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How formation time-scales affect the period dependence of the transition between rocky super-Earths and gaseous sub-Neptunes and implications for η_{\oplus}

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ABSTRACT

One of the most significant advances by NASA’s *Kepler* Mission was the discovery of an abundant new population of highly irradiated planets with sizes between those of the Earth and Neptune, unlike anything found in the Solar System. Subsequent analysis showed that at $\sim 1.5 R_{\oplus}$ there is a transition from a population of predominantly rocky super-Earths to non-rocky sub-Neptunes, which must have substantial volatile envelopes to explain their low densities. Determining the origin of these highly irradiated rocky planets will be critical to our understanding of low-mass planet formation and the frequency of potentially habitable Earth-like planets. These short-period rocky super-Earths could simply be the stripped cores of sub-Neptunes, which have lost their envelopes due to atmospheric photo-evaporation or other processes, or they might instead be a separate population of inherently rocky planets, which never had significant envelopes. We suggest an observational path forward to distinguish between these scenarios. Using models of atmospheric photo-evaporation, we show that if most bare rocky planets are the evaporated cores of sub-Neptunes, then the transition radius should *decrease* as surveys push to longer orbital periods, since on wider orbits only planets with smaller less massive cores can be stripped. On the other hand, if most rocky planets formed after their discs dissipate, then these planets will have formed without initial gaseous envelopes. In this case, we use N -body simulations of planet formation to show that the transition radius should *increase* with orbital period, due to the increasing solid mass available in their discs. Moreover, we show that distinguishing between these two scenarios should be possible in coming years with radial velocity follow-up of planets found by *Transiting Exoplanet Survey Satellite*. Finally, we discuss the broader implications of this work for current efforts to measure η_{\oplus} , which may yield significant overestimates if most rocky planets form as evaporated cores.

Key words: planets and satellites: atmospheres – planets and satellites: composition – planets and satellites: physical evolution.

1 INTRODUCTION

One of the key revelations from NASA’s *Kepler* Mission has been the discovery of an abundant new population of short-period planets with transit radii in between the radii of Earth and Neptune. These planets occupy a range of sizes and orbits, which is completely vacant in the Solar System, and so they present a key test of traditional models of planet formation. Below a transit radius

of $\sim 1.5 R_{\oplus}$, most planets with measured masses seem to be consistent with bare rocky compositions with an Earth-like mixture of silicates and iron (Dressing et al. 2015). However, at $\sim 1.5 R_{\oplus}$, there appears to be a transition between a primarily rocky and a primarily non-rocky planet population, where most planets above this size must have large volatile envelopes to explain their lower densities (Rogers 2015). Throughout this paper, we will refer to the radius at which this transition between these rocky super-Earths and non-rocky sub-Neptunes takes place as R_{trans} .

One key question is whether the rocky super-Earths below R_{trans} and the volatile-rich sub-Neptunes above R_{trans} represent a sin-

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gle continuous population, whose initial compositions have been sculpted by post-formation evolution, or two separate populations with distinct formation mechanisms. In the first case, both the rocky super-Earths and non-rocky sub-Neptunes would form quickly, while they were still embedded in their gaseous proto-stellar discs, which typically last up to ~ 10 Myr (Haisch, Lada & Lada 2001; Williams & Cieza 2011). The planets would quickly form their rocky or icy cores and would then accrete gaseous envelopes directly from the disc (e.g. Bodenheimer & Pollack 1986; Lissauer 1993; Pollack et al. 1996). If planets reach approximately half of their total mass in gas, then they will typically undergo run-away accretion to produce gas giants (e.g. Pollack et al. 1996; Alibert et al. 2005; Rafikov 2006). However, most short-period planets should remain well below this limit, and we would instead expect them to have a broad range of initial envelope mass fractions from <1 per cent gas to ~ 50 per cent by mass as determined primarily by disc mass and lifetime, the opacity of dust grains, planetary core mass, the local disc temperature, and orbit of the planets (e.g. Rogers et al. 2011; Ikoma & Hori 2012; Mordasini et al. 2012; Bodenheimer & Lissauer 2014; Mordasini et al. 2014; Lee & Chiang 2015, 2016). The short-period rocky planets, meanwhile, would represent those planets that lost their initial gaseous envelopes through processes like XUV-driven photo-evaporation (e.g. Lopez, Fortney & Miller 2012; Owen & Wu 2013; Jin et al. 2014; Chen & Rogers 2016; Lopez 2016) or atmospheric impact erosion (e.g. Catling & Zahnle 2013; Inamdar & Schlichting 2015; Liu et al. 2015; Schlichting, Sari & Yalinewich 2015; Inamdar & Schlichting 2016).

On the other hand, if the rocky planet population finished assembling their cores after their proto-stellar discs had already dissipated, then they would never have had initial gaseous envelopes and would represent a primordial rocky planet population. This is generally believed to be how the terrestrial planets of the inner solar system formed, where evidence from isotope ratios (e.g. Tera, Papanastasiou & Wasserburg 1974; Allègre, Manhès & Göpel 1995; Touboul et al. 2007; Kleine et al. 2009; Nemchin et al. 2009) suggest that the Earth mostly finished assembling with the moon-forming impact when the solar system was ~ 30 – 100 -Myr old. While there is a great deal of debate about the exact timing (e.g. see Kleine et al. 2009), the final phases of Earth's formation very likely took place long after the proto-solar nebula had already dissipated (Lissauer 1987; Haisch et al. 2001). This is not to say that a primordial rocky planet population would consist of completely airless bodies, they could have significant secondary atmospheres reaching pressures of hundreds of bars that are outgassed from their interiors (Elkins-Tanton & Seager 2008; Schaefer & Fegley 2010). However, any such secondary atmospheres would still only represent a small fraction of a planet's total mass, $\lesssim 0.1$ per cent. Although extremely important for habitability, a hydrogen envelope that is $< 10^{-3} M_{\oplus}$ would be far too thin to measurably change the bulk radius of a transiting planet (Lopez & Fortney 2014), and so from the point of view of current transit surveys any planets without large gaseous envelopes accreted from the disc, including the Earth, are indistinguishable from bare rocks.

Both these scenarios, a primordial rocky planet population and one that originated as the stripped cores of gaseous sub-Neptunes, are consistent with current evidence. This is because the vast majority of confirmed rocky exoplanets are on highly irradiated orbits where planets are extremely vulnerable to losing any primordial envelopes to atmospheric escape (e.g. Léger et al. 2009; Batalha et al. 2011; Pepe et al. 2013; Dressing et al. 2015). Indeed, from the overall distribution of short-period low-mass transiting planets, there is substantial evidence that this population has been significantly

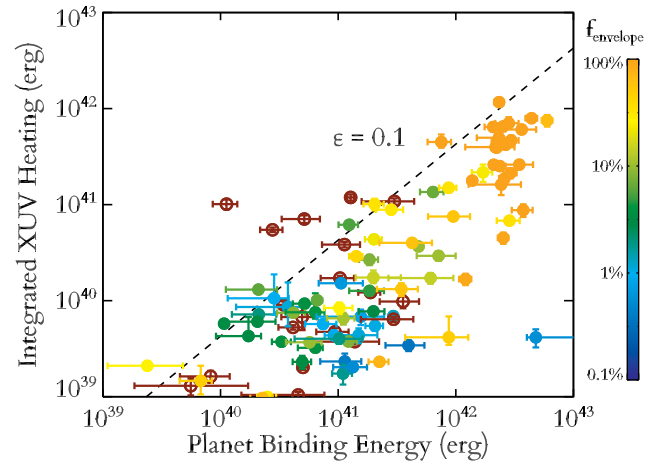


Figure 1. Updated from Lopez & Fortney (2014), this shows the total lifetime XUV heating planets receive, assuming their current orbits and radii, versus their current gravitational binding energy. Included are all currently known transiting planets below $100 M_{\oplus}$ with masses and radii measured to better than 50 per cent. Planets are colour-coded by their H/He envelope mass fractions, assuming Earth-like cores, with planets that are likely bare rock shown by rust-coloured open circles. The dashed line meanwhile corresponds to the photo-evaporation threshold predicted by Lopez et al. (2012), above which we find only bare rocky planets, which have likely been stripped of any primordial envelope.

sculpted by XUV-driven photo-evaporation, or another comparable process (e.g. Jackson, Davis & Wheatley 2011; Lopez, Fortney & Miller 2012; Owen & Wu 2013; Jin et al. 2014; Chen & Rogers 2016). For example, in Fig. 1, we compare the current gravitational binding energy of all low-mass transiting planets to the lifetime-integrated X-ray heating these planets have received. This shows a clear threshold, consistent with models of photo-evaporation, beyond which no known planets have retained gaseous H/He envelopes (Jackson et al. 2011; Lopez et al. 2012; Owen & Jackson 2012; Owen & Wu 2013; Jin et al. 2014; Lopez & Fortney 2014; Chen & Rogers 2016). Moreover, starting from a log-uniform initial distribution of envelope mass fractions, Chen & Rogers (2016) recently showed that photo-evaporation models can adequately reproduce the overall distribution of planet radii for planets with orbits out to ~ 0.25 au.

Critically, however, because of its relatively low XUV irradiation, the Earth could not have formed as the evaporated core of a sub-Neptune. Although very early studies by Hayashi, Nakazawa & Mizuno (1979) and Sekiya, Nakazawa & Hayashi (1980) suggested that Earth could have had a significant initial hydrogen envelope, these studies were conducted before we had good observational constraints on disc lifetimes and stellar XUV histories. Subsequent studies by Erkaev et al. (2013) and Lammer et al. (2014) estimate that the Earth could have only lost up to ~ 10 Earth Ocean equivalents of hydrogen, which corresponds to $\sim 3 \times 10^{-4} M_{\oplus}$ or a surface pressure of ~ 50 bar. A more recent study by Johnstone et al. (2015) found that the Earth could have lost up to 1 per cent of its mass; however, this required assuming both that the Sun was born as a very fast rotator and an extremely high core luminosity. Consequently, $10^{-3} M_{\oplus}$ is likely a reasonable estimate for the maximum primordial envelope that can be evaporated from a typical Earth-like planet, and, therefore, we consider Earth to have been ‘born-rocky’.

As a result, distinguishing between these two scenarios for the origin of the rocky exoplanet population will be critical in our efforts to constrain the frequency of Earth-like planets in the habitable

zones (HZs) of Sun-like stars, as discussed in Section 4.2. To do this, we propose a new observational test to constrain the origin of the rocky planet population by determining how the transition radius R_{trans} between the rocky and non-rocky planet populations scales with the incident flux a planet receives from its star. Using N -body simulations of planet assembly and models of atmospheric photo-evaporation, we show that these two scenarios make opposing predictions for how the transition should scale with flux, thereby providing us with a powerful diagnostic for the formation of rocky planets.

2 SCENARIO 1: ROCKY PLANETS AS THE STRIPPED CORES OF HOT NEPTUNES

2.1 Evaporation model

In order to make predictions for the impact of photo-evaporation on the rocky to non-rocky transition, we use the planet evolution model described in Lopez (2016). Lopez & Fortney (2014) and Lopez (2016) describe this model in greater detail; however, for the benefit of the reader, we briefly summarize the key features here. The evolution model consists of two main components. The first is an interior structure and thermal evolution model, which computes hydrostatic structure models for planets with gaseous envelopes atop rocky cores and then evolves them in time as a planet cools and contracts after formation (Lopez et al. 2012; Lopez & Fortney 2014). This allows us to predict planetary radii as a function of a planet’s core mass, envelope mass, irradiation, and age. The models presented here are computed for solar-composition H/He envelopes atop Earth-like rocky cores with one-third of their mass in and iron core and two-thirds in a silicate mantle orbiting Sun-like stars. We have made one small modification to the structure model from previous publications; to maintain consistency with the later discussion we now use the core mass-radius relationship from Zeng, Sasselov & Jacobsen (2016). This thermal evolution model is then coupled to a parametrized photo-evaporation model as described in Lopez et al. (2012) and Lopez & Fortney (2013), which allows us to predict how the envelope mass and therefore radius change as a planet loses mass to XUV-driven photo-evaporation. Here, we use the modified mass-loss prescription described in Lopez (2016), which accounts for the impact of radiation-recombination limited cooling (Murray-Clay, Chiang & Murray 2009) for the most extremely irradiated planets. Specifically, we compute the overall mass-loss rate by taking the minimum of the energy-limited mass-loss rate using the formalism of Erkaev et al. (2007), which is applicable for planets receiving XUV fluxes higher than Earth today, and the radiation-recombination limited rate, which is appropriate for more irradiated planets, following the formalism of Murray-Clay et al. (2009) and Chen & Rogers (2016). These two rates are described by

$$\dot{M}_{\text{EL}} = -\frac{\epsilon_{\text{XUV}}\pi F_{\text{XUV}}R_{\text{base}}^3}{GM_{\text{p}}K_{\text{tide}}}, \quad (1)$$

$$\begin{aligned} \dot{M}_{\text{RR}} = & -4\pi c_s R_s^2 \mu_{+, \text{wind}} m_{\text{H}} \left(\frac{F_{\text{XUV}} GM_{\text{p}}}{h\nu_0 \alpha_{\text{rec, B}} c_s^2 R_{\text{base}}^2} \right)^{1/2} \\ & \times \exp \left[\frac{GM_{\text{p}}}{R_{\text{base}} c_s^2} \left(\frac{R_{\text{base}}}{R_s} - 1 \right) \right]. \end{aligned} \quad (2)$$

Here, ϵ_{XUV} is a parametrization of the efficiency of photo-evaporation, generally taken to be ~ 10 per cent for solar composition atmospheres (e.g. Jackson et al. 2010; Valencia et al.

2010; Lopez et al. 2012; Jin et al. 2014; Chen & Rogers 2016). F_{XUV} is the XUV flux at a planet’s orbit, which we take as $F_{\text{XUV}} = 29.7(L_s/L_{\odot})(a/\text{au})^{-2}(\text{age}/\text{Gyr})^{-1.23} \text{ erg s}^{-1} \text{ cm}^{-2}$ for Sun-like stars older than 100 Myr based on Ribas et al. (2005), while at ages younger than 100 Myr we assume that the stellar XUV luminosity saturates at $\approx 10^{(-3.5)}L_{\text{bol}}$ based on Jackson et al. (2011). R_{base} and R_s are the radii of the XUV photosphere and the sonic point, respectively, computed following the method described in Lopez (2016), and c_s is the sound speed at the sonic point, typically $\sim 10 \text{ km s}^{-1}$. M_{p} is the total planet mass. K_{tide} is a slight geometric correction factor. Finally, $h\nu_0 \approx 20 \text{ eV}$ is the typically energy of the incoming ionizing radiation and $\alpha_{\text{rec, B}}$ is the case B recombination coefficient for hydrogen.

Taking the minimum of these two rates is a commonly used approximation (e.g. Jin et al. 2014; Chen & Rogers 2016), which approximates the predictions of hydrodynamic mass-loss models (e.g. Murray-Clay et al. 2009; Owen & Jackson 2012), and is generally applicable for planets with H/He envelopes and periods $\lesssim 100 \text{ d}$, where the evaporative wind should be fully collisional. For planets in the HZs of Sun-like stars, this model is not applicable, since there it is necessary to take into account the role of molecular coolants and conduction (e.g. Tian et al. 2008), along with non-collisional and non-thermal escape processes; however, modeling such planets is beyond the scope of this paper and at any rate they are not relevant to the observational predictions made here.

2.2 Evaporation results

Using this model, we then ran a large suite of approximately 20 000 evolution models on a grid covering a range of initial core masses, envelope fractions, and levels of irradiation. The points on this grid were spaced uniformly in log space with cores ranging from 1 to 20 M_{\oplus} , initial envelope fractions from 0.1 to 50 per cent, and bolometric incident fluxes from 10 to 1000 F_{\oplus} . We chose this log-uniform spacing primarily to fully explore the relevant parameter space. However, as we noted before, Chen & Rogers (2016) found that such an initial distribution was able to reproduce the observed radius distribution when photo-evaporation is included. In any case, the general predictions for the flux dependence of the transition radius presented here are insensitive to any of these choices. We allowed these models to start photo-evaporating at 10 Myr, shortly after the end of planet formation, and ended them once the planet reached 5 Gyr, at which point we recorded the final planet radius and envelope fraction.

Fig. 2 summarizes the results from this grid of models. At lower levels of irradiation $\lesssim 100 F_{\oplus}$ and larger radii $\gtrsim 1.5 R_{\oplus}$, we find the population of gas rich sub-Neptunes, which have resisted photo-evaporation. These are the most abundant population of exoplanets found by *Kepler* (Petigura, Marcy & Howard 2013b; Burke et al. 2015) and our model predicts that planets in this size range typically have moderate gas envelopes composing ~ 1 –10 per cent of their total mass (Lopez & Fortney 2014). Meanwhile, at higher levels of irradiation and smaller sizes, we find the population of bare rocky cores that have had their envelopes completely stripped away by photo-evaporation. These are the planets that we are interested in here, and we will discuss the features of this population more below. Finally, in between these two populations, there is a narrow ‘evaporation valley’ in which our model and other evaporation models (e.g. Owen & Wu 2013; Jin et al. 2014) predict that planets should be relatively rare. Although it is not the focus of this paper, the evaporation valley is also a key prediction of these models, which may be diagnostic in constraining the formation and composition of

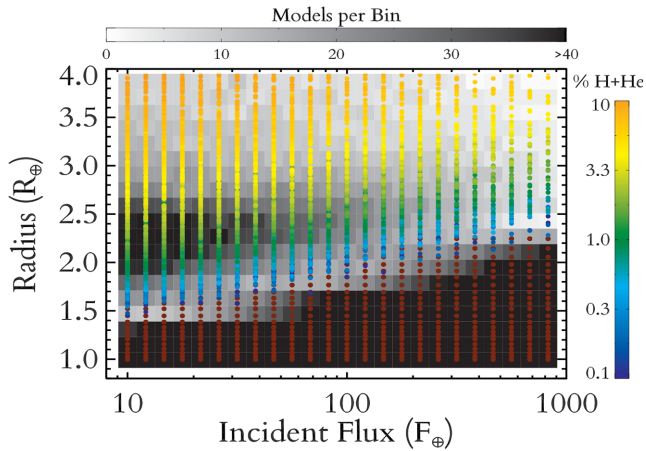


Figure 2. This plots the final planet radius predicted by the evolution model after 5 Gyr of thermal and photo-evaporative evolution versus the incident bolometric flux that a planet receives at its orbit, for planets with solar composition H/He envelopes atop Earth-like cores. $\gtrsim 20\,000$ individual model runs were performed to generate this figure. The results of individual runs are shown by the points, which have been colour-coded by their final H/He envelope mass fraction. Rust-coloured points in the bottom right-hand side indicate bare rocky planets that have completely lost their H/He envelopes. The grey-scale background meanwhile shows the number of models that ended up in each radius-flux bin, where darker shades correspond to a higher density of points, and clear regions correspond to areas devoid of models.

sub-Neptunes and super-Earths, particularly whether they contain large amounts of water or other volatiles ices that formed beyond the snow-line (Lopez & Fortney 2013; Owen & Wu 2013). Indeed, using updated stellar parameters from the CKS survey (Johnson et al. 2017; Petigura et al. 2017) of *Kepler* host stars, Fulton et al. (2017) recently found convincing evidence for the existence of this valley or gap in the radius distribution of *Kepler* planets with orbital periods of $\lesssim 50$ d around Sun-like stars. Moreover, Owen & Wu (2017) and Jin & Mordasini (2017) subsequently showed that the observed valley seen by Fulton et al. (2017) is well-reproduced by photo-evaporation similar to those used here.

In Fig. 3, we have isolated the sample of stripped rocky cores in order to focus-in on the features of this population. Here, we can clearly see that the maximum size of stripped rocky planets produced by photo-evaporation increases with the incident flux a planet receives, and therefore decreases with increasing orbital period. This is because planets with more massive rocky cores are more resistant to photo-evaporation and therefore require greater irradiation to lose their envelopes. Lopez & Fortney (2013) showed that the incident flux necessary for a planet to lose its envelope scales roughly as $F_p \propto M_{\text{core}}^{2.4}$. Therefore, at higher fluxes, planets with more massive cores can be stripped, and these more massive cores are correspondingly larger. At fixed composition, the radius of Earth-like rocky planets scales roughly as $R_p \propto M_p^{1/3.7}$ (Zeng et al. 2016). Combining these two factors, we would expect that if the rocky planet population is primarily produced by photo-evaporative stripping of sub-Neptunes, then the transition radius between rocky and non-rocky planets should scale roughly as $R_{\text{trans}} \propto F_p^{0.11}$, which is consistent with what is seen in Figs 2 and 3.

Note here that in this scenario what matters is the total lifetime XUV heating that a planet receives, $F_{\text{XUV, lifetime}}$. At a given stellar type, the average XUV heating will be proportional to a planet's current bolometric incident flux. However, at later stellar types,

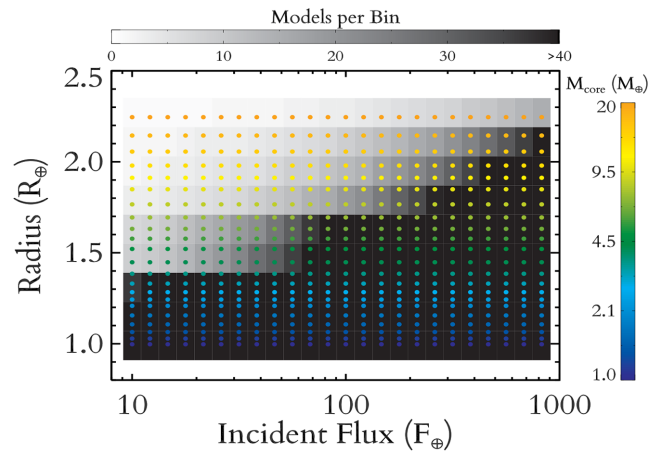


Figure 3. Similar to Fig. 2 except here we only show those rocky planets that have completely lost their gaseous envelopes. Again, the grey-scale shows the number of models that ended up in each radius-flux bin, where darker shades correspond to a higher density of models. The colour-coding meanwhile shows the core-mass for each of our models. Critically, if most rocky planets originate as the evaporated remnants of sub-Neptunes, then the maximum size of bare-rocky planets should *increase* with increasing incident flux (decreasing orbital period).

F_{XUV} increases relative to F_{bol} , so this must be taken into account when comparing planets across a range of host stellar types. Indeed, using X-ray observations of FGK stars from Jackson et al. (2011) and FUV observations of early to mid M dwarfs from Shkolnik & Barman (2014), McDonald et al. (submitted to ApJ) recently calculated the expected lifetime X-ray heating received by *Kepler* planets and found that at fixed present day bolometric flux this scales roughly as $E_{\text{XUV, lifetime}}/F_{\text{bol, current}} \propto M_s^{-3}$. X-ray, FUV, and XUV flux should generally be closely correlated, and so a rocky/non-rocky transition produced by photo-evaporation should produce a transition radius that scales as

$$R_{\text{trans, evap}} \propto F_p^{0.11} M_s^{-0.33}. \quad (3)$$

Alternatively, we can rewrite this in time of orbital period since the incident bolometric flux a planet receives is simply $F_p = L_s/(4\pi a^2)$, where $L_s \propto M_s^4$ for FGK and early M stars with $0.43 \leq M_s \leq 2.0$ (Kippenhahn & Weigert 1990). Therefore, using Kepler's law, $F_p \propto L_s M_s^{-2/3} P^{-4/3} \propto M_s^{10/3} P^{-4/3}$. Putting this in equation (3), we find that

$$R_{\text{trans, evap}} \propto P^{-0.15} M_s^{0.04}. \quad (4)$$

Although, given the observational uncertainties on $F_{\text{XUV, lifetime}}/F_{\text{bol, current}}$, this is also consistent with having no stellar mass dependence when the transition radius is written relative to period, so $R_{\text{trans, evap}} \propto P^{-0.15}$.

One key caveat with these results is that we have used just a single evolution track for the stellar XUV age relation for Sun-like stars. However, recent studies (e.g. Tu et al. 2015) have shown that among Sun-like stars the evolution of stellar XUV luminosity can depend strongly on a star's initial rotation rate, with stars with faster initial rotation staying in the saturation regime for a much larger fraction of their lifetimes. Comparing stellar evolution tracks for G stars at the 10th and 90th percentiles in initial rotation, we estimate that differences in initial rotation could lead to up to factor of ≈ 7 times in the total lifetime integrated XUV luminosity from a Sun-like star at an age of 5 Gyr. Consequently, the observed spread in stellar rotation rates may translate into spread in the location of

the rocky/non-rocky transition, with the largest stripped rocky cores likely occurring around the fastest rotating stars.

Fortunately, however, the slopes of the scaling relations derived above are independent of this scatter in initial rotation rates since these are set by the flux and mass dependence of the photo-evaporative escape process together with the rocky core mass distribution. Therefore, as long as there is no correlation between stellar rotation rates and the core masses of low-mass planets, our results should be unchanged. In short, changes in stellar rotation and XUV emission evolution will affect the location of the evaporation valley, but not its slope. This does, however, suggest that stellar rotation history may be a valuable extra dimension when examining the trends in the observed *Kepler* population, something which should be increasingly possible with dramatically improved stellar parameters, thanks to spectroscopic follow-up surveys (e.g. Johnson et al. 2017; Petigura et al. 2017), asteroseismology (e.g. Silva Aguirre et al. 2015; Van Eylen et al. 2017), and *Gaia* parallaxes (e.g. Berger et al. 2018).

2.3 Additional Caveats: Impact Erosion and Water Worlds

Before we move on to our second scenario, there are two other possible caveats that we wish to address here. First, as previously mentioned, some studies (e.g. Catling & Zahnle 2013; Inamdar & Schlichting 2015; Liu et al. 2015; Schlichting et al. 2015; Inamdar & Schlichting 2016) have suggested that atmospheric erosion by impacts may play a large role in sculpting the compositions of short-period planets in a manner comparable to mass-loss due to photo-evaporation. However, a population that has been stripped by impact erosion will have a different correlation between planet mass, orbital period, and stellar mass than one which has been sculpted by photo-evaporation. For example, in the limit of small impactors, Schlichting et al. (2015) found that to first order the final mass-loss fraction went as $X_{\text{loss}} \propto (m_{\text{imp}} v_{\text{imp}})/(M_{\text{p}} v_{\text{esc}})$, where m_{imp} and v_{imp} are the mass and velocity of the impactor and $v_{\text{esc}} = \sqrt{2GM_{\text{p}}/R_{\text{p}}}$ is the planet's escape velocity. Assuming circular orbits and that $v_{\text{imp}} \propto v_{\text{orb}} = 2\pi a/P$, we would expect that $M_{\text{p}}^{3/2} R_{\text{p}}^{-1/2} \propto M_{\text{s}}^{1/3} P^{-1/3}$. Again using $M_{\text{p}} \propto R_{\text{p}}^{3.7}$ for rocky planets from Zeng et al. (2016), this approximately gives $M_{\text{p}}^5 \propto (M_{\text{s}}/P)^{1/3}$ or

$$R_{\text{trans, impacts}} \propto (M_{\text{s}}/P)^{1/15}. \quad (5)$$

Again using $F_{\text{p}} \propto M_{\text{s}}^{10/3} P^{-4/3}$, this can then be re-written as

$$R_{\text{trans, impacts}} \propto F_{\text{p}}^{1/20} M_{\text{s}}^{-1/6}. \quad (6)$$

This is only an approximation and it neglects second-order terms that Schlichting et al. (2015) found were important for large impactors. None the less, this clearly predicts that the variation of the transition radius with irradiation and stellar type for planets that have been stripped by impacts should be weaker than what we found above for photo-evaporation in equations (3) and (4).

Finally, in addition to bare rocky planets with Earth-like compositions and planets with modest H/He envelopes atop Earth-like cores, it is also possible that, like Uranus and Neptune, some short-period exoplanets could accrete much of their mass from water and other volatile ices if they or their planetesimal building blocks migrated from beyond the snow-line (e.g. Rogers et al. 2011; Hansen & Murray 2012; Mordasini et al. 2012). For individual planets, it is not possible to rule out this possibility using planetary mass and radius alone (Rogers & Seager 2010). However, water-dominated envelopes should be much more resilient against photo-evaporation than solar composition envelopes, and so there should be very few

planets that are the stripped cores of former water-worlds and those should only be found on the most extreme ultra-short-period orbits (Lopez 2016).

3 SCENARIO 2: A PRIMORDIAL ROCKY POPULATION BORN AFTER DISC DISPERSAL

The second possibility is that the rocky planet population simply never had significant volatile envelopes in the first place and that they formed with their current essentially bare rocky compositions. This would make sense if these planets took $\gtrsim 10$ Myr to finish assembling, as by that point their proto-planetary gas discs will have already dissipated. As mentioned above, this is generally believed to be how the Earth and other terrestrial planets of the inner Solar System finished forming (e.g. Raymond et al. 2009; Morbidelli et al. 2012).

In this case, the rocky and non-rocky exoplanets would represent two separate populations, likely originating from different formation timescales, and the transition radius R_{trans} rocky and non-rocky planet would instead arise from the superposition of these two populations. Therefore, in this scenario, the transition radius will be set by the maximum typical mass of bare rocky planets as a function of distance from the star, which should simply be set by the available supply of solid materials that a planetary core can accrete by collisions. This is, in turn, will be set by the solid density profile in an average disc. To get an idea for the typical solid density profile from which exoplanetary systems formed, we turn to the analysis of Chiang & Laughlin (2013). By estimating the heavy element masses for the full sample of short-period planets found by *Kepler* as a function of their semi-major axes, Chiang & Laughlin (2013) were able to construct a typical minimum mass extrasolar nebula (MMEN). Specifically, they found that the *Kepler* sample implies a typical initial solid surface density profile of the form

$$\sigma_{\text{solid}} = 6.2 \times 10^2 \mathcal{F}_{\text{disc}} (a/0.2 \text{ au})^{-1.6} \text{ g cm}^{-2}. \quad (7)$$

Here, a is semi-major axis and $\mathcal{F}_{\text{disc}}$ is a normalization factor that can vary the surface density relative to the MMEN. This is quite similar in form to the standard overall surface density profile for the minimum mass solar nebula (MMSN) $\sigma(a) = \sigma_0 (a/\text{au})^{-3/2}$ (Hayashi 1981), except that the MMEN requires ~ 3 – 5 times more overall mass in solids than the MMSN (Chiang & Laughlin 2013).

Proto-planets grow their rocky cores by accreting solid material from within a feeding zone that is proportional to their Hill radius $r_{\text{H}} = a(M_{\text{p}}/3M_{\text{s}})^{1/3}$. Therefore, integrating $r_{\text{H}} \times \sigma_{\text{solid}}$ over the disc surface area should give us a good idea of the typical maximum mass for rocky planets that form by collisional growth as a function of semi-major axis. Integrating the Chiang & Laughlin (2013) profile in equation (7), this would predict that $M_{\text{p,max}} \propto a^{0.6} M_{\text{s}}^{-1/2}$ or using the Zeng et al. (2016) mass–radius relation for rocky planets that $R_{\text{trans}} \propto a^{0.16} M_{\text{s}}^{-0.14}$.

Indeed, this sort of simple calculation is backed up by the results of detailed N -body simulations. For example, Hansen & Murray (2013) performed simulated 100 planetary systems using their Monte Carlo N -body model, assuming a $\sigma \propto a^{-3/2}$ density profile. They found a mass distribution with a large amount of scatter but a general trend of increasing mass with increasing semi-major axis. Specifically, they found that the distribution of planet masses in their simulations is well fit by a Rayleigh distribution $f(m) = (m/\sigma_{\text{m}}^2) e^{-0.5(m/\sigma_{\text{m}})^2}$, with a dispersion $\sigma_{\text{m}} = 7 M_{\oplus} (a/\text{au})^{0.6}$, the same scaling relation we derived above.

At fixed stellar type then, this would predict that the transition radius for planets that are born rocky should scale as

$$R_{\text{trans, bornrocky}} \propto F_p^{-0.08}. \quad (8)$$

In terms of orbital period, this is equivalent to

$$R_{\text{trans, bornrocky}} \propto P^{0.11}, \quad (9)$$

which is in the opposite direction of the trend we predicted for a stripped rocky planet population in Section 2.

In terms of a dependence on stellar properties, naively, we would expect σ_{solid} to scale as $\propto M_{\text{disc}} Z_{\text{disc}} \propto M_s Z_s$ (e.g. Kokubo, Kominami & Ida 2006), accounting for the size of the Hill sphere this would predict that $M_{p, \text{max}} \propto M_s^{(1/2)} Z_s$ and therefore that $R_{\text{trans, bornrocky}}$ should increase with stellar mass as $\propto M_s^{0.14}$, again opposite to the predictions of photo-evaporation. However, there is currently debate as to whether there is any evidence for a correlation between planet mass and disc mass or metallicity for non-giant planets (e.g. Schlaufman & Laughlin 2011; Mann et al. 2013; Schlaufman 2015; Wang & Fischer 2015), which may be complicated by dispersion in the $Z_{\text{disc}}-Z_{\text{star}}$ relation (Liu, Zhang & Lin 2016). None the less, it seems clear that we would not expect the maximum size of rocky planets to decrease with increasing stellar mass in this scenario, unlike our predictions for a photo-evaporated population, presented in Section 2.

To illustrate our predictions for a primordial rocky planet population, we carried out our own large suite of N -body planetary growth simulations. Beginning with the solid surface density profile from equation (7) (Chiang & Laughlin 2013), we consider a disc that extends from $a_{\text{in}} = 0.03$ au to $a_{\text{out}} = 0.5$ au and randomly select $\mathcal{F}_{\text{disc}}$ between 0.1 and 1. We first generate our planetary embryos by selecting a random position near the inner edge of the disc and assuming that a planetary embryo grows from this material and that it accretes all of the mass within a feeding zone that is assumed to be 8 Hill radii wide. The next object is then assumed to form at a location such that it can also accrete all the material within its feeding zone, but without overlapping the feeding zone of the first. We then progress through the disc turning all the mass in the disc into planetary embryos, ultimately forming between about 20 and 40 proto-planetary bodies.

It is then assumed that these planetary bodies are initially all located within the radial extent of the initial disc (i.e. within $a_{\text{out}} \sim 0.5$ au), that they all have initial eccentricities of $e = 0$, and that they all lie in the same plane. We then carry out N -body integrations, using the hybrid symplectic integrator mercury6 (Chambers & Wetherill 1998) and evolve each simulation for 27 Myr, following the approach used by Dawson, Lee & Chiang (2016). The simulation time-step was set to 0.5 d, and the integrator was switched from the symplectic integrator to the Burlisch–Stoer integrator if two bodies, $M_{p,1}$ and $M_{p,2}$, have a close encounter that brings them within $1 R_H$, where R_H is the mutual Hill radius

$$R_H = \frac{a_1 + a_2}{2} \left(\frac{M_{p,1} + M_{p,2}}{3M_s} \right)^{1/3}. \quad (10)$$

If two planetary bodies collide, then we assume perfect accretion with no fragmentation. However, our general results should be insensitive to this assumption. Kokubo & Genda (2010) showed that using more realistic accretion conditions, including fragmentation and hit-and-run collisions, barely affects the mass and number of planets, or even the formation time-scale, produced by N -body simulations. In total, we carried out 200 N -body simulations and, after each had been evolved for 27 Myr, were left with 2052 planetary bodies. The results of these simulations are summarized in Fig. 4. As

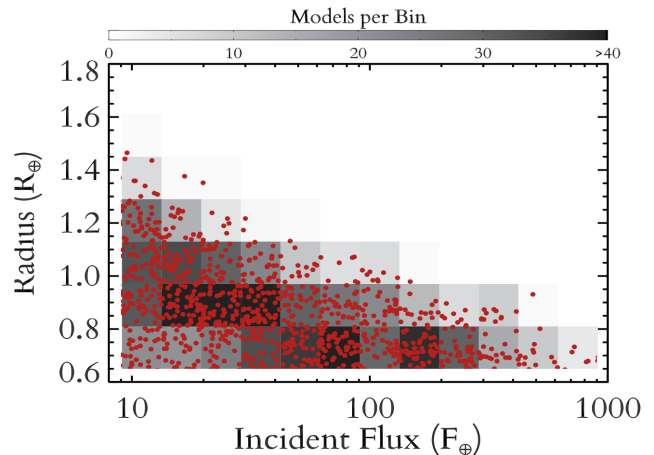


Figure 4. Similar to Fig. 3, except here we show the prediction from N -body simulations for planets that form without initial envelopes. Unlike in Fig. 3, here we find a maximum size of bare-rocky planets that *decreases* with increasing incident flux (Decreasing period).

expected, we find a maximum radius for the rocky planets formed by collisions that decreases with incident flux, consistent with our analytic predictions and the results of Hansen & Murray (2013), and in stark contrast to the results predicted by our photo-evaporation models in Section 2.

Before we move on, however, there is one more important caveat we need to address. In our photo-evaporation calculations in Section 2, we use log uniform core-mass and incident flux distributions, which were independent of each other. Yet, we have just shown in Section 3 that N -body calculations would predict that on average there should be a correlation between a planet's orbit and the typical core mass that it can reach. Our justification for this apparent contradiction is that while the MMEN found by Chiang & Laughlin (2013) represents the typical mass of an extrasolar nebula, scatter in extrasolar disc masses appears to be very large (Chiang & Laughlin 2013; Gaidos 2017). As long as the scatter solid mass in extrasolar discs is large enough to provide a large range of core masses at all orbital periods, than in our evaporation scenario, the effects of photo-evaporation should dominate over the increase in average core mass with period, and we should find a transition radius that scales as described in equations (3) and (4). This assumption is supported by the significant numbers of ultra-short-period rocky planets with radii $\approx 1.5 R_{\oplus}$ (Sanchis-Ojeda et al. 2014; Lundkvist et al. 2016). In our born rocky scenario meanwhile, the transition radius, which again is defined as the radius at which the exoplanet population switches from a majority rocky to a majority non-rocky population, would reflect the average MMEN, but with significant scatter above and below the transition. If, however, the typical scatter in disc masses is small, the correlations with orbital period or incident flux described in equations (3)/(4) and (8)/(9) would almost cancel out and we would only be left with a weak correlation that scales as $R_{\text{trans}} \propto F_p^{0.03}$ or $\propto P^{-0.04}$.

4 DISCUSSION

4.1 Observational tests

The results from Sections 2 and 3 present a clear test for exoplanet observers. A clear way to distinguish between a primordial rocky planet population and one that originated as the photo-evaporated

cores of sub-Neptunes is to obtain precise mass measurements, particularly with radial velocities, for planets near the observed $1.5 R_{\oplus}$ transition for planets receiving a wide range of orbital periods and stellar types. If most of these planets born-rocky like the Earth, then we would expect the radius of the rocky to non-rocky transition to *increase* with orbital period as $\propto P^{0.11}$, while if they are instead predominantly the evaporated cores of sub-Neptunes then would expect the rocky/non-rocky transition to *decrease* with orbital period as $\propto P^{-0.15}$.

Such an observational test should soon be feasible. In early 2018, NASA will launch the *Transiting Exoplanet Survey Satellite (TESS)* Mission, a 2-year all sky transit survey of $>200\,000$ stars with visual magnitudes of 4–13 (Ricker et al. 2014). Likewise, next year ESA plans to launch the *CHaracterising ExOPlanets Satellite (CHEOPS)* Mission (Broeg et al. 2013), which will search for transits around bright stars with known radial velocity planets. Together, these new surveys should yield a large new sample of planets around bright nearby stars that can more easily be followed up with radial velocity observations to determine planet masses and therefore constrain planet compositions.

Moreover, a characterization of the period dependence of the transition should be achievable with a reasonable sample size. For comparison, the current constraints on the rocky/non-rocky transition by Rogers (2015) used the sample of 42 transiting planets with RV follow-up described by Marcy et al. (2014). While, this sample is quite large, the majority of these targets provide little information about the transition radius. Only 14 of the planets in the Marcy et al. (2014) catalogue are between 1 and $2 R_{\oplus}$, and 6 of those have uncertainties too large to provide meaningful constraints, i.e. these 6 have only mass upper limits and those limits are not low enough to test whether those planets have rocky compositions. Therefore, it is only a small subset of approximately eight planets that are providing most of our current information on the transition. Accordingly, ~ 20 – 30 planets with radii near the current $1.5 R_{\oplus}$ transition and 2 – 3σ mass detections, or upper limits that are deep enough to rule out a rocky composition, should be sufficient to test the predictions made here.

Of particular interest will be planets that receive either very high or relatively low levels of irradiation, e.g. $\gtrsim 500 F_{\oplus}$ or $\lesssim 50 F_{\oplus}$, which for Sun-like stars corresponds to orbital periods $\lesssim 3.5$ or $\gtrsim 20$ d, respectively. Such a sample should be possible once *TESS* launches. Using the mock catalogue of simulated *TESS* detections from Sullivan et al. (2015), we estimate that *TESS* should find ~ 30 planets with radii 1.2 – $1.8 R_{\oplus}$ around stars with V-band magnitudes brighter than 10 and receiving $>500 F_{\oplus}$, and ~ 170 planets receiving $<50 F_{\oplus}$. Likewise, planets around later spectral types will also provide a valuable test of the predictions made here, and using the Sullivan et al. (2015) mock catalogue, we estimate that *TESS* should find ~ 80 planets in this size range around mid-M dwarfs with $T_{\text{eff}} < 3400$ K and V mag < 10 .

4.2 Implications for η_{\oplus}

One of the primary goals of recent transit surveys, including NASA’s *Kepler* Mission, has been to determine η_{\oplus} , the frequency of Earth-sized rocky planets in the HZs of Sun-like stars. Unfortunately, current surveys are still highly incomplete for planets $\lesssim 1.5 R_{\oplus}$ with orbital periods $\gtrsim 200$ d (e.g. Petigura, Marcy & Howard 2013b; Foreman-Mackey, Hogg & Morton 2014; Burke et al. 2015). Although *Kepler* has found several $\gtrsim 1.5 R_{\oplus}$ planets in the HZ of sun-like stars (Petigura et al. 2013b; Foreman-Mackey et al. 2014; Burke et al. 2015), these planets may be gas-rich sub-

Neptunes (Rogers 2015). So far, only one likely rocky planet *Kepler-452b* has been found that is potentially in the HZ of a Sun-like star (Jenkins et al. 2015). Meanwhile, the shorter period rocky planets, whose occurrence rate is well-constrained (e.g. Fressin et al. 2013; Petigura et al. 2013b; Foreman-Mackey et al. 2014; Burke et al. 2015), may be evaporated cores as described above in Section 2. Due to their greater detectability, many recent efforts to estimate the frequency of potentially habitable planets have focused on K and M dwarfs (e.g. Dressing & Charbonneau 2013; Kopparapu 2013; Morton & Swift 2014; Dressing & Charbonneau 2015; Gaidos et al. 2016), where *Kepler* has found many more potentially rocky planets in or near the HZ. However, many rocky planets in the HZs of late-type stars could also be the products of photo-evaporation, as their stars higher XUV activity means that they could lose initial H/He envelopes of up to ~ 1 per cent of their mass (Owen & Mohanty 2016).

All of this means that any efforts to constrain the frequency of true Earth-analogues by looking at the frequency of planets slightly larger and more irradiated than Earth (e.g. Petigura, Howard & Marcy 2013a; Silburt, Gaidos & Wu 2015), or around later stellar types, may significantly overestimate η_{\oplus} by including planets that either are not rocky or have only become rocky due to photo-evaporation. Unfortunately, this is a problem for studies like Traub (2012), which assume that the radius and period distributions can be described by separate uncorrelated power-laws, Burke et al. (2015), which use more complex separable broken power-laws, or indeed any study that assumes that planetary radii and irradiation/period distributions are uncorrelated and separable. For this reason, it is important to try and determine the origin of the irradiated rocky planet population to understand the significance of this bias. Some methods such as the Gaussian process model used by Foreman-Mackey et al. (2014) or the approximate Bayesian computation framework presented by Hsu et al. (2018) allow for the possible of complex non-monotonic correlations between the planetary radii and period distributions; however, thus far these methods have not been applied to the final DR25 *Kepler* sample (Thompson et al. 2018).

The need for updated planet occurrence rates is particularly pressing at the moment given current efforts to design future missions and plan observations to search for biomarkers on potentially Earth-like exoplanets. For example, there is currently a large effort to find Earth-sized rocky planets transiting nearby M dwarfs, so that their atmospheres can be characterized by *JWST* with transmission spectroscopy to look for oxygen and methane in their atmospheres (e.g. Deming et al. 2009; Berta, Irwin & Charbonneau 2013; Cowan et al. 2015; Barstow & Irwin 2016; Greene et al. 2016). Meanwhile, NASA is currently studying mission concepts for next generation direct imaging missions capable of observing Earth-like planets, including Large UV/Optical/IR Surveyor (LUVOIR; Croke et al. 2016) and the Habitable Exoplanet Imaging Mission (Habex; Mennesson et al. 2016), in preparation for the 2020 decadal survey. Obtaining an accurate estimate for η_{\oplus} , which accounts for complex correlations between the radii and period distributions, and an understanding of how this is affected by photo-evaporation will be critical to the success of these efforts.

5 SUMMARY

Using models of planet evolution with atmospheric photo-evaporation and Monte Carlo simulations of rocky planet growth, we have examined two possible scenarios for the origin of the highly irradiated rocky planet population recently found by NASA’s *Kepler* Mission. Specifically, we considered the possibility that this rocky

planet population is formed from the rocky cores of gas-rich planets that have had their gaseous envelopes stripped by photo-evaporation or impact erosion. We then compared this to a population that initially formed rocky and examined what each of these scenarios predicts for the period dependence of the observed transition between rocky and non-rocky planets. The key points of this study are summarized below.

(i) Recent studies have identified a new population of highly irradiated rocky planets with a transition to non-rocky volatile rich sub-Neptunes that occurs around $\sim 1.5 R_{\oplus}$.

(ii) If this rocky planet population originated as the evaporated cores of gas-rich sub-Neptunes, then this transition radius should decrease and rocky planets should become less common at longer orbital periods.

(iii) If, on the other hand, these planets formed rocky after their discs had already dissipated, then the transition radius should increase with orbital period.

(iv) With the upcoming launch of new transit missions like *TESS* and *CHEOPS*, it should be soon be possible to determine the period dependence of the transition radius with radial velocity follow-up.

(v) Of particular interest will be planets that receive either very high or relatively low levels of irradiation, e.g. $\gtrsim 500 F_{\oplus}$ or $\lesssim 50 F_{\oplus}$, which for Sun-like stars corresponds to orbital periods $\lesssim 3.5d$ or $\gtrsim 20d$, respectively.

(vi) The difference between these two scenarios has important implications for current efforts to measure η_{\oplus} , and will therefore be essential for future missions to study Earth-like planets.

(vii) In particular, these results show that it is important that studies of planet occurrence rates account for the possibility of complex non-monotonic correlations between planetary radius and period distributions and do not assume that these distributions are separable.

In reality of course, the two scenarios considered here represent two end member possibilities, and the real hot rocky exoplanet population is likely to be a mix of both that were born rocky and of stripped sub-Neptunes. However, since understanding the relative importance of these two formation channels will be critical to our efforts to search for Earth-like planets, it is important to understand the statistical predictions of these mechanisms so that we can estimate their relative contributions to the observed population.

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REFERENCES

Alibert Y., Mordasini C., Benz W., Winisdoerffer C., 2005, *A&A*, 434, 343
 Allègre C. J., Manhès G., Göpel C., 1995, *Geochim. Cosmochim. Acta*, 59, 1445
 Barstow J. K., Irwin P. G. J., 2016, *MNRAS*, 461, L92
 Batalha N. M. et al., 2011, *ApJ*, 729, 27

Berger T. A., Huber D., Gaidos E., van Saders J. L., 2018, preprint ([arXiv: 1805.00231](https://arxiv.org/abs/1805.00231))
 Berta Z. K., Irwin J., Charbonneau D., 2013, *ApJ*, 775, 91
 Bodenheimer P., Lissauer J. J., 2014, *ApJ*, 791, 103
 Bodenheimer P., Pollack J. B., 1986, *Icarus*, 67, 391
 Broeg C. et al., 2013, in Saglia R., ed., EPJ Web Conf. Vol. 47, Hot Planets and Cool Stars. p. 03005
 Burke C. J. et al., 2015, *ApJ*, 809, 8
 Zahnle K. J., Catling D. C., 2017, *ApJ*, 843, 122
 Chambers J. E., Wetherill G. W., 1998, *Icarus*, 136, 304
 Chen H., Rogers L. A., 2016, *ApJ*, 831, 180
 Chiang E., Laughlin G., 2013, *MNRAS*, 431, 3444
 Cowan N. B. et al., 2015, *PASP*, 127, 311
 Croke J. A., Roberge A., Domagal-Goldman S. D., Mandell A. M., Bolcar M. R., Rioux N. M., Perez M. R., Smith E. C., 2016, Proc. SPIE Conf. Ser. Vol. 9904, Status and Path Forward for the Large Ultraviolet/Optical/Infrared Surveyor (LUVOIR) Mission Concept Study. SPIE, Bellingham, p. 99044R
 Dawson R. I., Lee E. J., Chiang E., 2016, *ApJ*, 822, 54
 Deming D. et al., 2009, *PASP*, 121, 952
 Dressing C. D., Charbonneau D., 2013, *ApJ*, 767, 95
 Dressing C. D., Charbonneau D., 2015, *ApJ*, 807, 45
 Dressing C. D. et al., 2015, *ApJ*, 800, 135
 Elkins-Tanton L. T., Seager S., 2008, *ApJ*, 685, 1237
 Erkaev N. V., Kulikov Y. N., Lammer H., Selsis F., Langmayr D., Jaritz G. F., Biernat H. K., 2007, *A&A*, 472, 329
 Erkaev N. V. et al., 2013, *Astrobiology*, 13, 1011
 Foreman-Mackey D., Hogg D. W., Morton T. D., 2014, *ApJ*, 795, 64
 Fressin F. et al., 2013, *ApJ*, 766, 81
 Fulton B. J. et al., 2017, *AJ*, 154, 109
 Gaidos E., 2017, *MNRAS*, 470, L1
 Gaidos E., Mann A. W., Kraus A. L., Ireland M., 2016, *MNRAS*, 457, 2877
 Greene T. P., Line M. R., Montero C., Fortney J. J., Lustig-Yaeger J., Luther K., 2016, *ApJ*, 817, 17
 Haisch K. E. Jr, Lada E. A., Lada C. J., 2001, *ApJ*, 553, L153
 Hansen B. M. S., Murray N., 2012, *ApJ*, 751, 158
 Hansen B. M. S., Murray N., 2013, *ApJ*, 775, 53
 Hayashi C., 1981, *Prog. Theor. Phys. Suppl.*, 70, 35
 Hayashi C., Nakazawa K., Mizuno H., 1979, *Earth Planet. Sci. Lett.*, 43, 22
 Hsu D. C., Ford E. B., Ragozzine D., Morehead R. C., 2018, *AJ*, 155, 205
 Ikoma M., Hori Y., 2012, *ApJ*, 753, 66
 Inamdar N. K., Schlichting H. E., 2015, *MNRAS*, 448, 1751
 Inamdar N. K., Schlichting H. E., 2016, *ApJ*, 817, L13
 Jackson B., Miller N., Barnes R., Raymond S. N., Fortney J. J., Greenberg R., 2010, *MNRAS*, 407, 910
 Jackson A. P., Davis T. A., Wheatley P. J., 2011, *MNRAS*, 422, 2024
 Jenkins J. M. et al., 2015, *AJ*, 150, 56
 Jin S., Mordasini C., 2017, *ApJ*, 853, 163
 Jin S., Mordasini C., Parmentier V., van Boekel R., Henning T., Ji J., 2014, *ApJ*, 795, 65
 Johnson J. A. et al., 2017, *AJ*, 154, 108
 Johnstone C. P. et al., 2015, *ApJ*, 815, L12
 Kippenhahn R., Weigert A., 1990, Stellar Structure and Evolution, XVI, Springer-Verlag, Berlin, Heidelberg, p. 468
 Kleine T. et al., 2009, *Geochim. Cosmochim. Acta*, 73, 5150
 Kokubo E., Genda H., 2010, *ApJ*, 714, L21
 Kokubo E., Kominami J., Ida S., 2006, *ApJ*, 642, 1131
 Kopparapu R. K., 2013, *ApJ*, 767, L8
 Lammer H. et al., 2014, *MNRAS*, 439, 3225
 Lee E. J., Chiang E., 2015, *ApJ*, 811, 41
 Lee E. J., Chiang E., 2016, *ApJ*, 817, 90
 Léger A. et al., 2009, *A&A*, 506, 287
 Lissauer J. J., 1987, *Icarus*, 69, 249
 Lissauer J. J., 1993, *ARA&A*, 31, 129
 Liu B., Zhang X., Lin D. N. C., 2016, *ApJ*, 823, 162
 Liu S.-F., Hori Y., Lin D. N. C., Asphaug E., 2015, *ApJ*, 812, 164
 Lopez E. D., 2016, *MNRAS*, 472, 245
 Lopez E. D., Fortney J. J., 2013, *ApJ*, 776, 2

- Lopez E. D., Fortney J. J., 2014, *ApJ*, 792, 1
- Lopez E. D., Fortney J. J., Miller N., 2012, *ApJ*, 761, 59
- Lundkvist M. S. et al., 2016, *Nature Communications*, 7, 11201
- Mann A. W., Gaidos E., Kraus A., Hilton E. J., 2013, *ApJ*, 770, 43
- Marcy G. W. et al., 2014, *ApJS*, 210, 20
- Mennesson B. et al., 2016, Proc. SPIE Conf. Ser. Vol. 9904, SPIE, Bellingham, p. 99040L
- Morbidelli A., Lunine J. I., O'Brien D. P., Raymond S. N., Walsh K. J., 2012, *Annu. Rev. Earth Planet. Sci.*, 40, 251
- Mordasini C., Alibert Y., Klahr H., Henning T., 2012, *A&A*, 547, A111
- Mordasini C., Klahr H., Alibert Y., Miller N., Henning T., 2014, *A&A*, 566, A141
- Morton T. D., Swift J., 2014, *ApJ*, 791, 10
- Murray-Clay R. A., Chiang E. I., Murray N., 2009, *ApJ*, 693, 23
- Nemchin A., Timms N., Pidgeon R., Geisler T., Reddy S., Meyer C., 2009, *Nature Geoscience*, 2, 133
- Owen J. E., Jackson A. P., 2012, *MNRAS*, 425, 2931
- Owen J. E., Mohanty S., 2016, *MNRAS*, 459, 4088
- Owen J. E., Wu Y., 2013, *ApJ*, 775, 105
- Owen J. E., Wu Y., 2017, *ApJ*, 847, 29
- Pepe F. et al., 2013, *Nature*, 503, 377
- Petigura E. A., Howard A. W., Marcy G. W., 2013a, *Proc. Natl. Acad. Sci.*, 110, 19273
- Petigura E. A., Marcy G. W., Howard A. W., 2013b, *ApJ*, 770, 69
- Petigura E. A. et al., 2017, *AJ*, 154, 107
- Pollack J. B., Hubickyj O., Bodenheimer P., Lissauer J. J., Podolak M., Greenzweig Y., 1996, *Icarus*, 124, 62
- Rafikov R. R., 2006, *ApJ*, 648, 666
- Raymond S. N., O'Brien D. P., Morbidelli A., Kaib N. A., 2009, *Icarus*, 203, 644
- Ribas I., Guinan E. F., Güdel M., Audard M., 2005, *ApJ*, 622, 680
- Ricker G. R. et al., 2014, Proc. SPIE Conf. Ser. Vol. 9143, Space Telescopes and Instrumentation 2014: Optical, Infrared, and Millimeter Wave. SPIE, Bellingham, p. 914320
- Rogers L. A., 2015, *ApJ*, 801, 41
- Rogers L. A., Seager S., 2010, *ApJ*, 712, 974
- Rogers L. A., Bodenheimer P., Lissauer J. J., Seager S., 2011, *ApJ*, 738, 59
- Sanchis-Ojeda R., Rappaport S., Winn J. N., Kotson M. C., Levine A., El Mellah I., 2014, *ApJ*, 787, 47
- Schaefer L., Fegley B., 2010, *Icarus*, 208, 438
- Schlaufman K. C., 2015, *ApJ*, 799, L26
- Schlaufman K. C., Laughlin G., 2011, *ApJ*, 738, 177
- Schlichting H. E., Sari R., Yalinewich A., 2015, *Icarus*, 247, 81
- Sekiya M., Nakazawa K., Hayashi C., 1980, *Prog. Theor. Phys.*, 64, 1968
- Shkolnik E. L., Barman T. S., 2014, *AJ*, 148, 64
- Silburt A., Gaidos E., Wu Y., 2015, *ApJ*, 799, 180
- Silva Aguirre V. et al., 2015, *MNRAS*, 452, 2127
- Sullivan P. W. et al., 2015, *ApJ*, 809, 77
- Tera F., Papanastassiou D. A., Wasserburg G. J., 1974, *Earth Planet. Sci. Lett.*, 22, 1
- Thompson S. E. et al., 2018, *ApJS*, 235, 38
- Tian F., Kasting J. F., Liu H.-L., Roble R. G., 2008, *J. Geophys. Res.*, 113, E05008
- Touboul M., Kleine T., Bourdon B., Palme H., Wieler R., 2007, *Nature*, 450, 1206
- Traub W. A., 2012, *ApJ*, 745, 20
- Tu L., Johnstone C. P., Güdel M., Lammer H., 2015, *A&A*, 577, L3
- Valencia D., Ikoma M., Guillot T., Nettelmann N., 2010, *A&A*, 516, A20
- Van Eylen V. et al., 2018, *MNRAS*, 478, 4866
- Wang J., Fischer D. A., 2015, *AJ*, 149, 14
- Williams J. P., Cieza L. A., 2011, *ARA&A*, 49, 67
- Zeng L., Sasselov D. D., Jacobsen S. B., 2016, *ApJ*, 819, 127

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