









# Plasma Evolution within an Erupting Coronal Cavity

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## Abstract

Coronal cavities have previously been observed to be associated with long-lived quiescent filaments and are thought to correspond to the associated magnetic flux rope. Although the standard flare model predicts a coronal cavity corresponding to the erupting flux rope, these have only been observed using broadband imaging data, restricting an analysis to the plane-of-sky. We present a unique set of spectroscopic observations of an active region filament seen erupting at the solar limb in the extreme ultraviolet. The cavity erupted and expanded rapidly, with the change in rise phase contemporaneous with an increase in nonthermal electron energy flux of the associated flare. Hot and cool filamentary material was observed to rise with the erupting flux rope, disappearing suddenly as the cavity appeared. Although strongly blueshifted plasma continued to be observed flowing from the apex of the erupting flux rope, this outflow soon ceased. These results indicate that the sudden injection of energy from the flare beneath forced the rapid eruption and expansion of the flux rope, driving strong plasma flows, which resulted in the eruption of an under-dense filamentary flux rope.

*Key words:* Sun: activity – Sun: corona – Sun: filaments, prominences – Sun: flares

*Supporting material:* animation

## 1. Introduction

Solar eruptions are the most energetic and spectacular events that occur in the solar system. However, the processes leading to their initiation and how they subsequently evolve remain areas of interest to the solar physics community. The currently accepted model of a solar eruption (commonly called the “standard flare model”) was originally proposed by Carmichael (1964), Sturrock (1966), Hirayama (1974), and Kopp & Pneuman (1976). This model describes a solar flare as a brightening driven by magnetic reconnection of coronal loops during the eruption of a magnetic flux rope that is subsequently observed in the corona as a coronal mass ejection (CME). While this model has begun to be supplanted by more physically realistic, 3D interpretations (e.g., Janvier et al. 2014), the basic configuration remains the same. The origin of erupting magnetic flux ropes also continues to be a source of investigation, with a debate as to whether they are pre-existing magnetic structures or are formed “on-the-fly” during an eruption (see, e.g., Forbes 2000; Chen 2011; Patsourakos et al. 2013).

As magnetic structures in the solar corona, flux ropes are difficult to observe directly. Their involvement in eruptions is typically inferred by either extrapolating the pre-eruption photospheric magnetic field or via direct measurement of the magnetic field of the associated CME in situ following the eruption. However, the existence of a pre-eruptive flux rope configuration can also be inferred through the observations of dark cavities in white light, extreme ultraviolet (EUV) and X-ray observations when located at or near the solar limb

(cf. Gibson et al. 2006; Habbal et al. 2010; Kucera et al. 2012; Reeves et al. 2012; Karna et al. 2015). These cavities are generally associated with quiescent prominences at the limb, leading to the conclusion that the cavity is a magnetic flux rope supporting cool filamentary material in the lower apex (cf. Régnier et al. 2011; Berger et al. 2012; Forland et al. 2013). Although long-lived, these cavity structures can ultimately become unstable and erupt, producing CMEs observable in white light coronagraph data (e.g., Sterling & Moore 2004; Gibson et al. 2006).

While numerous observations exist of coronal cavities associated with quiescent filaments, coronal cavities associated with eruptive events are much rarer (although some observations have been presented by, e.g., Kleint et al. 2015; McCauley et al. 2015). Limb observations of eruptive events, particularly using emission measure analysis techniques, have revealed magnetic flux ropes with a hot, bright sheath surrounding a cool, dark cavity that contains a hot, bright core (cf. Hannah & Kontar 2013; Kumar & Cho 2014; Lee et al. 2017). These observations have also been used to try and probe the characteristics of the erupting plasma, allowing the density and the plane-of-sky velocity of the plasma to be estimated. Hannah & Kontar (2013) also found that the kinetic energy of the core and the current sheet continued to grow during the eruption, indicating that there was a continuous input of energy. However, spectroscopic observations of erupting coronal cavities, which provide valuable information on how the plasma contained within these cavities is moving in 3D, remain frustratingly rare. This is primarily due to the very small field of view of the available instruments and the long timescales required to obtain the observations, both of which require precise advance knowledge of the location of the eruption.

In this paper, we report on a unique set of spectroscopic observations of a large eruptive event with an associated



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coronal cavity across a broad temperature range. This is the first time that such a cavity eruption has been observed spectroscopically to be erupting from an active region, with the observations providing a unique physical insight into this phenomenon.

## 2. Observations and Data Analysis

The event studied here erupted from the west limb of the Sun on 2017 September 10 and was associated with a *Geostationary Operational Environmental Satellite (GOES)* X8.2 class flare, which began at 15:35 UT and peaked at 16:06 UT. The eruption came from NOAA active region AR 12673 and was one of a series of major flares produced by the active region following emergence. The evolution of the active region and the eruption discussed here were well observed by the Atmospheric Imaging Assembly (AIA; Lemen et al. 2012) onboard the *Solar Dynamics Observatory (SDO)*; Pesnell et al. 2012) spacecraft. However, AR 12673 was a highly energetic active region and as a result was also well observed by multiple space-based and ground-based instruments as it evolved and transited the solar limb.

Following a series of major flares, the active region became the focus of a major flare watch campaign, with regular observations from the Extreme ultraviolet Imaging Spectrometer (EIS; Culhane et al. 2007) onboard the *Hinode* spacecraft (Kosugi et al. 2007). The primary observing plan involved a scanning raster campaign designed to study post-eruption supra-arcade plasma (*Hinode* Observing Plan 244). This campaign, which used EIS study *FlareResponse01*, rastered the 2" slit from right to left across a field of view of  $239'' \times 304''$ , and observed a range of emission lines from He II (at  $\log T = 4.7$ ) to Fe XXIV (at  $\log T = 7.2$ ). As the active region transited the west limb, the campaign was designed to begin off-limb and raster toward the solar disk, taking  $\approx 8$  minutes 52 s to complete each raster. As a result, the erupting cavity described here was only observed by the rasters that began at 15:42:26 UT and 15:51:18 UT, respectively, with both rasters used in this analysis. Before the analysis, the EIS data in each case were first aligned to the *SDO/AIA* field of view to allow a direct comparison between the datasets used.

These narrowband spectroscopic observations from *Hinode/EIS* were complemented using the broadband full-Sun observations from *SDO/AIA*. This allowed the erupting feature to be tracked at high cadence using multiple passbands at multiple temperatures, with the 131, 171, 193, and 211 Å passbands providing the clearest observations of the erupting cavity feature. The data were processed using the standard *aia\_prep.pro* routine in SolarSoftWare, with the data also deconvolved with the relevant point spread functions using the *aia\_deconvolve\_richardsonlucy.pro* routine for the differential emission measure (DEM) analysis outlined in Section 4.

## 3. Results

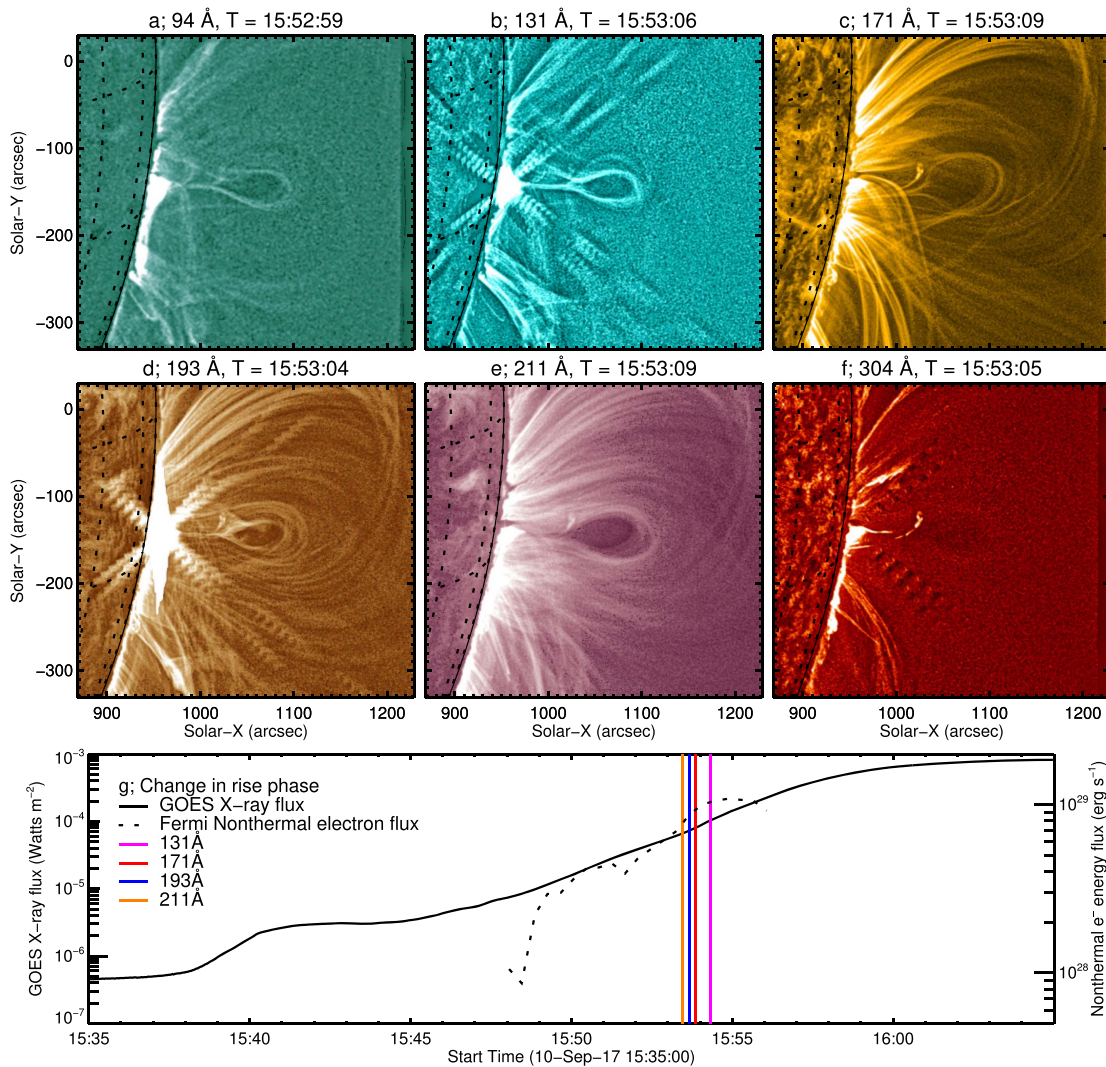
A snapshot of the erupting cavity, as observed by the six EUV passbands onboard *SDO/AIA* at  $T \sim 15:53:06$  UT, is shown in Figure 1. It is clear (particularly from the 94 Å, 131 Å, 171 Å, 193 Å, and 211 Å passbands in panels (a)–(e), respectively) that the eruption appears to be a textbook example of the “standard flare model,” with a thin bright feature (consistent with a current sheet) connecting the bright flare loops below to the dim cavity above. Note that these images

have been processed using the Multiscale Gaussian Normalization (MGN) technique of Morgan & Druckmüller (2014) to highlight the fine structure of the erupting cavity and the associated current sheet. Although both features were observed by *Hinode/EIS*, only the cavity is discussed here, with the evolution of the current sheet discussed by Warren et al. (2017). While the cavity is less clear in the 304 Å passband image shown in panel (f), a thin bright curved structure corresponding to the filamentary plasma is visible at the lower apex of the teardrop-shaped cavity. The evolution of this material is clear in the movie attached to Figure 1; as the cavity forms and rises, bright material can be seen rising with it at the lower apex of the teardrop. This material then disappears in all passbands almost simultaneously at  $\sim 15:51:30$  UT.

The temporal evolution of the cavity was tracked using the 131, 171, 193, and 211 Å passbands by fitting an ellipse to the bright edge of the cavity at each time step in each passband. Although this is shown in Figure 2(a) solely for the 211 Å passband, the evolution of the cavity was comparable in each passband studied. In each case, the data were processed using the MGN technique, with the edge of the cavity then manually identified and fitted using an ellipse. The temporal evolution of both the centroid (denoted by the diamonds) and width (denoted by the cross) of the fitted ellipses are shown in the right column of Figure 2 for each passband. A similar evolution was observed in each passband, with the cavity found to exhibit two distinct rise phases, particularly when considering the evolution of cavity width. The vertical dashed lines in each of the panels on the right-hand panel of Figure 2 indicate the point at which the change in rise phase occurs in each passband. The data were best fitted using two independent quadratic functions, with the resulting fit parameters given in the right-hand side of each panel. These measurements indicate that the cavity initially rose slowly, then more rapidly, which is an observation consistent with previous analyses of the initial stages of solar eruptive events (cf. Régnier et al. 2011; Byrne et al. 2014).

The points at which the rise phase changed in each passband are also shown in panel (g) of Figure 1, which shows the evolution of the soft X-ray flux from the *GOES* spacecraft and the nonthermal electron energy flux derived from the *Fermi* Large-area Telescope (LAT; Atwood et al. 2009). Although the flare was seen by all of the *Fermi/LAT* detectors as a result of its whole Sun field of view, only the NAL04 detector was used here, as it showed the least saturation at the lower energies. The *Fermi* spectra were accumulated over 24 s (because NAL04 was not the most sunward facing detector) and were fitted with a combination of a thermal + thick target function over the energy range 10–100 keV. The nonthermal energy was then calculated at that cadence. It is clear from both Figure 1(g) and the right-hand panel of Figure 2 that the increased upward acceleration of the cavity occurred during the rise phase of the X-ray flare and as the nonthermal electron energy flux measured by *Fermi* approached its peak. This indicates that the upward acceleration of the cavity was driven by the erupting flare below (which is consistent with the “standard” flare model) and suggests that the observed cavity was an erupting flux rope.

Although the cavity was best observed using the images from *SDO/AIA*, it was first observed by *Hinode/EIS*. Due to its west-to-east rastering approach, EIS first observed the cavity between 15:48:00 and 15:49:20 UT (i.e., during the raster that began at 15:42 UT) as it was first developing and again at



**Figure 1.** Eruption from 2017 September 10 at 15:53:06 UT, as observed by *SDO/AIA* in the 94 Å (panel a), 131 Å (panel b), 171 Å (panel c), 193 Å (panel d), 211 Å (panel e), and 304 Å (panel f) passbands. Each image has been processed using the MGN technique of Morgan & Druckmüller (2014) to highlight the fine structure. Panel (g): *GOES* X-ray flux showing the flare (solid black line) and nonthermal electron energy flux from *Fermi* (dashed black line). Vertical colored lines show when the rise phase of the erupting cavity changed in each passband from slow to fast (see Figure 2). Note that the diagonal patterns in panels (b), (d), and (f) emanating from the site of the flare are diffraction patterns and are unphysical. The associated animation shows the restructuring of the active region prior to the eruption and formation of the cavity, which then leaves the field of view as the flare emission starts to saturate the detector.

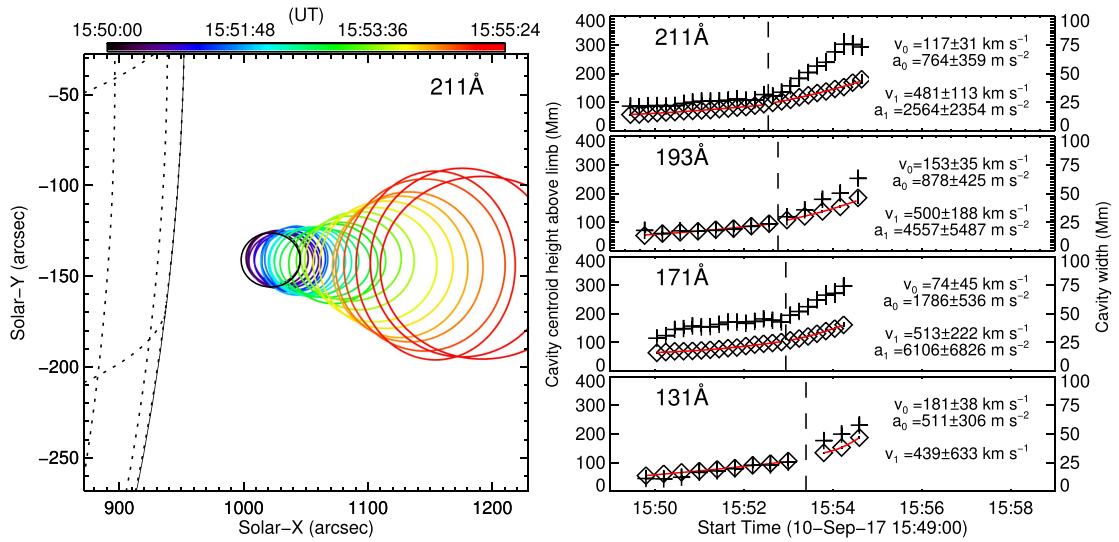
(An animation of this figure is available.)

15:53:20–15:54:40 UT as it erupted (during the raster that began at 15:51 UT). The cavity is most clearly seen in the Fe XV emission line shown in Figure 3 for the raster beginning at 15:42 UT and the Fe XXIV emission line shown in Figure 4 for the raster beginning at 15:51 UT. By comparison, *SDO/AIA* first observed the cavity at  $\approx 15:50:00$  UT (see Figure 2). Although the emission lines observed by EIS are also observed by *SDO/AIA* (cf. O’Dwyer et al. 2010), we believe that the *SDO/AIA* observations of the cavity were masked by the broadband nature of the passbands, with the result that it was not observed until the drop in intensity was sufficiently strong across the wavelength range of the given passband.

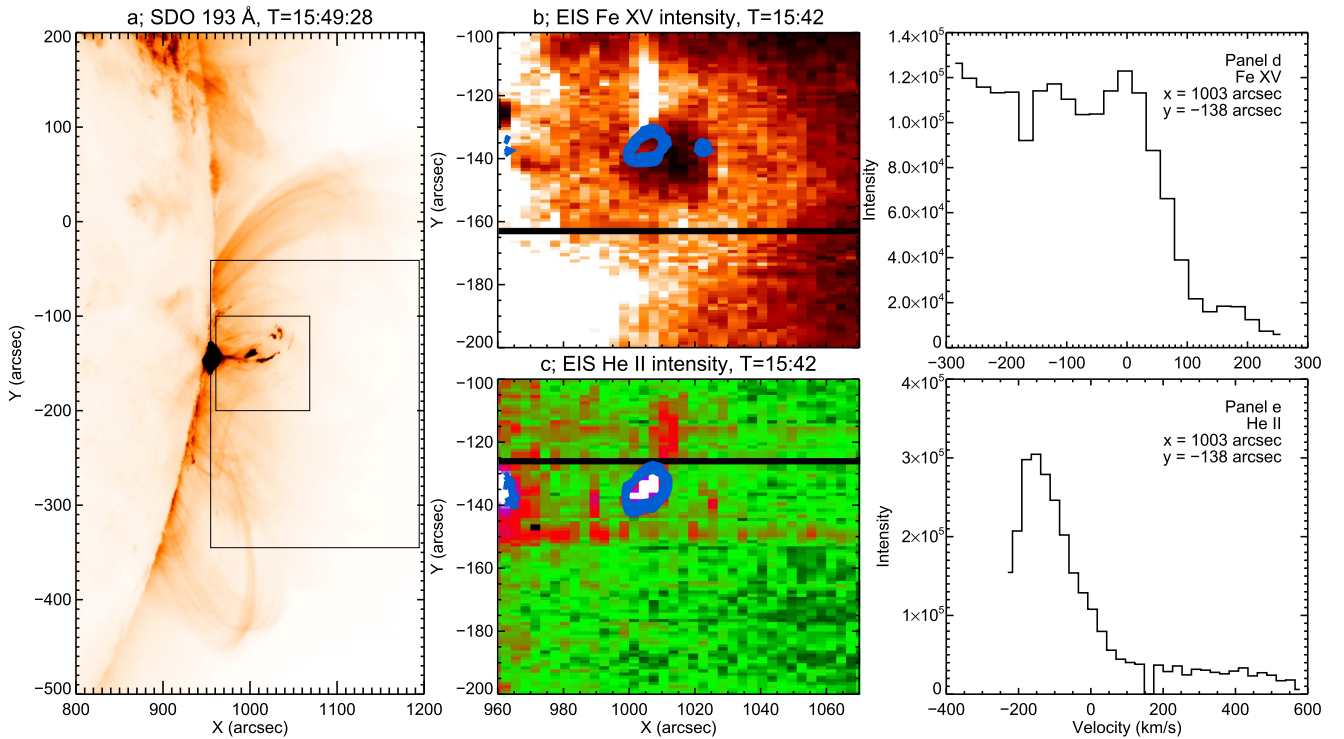
While no clear filamentary material can be observed in the Fe XV image at 15:42 UT shown in panel (b) of Figure 3, strongly blueshifted emission was observed in the region enclosed by the blue contours, corresponding to the bright filament emission observed in He II shown in Figure 3(c). The line spectra observed were incredibly complex, with an

example of the blueshifted emission shown in Figure 3(d). While the Fe XV emission line is the dominant emission line in this portion of the spectrum under normal circumstances, the anomalous nature of the line spectrum shown in panel (d) may suggest that this is not the case here. Other nearby lines that could be contributing include the nearby Al IX emission line at 284.03 Å that corresponds to a  $\log T = 6.1$  and a Fe XVII emission line at 283.9 Å that corresponds to a  $\log T = 6.9$  (cf. Brown et al. 2008; Landi et al. 2013). However, a comparable blueshift to that seen in the Fe XV emission line was also observed in the Fe XVI emission line, indicating that the blueshift is indeed due to plasma motion rather than a contribution from the nearby Al IX and Fe XVII emission lines.

As noted above (and seen in the online movie associated with Figure 1), the filamentary material associated with the erupting cavity can be observed rising at the lower apex of the teardrop-shaped cavity in each *SDO/AIA* passband. It was also clearly observed in the EIS He II data as a bright, very strongly



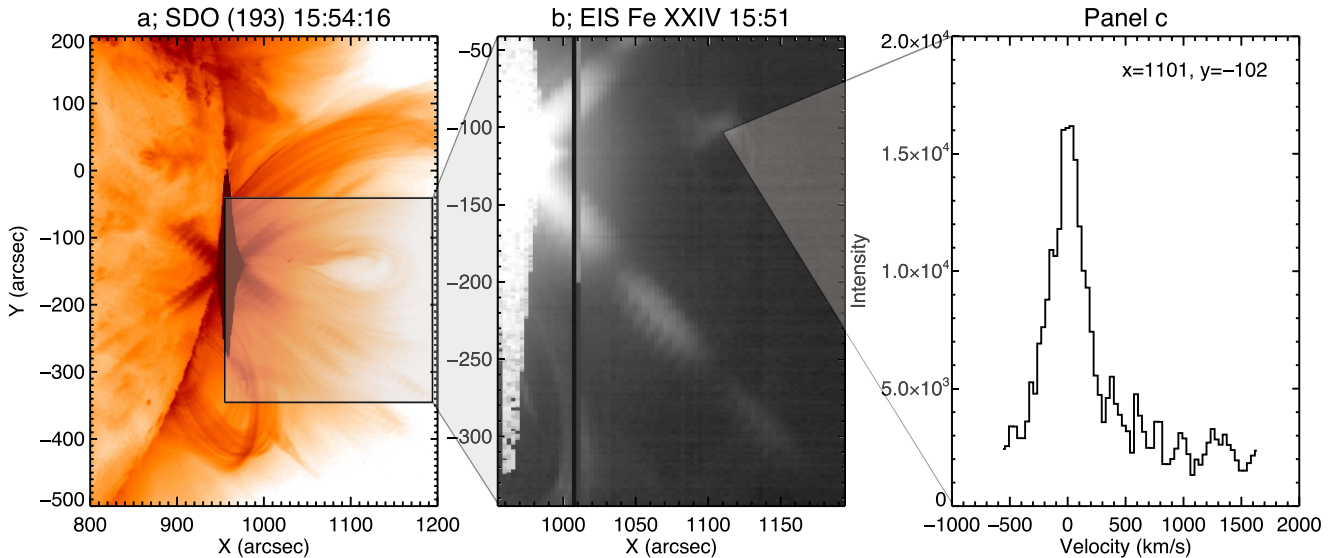
**Figure 2.** Left panel: ellipses fitted to the evolving cavity identified using the 211 Å passband. Right panels: temporal evolution of the height above the limb of the centroid of the ellipse fitted to the cavity (diamonds) and full length of the minor axis of the ellipse fitted to the cavity (crosses) in each passband studied. The height–time evolution was best fitted using two separate quadratic functions that are consistent with two distinct rise phases (i.e., a slow then fast rise phase). The point at which the rise phase changed is indicated by the vertical dashed line and was identified using the variation in ellipse width. Note that the separation of the vertical dashed line corresponding to the 131 Å passband and the missing acceleration value for the second rise phase are due to data gaps as a result of images affected because the Automatic Exposure Control was ignored.



**Figure 3.** Panel (a): SDO/AIA 193 Å image at 15:49:28 UT showing the full *Hinode*/EIS field of view (large solid box) and the field of view used in panels (b) and (c) (small solid box). Panel (b): Fe XV intensity image from *Hinode*/EIS. Panel (c): same as panel (b), but for the He II emission line. Blue contours in both panels (b) and (c) indicate where the plasma is blueshifted by 214 km s $^{-1}$ . Note that the cavity is most apparent in the Fe XV image (panel b), while the filamentary material is most apparent in the He II image (panel c). Panels (d) and (e) show sample Doppler-shifted spectra from the Fe XV and He II emission lines, respectively, illustrating their complex nature. These Doppler-shifted spectra were constructed by converting the spectrum in the pixel defined in the legend from wavelength to velocity space using a rest wavelength estimated by averaging the spectra in a nearby portion of quiet activity.

blueshifted blob (see Figure 3(c) and the spectrum in panel (e)). In fact, the spectra in Figure 3(d) show that some of the spectra were nearly blueshifted out of the spectral window, suggesting Doppler velocities greater than 200 km s $^{-1}$  (consistent with the velocity of 214 km s $^{-1}$  shown by the blue contours in panels (b) and (c)). As a cross-section of the cavity is observed

here, this indicates that as the cavity expands and erupts, the plasma contained within it is flowing rapidly toward the observer and most likely draining down the legs of the flux rope defined by the cavity. There is also some diffuse Fe XXIV emission around and across the cavity, which is consistent with hot, low density plasma in the core of the cavity.



**Figure 4.** Panel (a): *SDO/AIA* 193 Å image at 15:54:16 UT using a reversed color table with the full EIS field of view highlighted. Panel (b): full field of view EIS Fe XXIV intensity image from the raster beginning at 15:51 UT. Panel (c): a sample Fe XXIV profile for the pixel at (1101, -102), as illustrated in panel (b).

The subsequent EIS raster began at 15:51:18 UT, with the corresponding Fe XV image shown in Figure 4(b). At this time, the cavity was observable in all *SDO/AIA* passbands, indicating a significant drop in density and/or temperature compared to the observations at the time of the earlier EIS raster shown in Figure 3. Although the intensity of the Fe XV emission line had dropped too much by this time to provide any usable observations, the Fe XXIV emission line (at  $\text{Log } T = 7.1$ ) shows a bright edge around the cavity (shown in Figure 4(b)). However, the corresponding spectra show that the strong plasma flow had mainly stopped by this time, with the spectra shown in Figure 4(c) showing a plasma velocity peaking at  $0 \text{ km s}^{-1}$ , with some slight broadening of the profile. This indicates that while the significant downward plasma flow associated with the eruption of the flux rope had mostly ceased by this time, there continued to be some draining of hot material from the cavity.

#### 4. DEM Analysis

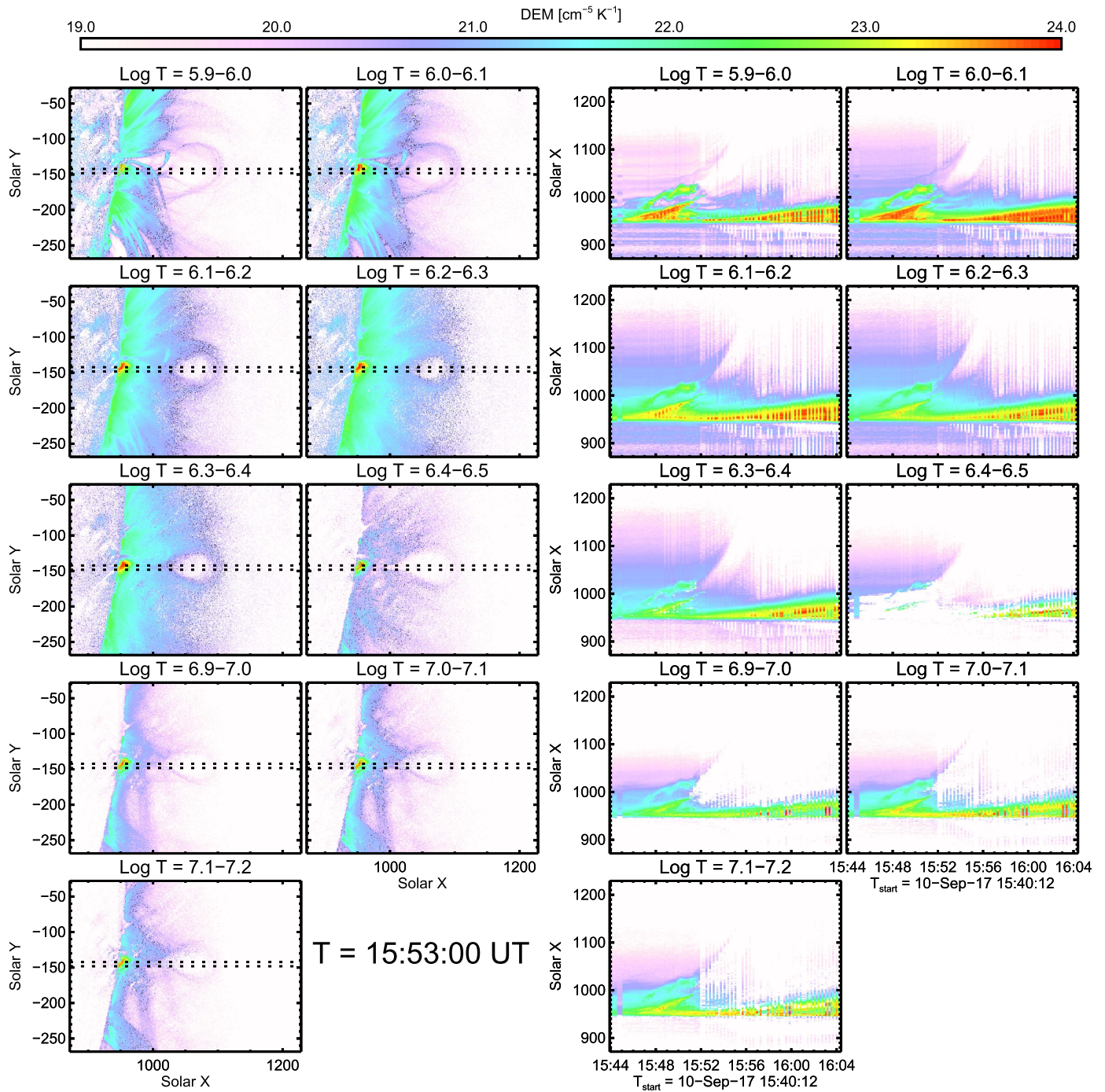
While *Hinode/EIS* gives a unique insight into the evolution and properties of the plasma during the eruption of the cavity, the very low temporal resolution complicates a complete analysis. The high temporal and spatial resolution data provided by *SDO/AIA* was therefore processed using the regularized inversion technique of Hannah & Kontar (2013) to determine the DEM of the plasma within and surrounding the erupting cavity. This offers a complementary dataset, which can be used to explore the properties and behavior of the plasma at much higher cadence. Although the Hannah & Kontar technique was used to produce a DEM with 20 temperature bins of width  $\text{Log } T = 0.1$  from  $\text{Log } T = 5.7\text{--}7.7$ , the cavity was best observed using the temperature bins from  $\text{Log } T = 5.9\text{--}6.5$  and  $\text{Log } T = 6.9\text{--}7.2$  (shown in the left two columns of Figure 5). The clear observations of the erupting cavity in the range  $\text{Log } T = 5.9\text{--}6.5$  match the cavity seen in the Fe XV emission line in Figure 3, while the edge of the cavity observed at  $\text{Log } T = 6.9\text{--}7.2$  is consistent with the hot diffuse emission around the cavity observed in the Fe XXIV emission line at  $\text{Log } T = 7.1$ .

To investigate the temporal evolution of the filament plasma and the erupting cavity, a cut was taken along the black dashed line shown in each panel of the left two columns of Figure 5 and was used to produce stack plots showing the temporal variation in DEM shown in the right two columns of Figure 5. These stack plots show the gradual increase in the height of the filamentary material before it suddenly disappears in all temperature bins at  $\sim 15:50:00$  UT. At this point, a cavity appears in most of the temperature bins before rising out of the field of view by  $\sim 15:58:00$  UT (which is consistent with the cavity evolution shown in the left-hand panel of Figure 2). These stack plots of DEM also indicate that both the disappearing filamentary material and the surrounding envelope had a combination of cool and hot components.

#### 5. Discussion and Conclusions

The combination of spectroscopic EIS and broadband *SDO/AIA* observations presented here describe the very fast eruption of a coronal cavity with a mixed temperature outer sheath and containing a combination of hot and cool plasma that exhibited very strong blueshifts. This is a unique set of observations, and as such requires some examination to determine the physical processes at work.

The temporal evolution of the coronal cavity obtained using the images from *SDO/AIA* indicates that the cavity initially began to rise slowly (with a velocity  $v \sim 74\text{--}181 \text{ km s}^{-1}$ ). However, it then began to rise much more rapidly (at a velocity of  $v \sim 439\text{--}513 \text{ km s}^{-1}$ ), with the *Fermi* nonthermal electron energy flux indicating that this was driven by continuous energy input from the flare below. The strongly blueshifted plasma observed by *Hinode/EIS* suggest that the rapid increase in height of the flux rope driven by the flare forced the plasma from the apex of the flux rope toward the legs, decreasing the density at the apex of the flux rope. As the flux rope initially rose gradually, this drop in plasma density would have been matched by a rise in the plasma temperature, giving the observed combination of hot and cool plasma in the core of the flux rope. However, the subsequent rapid rise of the flux rope and the drop in its density also forced the expansion of the flux rope volume, producing the sudden appearance of the cavity in



**Figure 5.** Left two columns show the DEM calculated from *SDO*/*AIA* EUV observations at  $\sim 15:53:00$  UT. The right two columns show stack plots of the temporal variation in DEM along the black dashed lines shown in the left two columns. Although the DEM analysis produced 20 temperature bins of width  $\text{Log } T = 0.1$  from  $\text{Log } T = 5.7-7.7$ , the cavity is best seen in the temperature bins given here (i.e.,  $\text{Log } T = 5.9-6.5$  and  $\text{Log } T = 6.9-7.2$ ).

all wavelengths. As the cavity erupted and increased in height, the drop in pressure would have stopped the flow of plasma toward the legs of the flux rope, which is consistent with the lack of any clear Doppler motion in the Fe XXV emission line.

While the sudden off-loading of material from an erupting filament or a flux rope has been previously observed (e.g., Jenkins et al. 2017), this is the first time it has been observed spectroscopically for an eruption associated with an active region. These observations indicate that the eruption of the cavity was initially driven by the impulsive phase of the solar flare, with the increases in nonthermal electron energy

matching the changing rise phase of the cavity. This rapid injection of a significant amount of energy forced a dramatic downflow of plasma from the apex of the erupting flux rope and meant that the erupting flux rope was much less dense and exhibited a different structure to other flux rope observations (cf. Hannah & Kontar 2013). These observations are consistent with the “standard flare model” and highlight the vital insight provided by spectrometers such as *Hinode*/EIS.

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### References

- Atwood, W. B., Abdo, A. A., Ackermann, M., et al. 2009, *ApJ*, 697, 1071  
 Berger, T. E., Liu, W., & Low, B. C. 2012, *ApJL*, 758, L37  
 Brown, C. M., Feldman, U., Seely, J. F., Korendyke, C. M., & Hara, H. 2008, *ApJS*, 176, 511  
 Byrne, J. P., Morgan, H., Seaton, D. B., Bain, H. M., & Habbal, S. R. 2014, *SoPh*, 289, 4545  
 Carmichael, H. 1964, *NASSP*, 50, 451  
 Chen, P. F. 2011, *LRSP*, 8, 1  
 Culhane, J. L., Harra, L. K., James, A. M., et al. 2007, *SoPh*, 243, 19  
 Forbes, T. G. 2000, *JGR*, 105, 23153  
 Forland, B. C., Gibson, S. E., Dove, J. B., Rachmeler, L. A., & Fan, Y. 2013, *SoPh*, 288, 603  
 Gibson, S. E., Foster, D., Burkepile, J., de Toma, G., & Stanger, A. 2006, *ApJ*, 641, 590  
 Habbal, S. R., Druckmüller, M., Morgan, H., et al. 2010, *ApJ*, 719, 1362  
 Hannah, I. G., & Kontar, E. P. 2013, *A&A*, 553, A10  
 Hirayama, T. 1974, *SoPh*, 34, 323  
 Janvier, M., Aulanier, G., Bommier, V., et al. 2014, *ApJ*, 788, 60  
 Jenkins, J., Long, D. M., van Driel-Gesztelyi, L., & Carlyle, J. 2017, *SoPh*, 293, 1  
 Karna, N., Pesnell, W. D., & Zhang, J. 2015, *ApJ*, 810, 123  
 Kleint, L., Battaglia, M., Reardon, K., et al. 2015, *ApJ*, 806, 9  
 Kopp, R. A., & Pneuman, G. W. 1976, *SoPh*, 50, 85  
 Kosugi, T., Matsuzaki, K., Sakao, T., et al. 2007, *SoPh*, 243, 3  
 Kucera, T. A., Gibson, S. E., Schmit, D. J., Landi, E., & Tripathi, D. 2012, *ApJ*, 757, 73  
 Kumar, P., & Cho, K.-S. 2014, *A&A*, 572, A83  
 Landi, E., Young, P. R., Dere, K. P., Del Zanna, G., & Mason, H. E. 2013, *ApJ*, 763, 86  
 Lee, J.-Y., Raymond, J. C., Reeves, K. K., Moon, Y.-J., & Kim, K.-S. 2017, *ApJ*, 844, 3  
 Lemen, J. R., Title, A. M., Akin, D. J., et al. 2012, *SoPh*, 275, 17  
 McCauley, P. I., Su, Y. N., Schanche, N., et al. 2015, *SoPh*, 290, 1703  
 Morgan, H., & Druckmüller, M. 2014, *SoPh*, 289, 2945  
 O'Dwyer, B., Del Zanna, G., Mason, H. E., Weber, M. A., & Tripathi, D. 2010, *A&A*, 521, A21  
 Patsourakos, S., Vourlidas, A., & Stenborg, G. 2013, *ApJ*, 764, 125  
 Pesnell, W. D., Thompson, B. J., & Chamberlin, P. C. 2012, *SoPh*, 275, 3  
 Reeves, K. K., Gibson, S. E., Kucera, T. A., Hudson, H. S., & Kano, R. 2012, *ApJ*, 746, 146  
 Régnier, S., Walsh, R. W., & Alexander, C. E. 2011, *A&A*, 533, L1  
 Sterling, A. C., & Moore, R. L. 2004, *ApJ*, 613, 1221  
 Sturrock, P. A. 1966, *Natur*, 211, 697  
 Warren, H. P., Brooks, D. H., Ugarte-Urra, I., et al. 2017, arXiv:1711.10826