59	0.184	0.146	2.8	0.00580	
f_{524}	6.1	0.84	11.57707	0.59	-0.202	0.146	3.5	0.00636	
f_{525}	5.8	0.84	2.91693	0.59	-0.302	0.146	2.7	0.00580	
f_{526}	6.1	0.84	13.17660	0.59	0.356	0.147	3.6	0.00665	
f_{527}	6.1	0.84	13.15166	0.59	0.166	0.146	3.6	0.00484	
f_{528}	5.9	0.84	5.71110	0.59	0.375	0.147	2.7	0.00478	
f_{529}	5.7	0.84	2.40494	0.59	0.164	0.147	2.7	0.00526	
f_{530}	5.9	0.84	2.60461	0.59	0.062	0.147	2.7	0.00417	
f_{531}	6.0	0.84	13.19812	0.59	0.066	0.147	3.6	0.00288	
f_{532}	5.8	0.84	4.78435	0.59	0.190	0.147	2.8	0.00334	
f_{533}	5.7	0.84	3.40096	0.59	-0.202	0.147	2.8	0.00372	
f_{534}	5.8	0.84	0.00794	0.59	0.451	0.146	2.7	0.00388	Instrumental
f_{535}	5.8	0.84	5.51550	0.60	-0.234	0.149	2.9	0.00423	
f_{536}	5.8	0.84	4.25822	0.60	-0.284	0.149	2.8	0.00375	
f_{537}	5.7	0.84	7.26256	0.60	-0.146	0.149	3.1	0.00348	
f_{538}	5.6	0.84	15.51801	0.60	0.393	0.149	4.1	0.00388	
f_{539}	5.7	0.84	4.35084	0.60	0.217	0.150	2.7	0.00358	
f_{540}	5.7	0.84	2.24558	0.60	-0.489	0.150	2.8	0.00350	$f_{540} = 3f_{182} + 0.057/T$
f_{541}	5.6	0.84	8.46660	0.60	0.402	0.150	3.2	0.00294	
f_{542}	5.7	0.84	5.61233	0.60	0.414	0.149	2.9	0.00166	
f_{543}	5.6	0.84	3.72443	0.61	-0.046	0.150	2.8	0.00222	
f_{544}	5.6	0.84	5.20780	0.60	-0.446	0.150	2.9	0.00136	$f_{544} = 2f_{157} - 0.154/T$
f_{545}	5.6	0.84	8.33646	0.61	0.089	0.151	3.5	0.00160	
f_{546}	5.6	0.84	10.96582	0.61	-0.115	0.151	3.5	0.00117	
f_{547}	5.6	0.84	5.82037	0.61	0.168	0.151	2.9	0.00047	
f_{548}	5.5	0.84	1.43874	0.61	0.322	0.152	2.8	0.00026	
f_{549}	5.5	0.84	1.71109	0.61	-0.028	0.152	2.8	0.00012	
f_{550}	5.5	0.84	0.94914	0.61	-0.229	0.153	2.8	0.00028	
f_{551}	5.5	0.84	3.25768	0.61	0.169	0.153	2.8	0.00106	
f_{552}	5.6	0.84	10.33750	0.62	0.424	0.153	3.4	0.00079	
f_{553}	5.4	0.84	14.22315	0.63	0.108	0.155	3.6	0.00030	
f_{554}	5.4	0.84	3.92394	0.63	-0.056	0.156	2.7	0.00066	$f_{554} = 3f_{092} - 0.121/T$
f_{555}	5.4	0.84	6.25869	0.63	-0.371	0.156	2.9	0.00000	
f_{556}	5.4	0.84	10.43598	0.63	0.260	0.156	3.4	0.00088	

1 σ denotes the standard deviation

2 Frequencies in boldface are related to the period spacing, frequencies in italics denote the pressure mode candidates

3 SNR calculated in the residual periodogram over a 6 d⁻¹ interval

4 Normalised BIC

5 T denotes the total time span