

Helium enhanced stars and multiple populations along the horizontal branch of NGC 2808: direct spectroscopic measurements^{*}

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ABSTRACT

We present an abundance analysis of 96 horizontal branch (HB) stars in NGC 2808, a globular cluster exhibiting a complex multiple stellar population pattern. These stars are distributed in different portions of the HB and cover a wide range of temperature. By studying the chemical abundances of this sample, we explore the connection between HB morphology and the chemical enrichment history of multiple stellar populations. For stars lying on the red HB, we use GIRAFFE and UVES spectra to determine Na, Mg, Si, Ca, Sc, Ti, Cr, Mn, Fe, Ni, Zn, Y, Ba and Nd abundances. For colder, blue HB stars, we derive abundances for Na, primarily from GIRAFFE spectra. We were also able to measure direct non-local thermodynamic equilibrium He abundances for a subset of these blue HB stars with temperature higher than ~ 9000 K. Our results show that: (i) HB stars in NGC 2808 show different content in Na depending on their position in the colour–magnitude diagram, with blue HB stars having higher Na than red HB stars; (ii) the red HB is not consistent with a uniform chemical abundance, with slightly warmer stars exhibiting a statistically significant higher Na content; and (iii) our subsample of blue HB stars with He abundances shows evidence of enhancement with respect to the predicted primordial He content by $\Delta Y = +0.09 \pm 0.01 \pm 0.05$ (internal plus systematic uncertainty). Our results strongly support theoretical models that predict He enhancement among second-generation(s) stars in globular clusters and provide observational constraints on the second-parameter governing HB morphology.

Key words: techniques: spectroscopic – stars: abundances – stars: Population II – globular clusters: general – globular clusters: individual: NGC 2808.

1 INTRODUCTION

Globular clusters (GCs) exhibit a range of horizontal branch (HB) morphology, with some clusters showing red, blue, extended blue, or even multimodal HBs. While variations in metallicity are proposed

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as the first parameter governing the HB, these alone cannot explain the wide range of observed structure. Many GC properties, such as age, mass-loss and He content, have been proposed to be the second parameter, but none has been fully successful (see Catelan 2009, for a review).

Observations of multiple stellar populations in GCs allow us to look at the second-parameter problem from a new perspective. Recent studies showed that the colour–magnitude diagram (CMD) of GCs, from the main sequence (MS) up to the red giant branch (RGB), if analysed in appropriate filters, are composed of different sequences, which correspond to different stellar populations with a range of light element and (possibly) He abundances. Furthermore, the multiple HB components revealed in the CMD of several GCs have been tentatively assigned to different stellar generations by many authors (e.g. D’Antona et al. 2005; Busso et al. 2007; Piotto et al. 2007; Milone et al. 2008; Milone 2013).

The association of chemical inhomogeneity and HB structure was originally advanced by Norris (1981) and Smith & Norris (1983) in the context of CN variations, and by Catelan & de Freitas Pacheco (1995) in the context of extramixed O-poor stars. He-enhanced stars are expected to populate the blue extreme of the HB (e.g. D’Antona et al. 2002), and to be O depleted and Na enhanced, as observed in the surface abundances of second-generation(s) GC stars. In some cases, the connection between chemistry and HB morphology is supported by the agreement between the fraction of stars on individual GC CMD sequences, associated with different chemical content, and the fraction in different HB segments (e.g. D’Antona et al. 2005; Piotto et al. 2007; Milone et al. 2008, 2012d).

A clear and direct confirmation of the chemistry/HB-morphology connection in GCs was presented in recent work on M 4 (Marino et al. 2011b). Spectroscopic studies of the RGB revealed that two groups of stars are present in this cluster, one Na/CN-band strong (O-poor) and the other Na/CN-band weak (O-rich). These stellar populations define two different RGB sequences in the B versus $(U - B)$ CMD (see Marino et al. 2008). Analogous to the RGB stellar distribution, the HB is also bimodal, i.e. is well populated on both sides of the instability strip. Stars located on the blue side of the RR-Lyrae gap are Na rich, and stars on the red side are Na poor (see fig. 1 a from Marino et al. 2011b). Successively, the connection between stellar populations and HB morphology in GCs has been confirmed for NGC 2808 (Gratton et al. 2011), NGC 1851 (Gratton et al. 2012), 47 Tucanae (Milone et al. 2012d; Gratton et al. 2013), NGC 6397 (Lovisi et al. 2012), M 5 (Gratton et al. 2013) and M 22 (Marino, Milone & Lind 2013).

The interplay of He-enhanced stars in GCs seems to be the only viable way to explain multiple MSs in GCs. In some cases, like ω Centauri and NGC 2808, the He enrichment in second stellar generations is expected to be large (up to $Y \sim 0.40$; Bedin et al. 2004; Norris 2004; Piotto et al. 2005, 2007; King et al. 2012), in other cases, like NGC 6752 and NGC 6397, the minimal separation in colour between the sequences is consistent with much smaller He enhancements (Milone et al. 2010, 2012c, 2013).

In this observational and theoretical framework, strong spectroscopic evidence of He-enhanced stars in a GC will both corroborate the proposed generational scenario to explain multiple MSs, and shed light on the second-parameter problem. Unfortunately, while abundance variations in p -capture elements, such as O and Na, have been widely observed, direct spectroscopic evidence of He enhancement in GCs is scarce, as these measurements are very difficult. Reliable He abundances can be measured for only a small fraction of HB stars with surface temperatures between ~ 8000 and $\sim 11\,500$ K. Hotter stars experience He settling, which results in

abundances in the stellar atmosphere that are not representative of the initial He content (e.g. Behr 2003; Moehler et al. 2004; Fabian et al. 2005). Cool stars do not typically have strong enough He lines in the optical. When they do, these He lines are chromospheric rather than photospheric; to determine a reliable abundance from these features requires complex models that take into account the chromospheric activity (Dupree, Strader & Smith 2011).

Pasquini et al. (2011) and Dupree et al. (2011) used near-infrared He lines in RGB stars of NGC 2808 and ω Centauri, respectively, and found deeper and/or detectable lines for Na-rich stars. These results strongly suggest a higher He abundance for the Na-rich stars hosted in these two clusters. Pasquini et al. (2011) suggested a difference of $\Delta(Y) \gtrsim 0.17$ in He for the two RGBs in NGC 2808, although they cannot provide the absolute abundances. A similar difference has also been estimated for two RGBs in ω Centauri, with absolute He abundances inferred from one near-infrared line analysed using appropriate chromospheric models (Dupree & Avrett 2013).

Villanova, Piotto & Gratton (2009) presented the first direct He abundance determinations from spectral lines not significantly affected by chromospheric activity for four blue horizontal branch (BHB) stars in NGC 6752 and six in M 4. For three of the four stars in NGC 6752, Na, O, and He were compatible with first-generation GC stars (primordial Y and Na-poor/O-rich). For M 4, the He content appeared enhanced with respect to the primordial value, by $\Delta(Y) \sim 0.04$. Note however that these two clusters are not ideal to detect He enhancements. NGC 6752 shows only a BHB (HB ratio, $HBR = 1.0$, Harris 1996, updated as in 2010). This suggests that the different stellar populations observed on the RGB and MS (including the He-primordial first generation) are expected to distribute along the BHB. In fact, first-generation stars have been found to populate the coldest segment of the BHB of NGC 6752 (Villanova et al. 2009). Furthermore, according to the MS splits observed by Milone et al. (2013), second-generation stars in NGC 6752 and M 4 will be enhanced in helium by only small amounts, which would be very difficult to detect spectroscopically.

The GC NGC 2808 represents the ideal target to detect the presence of HB He enhanced stars in GCs. While its uniformity in metallicity suggests that this GC could be classified among *normal*¹ GCs, i.e. those showing multiple stellar populations by means of variations in the chemical content of light elements (e.g. C, N, O, Na, Mg and Al) alone, photometric studies show that it is more complex. *Hubble Space Telescope* (HST) photometry has discovered a spectacular triple MS, unambiguously revealing that NGC 2808 hosts at least three different stellar generations (Piotto et al. 2007; Milone et al. 2012a,e). The RGB also exhibits evidence of non-singular populations, with large colour spreads (Yong et al. 2008; Lee et al. 2009; Monelli et al. 2013). NGC 2808 exhibits complex HB morphology, with multiple components and stars distributed along an extended blue tail (Sosin et al. 1997; Bedin et al. 2000, hereafter B00; Iannicola et al. 2009). Helium enhanced stars are expected to lie on the blue side of the instability strip, including the temperature range suitable for He measurements from spectral lines, and theoretical models require a large He enhancement, up to $Y \sim 0.38$,

¹ Based on Marino et al. (2012a), we consider GCs exhibiting internal variations in iron and heavy elements to be anomalous. Examples of these GCs are M22 (Marino et al. 2009, 2011a), NGC 1851 (Yong & Grundahl 2008; Carretta et al. 2011) and the more extreme ω Centauri (e.g. Norris & Da Costa 1995; Pasquini et al. 2002; Johnson & Pilachowski 2010; Marino et al. 2011c, 2012b).

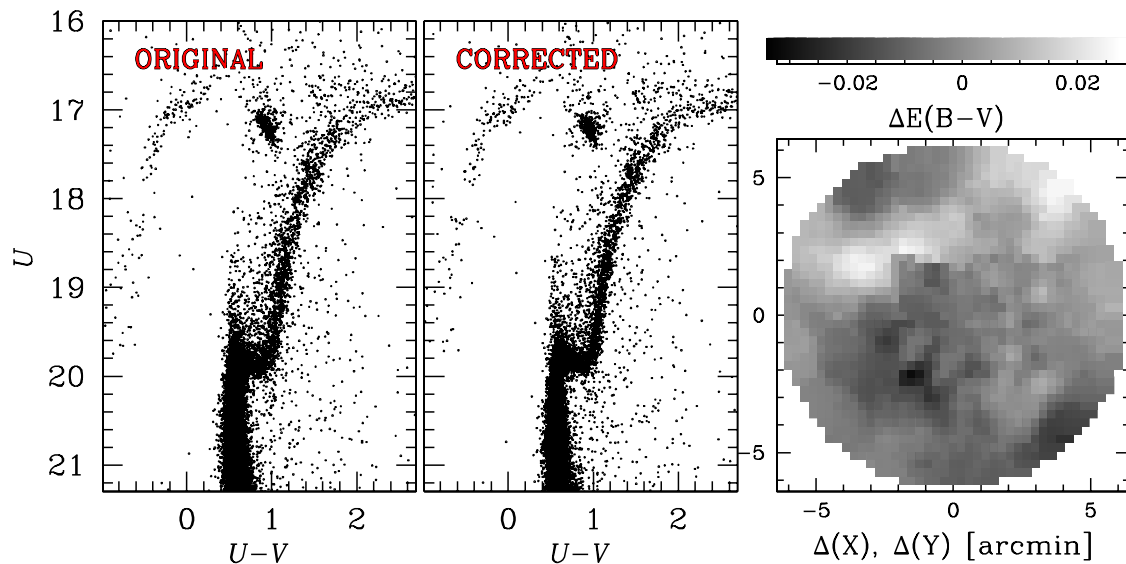


Figure 1. Left-hand panels: comparison between the original U versus $(U - V)$ CMD of NGC 2808 (left hand) and the CMD corrected for differential reddening (right hand). Right-hand panel: map of the differential reddening and zero-points effects in the $11.25 \text{ arcmin} \times 11.25 \text{ arcmin}$ field of view centred on NGC 2808. The grey-scale reflects the magnitude of the variation in local reddening.

to reproduce the CMD of this cluster (D’Antona et al. 2005; Piotto et al. 2007; Milone et al. 2012a). On the spectroscopic side, NGC 2808 shows an extended Na–O anticorrelation (Carretta et al. 2006). In addition, Bragaglia et al. (2010) analysed one red and one blue star on the MS, finding that the blue was enhanced in N and Al, while being depleted in C. This chemical information suggests that the blue MS stars in NGC 2808 may be helium enhanced.

In this paper, we analyse chemical abundances, including He, along the HB of NGC 2808. The organization of this paper is as follows: in Section 2, we describe our sample and observations; Section 3 explains our atmospheric parameter determination; Section 4 summarizes our derived chemical abundances, with detailed results for Na and He described in Sections 5 and 6, respectively. We then discuss our results and repercussions for the second-parameter problem, in Section 7. A summary of our results is presented in Section 8.

2 DATA

2.1 The photometric catalogues

To identify our stellar sample, we use the photometric catalogue of Momany et al. (2004), which has been obtained from U , B and V images collected with the Wide-Field Imager (WFI) mounted at the 2.2m ESO-MPI (Max-Planck-Institut) telescope at La Silla observatory, Chile. This catalogue has been used in previous spectroscopic studies of the HB of NGC 2808 (Pace et al. 2006; Gratton et al. 2011). To estimate systematic errors related to the adopted photometric catalogue, we also use photometry from B00, which was obtained from U , B and V images taken with the Danish Faint Object Spectrograph and Camera mounted at the 1.54 ESO-Danish telescope at La Silla.

The average reddening of NGC 2808 is $E(B - V) = 0.19$ (B00) and is not uniform across the field of view analysed in this paper. Significant star-to-star reddening variations have been observed in even a small $\sim 3 \text{ arcmin} \times 3 \text{ arcmin}$ field around the cluster centre (Milone et al. 2012a,e, Section 2).

A visual inspection of the CMD shown in the left-hand panel of Fig. 1 reveals that the sequences, from MS to RGB, are spread

by the interplay of several effects such as differential reddening, zero-point variations and sky concentration.² To correct for these effects, we used the procedure described by Milone et al. (2012b). Briefly, we first draw an MS ridge line by putting a spline through the median colours found in successive short intervals of magnitude, and iterating with sigma clipping. For each star, we then estimate how the observed stars in its vicinity may systematically lie to the red or the blue of the fiducial sequence; this systematic colour and magnitude offset, measured along the reddening line, is indicative of the local differential reddening. This procedure treats the CMD as a single sequence. In the case of CMDs characterized by multiple sequences, like in NGC 2808, this approximation results in larger internal uncertainties, but it does not systematically affect the correction (see Milone et al. 2012b). The final CMD, corrected for differential reddening and zero-points is shown in the middle panel of Fig. 1. Our corrections significantly decrease the internal photometric uncertainty, as shown by the comparison between the original and corrected CMDs of Fig. 1. To give an idea of the size of our corrections, we show the spatial reddening variations in the field of NGC 2808 in the right-hand panel of Fig. 1. We divide the whole field of view into 64×64 boxes of $11.25 \text{ arcsec} \times 11.25 \text{ arcsec}$ and calculate the average $\Delta E(B - V)$ within each of them. In the resulting reddening map, the grey-scale reflects the variation over the field. The scale is reported at the top of the right-hand panel.

2.2 The spectroscopic data set

Our spectroscopic data consist of FLAMES/GIRAFFE and FLAMES/UVES data collected under the ESO programme 086.D-0141 (PI: Marino). The GIRAFFE fibres were used with the HR12 setup, covering the spectral range from ~ 5820 to $\sim 6140 \text{ \AA}$ with a resolution of $\sim 18\,700$. The higher resolution UVES spectrograph was used in the RED 580 configuration, providing spectra centred at $\lambda 5800 \text{ \AA}$ (with a spectral coverage of $\sim 2000 \text{ \AA}$) at a resolution of

² Sky concentration is a spurious re-distribution of light in the focal plane due to the reflections of light at discontinuities in the optics (Manfroid, Selman & Jones 2001). The WFI@2.2m camera is significantly affected by this and by zero-point variations, as discussed in Bellini et al. (2009).

Table 1. Basic data for the analysed UVES and GIRAFFE stars. Identification numbers and photometric data are from Momany et al. (2004).

ID	$\alpha(2000)$	$\delta(2000)$	V	$(B - V)_{\text{ori}}$	$(B - V)_{\text{corr}}$	RV (km s ⁻¹)	ID(2MASS)	Type
UVES								
21854	09:12:04.408	-64:50:47.662	16.256	0.829	0.843	107.1	–	RHB
18106	09:12:02.698	-64:52:31.346	16.249	0.815	0.814	109.4	–	RHB
20577	09:11:55.598	-64:51:21.500	16.100	0.781	0.789	102.4	–	RHB
716	09:12:03.965	-64:49:26.226	16.705	0.317	0.298	106.0	–	BHB
19229	09:12:23.239	-64:51:59.650	16.290	0.212	0.232	110.5	09122323-6451595	BHB
509	09:12:18.219	-64:49:37.130	16.768	0.189	0.180	116.8	09121814-6449371	BHB
11715	09:11:58.235	-64:57:13.035	16.917	0.132	0.121	105.5	09115825-6457129	BHB
GIRAFFE								
2771	09:11:54.180	-64:47:10.765	16.236	0.844	0.828	95.3	–	RHB
2648	09:11:44.296	-64:47:23.509	16.296	0.845	0.827	100.5	09114429-6447236	RHB
4036	09:11:39.084	-64:49:51.110	16.290	0.925	0.869	112.8	09113906-6449509	RHB
4491	09:11:40.148	-64:49:04.720	16.356	0.808	0.769	111.9	09114013-6449046	RHB
3093	09:12:09.296	-64:46:32.395	16.289	0.845	0.835	94.4	09120929-6446324	RHB
1317	09:12:02.282	-64:48:55.849	16.334	0.872	0.849	107.8	09120228-6448558	RHB

~45 000. All our target stars were observed on the same FLAMES plate for a total exposure time of 20 h. At the wavelength of the Na D lines, the typical signal-to-noise ratio (S/N) of the fully reduced and combined GIRAFFE spectra is ~100–120 and ~40–80 for red and BHB stars, respectively. UVES combined spectra have much lower S/N, ranging from ~35 (red horizontal branch, RHB) to ~15 (BHB) at Na D lines. Data reduction involving bias subtraction, flat-field correction, wavelength calibration and sky subtraction is done using the dedicated pipelines for GIRAFFE³ and UVES (Ballester et al. 2000).

In total we observed 108 candidate cluster members on the HB of NGC 2808, covering the observed colour range $0.05 < (B - V) < 0.90$. This sample spans the HB from the reddest stars, corresponding to temperatures ~5300–5600 K, to BHB stars, with temperatures ~9000–12 000 K.

The HB of NGC 2808 is contaminated by field stars. We determine cluster membership from radial velocities (RVs) obtained using the IRAF@FXCOR task, which cross-correlates the object spectrum with a template. For the template, we used a synthetic spectrum obtained through the spectral synthesis code SPECTRUM⁴ (Gray & Corbally 1994). After applying a heliocentric correction, we find a mean RV of $+104.5 \pm 0.8$ km s⁻¹ ($\sigma = 8.2$ km s⁻¹) for the whole sample. This value agrees within 3σ with the values in the literature (e.g. $+101.6 \pm 0.7$ km s⁻¹, $\sigma = 13.4$ km s⁻¹; Harris 1996, updated as in 2010). We then reject individual stars with values deviating by more than 3σ from this average velocity, deeming them to be probable field stars.

Our sample of bona fide cluster stars is composed of 96 HBs: 65 distributed on the red side of the instability strip and 31 on the blue side. All the UVES targets have RVs compatible with being cluster members, providing a sample of 3 RHB and 4 BHB of NGC 2808 observed at higher resolution and with larger spectral coverage. Basic photometric data plus RVs for the observed stars are listed in Table 1. Each star is flagged as red HB (RHB), blue HB (BHB) or field (F), depending on its position along the HB and its membership status. A few BHB stars have been flagged as Grundahl jump (GJ; see Grundahl et al. 1999). These GJ stars suffer from levitation of

metals and sedimentation of He; consequently, their atmospheric chemical contents are not representative of the cluster abundances (see Section 6).

3 ATMOSPHERIC PARAMETERS

Given the large differences between red and BHB spectra, they require different approaches to determine their atmospheric parameters. In the following Section 3.1, we describe the procedure for BHB stars. We discuss our analysis of RHB stars in Section 3.2.

3.1 Blue horizontal branch

Colour–temperature calibrations for BHB stars at a range of gravities are provided by the Castelli website.⁵ These relationships require precise photometry. The typical photometric error of our data is ~0.02 mag. At 10 000 K, this translates to a temperature uncertainty of ~500 K. For the hottest BHB stars, with temperatures up to ~12 000 K, this uncertainty is as high as 1000 K. Clearly, correcting for differential reddening and zero-point effects is critical for this cluster. Therefore, instead of assigning each star the temperature corresponding to its $(B - V)$ colour in the colour–temperature calibration, we project the target BHB stars on the Zero-Age HB (ZAHB) that best fits the CMD, thus minimizing the impact of the photometric uncertainties. We use the CMD in V versus $(U - V)$ instead of V versus $(B - V)$; with $(B - V)$ the BHB is almost vertical, and a small colour error translates into a large uncertainty in atmospheric parameters.

We estimate effective temperature (T_{eff}) and gravity ($\log g$) for the BHB stars by comparing the observed CMD with theoretical models. To this aim, we considered isochrones and ZAHB loci from BaSTI⁶ (Pietrinferni et al. 2004, 2006). For comparison purposes, we also used isochrones from PGPUC⁷ data base (Valcarce, Catelan & Sweigart 2012).

In Fig. 2 (bottom panel), we show the V versus $(U - V)$ Hess diagram for NGC 2808 with a set of best-fitting isochrones, assuming a distance modulus $(m - M)_V = 15.67$, reddening $E(B - V) = 0.19$, $[\text{Fe}/\text{H}] = -1.15$, $[\alpha/\text{Fe}] = +0.3$ and age 10.0 Gyr. These values

³ See <http://girbldr.sourceforge.net>

⁴ See <http://www.appstate.edu/grayro/spectrum/spectrum.html> for more details.

⁵ <http://wwwuser.oat.ts.astro.it/castelli/>

⁶ <http://albione.ia-teramo.inaf.it>

⁷ <http://www2.astro.puc.cl/pgpuc/iso.php>

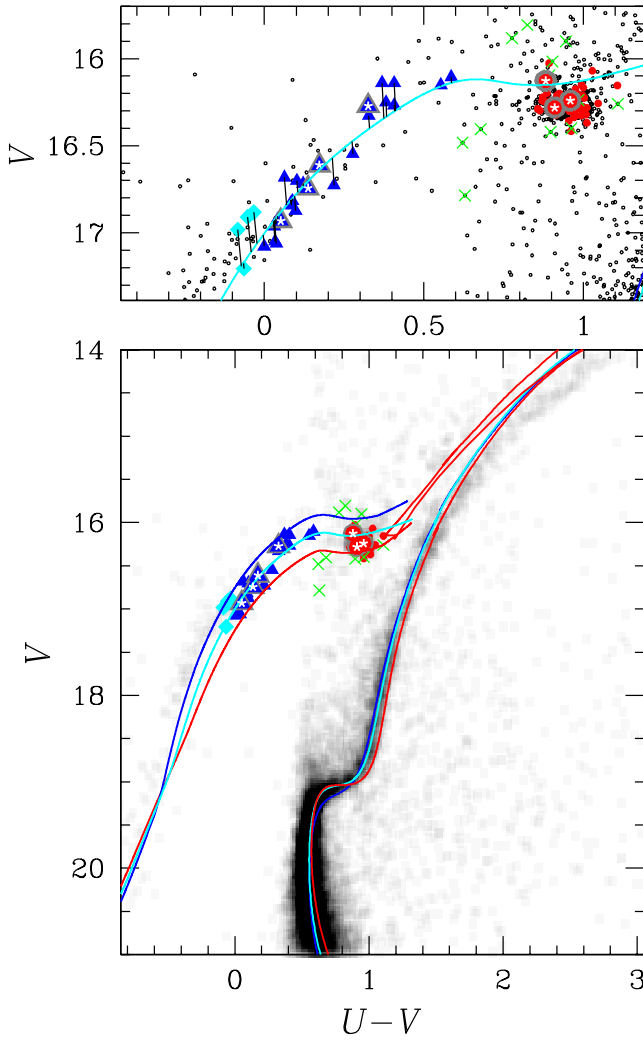


Figure 2. Lower panel: V versus $(U - V)$ Hess diagram for NGC 2808. Spectroscopic targets are represented with different symbols: red circles are RHB, blue triangles represent the BHB, cyan diamonds GJ stars and green crosses probable field stars. The seven UVES targets are identified with grey open symbols with a white star. The superimposed theoretical models with $Y = 0.246$ (red), $Y = 0.320$ (cyan) and $Y = 0.380$ (blue) have been extracted from the BaSTI α -enhanced isochrones with $[\text{Fe}/\text{H}] = -1.15$. Upper panel: zoom of the V versus $(U - V)$ CMD in the region of the HB. Black vertical lines show the projection of the BHB targets on the ZAHB locus with $Y = 0.320$ (see the text for details).

are all in agreement with estimates available in the literature (Harris 1996; B00; Marín-Franch et al. 2009; Dotter et al. 2010). The red, cyan and blue isochrones correspond to three different choices of helium, namely $Y = 0.25$, 0.32 and 0.38 . The positions of our spectroscopic targets on this diagram are represented with different colours and symbols, according to their position along the HB.

The region of the CMD populated by our BHB targets is best fitted by assuming $Y = 0.32$ (cyan model in Fig. 2), and, as such, we use a ZAHB corresponding to this helium abundance to derive the atmospheric parameters for our BHB sample. We estimate effective temperatures by projecting the observed position of BHB stars in the $V - (U - V)$ diagram on the ZAHB locus with $Y = 0.32$ (as schematically shown in Fig. 2). The projection is computed as in Gallart et al. (2003, see their section 4), i.e. by enhancing the difference in colour by a factor of 7, which is determined empirically,

with respect to the difference in magnitude. For a given isochrone, a star's colour is better constrained than the magnitude, as any uncertainty in distance, gravity, or reddening corresponds to a greater difference in magnitude than in colour. We obtain surface gravities from the observed V magnitude and temperatures from theoretical models. We then use each star's position on the ZAHB locus to determine the corresponding stellar mass.

An inspection of the CMD suggests that most of the BHB targets are spread around the ZAHB with $Y = 0.32$. The broadening of the HB is due to both measurement uncertainties (like photometric errors and residual differential reddening) and intrinsic spread, perhaps from star-to-star helium variations and/or the presence of stars leaving the ZAHB. Our best-fitting model used to derive the atmospheric parameters consists of only the ZAHB, while the real BHB is also populated by stars on their evolutionary tracks, with gravities lower than predicted from the ZAHB.

In Fig. 3, we investigate the impact of internal photometric errors on our temperature and gravity determinations. We have compared the ZAHB loci from the PGPUC data base (red lines) in the absolute $T_{\text{eff}} - (U - V)$ and $\log g - (U - V)$ planes for two choices of Y , namely $Y = 0.245$ and 0.320 . Internal photometric errors in $(U - V)$ of ~ 0.04 mag affect the temperature determinations from isochrones by ~ 250 K. The impact of the photometric errors on $\log g$ is at most ~ 0.05 dex.

Microturbulent velocities (ξ_t) come from the $T_{\text{eff}} - \xi_t$ empirical relation derived in Pace et al. (2006) from the least-squares fit of the HB data analysed by Behr, Cohen & McCarthy (2000) and Behr (2003). For the warmer stars, which have a very stable atmosphere where helium sinks and heavy metals levitate, we adopted $\xi_t = 0$, as suggested by Behr (2003) in NGC 288. The dispersion of ξ_t values around the relation found by Behr is 1 km s^{-1} .

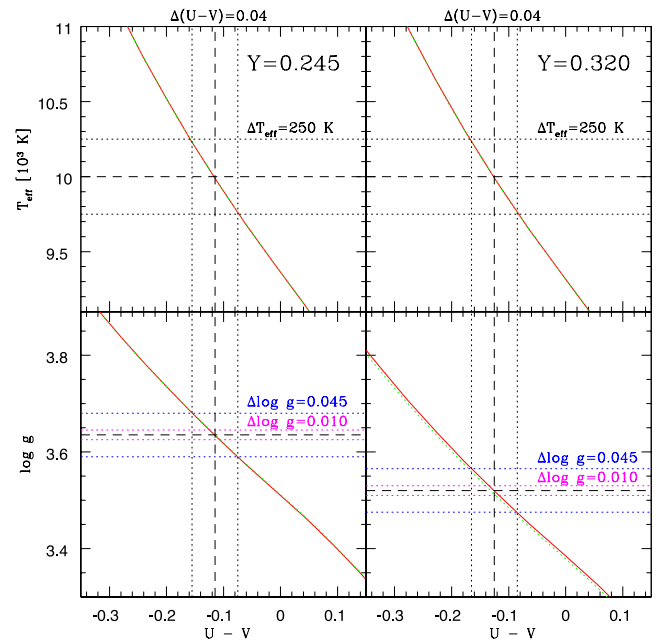


Figure 3. ZAHB loci extracted from the PGPUC data base (red lines) in the absolute T_{eff} versus $(U - V)$ plane (upper panels) and in the absolute $\log g$ versus $(U - V)$ plane (bottom panels) for $Y = 0.245$ (left-hand panels) and $Y = 0.320$ (right-hand panels). Dashed lines correspond to standard BHB stars at $T_{\text{eff}} \sim 10\,000$ K and $\log g = 3.64$ dex. Dotted lines correspond to an uncertainty in $(U - V)$ of ~ 0.04 mag. The green dotted line is the ZAHB locus with the same properties as red lines, except with a progenitor mass of $0.7 M_{\odot}$.

From the discussion above, we consider ± 250 K, ± 0.05 dex and ± 1 km s⁻¹ as estimates of internal errors associated with our adopted T_{eff} , $\log g$ and ξ_t values for the BHB sample. To accurately determine absolute abundances for the BHB stars, we must also consider possible systematic uncertainties affecting our atmospheric parameter determinations.

Fig. 3 displays that a different choice of Y does not affect the temperature determinations. From the theoretical point of view, models predict that HB stars near the ZAHB locus with the same $(U - V)$ colour but different helium have the same effective temperature. As such, the choice of Y is not critical for the determination of T_{eff} from isochrones.

In contrast, canonical models with He enhancement predict $\log g$ lower than that from He normal models. If stars on the BHB of NGC 2808 are not He enhanced, our gravity estimates will be underestimated by ~ 0.15 dex. In addition, stars leaving the ZAHB track appear more luminous and have lower gravities, mimicking a ZAHB loci with higher He (of at most $Y = 0.38$). When the ZAHB loci have Y higher by ~ 0.08 , this changes the estimated $\log g$ values by ~ 0.12 dex (see also Moni Bidin et al. 2007). Age does not affect ZAHB models over the range of temperatures studied here, as suggested by the comparison between ZAHBs with a difference in age of several Gyr (red and green dotted lines in Fig. 3).

Although our data do not allow independent measurements of atmospheric parameters to be compared with the adopted ones, we note the following: (i) on the theoretical side, we do not expect significant systematics in T_{eff} due to the treatment of Y , but we do expect systematic uncertainties in $\log g$ by at most ~ 0.15 dex; (ii) the best fitted ZAHB with the BHB has a temperature of $\sim 11\,500$ K at the so-called GJ (Grundahl et al. 1999), which matches the expected value for stars in this region of the HB within ± 500 K; (iii) comparison of the T_{eff} and $\log g$ values for RHB stars from the technique described in Section 3.2 and the methodology used for BHB stars shows agreement within 170 K in T_{eff} and ~ 0.15 dex in $\log g$ with the BHB scale returning higher T_{eff} and $\log g$ values. Based on these arguments, we conclude that our T_{eff} and $\log g$ values do not seem to be affected by systematic effects larger than ~ 200 K and ~ 0.20 dex, respectively. The impact of these possible systematics on our results is fully discussed in Section 6.

As a final test, we verify that our choice of stellar models does not significantly change our adopted stellar parameters. We compare our adopted T_{eff} and $\log g$ values derived from the BaSTI data base with those derived from the PGPUC data bases. The two sets of atmospheric parameters show a satisfactory agreement, with mean differences of $(T_{\text{eff BaSTI}} - T_{\text{eff PGPUC}}) = -23 \pm 11$ K (rms = 67 K) and $(\log g_{\text{BaSTI}} - \log g_{\text{PGPUC}}) = +0.15 \pm 0.03$ dex (rms = 0.21).

3.2 Red horizontal branch

The GIRAFFE HR12 setup covers a relatively small spectral range and the number of isolated Fe lines (~ 10 Fe I and 2 Fe II) is not enough for accurate spectroscopic determination of the atmospheric parameters. To derive the atmospheric parameters of RHB targets observed with GIRAFFE, we use the Momany et al. (2004) photometry, corrected for differential reddening, in conjunction with $(B - V) - T_{\text{eff}}$ relations⁸ (Alonso, Arribas & Martínez-Roger 1999,

Table 2. Atmospheric parameters from Fe lines, and T_{eff} from photometry for the RHB stars observed with UVES.

ID	T_{eff}	$\log g$ Fe lines	ξ_t	T_{eff} (Alonso photometry)
20577	5390	2.26	1.66	5501
18106	5480	2.55	1.48	5422
21854	5350	2.20	1.73	5336

2001). The use of these relations introduces some internal and systematic uncertainties in T_{eff} due to photometric errors and reddening effects.

NGC 2808 has an absolute reddening estimated in the literature from $E(B - V) = 0.13$ (Castellani et al. 2006) to $E(B - V) = 0.23$ (Piotto et al. 2002). The use of these two different absolute reddening values corresponds to a shift in the derived T_{eff} of ~ 300 K. To establish the reddening to be used in the colour- T_{eff} relation, and hence the T_{eff} scale, we first determined atmospheric parameters for three RHB stars observed with UVES, independent of photometry. As previously mentioned, these spectra have higher resolution and a larger spectral coverage. In the typical stellar parameter space of the RHB sample, we measure ~ 25 – 30 Fe I and ~ 10 Fe II lines at the S/N of our spectra, such that we can derive the atmospheric parameters from the Fe lines with sufficient confidence. We determine T_{eff} by imposing the excitation potential equilibrium of the Fe I lines and gravity with the ionization equilibrium between Fe I and Fe II lines. Note that for $\log g$ we impose Fe II abundances that are slightly higher (by 0.07–0.08 dex) than the Fe I ones to adjust for non-local thermodynamic equilibrium (NLTE) effects (Bergemann et al. 2012; Lind, Bergemann & Asplund 2012). For this analysis, microturbulent velocities, ξ_t were set to minimize any dependence on Fe I abundances as a function of EWs.

In Table 2, we compare effective temperatures for UVES RHB stars derived from spectroscopy and those obtained from the Alonso $(B - V) - T_{\text{eff}}$ calibrations, assuming a reddening of $E(B - V) = 0.19$. We find that this reddening value best reproduces the spectroscopic T_{eff} measurements listed in Table 2, with two stars having the same T_{eff} within ~ 30 K, and one within ~ 100 K. This value of $E(B - V)$ corresponds with that provided by B00 and agrees within errors with the value determined from RR-Lyrae by Kunder et al. (2013, $E(B - V) = 0.17 \pm 0.02$). A reddening of $E(B - V) = 0.22$ (as listed in the Harris catalogue) gives too high a T_{eff} with respect to the spectroscopic values.

The approach adopted for UVES spectra, independent of photometry, allows us to evaluate systematic differences with the atmospheric parameters determined for the GIRAFFE targets via the colour- T_{eff} calibrations, and adopt the same T_{eff} scale for UVES and GIRAFFE targets. Of course, un-accounted sources of systematic errors, such as the local thermodynamic equilibrium (LTE) approximation, could exist in the spectroscopic determination of T_{eff} . However, as our T_{eff} scale agrees with recent reddening estimates for the cluster, we are confident in our adopted temperature scale.

In our study of RHB stars, we are mainly interested in the star-to-star internal abundance variations. We are thus sensitive to the internal errors associated with the atmospheric parameters. Internal uncertainties on the T_{eff} derived from the Alonso et al. (1999) calibrations are mainly due to photometric errors and differential reddening, which is expected to be a significant effect for a cluster with reddening comparable to NGC 2808. Indeed, the map of Fig. 1 shows that star-to-star colour variations in the

⁸ We prefer not to use the 2MASS (Skrutskie et al. 2006) magnitudes to derive T_{eff} from other colours [such as the $(V - K)$] because some of our targets are in the crowded central region of the cluster where the 2MASS magnitudes are either not available or too uncertain.

11.25 arcmin \times 11.25 arcmin field around the cluster centre can be as large as ~ 0.06 mag, which corresponds to variations in T_{eff} of ~ 200 K. To increase the precision on the temperature estimate, we correct the photometry for these effects. After this correction, for bright HB stars the internal errors associated with our $(B - V)$ colours are typically around 0.02 mag, corresponding to errors in T_{eff} of the order of ~ 50 –60 K.

As a check, we re-determine temperatures using the B00 photometry. The mean difference between the T_{eff} derived from the two different photometries is 75 ± 5 K (rms = 37 K, with the B00 T_{eff} lower). Comparison of the two different photometric catalogues suggests that, after the absolute reddening value for the cluster has been fixed, we expect an accuracy for the RHB T_{eff} of ~ 100 K. Uncertainty due to internal photometric errors, as well as the rms associated with the mean difference between the T_{eff} derived from the two photometric data sets, is thus less than this value. This suggests that internal errors associated with our photometric T_{eff} are no larger than ~ 50 K.

Surface gravities for the GIRAFFE RHB stars are obtained from the apparent V magnitudes from Momany et al. (2004), corrected for differential reddening, the T_{eff} , bolometric corrections from Alonso et al. (1999) and an apparent distance modulus of $(m - M)_V = 15.67$, which is obtained from isochrone fitting and consistent with the value obtained in B00. We assume masses taken from isochrones, which range over an interval of $\sim 0.1 M_{\odot}$ ($0.61 \lesssim M \lesssim 0.75 M_{\odot}$). Our $\log g$ determinations for stars observed with GIRAFFE are affected by intrinsic uncertainty in mass owing to stochastic mass-loss in the RGB phase. Internal errors in T_{eff} values of $\sim \pm 50$ K and of $\sim \pm 0.1$ in mass, affect the $\log g$ values by $\sim \pm 0.02$ and ± 0.06 dex, respectively. The internal photometric uncertainty associated with our V mag modifies our $\log g$ by ~ 0.01 dex. All these effects, added in quadrature, contribute to an internal error in $\log g \lesssim 0.10$ dex.

Microturbulent velocities for GIRAFFE data cannot be independently determined from the spectral lines or photometry. We utilize a $\xi_t - T_{\text{eff}} - \log g - \text{metallicity}$ relation developed for the *Gaia*-ESO survey (GES; Gilmore et al. 2012) which was obtained from different literature sources.⁹ For this relation, we adopt a mean $[\text{Fe}/\text{H}] = -1.14$ (Harris catalogue), and our estimates of T_{eff} and $\log g$ explained above. The range in our estimated ξ_t is quite small, from 1.50 to 1.56 km s^{-1} , with a mean value of 1.54 km s^{-1} ($\sigma = 0.01$). This range is much smaller than that observed for the UVES RHB stars (see Table 2). We note that this relation for ξ_t has not been calibrated for HB stars and is formally applicable only to giants and dwarfs. However, comparison with ξ_t values derived from UVES spectra using Fe I spectral lines suggests that our values are accurate to $\sim 0.2 \text{ km s}^{-1}$.

As a further test, we derive ξ_t values from Fe I spectral lines from GIRAFFE data, despite the small number of features available. We obtain a median value of $\xi_t = 1.64 \pm 0.13 \text{ km s}^{-1}$. We cannot establish if our sample of stars has an intrinsic dispersion in ξ_t , or if this higher dispersion obtained from spectroscopy is due to observational uncertainties. Given the limited number of spectral lines in the GIRAFFE data, we use the ξ_t values from the GES empirical relations for our analysis. We also assume an internal error of 0.15 km s^{-1} , similar to the dispersion associated with spectroscopic ξ_t .

The photometric atmospheric parameters have been used to derive the final model atmospheres by interpolating in the Castelli & Kurucz (2004) grid.

3.3 Tests on the atmospheric parameters

We test our adopted atmospheric parameters for RHB stars by comparing them with those available in the literature, and those obtained with different techniques.

Gratton et al. (2011; hereafter G11) analysed GIRAFFE spectra of red and BHB stars in NGC 2808. These authors estimated effective temperatures of RHB stars from their $(B - V)$ and $(V - K)$ colours (with higher weight to $(B - V)$), using the calibration of Alonso et al. (1999). For all of the stars, they assumed the average reddening of NGC 2808 given by Harris (1996) and the $E(B - V)/E(V - K)$ value from Cardelli, Clayton & Mathis (1989). Gravities were estimated from masses, luminosities and effective temperature by assuming a constant mass of $0.7 M_{\odot}$ for RHB stars. Metallicity values are obtained from the analysis of Fe I lines. For further details, we refer the reader to Section 3 of G11.

17 RHB stars analysed by G11 are included in our work. We compare the atmospheric parameters obtained in these two works in Fig. 4, where we show $\Delta T_{\text{eff}} = T_{\text{eff G11}} - T_{\text{eff this paper}}$ (lower panel), $\Delta[\text{Fe}/\text{H}] = [\text{Fe}/\text{H}]_{\text{G11}} - [\text{Fe}/\text{H}]_{\text{this paper}}$ (middle panel) and $\Delta \log g = \log g_{\text{G11}} - \log g_{\text{this paper}}$ (upper panel) as a function of the $(B - V)$ colour corrected for differential reddening.

On average, T_{eff} estimates by G11 are systematically higher than those measured in our paper by 160 K with a scatter of ~ 55 K. This offset is in part due to the different $E(B - V)$ we use with the Alonso et al. (1999) calibration, as our value provides T_{eff} that are

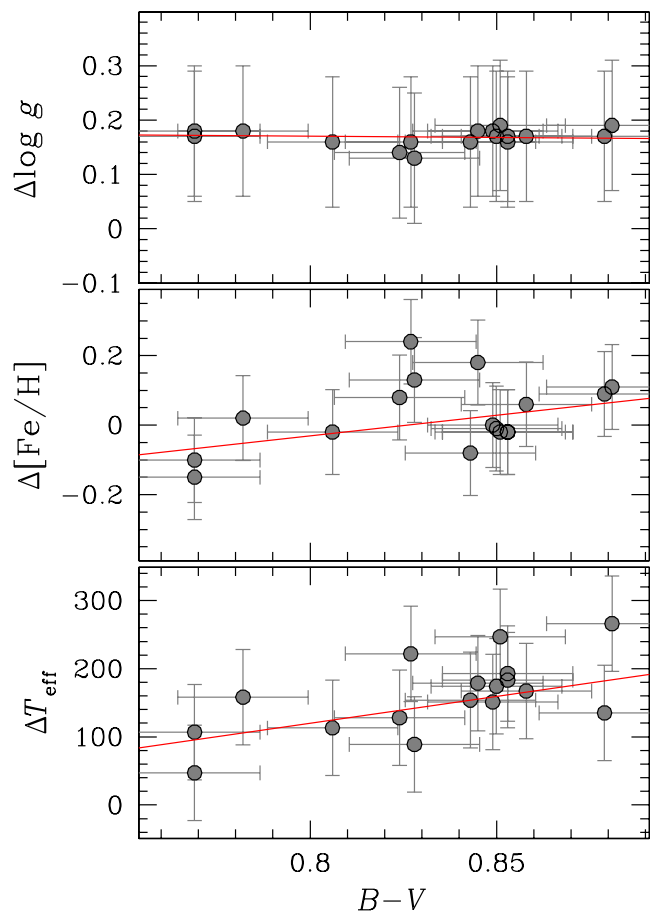


Figure 4. Upper, middle and lower panels show the difference between gravity ($\Delta \log g$), iron ($\Delta[\text{Fe}/\text{H}]$) and effective temperature (ΔT_{eff}) measurements derived by G11 and those of this paper as a function of the $(B - V)$ colour. Red lines are the best-fitting straight lines.

⁹ <http://great.ast.cam.ac.uk/GESwiki/GesWg/GesWg11/Microturbulence>

lower by ~ 100 K. The temperature difference is mildly correlated with the colour, with a Spearman's coefficient equal to 0.45. This may be due to our corrections in $(B - V)$ for differential reddening, zero-point effects and/or the larger photometric uncertainties and reddening variations in $(V - K)$.

Iron abundances from G11 and from this paper differ by 0.03 dex with a scatter of 0.1 dex. There is a mild trend between $\Delta[\text{Fe}/\text{H}]$ and the colour, which is a consequence of the $T_{\text{eff}}-(B - V)$ trend. The difference in $[\text{Fe}/\text{H}]$ ranges from ~ -0.1 dex for stars with $(B - V) \sim 0.77$ to $\sim +0.1$ dex for stars with $(B - V) \sim 0.88$.

Gravities from G11 are systematically higher by ~ 0.17 dex. This difference appears to be constant over the analysed colour interval, and could be due, in part, to the assumed distance modulus, the T_{eff} scale and the different choice for the stellar masses. Indeed, the distance modulus, used in G11, is lower ($(m - M)_V = 15.59$) than the value adopted here and, as such, gives higher $\log g$ values by 0.08 dex, on average. A systematic error in temperature of 160 K will change gravities by another 0.06 dex. Finally, rather than a fixed mass of $0.7 M_{\odot}$ as in G11, we use masses from isochrones that range towards lower stellar mass values.

As a further test of our effective temperatures, we compute T_{eff} values from the infrared flux method (IRFM; see Casagrande et al. 2010), using our B , V , JHK from 2MASS, and our adopted $\log g$. Fig. 5 shows the comparison between the T_{eff} values adopted here with those derived from the IRFM [$T_{\text{eff}}(\text{IRFM})$]. The error bars associated with the $T_{\text{eff}}(\text{IRFM})$ are set equal to the scatter in temperature derived from J , H and K . Additional uncertainties in these temperatures come from the errors associated with the adopted $\log g$, metallicities, zero-points in the T_{eff} scale and reddening effects. The T_{eff} from the IRFM (listed in the last column of Table 3)

are systematically higher by 87 ± 16 K, with an rms of 121 K. The difference between the two T_{eff} scales may be caused by an offset in reddening. The contribution to the $T_{\text{eff}}(\text{IRFM})$ error introduced by the adopted $\log g$ is negligible, as a relatively large error in $\log g$ of ~ 0.5 dex corresponds to an error in T_{eff} of $\lesssim 50$ K. As shown in the upper panel of Fig. 5, there is no significant trend between the difference $\Delta T_{\text{eff}} = T_{\text{eff}}(\text{IRFM}) - T_{\text{eff}}(\text{this paper})$ with the $(B - V)$ colour.

4 CHEMICAL ABUNDANCES

Chemical abundances for all elements, with the exception of helium, were obtained from the equivalent widths (EWs) of isolated spectral lines, whose profiles have been fitted with a Gaussian. The spectra of RHB stars are very different from those of BHB stars. For RHB stars, in the GIRAFFE HR12 spectral range, we measure Fe I and Fe II from 6–8 and 1–2 spectral lines, respectively. For our measurements of Na, Si, Ca, Ti, Mn and Ba, there are one to two transitions for each element. The UVES spectra have larger wavelength coverage and higher resolution; thus, we can determine chemical abundances in the RHB stars for additional elements, i.e. Fe, p -capture elements Na and Mg, α elements Si, Ca and Ti, iron-peak elements Sc, Cr, Ni and Zn, and n -capture elements Y, Ba and Nd. In HR12 observations of cooler BHB stars, we are able to measure only the Na resonance doublet. Thus, for the BHB stars with $T_{\text{eff}} \gtrsim 9000$ K, we cannot derive reliable Na abundances because of the very weak Na lines. For hotter BHB stars, the He feature at $\sim 5875.6 \text{ \AA}$ is detectable in the GIRAFFE spectra. For the four BHB observed with UVES, we determine iron abundances from 6–7 Fe II lines, and magnesium from the lines $\lambda 5173 \text{ \AA}$ and $\lambda 5184 \text{ \AA}$. Given the low S/N of the UVES spectra, the Na D lines are not well detected, and we are able to estimate only upper limits to the Na content.

Our atomic data are from the National Institute of Standards and Technology¹⁰ (NIST) compilation. For the few lines for which NIST does not provide the oscillator strengths ($\log gf$), we use $\log gf$ from Vienna Atomic Line Database (VALD; Kupka et al. 2000) and Meléndez & Barbuy (2009) for Fe II lines.

We derive chemical abundances from a LTE analysis using the latest version of the spectral analysis code MOOG (Snedden 1973). For the helium-line analysis, we employ a hybrid NLTE approach as described by Przybilla et al. (2006a,b, see Section 4.2 for further details). For both RHB and BHB stars, we apply NLTE corrections to Na abundances, as prescribed by Lind et al. (2011).

4.1 Red horizontal branch stars

4.1.1 UVES sample

The derived chemical abundances for the three RHB stars with UVES spectra are listed in Table 4. For each element, we report the abundance relative to Fe I or Fe II according to their ionization state; this minimizes the dependence of the derived abundance on the model atmosphere.

To determine the influence of atmospheric parameter uncertainties on our chemical contents, we re-derive abundances changing $T_{\text{eff}}/\log g/\xi_t/[A/H]$ by $\pm 100 \text{ K}/\pm 0.15/\pm 0.20 \text{ km s}^{-1}/0.10$ dex, respectively. Neutral iron is primarily sensitive to T_{eff} and ξ_t ; it increases by 0.09 dex with a larger temperature and decreases by 0.06 dex with a larger microturbulence. As expected, $\log g$ has the largest effect on singly ionized iron; a larger gravity value increases

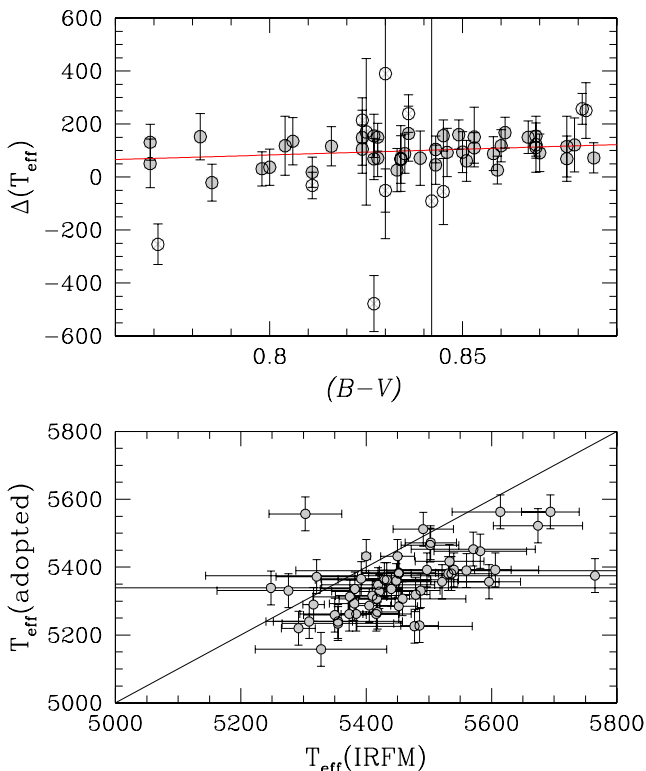


Figure 5. Lower panel: adopted T_{eff} as a function of T_{eff} derived from the IRFM. Upper panel: difference between T_{eff} estimates from the IRFM and the adopted values as a function of the $(B - V)$ colour. The empty circles are 2σ outliers and have been rejected from the least-squares fit.

¹⁰ http://physics.nist.gov/PhysRefData/ASD/lines_form.html

Table 3. Atmospheric parameters and chemical abundances for the RHB stars and colder BHB stars observed with GIRAFFE. For Na and Mn we list the line-to-line scatter.

ID	T_{eff}	$\log g$	ξ_t	[Fe/H] _I	[Fe/H] _{II}	[Na/Fe] _{LTE}	[Na/Fe] _{NLTE}	err _[Na/Fe]	[Si/Fe]	[Ca/Fe]	[Ti/Fe]	[Mn/Fe]	err _[Mn/Fe]	[Ba/Fe] _{II}	T_{eff} (IRFM)
2771	5381	2.34	1.54	-1.32	-1.24	0.49	0.04	0.04	0.32	0.60	0.10	-0.28	-	0.22	-
2648	5384	2.36	1.54	-1.07	-0.92	0.39	0.01	0.04	0.34	0.40	-0.06	-0.37	-	0.15	5452
4036	5262	2.30	1.55	-1.22	-1.14	0.21	-0.20	0.15	0.17	0.25	-	-0.55	-	-0.01	5384
4491	5563	2.40	1.53	-1.11	-0.98	0.76	0.39	0.10	0.41	0.51	-	-0.55	-	0.55	5614
3093	5360	2.35	1.54	-1.19	-1.17	0.34	-0.09	0.00	0.32	0.44	-	-0.45	0.09	0.35	5448
1317	5319	2.34	1.54	-1.25	-1.13	0.57	0.23	0.05	0.38	0.49	0.12	-0.46	-	0.22	5479

Table 4. Atmospheric parameters and individual abundances, with the associated line-to-line scatter, for the RHB stars observed with UVES.

ID	T_{eff}	$\log g$	ξ_t	[Fe/H]	[Fe/H] _I	[Fe/H] _{II}	[Na/Fe] _{LTE}	[Na/Fe] _{NLTE}	[Mg/Fe]	[Si/Fe]	[Ca/Fe]			
20577	5390	2.26	1.66	-1.21	-1.24 ± 0.03	-1.17 ± 0.12	0.23 ± 0.19	-0.13 ± 0.08	0.23 ± 0.10	0.24 ± 0.13	0.20 ± 0.06			
18106	5480	2.55	1.48	-1.14	-1.18 ± 0.03	-1.10 ± 0.07	0.35 ± 0.16	0.03 ± 0.04	0.29 ± 0.11	0.18 ± 0.10	0.15 ± 0.06			
21854	5350	2.20	1.73	-1.17	-1.21 ± 0.03	-1.12 ± 0.08	0.11 ± 0.08	-0.27 ± 0.13	0.31 ± 0.08	0.12 ± 0.12	0.24 ± 0.06			
					[Sc/Fe] _{II}	[Ti/Fe] _I	[Ti/Fe] _{II}	[Cr/Fe] _I	[Cr/Fe] _{II}	[Ni/Fe]	[Zn/Fe]	[Y/Fe] _{II}	[Ba/Fe] _{II}	[Nd/Fe] _{II}
					-0.01 ± 0.05	0.34 ± 0.01	0.20 ± 0.06	-0.19 ± 0.04	0.12 ± 0.00	-0.07 ± 0.07	-0.18 ± 0.00	-0.20 ± 0.04	-0.14 ± 0.00	0.15 ± 0.00
					0.17 ± 0.08	0.39 ± 0.11	0.22 ± 0.15	-0.13 ± 0.06	-0.18 ± 0.00	-0.04 ± 0.12	0.16 ± 0.00	-0.11 ± 0.04	0.12 ± 0.00	0.01 ± 0.00
					0.07 ± 0.08	0.25 ± 0.08	0.13 ± 0.10	-0.11 ± 0.11	-0.04 ± 0.00	-0.03 ± 0.05	-0.02 ± 0.00	-0.22 ± 0.07	-0.04 ± 0.00	-0.03 ± 0.00

the measured abundances by 0.06 dex. In most cases, the corresponding variations in the abundances relative to Fe are small (a few hundredths of a dex).

Changes in T_{eff} by ± 100 K result in the largest variations in [Si, Ba, Nd/Fe]; these abundances change by ∓ 0.06 , ± 0.07 and ± 0.05 dex, respectively. Barium abundances are also sensitive to ξ_t . Variation in this parameter by ± 0.20 km s⁻¹ corresponds to a change in [Ba/Fe] by ∓ 0.12 dex. The other parameters have small impact on the derived abundances for these elements over Fe.

The analysed elements show one of the typical chemical patterns of a GC, e.g. α enhancement indicated by the abundances of Si, Ca and Ti. Mg is not considered to be a pure α element, as it is affected by the p -capture reactions that convert Mg to Al. Unfortunately, the quality of our spectra is not good enough to infer Al abundances; however, the Mg content appears to be un-depleted in all three stars.

We determine the sodium abundance from both the resonance doublet at ~ 5890 Å and the feature at ~ 5688 Å. The resonance doublet is heavily affected by deviations from LTE. We apply the NLTE corrections from Lind et al. (2011) to all our Na spectral lines, significantly decreasing the abundance from the Na resonance doublet by ~ 0.4 – 0.5 dex. The correction for the other Na feature is smaller, ~ 0.10 dex. After the NLTE correction, the Na abundances from all three lines are consistent within errors, and the standard deviation associated with the mean Na value of each star decreases, except for star #21854, for which the standard deviation slightly increases. The agreement between the NLTE-corrected abundances derived from the Na resonance doublet and the 5688 Å line suggests that, with the proper NLTE corrections, abundances derived from the very strong Na resonance lines are reliable.

4.1.2 GIRAFFE sample

In Table 3, we list the adopted atmospheric parameters and corresponding abundances determined from the GIRAFFE spectra of the RHB stars. The only feature available to determine Na abundances in these data is the resonance doublet. Results obtained on UVES data suggest that we can confidently use these strong

lines for abundance determinations, once NLTE corrections have been applied. As done for the UVES data, we apply NLTE corrections from Lind et al. (2011), decreasing the mean [Na/Fe] content for our RHB sample from $[\text{Na/Fe}]_{\text{LTE}} = 0.48 \pm 0.02$ ($\sigma = 0.15$) to $[\text{Na/Fe}]_{\text{NLTE}} = 0.09 \pm 0.02$ ($\sigma = 0.15$). This correction decreases the mean difference between the abundances given by the two single lines of the doublet from $[\text{Na/Fe}]_{5889-5895} = -0.09$ to $[\text{Na/Fe}]_{5889-5895} = 0.01$ dex, indicating that accounting for NLTE effects improves our Na values.

The mean abundance ratios relative to Fe for the other elements measured from GIRAFFE spectra of RHB stars are: $[\text{Si/Fe}]_I = 0.44 \pm 0.01$ ($\sigma = 0.09$); $[\text{Ca/Fe}]_I = 0.52 \pm 0.01$ ($\sigma = 0.09$); $[\text{Ti/Fe}]_I = 0.21 \pm 0.02$ ($\sigma = 0.12$); $[\text{Mn/Fe}]_I = -0.42 \pm 0.01$ ($\sigma = 0.10$); $[\text{Fe/H}]_I = -1.22 \pm 0.01$ ($\sigma = 0.07$); $[\text{Fe/H}]_{II} = -1.10 \pm 0.01$ ($\sigma = 0.08$); and $[\text{Ba/Fe}]_{II} = 0.14 \pm 0.02$ ($\sigma = 0.19$). We notice that these mean chemical abundances are similar to those observed in other clusters. Silicon and titanium are enhanced, reflecting the α enhancement observed in field and GCs stars at similar metallicities, and manganese is subsolar, with a mean value consistent with that observed in other clusters (Sobeck et al. 2006). The quite high average value for [Ca/Fe] is due to the limited spectral features available; we have only the transition $\lambda 6122$ Å which is on the flat part of the curve of growth. As a consequence, it returns abundances systematically higher by ~ 0.20 dex than other transitions, as verified on UVES spectra.

To derive estimates of the internal uncertainties associated with our determinations, we calculate the sensitivity of the abundances to various sources of internal errors, e.g. uncertainties in the model atmosphere and in the EWs. The stellar parameter internal uncertainties are ($T_{\text{eff}}/\log g/[\text{Fe/H}]/\xi_t$): ± 50 K/0.1 dex/0.07 dex/0.15 km s⁻¹ (see Section 3.2). We set the uncertainty associated with the [Fe/H] values equal to the dispersion of the obtained Fe I abundances, assuming that the cluster is monometallic. To quantify the effect on the abundances, we run a series of models in which each parameter is varied by its corresponding uncertainty one at a time.

The contribution to the errors given by EW uncertainties has been calculated by varying the EWs of spectral lines by ± 4 mÅ. This is the typical error associated with our EW measurements,

Table 5. Sensitivity of derived GIRAFFE RHB abundances to the atmospheric parameters and EWs. We reported the total error due to the atmospheric parameters and the EW measurement, the squared sum of these contributions (σ_{total}) and the observed dispersion (σ_{obs}) for each element.

	ΔT_{eff} $\pm 50 \text{ K}$	$\Delta \log g$ ± 0.10	$\Delta \xi_t$ $\pm 0.15 \text{ km s}^{-1}$	$\Delta [A/H]$ 0.07 dex	σ_{EW} $\pm 4 \text{ m}\text{\AA}$	σ_{total}	σ_{obs}
[Na/Fe] _I	± 0.03	∓ 0.03	∓ 0.01	± 0.01	± 0.05	0.07	0.15
[Si/Fe] _I	∓ 0.01	± 0.00	± 0.01	± 0.00	± 0.08	0.08	0.09
[Ca/Fe] _I	± 0.01	∓ 0.02	∓ 0.06	∓ 0.01	± 0.07	0.10	0.09
[Ti/Fe] _I	± 0.02	± 0.00	± 0.01	± 0.00	± 0.11	0.11	0.12
[Mn/Fe] _I	± 0.01	± 0.00	± 0.02	± 0.00	± 0.11	0.11	0.10
[Fe/H] _I	± 0.04	± 0.00	∓ 0.03	± 0.00	± 0.04	0.07	0.07
[Fe/H] _{II}	± 0.00	± 0.04	∓ 0.01	± 0.01	± 0.11	0.12	0.08
[Ba/Fe] _{II}	± 0.03	∓ 0.01	∓ 0.10	± 0.00	± 0.16	0.19	0.19

determined by comparing EWs for stars with similar atmospheric parameters and abundances. The variations in the Fe I abundances from EW uncertainty are then divided by the square root of the number of available spectral lines. Hence, since the EW measurement errors are random, the uncertainty is lower for those elements with a large number of lines. For the other elements, we have only one or two lines, and the error contribution by EW uncertainties is higher. Given the importance of the Na abundances for the discussion (see Section 5), we perform an additional test to estimate the error introduced by uncertainties in EW, measuring the Na EWs from each single exposure. The associated $\text{rms}/\sqrt{N-1}$ (N = number of exposures) is $\sim 4 \text{ m}\text{\AA}$, agreeing well with the EWs error estimate used to determine the abundance uncertainties earlier.

The variations in abundances obtained by varying the atmospheric parameters and EWs are listed in Table 5. In this table, we also list the squared sum of these different contributions (σ_{total}) and the observed dispersion (σ_{obs}) for each element. In some cases, the dispersions are large, as we expect from an analysis limited to only one or two spectral lines. The expected error values σ_{total} are only rough estimates of the internal uncertainties we expect for our abundances, but they suggest that, in our analysis, the main contributors to the errors are EW uncertainties. Overall, the observed dispersions agree well with the expected values, apart for Fe II, where the expected errors are slightly overestimated, and Na, where the observed dispersion is more than twice as large as expected. We discuss the Na spread further in Section 5.

4.2 Blue horizontal branch stars

4.2.1 UVES sample

The UVES sample of BHB stars consists of four stars, with $9200 < T_{\text{eff}} < 11\,000 \text{ K}$. Despite the larger available spectral range, the only measurable iron lines are 6–7 singly ionized lines. Hence, as with the GIRAFFE data, we use T_{eff} and $\log g$ values from isochrones, as explained in Section 3.1. This assures uniformity

Table 6. Atmospheric parameters, chemical abundances and line-to-line scatter (when applicable) for the BHB UVES stars.

ID	T_{eff}	$\log g$	ξ_t	[Fe/H] _{II}	[Na/Fe] _{NLTE}	[Mg/Fe]
19229	9268	3.25	2.26	-1.04 ± 0.06	0.44	0.12 ± 0.05
716	10144	3.51	2.07	-1.23 ± 0.10	< 0.85	0.21 ± 0.08
509	10387	3.59	2.02	-1.15 ± 0.10	< 0.86	0.16 ± 0.04
11715	11036	3.74	1.90	-1.07 ± 0.11	< 0.80	0.39 ± 0.07

in the T_{eff} and $\log g$ scales used for UVES and GIRAFFE BHB stars.

Atmospheric parameters and chemical abundances obtained for these stars are listed in Table 6. Given that our information on the metallicity of these stars comes only from singly ionized Fe lines, we choose [Fe/H]_{II} as our metallicity estimate. The NLTE effects for these singly ionized lines are smaller (Bergemann et al. 2012; Lind et al. 2012). We are able to infer the Na content for only the coldest star. For the other three, we provide an upper limit to the Na abundances.

We estimate the internal abundance errors for these stars by altering the atmospheric parameters, changing T_{eff} by 250 K and ξ_t by 1 km s^{-1} , and redetermining the abundances. Iron content derived from singly ionized lines increases by $\sim 0.06 \text{ dex}$ with increasing T_{eff} and by $\sim 0.12 \text{ dex}$ with decreasing ξ_t . A variation in T_{eff} by $\pm 250 \text{ K}$ changes magnesium and sodium abundances relative to Fe II by ± 0.12 and $\pm 0.10 \text{ dex}$, respectively, while a $\pm 1 \text{ km s}^{-1}$ change in ξ_t varies both these abundances by roughly $\pm 0.10 \text{ dex}$. Formally, $\log g$ is not affected by high internal errors (Section 3.1), and should not significantly contribute to the internal abundance uncertainties. Systematics in $\log g$ of $\pm 0.20 \text{ dex}$ will change Fe II by $\sim \pm 0.05 \text{ dex}$, and Na and Mg relative to Fe II by $\mp 0.10 \text{ dex}$. In addition to uncertainties from the model atmosphere, the relatively low S/N of these spectra translates into a non-negligible error contribution ($\sim 0.10 \text{ dex}$ for all the species) from EW measurement errors to the *real* abundance uncertainties.

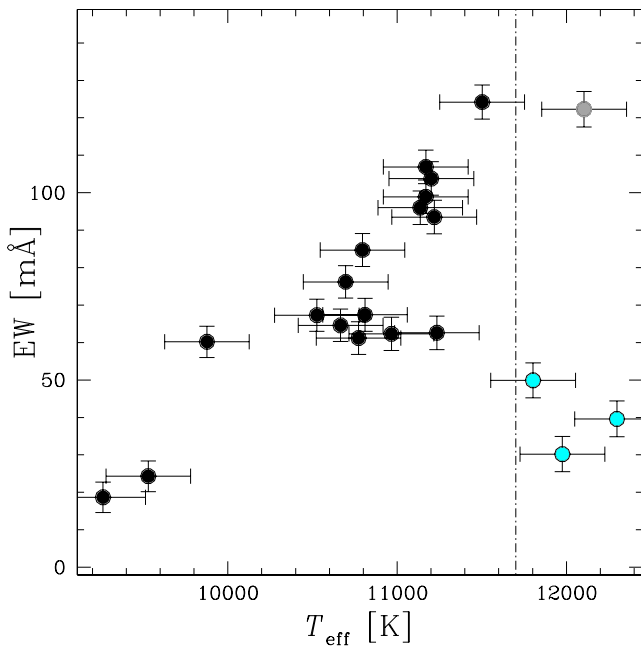
From the analysis of 12 RGB stars in NGC 2808, Carretta et al. (2009) found three of them with undersolar [Mg/Fe] abundance. The low Mg abundance in these three stars is accompanied by a higher Na abundance, typical of second-generation(s) stars. None of our 4 BHB stars analysed with UVES shows a similar low Mg abundance, but the [Na/Fe] abundance, measured for just one of the UVES sample, resembles the values of the Na-rich stars observed on the RGB. Given the large internal uncertainties associated with our [Mg/Fe] measurements, that are $\sim 0.20 \text{ dex}$, it is very difficult to draw firm conclusions here on the basis of our results on a sample of four stars. We can just not exclude that these stars, as suggested by theoretical studies of the HB (D’Antona et al. 2005), may not be the counterpart of the extremely Na-rich (O and Mg poor) stars observed on the RGB, but an intermediate population with mild Mg depletion. Three out four BHB stars in our sample show in fact slightly lower Mg abundances than the RHB stars.

4.2.2 GIRAFFE sample

The chemical abundances for Na and He obtained for the BHB stars are listed in Tables 3 and 7, respectively. For the hotter stars in our

Table 7. Atmospheric parameters, EWs for the He line and He mass fractions for the BHB stars observed with GIRAFFE.

ID	T_{eff} (K)	$\log g$ (cgs)	ξ_t (km s^{-1})	EW ($\text{m}\text{\AA}$)	$v \sin i$ (km s^{-1})	Y_{NLTE} mass fraction	$T_{\text{eff}} (HST)$ (K)	$\log g (HST)$ (cgs)
17132	9263	3.27	2.26	18.7	0	0.305	9312	3.27
17826	9525	3.39	2.20	24.3	10	0.308	9708	3.43
10342	9862	3.51	2.13	60.2	0	0.441		
3030	10500	3.59	2.00	67.3	0	0.357		
231	10639	3.60	1.97	64.6	8	0.333	10697	3.64
17336	10669	3.68	1.97	76.2	0	0.384		
11562	10744	3.66	1.95	61.2	0	0.271		
23393	10760	3.68	1.95	84.7	0	0.378		
14629	10783	3.67	1.95	67.4	0	0.338		
16837	10939	3.63	1.92	62.3	0	0.250	10003	3.49
14968	11109	3.75	1.89	96.0	15	0.329		
10883	11143	3.81	1.88	106.9	0	0.382		
15081	11143	3.76	1.88	98.9	5	0.353		
1129	11184	3.81	1.87	103.8	10	0.400		
23610	11184	3.77	1.87	93.5	10	0.324		
20630	11210	3.78	1.87	62.6	10	0.234	11189	3.76
1169	11478	3.85	1.82	124.2	0	0.366		
GJ stars								
13305	11779	3.81	0.00	49.9	10	0.052		
18640	11951	3.83	0.00	30.2	0	0.035	11604	3.84
4129	12068	3.96	0.00	122.3	10	0.218		
13132	12269	3.89	0.00	39.6	10	0.031		


Figure 6. EWs of the He line as a function of T_{eff} . Stars with $T_{\text{eff}} \gtrsim 11\,500$ K, likely affected by helium sedimentation, are represented by cyan circles. One of these stars, shown in grey, does not exhibit the extreme drop in the EW of the He line shown by the other three stars at similar temperature.

sample (21 stars in total with $T_{\text{eff}} > 9000$ K) we are able to detect the He I $\lambda 5875.6$ Å line, which has been used in previous studies (e.g. Behr 2003; Villanova et al. 2009, 2012).

In Fig. 6 we illustrate the behaviour of the EWs of the He spectral line with effective temperature. The EW steeply increases from ~ 10 up to ~ 130 mÅ over the T_{eff} interval ~ 9000 – 11500 K, and then falls

down to ~ 30 – 50 mÅ at higher temperature. This drop is likely a consequence of helium settling and metal levitation, which results in a depletion of the surface abundance of He and an overabundance of some heavy elements (Behr 2003; Fabbian et al. 2005). Since these phenomena cause the surface abundances to dramatically deviate from those of the star as a whole, we exclude from the following discussion the three stars with $T_{\text{eff}} > 11\,500$ K and anomalous values of the helium-line EW (stars represented in cyan in Fig. 6). We also exclude one star (#4129, shown in grey in Fig. 6) with a temperature hotter than the GJ that does not show an extreme drop in the EW. On the basis of its iron abundance (Pace et al. 2006), this object appears to be suffering from mild metal levitation, as its metallicity is higher than the cluster mean metallicity, but sensibly lower than that of hotter GJ stars (see table 3 in Pace et al. 2006).

To estimate the abundance of He, we perform a spectral synthesis of the observed line $\lambda 5875$ Å. The best fits to observations are shown in Figs 7 and 8 for BHB stars with temperatures lower and higher than $T_{\text{eff}} = 11\,500$ K, respectively. For 10 stars in the sample, we must account for a non-zero projected rotational velocity ($v \sin i$) to match the observed He line profile (see Table 7). The adopted values agree with the $v \sin i$ range derived for BHB stars in NGC 2808 by Recio-Blanco et al. (2004, $5 \lesssim v \sin i \lesssim 15$ km s^{-1}). The synthetic spectra, for the $T_{\text{eff}}/\log g/\xi_t$ values given by the isochrones and metallicity equal to -1.15 (see Section 3.1), were calculated employing NLTE line-formation computations with DETAIL/SURFACE (Giddings 1981; Butler & Giddings 1985) on ATLAS12 LTE model atmospheres, using the He I model atom of Przybilla (2005). The most relevant extension here is the implementation of opacity sampling (instead of opacity distribution functions) to account for the line blocking, adopting the technique and data discussed by Kurucz (1996). The approach has been successfully applied to spectrum synthesis analyses of Population II BHB and blue straggler stars before (Tillich et al. 2010; Przybilla, Nieva & Butler 2011).

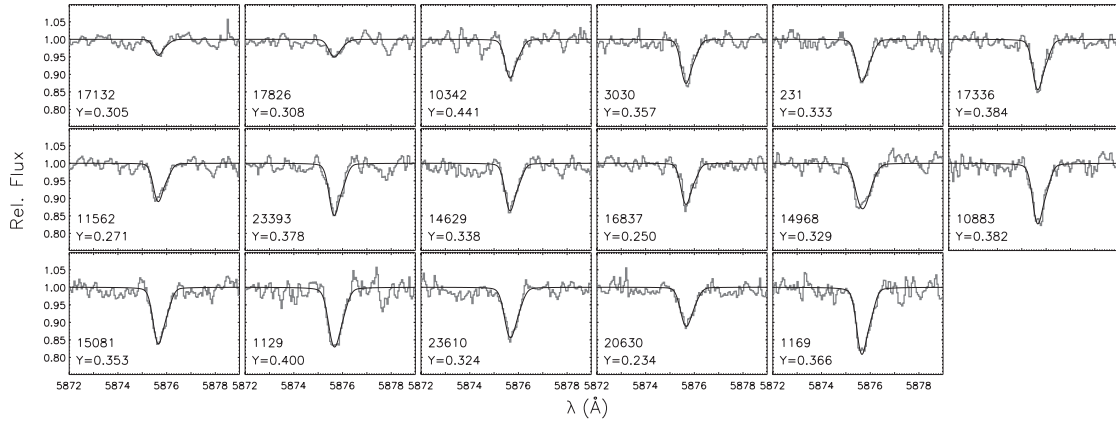


Figure 7. Comparison of synthetic NLTE spectra (black lines) with observed spectra (grey lines) for 17 BHB stars cooler than the GJ. Stellar identifications and the best-fitting helium mass fractions Y are indicated. The individual panels are ordered by increasing effective temperature of the stars.

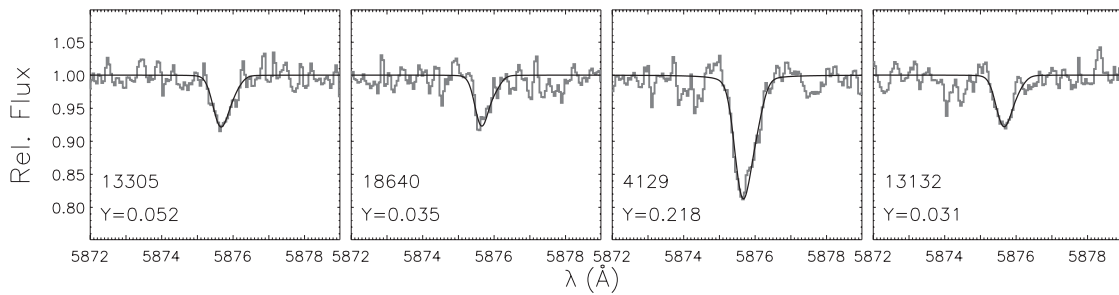


Figure 8. Same as Fig. 7, but for the four sample stars beyond the GJ.

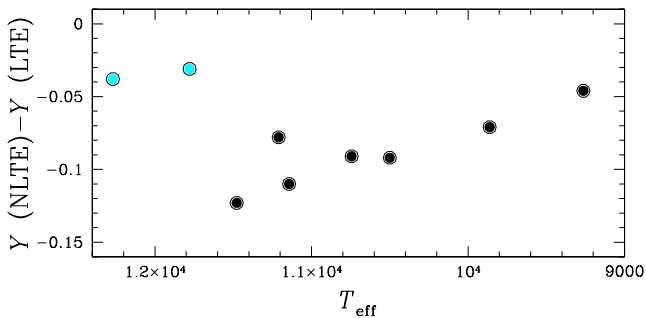


Figure 9. Difference between NLTE and LTE helium mass fractions as a function of the effective temperature. Colours are the same as in Fig. 6.

Uncertainties associated with the He abundances were determined for seven BHB stars (#17132, #10342, #3030, #14968 and #1169, and GJ stars #13305 and #4129), at characteristic points of the parameter range. Considering typical uncertainties in $T_{\text{eff}}/\log g/\xi_v/[A/H]$ of 250 K/0.05 dex/1.00 km s⁻¹/0.07 dex, and those due to continuum placement, we find a *typical* total uncertainty (sum over the squared contributions) of 0.05–0.06 and 0.02–0.03 in mass fraction for stars below and above the GJ. The main source of uncertainty stems from inaccuracies in the T_{eff} estimation; all other factors are relatively minor.

We find that NLTE effects on the helium abundance estimates are substantial. The difference between LTE and NLTE helium abundances depends on both the atmospheric parameters and helium content of the star. This difference is plotted in Fig. 9 as a function of T_{eff} for a representative sample of stars. It varies between 0.03 and 0.12 in mass fraction from star to star, with higher values derived for LTE. The NLTE effects on He I $\lambda 5875.6$ Å tend to increase when the

line gets stronger, as the feature is not saturated in our hotter BHB stars. This results from a slight NLTE overpopulation of the lower level $2p^3P^o$ and a slight NLTE underpopulation of the upper level $3d^3D$, both by less than 5 per cent relative to detailed equilibrium values throughout the line-formation region. Both levels play an important role in the recombination cascade, which leads to the NLTE overpopulation of the metastable lowest state in the helium triplet spin system. Neglect of NLTE effects systematically biases the He abundances, exceeding the combined effect of all other sources of uncertainty (internal and external).

We do not expect any significant uncertainty due to inaccuracies in the atomic data employed for the model atom construction of He I. The atomic data for He I measured in experiments and provided by ab initio computations has the highest accuracy and precision next to that of hydrogen. Synthetic spectra obtained with the He I model atom adopted here compare well to those obtained with other NLTE codes (see e.g. Nieva & Przybilla 2007; Przybilla et al. 2011). Moreover, and even more importantly, this model atom can reproduce the *entire* observed visual/near-IR He I line spectra in a large variety of star types: in supergiants with similar T_{eff} than the BHB stars investigated here (Przybilla et al. 2006b), in B-type MS stars (Nieva & Przybilla 2012), in extreme HB stars (Przybilla, Nieva & Edelmann 2006a) and in extreme helium stars (Przybilla et al. 2005).

5 THE ABUNDANCE DISTRIBUTION OF STARS ALONG THE HORIZONTAL BRANCH

The distribution of [Na/H] for HB stars observed with GIRAFFE is shown in Fig. 10. There is a large star-to-star variation in sodium abundance, with [Na/H] ranging from ~ -1.4 up to ~ -0.5 dex. In

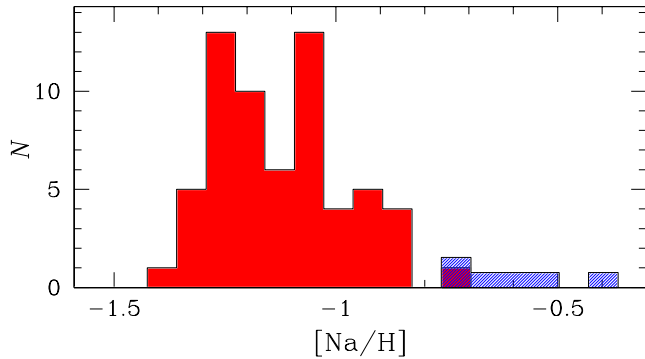


Figure 10. Histogram distribution of $[\text{Na}/\text{H}]$ for the HB stars in NGC 2808. RHB and BHB stars have been indicated with red and blue colours, respectively.

this section, we investigate the relation between the sodium abundance of a star and its position along the HB. To this aim, we consider two groups of BHB and RHB stars, represented with different colours in Fig. 10.

Following B00, the BHB is formed by three segments separated by different gaps: extended blue tail 1, 2 and 3 (EBT1, EBT2 and EBT3) in order of increasing temperature. Specifically, the EBT1 includes hot HB stars located between the blue edge of the RR-Lyrae instability strip and the first gap; the EBT2 includes extreme HB stars located between the first and the second gap; and stars in the EBT3 are extreme HB stars hotter than the second gap.

All our analysed BHB stars (with the exception of GJ stars) belong to the EBT1. The Na richest stars (with $[\text{Na}/\text{H}] \gtrsim -0.8$) are EBT1 stars, while stars on the red side of the instability gap show lower sodium values. This distinction suggests that the HB morphology in GCs is closely connected with the present stellar populations and demonstrates that the EBT1 of NGC 2808 is made of second-generation stars. In the following subsection, we discuss the Na abundance distribution for RHB stars.

5.1 The sodium content of red horizontal branch stars

Fig. 10 reveals that the $[\text{Na}/\text{H}]$ distribution of RHB stars is broad with some hints of bimodality. Similar conclusions can be inferred from the distribution of $[\text{Na}/\text{Fe}]$ shown in the upper-left panel of Fig. 11. These stars span an interval in $[\text{Na}/\text{Fe}]$ of ~ 0.6 dex, which is larger than that expected from observational errors alone (~ 0.07 dex, Table 5).

To further investigate the presence of multiple stellar populations along the RHB of NGC 2808, we define two groups of stars on the basis of the $[\text{Na}/\text{Fe}]$ distribution of Fig. 11: (i) Na-rich RHB stars with $[\text{Na}/\text{Fe}] > 0.10$ dex; and (ii) Na-poor RHB stars with $[\text{Na}/\text{Fe}] < 0.10$ dex. These two groups of stars have been coloured in magenta and green, respectively. These same colour codes are used consistently hereafter. In the other panels of Fig. 11, we show the B versus $(B - V)$ (upper-right panel), U versus $(U - V)$ (lower-left panel) and U versus $(U - B)$ (lower-right panel) CMDs in the RHB region. The position of spectroscopic targets in the CMD suggests that the Na-rich RHB stars have on average bluer colours with respect to the Na-poor stars. This behaviour is seen in the $(B - V)$, $(U - V)$ and $(U - B)$ colours. The mean differences (Na rich–Na poor) observed in various colours are: $\Delta(B - V) = -0.035 \pm 0.007$, $\Delta(U - V) = -0.048 \pm 0.011$ and $\Delta(U - B) = -0.012 \pm 0.005$.

Considering the working hypothesis of two groups of stars separated at $[\text{Na}/\text{Fe}] = 0.10$ dex, the average sodium abundances for

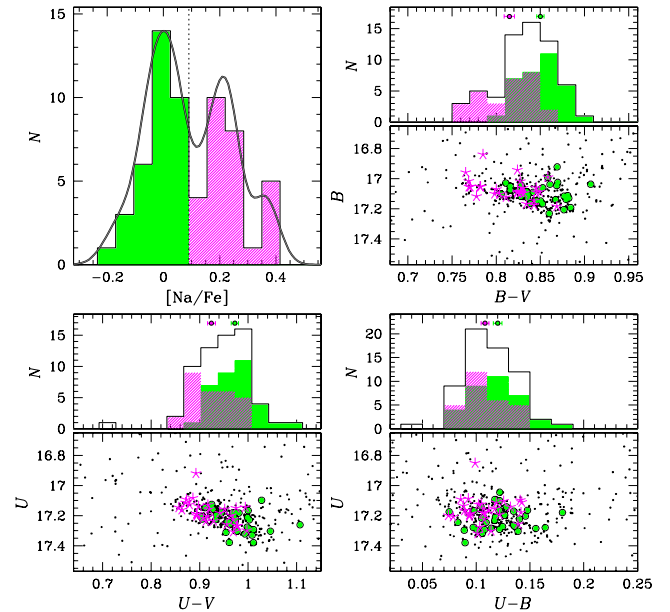


Figure 11. Upper-left panel: histogram distribution of $[\text{Na}/\text{Fe}]$ for RHB stars. Green and magenta colours represent the two samples of Na-poor and Na-rich RHB stars, respectively. A kernel density distribution has been superimposed to the observed distribution. Upper-right and lower panels: the B versus $(B - V)$, U versus $(U - V)$, and U versus $(U - B)$ CMD of NGC 2808 around the RHB. Na-poor and Na-rich RHB stars are represented with green dots and magenta asterisks, respectively. We also show the colour histogram distributions of the two Na subsamples of the RHB stars, using the same colour scheme.

these groups are: $[\text{Na}/\text{Fe}] = 0.24 \pm 0.01$ dex (rms = 0.08) and $[\text{Na}/\text{Fe}] = -0.02 \pm 0.01$ dex (rms = 0.07). This difference does not originate in our treatment of NLTE effects, as these corrections are comparable for the two groups. NLTE corrections for stars at higher Na are larger than those for low Na stars by only 0.02 dex. Since the NLTE corrections are fairly line-strength dependent, the small difference with respect to Na reflects that temperature and abundance influence line strength in opposite directions, partially cancelling with each other. In addition to this, the fact that Na-poor stars have, on average, redder colours than Na-rich ones further supports the idea that the RHB is not consistent with a simple stellar population. We note that systematics in the adopted T_{eff} values are unlikely to produce the observed difference in $[\text{Na}/\text{Fe}]$. To cancel the observed variation in $[\text{Na}/\text{Fe}]$ we would need T_{eff} values for the Na-rich stars colder by ~ 350 K (see Table 5). This requires these stars to be redder by ~ 0.15 in $(B - V)$, several times (seven) the observational error. As shown in Fig. 11, such a large systematic in the $(B - V)$ should also affect the other investigated colours, which we do not observe.

As a final test, we have re-determined the Fe and Na abundances using temperatures derived from the IRFM and $\log g$ consistent with these alternative T_{eff} . In Fig. 12, we compare the NLTE $[\text{Na}/\text{Fe}]$ abundances derived using the two different sets of atmospheric parameters, e.g. the adopted ($[\text{Na}/\text{Fe}]_{T_{\text{eff}}(\text{adopted})}$) and the IRFM ones ($[\text{Na}/\text{Fe}]_{T_{\text{eff}}(\text{IRFM})}$). There is no significant systematic in the derived $[\text{Na}/\text{Fe}]$, with a mean difference $[\text{Na}/\text{Fe}]_{T_{\text{eff}}(\text{IRFM})} - [\text{Na}/\text{Fe}]_{T_{\text{eff}}(\text{adopted})} = +0.01 \pm 0.01$ dex (rms = 0.04), consistent with zero. More importantly, there is no evidence for systematics that affect different ranges in the observed $[\text{Na}/\text{Fe}]$ distribution in different ways.

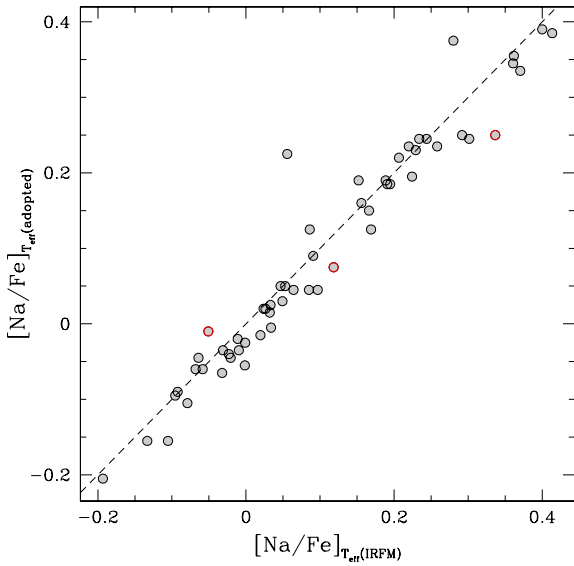


Figure 12. Comparison of the Na abundances relative to iron derived using our adopted atmospheric parameters and those derived using T_{eff} obtained from the IRFM. Both sets of Na abundances are corrected for NLTE effects. The dashed line represents perfect agreement. Stars represented with red open circles have the largest uncertainties in the T_{eff} derived from the IRFM ($\gtrsim 100$ K).

We conclude that the RHB stars Na distribution is not consistent with a single stellar population of stars homogeneous in Na. Although there are hints of a bimodal distribution in Na in the RHB, present uncertainties do not allow us to determine whether or not the observed distribution translates into a continuous or discrete star formation history. Apart from Na, the chemical composition of the two stellar populations along the RHB is uniform within observational errors. When we consider the groups as defined in the upper-left panel of Fig. 11, there is no evidence of separate distributions in other chemical abundances. This is shown in Fig. 13 where we plot the histograms of the silicon, calcium, titanium, manganese and barium abundance distribution for Na-rich and Na-poor RHB stars.

17 RHB targets analysed in this paper are in common with the sample of G11. In Fig. 14, we compare the $[\text{Na}/\text{Fe}]$ inferred in this paper and those of G11 for the same stars. The stars have been divided into Na-poor and Na-rich subsamples according to our

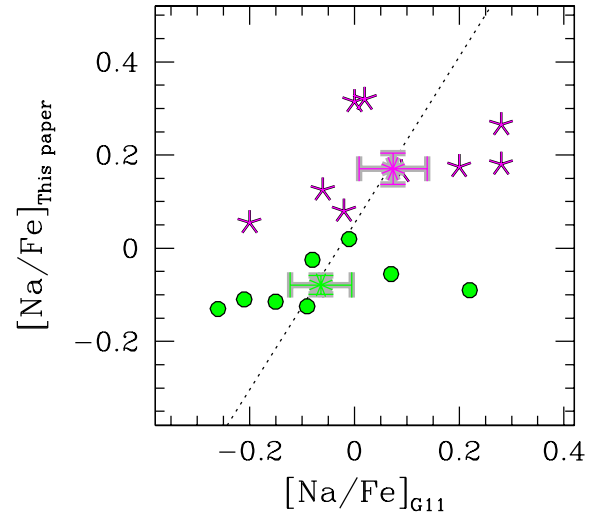


Figure 14. $[\text{Na}/\text{Fe}]$ derived in this analysis compared to that from G11 for the 17 stars in common. Colours and symbols are the same as in Fig. 11. Superimposed on the data are the average values for Na-rich and Na-poor stars, and the associated errors. The dotted line represents the least-squares linear fit.

criterion; eight are Na poor and nine Na rich. The least-squares fit with the data is shown as a dotted line. Both analyses exhibit similar dispersion, e.g. $\text{rms} = 0.16$ (this paper) and $\text{rms} = 0.17$ (G11), and a mean difference of $\Delta([\text{Na}/\text{Fe}]_{\text{this paper-G11}}) = 0.13 \pm 0.04$ ($\text{rms} = 0.16$). Although there is a positive correlation between our Na values and those from G11, in the G11 data, the stars we assign to the Na-rich and Na-poor group have, on average, similar $[\text{Na}/\text{Fe}]$ (within $\sim 1\sigma$) with a large overlap. We do not know the reason for this difference, and several effects may contribute.

To further strengthen our results, we examine the oxygen abundances for these common stars available from G11. Fig. 15 shows $[\text{Na}/\text{Fe}]$ as a function of $[\text{O}/\text{Fe}]$ for these stars, using sodium abundance determinations from this paper (left-hand panel) and from G11 (right-hand panel). Na-rich stars are, on average, depleted in oxygen by 0.12 ± 0.06 dex with respect to Na-poor stars. This difference, which is marginally significant (at the 2σ level), suggests that the RHB sample may have a mild Na–O anticorrelation, providing further evidence that it is not consistent with a simple stellar population.

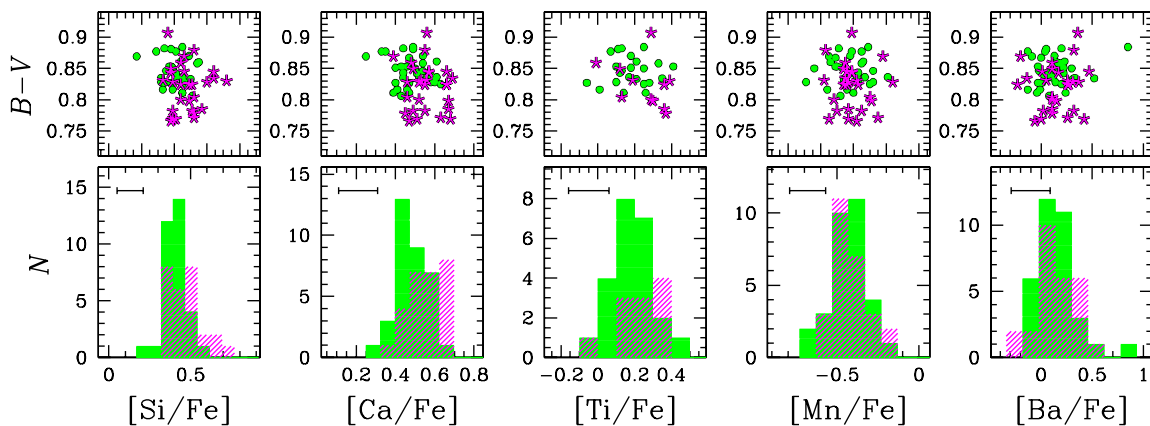


Figure 13. Lower panels: distribution of the derived chemical abundances for the Na-rich (magenta histograms) and Na-poor RHB stars (green histograms) analysed with GIRAFFE. Upper panels: $(B - V)$ as a function of the chemical abundance. Colours and symbols are the same as in Fig. 11.

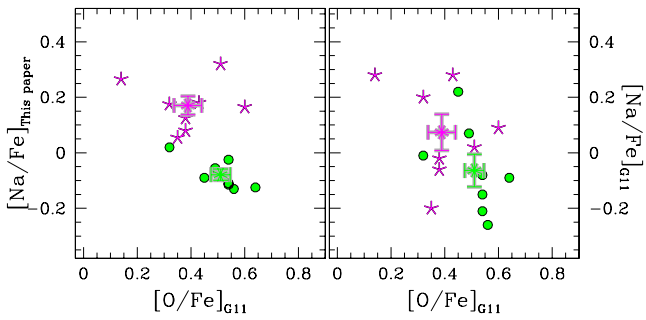


Figure 15. Sodium as a function of oxygen abundance for 17 stars in common with G11. Colours and symbols are the same as in Fig. 11. The $[\text{Na}/\text{Fe}]$ determinations shown in the left- and right-hand panel are from this work and G11, respectively. Superimposed on the data are the average values for Na-rich and Na-poor stars, and the associated errors.

6 HELIUM

Spectral synthesis of the He lines, as explained in Section 4.2, returns the He mass fractions corrected for NLTE effects listed in Table 7. Fig. 16 shows NLTE values for Y as a function of the effective temperature. Most stars with $T_{\text{eff}} < 11\,500$ K are enhanced in He with respect to the primordial value. Stars hotter than 11500 K, for which we assume solar metallicity due to radiative levitation of metals, have $Y \lesssim 0.05$. We attribute this drop in the Y value to helium sedimentation that occurs at these temperatures. We set the primordial Y level (dashed red line) to 0.2516, based on Izotov, Thuan & Stasińska (2007).¹¹ The average He abundance for stars redder than the GJ is $Y = 0.34 \pm 0.01$ (rms = 0.05).

To estimate the internal error associated with the individual measurements of helium contents we re-compute the abundance determinations for a set of BHB stars that span the observed range in temperature, varying the atmospheric parameters by their expected errors (as explained in Section 3.1). The sensitivities of Y to variations in stellar parameters, plus the errors introduced by uncertainties in the continuum placement, for these re-analysed stars are listed in Table 8. By summing in quadrature the various contributions, we estimate the total internal uncertainty in Y as 0.05–0.06, dominated by errors from temperature. Errors in microturbulence and metallicity are negligible, while gravities only marginally contribute to the abundance uncertainty.

Three stars in the BHB sample are consistent with primordial He. However, due to the relatively large internal uncertainty associated with the individual helium measurements, their Y values are still consistent (within 2σ level) with the mean enhanced He derived for the whole sample. Hence, our data do not allow us to establish if these stars truly belong to a population with *normal* helium content. As a check, we determined T_{eff} and $\log g$ by using the photometric catalogues obtained from Advanced Camera for Surveys (ACS)/*HST* (Anderson et al. 2008), coupled with theoretical models, similarly to what was done by using WFI photometry (see Section 3.1). Photometry from space is more accurate but is available for a limited region of ~ 3 arcmin \times 3 arcmin, which includes only six of the BHB stars in our sample. Temperatures and gravities obtained with this photometry are listed in Table 7 (last two columns). Atmospheric parameters from the two photometries are in agreement within expected errors, except for one star (#16837) for which T_{eff} from *HST* photometry is significantly lower. Note

that, for this star, the T_{eff} from *HST* agrees better with the EW of the He line, which lies among the values obtained for colder stars. Adopting a lower T_{eff} would significantly increase the He abundance of this star; the Y of the others would only marginally change.

Despite the relatively large internal uncertainty associated with the individual helium measurements, the statistic of our analysed sample provides some constraints on the Y content of the analysed BHB stars. Indeed, the uncertainty associated with the average Y (of ± 0.01) is relatively small, suggesting that BHB stars in NGC 2808, in the portion of the HB studied here, are on average He enhanced.

To examine the reliability of this result, we investigate how sensitive it is to systematics that may affect the atmospheric parameters. As discussed in Section 3.1, from the theoretical side, we expect that some stars have a $\log g$ lower than that adopted here. These differences in $\log g$, however, are expected to be at most ~ 0.15 dex, and would lower the Y measurements by no more than ~ 0.02 (see Table 8). Note that the Y used to derive the atmospheric parameters is 0.02 lower than the mean value derived here. However, the resulting difference in $\log g$ is relatively small and has a negligible impact on the Y measurements, given our uncertainties. These arguments suggest that, even when we consider systematics in $\log g$, the average Y obtained for the analysed BHB stars is still consistent with being higher than the primordial value.

In comparison to $\log g$, systematics in T_{eff} will cause larger variations in the derived Y abundances. While we do not expect that the assumption on Y systematically changes T_{eff} values (see Section 3.1), there could be a systematic of $\lesssim 200$ K with respect to the T_{eff} scale adopted for the RHB, that is colder. The values listed in Table 8 (systematic), suggest that this systematic uncertainty would increase Y by ~ 0.04 .

In summary, we conclude that our derived mean Y for the analysed BHB stars colder than ~ 11500 K is $Y = 0.34 \pm 0.01 \pm 0.05$ (internal plus systematic uncertainty), higher than the primordial value. This result is robust within both our estimated internal and systematic errors. To shift the mean Y value to a primordial value of $Y = 0.25$ requires either the T_{eff} scale to be underestimated by ~ 400 K or the $\log g$ -scale to be overestimated by ~ 0.60 dex. This result provides a direct confirmation of the predictions by D’Antona et al. (2005) that stellar populations enhanced in helium are present in the BHB of NGC 2808 (see also Piotto et al. 2007; Dalessandro et al. 2011; Milone et al. 2012a). As mentioned above, our observational errors do not allow us to establish if the analysed sample has a uniform content of helium.

As discussed in Section 4.2.2, four stars show a significant drop in the EW of the helium line as a function of T_{eff} , due to helium sedimentation. The helium content of these stars is sensibly low, ranging from $Y \sim 0.03$ to 0.21.

7 DISCUSSION

The presence of multiple stellar groups in NGC 2808 has been inferred along all the evolutionary stages of the CMD, both spectroscopically and photometrically, implying the co-existence of at least three populations of stars. The observed triple MS (Piotto et al. 2007) provides indirect information on the He content, implying the presence of:

- (i) a population defining the red MS characterized by primordial He, hereafter first generation;

¹¹ Izotov et al. (2007) report a helium content of $Y = 0.2472 \pm 0.0012$ or 0.2516 ± 0.0011 , depending on the method used.

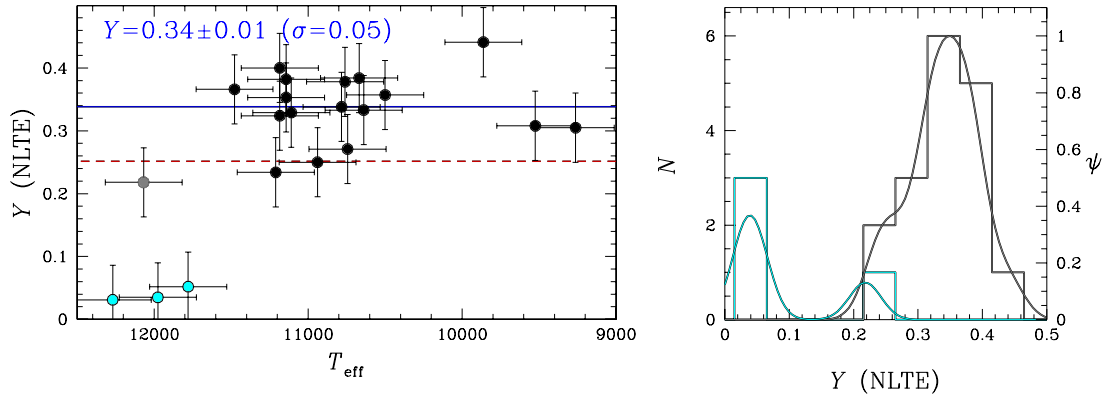


Figure 16. Left-hand panel: helium mass fraction values, corrected for NLTE effects, as a function of effective temperature. The red dashed line corresponds to primordial Y , while the blue solid line is the averaged value for the analysed BHB stars, excluding those hotter than the GJ. Symbols are the same as in Fig. 6. Right-hand panel: histogram and kernel density distribution of the Y mass fractions, for stars colder (black) and warmer (cyan) than GJ.

Table 8. Sensitivity of NLTE helium abundances (mass fractions Y) to variations of atmospheric parameters and continuum setting, exemplarily for several BHB sample stars.

	17132	10342	3030	14968	1169	13305	4129
Internal							
$\Delta T_{\text{eff}} = +250$ K	-0.055	-0.055	-0.054	-0.048	-0.049	-0.013	-0.044
$\Delta \log g = +0.05$ dex	+0.007	+0.008	+0.008	+0.007	+0.009	+0.004	+0.011
$\Delta \xi_t = +1.00$ km s $^{-1}$	-0.002	-0.001	-0.002	-0.002	-0.002	+0.000	-0.002
$\Delta [A/H] = +0.07$ dex	+0.000	+0.000	+0.000	+0.000	+0.000	+0.000	+0.000
σ_{cont}	± 0.015	± 0.015	± 0.015	± 0.015	± 0.015	± 0.015	± 0.015
σ_{total}	0.057	0.058	0.057	0.051	0.052	0.020	0.048
Systematic							
$\Delta T_{\text{eff}} = +200$ K	-0.045	-0.045	-0.043	-0.039	-0.039	-0.010	-0.036
$\Delta \log g = +0.15$ dex	+0.020	+0.022	+0.022	+0.019	+0.025	+0.012	+0.033
σ_{total}	0.049	0.050	0.048	0.043	0.046	0.017	0.049

(ii) an intermediate He population defining the middle MS associated with $Y \sim 0.32$, hereafter intermediate second generation; and

(iii) a highly He-enhanced population distributed on the blue MS with $Y \sim 0.38$, hereafter extreme second generation.

The different He content of the three NGC 2808 populations was first predicted by D’Antona et al. (2005) from their analysis of HB and MS stars. According to this work, stars with different He content populate different HB segments. The distribution of stars along the HB in this cluster shows three significant gaps and distinct segments, RHB, EBT1, EBT2 and EBT3 (Sosin et al. 1997; B00). Consistent with the predictions of D’Antona et al. (2005), stars occupying the RHB of NGC 2808 are associated with first-generation stars, and stars in the EBT1 and EBT2+EBT3 HB segments are representative of the intermediate second generation and the extreme second generation, respectively. The Na–O anticorrelation is also consistent with three groups in oxygen (Carretta et al. 2006), corresponding to three RGBs (Monelli et al. 2013) and providing indirect evidence of He enrichment in NGC 2808. Note that an alternative scenario for the hottest stars of the HB, the EBT3 stars, was proposed by Moehler et al. (2011). They found that the hot end of the HB in ω Centauri (the so-called blue hook) is not solely populated by He-rich stars and suggest that stars in this extremely hot HB segment may be formed as a consequence of independent evolutionary channels (like hot He flashers; see also Cassisi et al. 2009).

To explain the three stellar populations in NGC 2808, D’Ercole et al. (2008, 2010) propose that a highly He-rich and O-poor extreme second generation of stars, corresponding to the blue MS, formed directly from the pure ejecta of a first generation of asymptotic giant branches (AGBs). This happened after the protocluster was cleared from residual gas by SN feedback or by dark remnant accretion (Krause et al. 2012). After this gas expulsion epoch, a cooling flow sets in and brings to the core the low velocity ejecta of AGB stars. Later on, the massive AGB ejecta is mixed with pristine gas (re-accreted in the cluster core), giving origin to the ‘intermediate’ SG stars, with milder O depletion and He enhancement.

According to this picture, our results suggest that the evolution of this cluster could be even more complex than a three population scenario. We find that the RHB is consistent with stars having different Na and little difference in O. If we assume that the RHB stars are the progeny of the red MS stars, our results imply that the red MS is not composed of a single stellar population, but includes stars with some spread in the light element abundances. Interestingly, by analysing high precision *HST* data for the MS of NGC 2808, Milone et al. (2012a) noted that the red MS is broader than expected from observational errors and/or binaries, and its spread in colour at various magnitudes is larger than that associated with the intermediate and blue MS. This broadening in the ($F475W - F814W$) ACS/WFC (Wide Field Camera) filters (Piotto et al. 2007; Milone et al. 2012a) is consistent with a small spread in Y among red MS stars. The mean higher luminosity of Na-rich RHB stars we observe (see Fig. 11) is consistent with a small difference (spread) in the He content in

the RHB of NGC 2808, e.g. He rich and Na-rich RHB stars are brighter.

This observational scenario qualitatively agrees with results on the RHB of the metal-rich GC 47 Tucanae. The HB stars of this cluster ($[\text{Fe}/\text{H}] = -0.72$) are distributed on the red side of the instability strip, consistent with a single segment. However, CN strong stars observed in the GC, which are expected to be enriched in elements produced by high temperature H burning (like Na), are slightly more luminous and consistent with having higher He. The analysis of the MS for this cluster suggests that the maximum variation in Y for different stellar populations is small, at most ~ 0.02 (Milone et al. 2012d).

Our findings on NGC 2808 present a difficult observational puzzle. The RHB, expected to be the progeny of the red MS, hosts stars that underwent an enrichment in Na, along with first-generation Na-poor RHB stars. Given that no metallicity variations have been observed for NGC 2808, the spread on the red MS, observed in photometric filters not seriously affected by molecular bands, may be caused by a small Y spread (Sbordone et al. 2011; Cassisi et al. 2013). This prompts the question of when these mildly polluted stars formed in the evolution of the cluster. Following the D’Ercole et al. scenario, we suppose that they are the latest stars to have been formed from highly diluted material, such that their abundances in light elements and Y approach the primordial values of the first generation. Hence, these stars could have formed after the intermediate second-generation stars, which show evidence for a higher degree of AGB pollution. If this prediction is correct, the abundance pattern of light elements of these stars is dominated by dilution with pristine gas, and the red MS and the RHB contain the first and last stars formed in the cluster.

A bi-modal or broad distribution in the chemical abundances of light elements, such as Na, among RHB stars would provide support to this scenario. Stars belonging to the first generation and the most recently formed stellar population should record the abundances of the interstellar medium from which they were born, and thus, are expected to be different from one another. Some hints of possible bimodality in the Na content of our sample of RHB stars are present, but the relatively poor statistics and internal uncertainties of our data prevent us from making definitive conclusions on the presence of two discrete Na groups or an intrinsically broad distribution. In the latter case, the star formation in the final stages of the cluster evolution may have occurred continuously, not in discrete bursts.

As confirmed by the higher Na content, second-generation(s) stars, which formed from material enriched in the high-temperature H-burning products, populate the BHB. This is analogous to observations of M 4 (Marino et al. 2011b). The He abundances of the hottest BHB stars in our sample suggest that they are enhanced in the He mass fraction Y by ~ 0.09 with respect to the primordial value. This is a strong direct spectroscopic measurement of an enhanced He content for second-generation stars in a GC. Our analysis takes NLTE effects into account, which is an imperative to determine an accurate abundance of He in BHB.

The mean Y found here agrees with the value predicted from models for the middle MS. The extreme population of NGC 2808 (corresponding to the blue MS) is predicted to have even larger He enhancement, up to $Y \sim 0.40$, and is expected to populate the EBT3 segment on the HB (D’Antona et al. 2005; Dalessandro et al. 2011). The He abundances determined in this work are sensitive to systematics in our T_{eff} and $\log g$ scale. We pay close attention to the calibration of our atmospheric parameters, which seem to be correct within ~ 200 K in temperature and ~ 0.20 dex in gravity, and we remind the reader that the He enhancement of BHB stars

in NGC 2808 is valid if our T_{eff} and $\log g$ scales are correct within 400 K and 0.60 dex, respectively.

8 CONCLUSIONS

We have reported chemical abundance analysis for HB stars in the GC NGC 2808. From our study, we conclude that:

- (i) the presence of second-generation(s) helium-enhanced stars is strongly supported by the mean He abundance derived for BHB stars;
- (ii) in NGC 2808, He enriched stars occupy the BHB, corroborating the idea that He is a fundamental parameter for the HB morphology;
- (iii) BHB stars have higher Na than RHB stars;
- (iv) RHB stars show internal variations in Na, with hints of a possible bi-modality, demonstrating that not even this subpopulation is consistent with a single stellar population.

Our findings on NGC 2808 RHB stars fit within the scenario of multiple stellar populations in GCs, where Na-rich/O-poor stars are polluted by processed gas from a previous generation of stars and are representative of the He enhanced population(s).

That the RHB is not consistent with a simple stellar population suggests that NGC 2808 could have experienced a very complex star formation history with more than three stellar generations. This supports photometric studies that show both red and blue MS exhibiting internal $m_{\text{F475W}} - m_{\text{F814W}}$ colour spread, as expected if their stars are not chemically homogeneous (Piotto et al. 2007, Milone et al. 2012a).

Stars with normal O and Na values are representative of the unpolluted He normal population. Variation in helium is expected to manifest as peculiar observed features of the CMD, such as multiple MSs and HB morphology. Empirical detection of He enhancement in GC stars constitutes the essential confirmation of this scenario. Pasquini et al. (2011) found spectroscopic evidence for a possibly high difference in He for two stars on the RGB of NGC 2808. A large difference in He has also been inferred for two RGBs in the peculiar cluster ω Centauri (Dupree & Avrett 2013). This paper provides the first direct spectroscopic measurement of highly ($Y \sim 0.34$) He-enhanced stars in the BHB of a GC, confirming that He-enrichment occurred in GCs, and that He content guides the distribution of stars along the HB.

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SUPPORTING INFORMATION

Additional Supporting Information may be found in the online version of this article:

Table 1. Basic data for the analysed UVES and GIRAFFE stars. Identification numbers and photometric data are from Momany et al. (2004).

Table 3. Atmospheric parameters and chemical abundances for the RHB stars and colder BHB stars observed with GIRAFFE. For Na and Mn we list the line-to-line scatter (<http://mnras.oxfordjournals.org/lookup/suppl/doi:10.1093/mnras/stt1993/-/DC1>).

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