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## A DETECTION OF [C II] LINE EMISSION IN THE $Z = 4.7$ QSO BR1202-0725

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### ABSTRACT

We present  $\sim 3''$  resolution imaging of the  $z = 4.7$  QSO BR1202-0725 at  $900 \mu\text{m}$  from the Submillimeter Array. The two submillimeter continuum components are clearly resolved from each other, and the positions are consistent with previous lower frequency images. In addition, we detect [C II] line emission from the northern component. The ratio of [C II] to far-infrared luminosity is 0.04% for the northern component, and an upper limit of  $< 0.03\%$  is obtained for the southern component. These ratios are similar to the low values found in local ultraluminous galaxies, indicating that the excitation conditions are different from those found in local field galaxies. X-ray emission is detected by *Chandra* from the southern component at  $L_{0.5-2\text{keV}} = 3 \times 10^{45} \text{ erg s}^{-1}$ , and detected at 99.6% confidence from the northern component at  $L_{0.5-2\text{keV}} \sim 3 \times 10^{44} \text{ erg s}^{-1}$ , supporting the idea that BR1202-0725 is a pair of interacting galaxies at  $z = 4.7$  that each harbor an active nucleus.

*Subject headings:* galaxies: formation, galaxies: starburst, cosmology: observations, galaxies: high redshift, submillimeter, galaxies: individual (BR1202-0725)

### 1. INTRODUCTION

The far infrared (FIR) fine structure lines of abundant species (e.g. [C I], [C II], [O I]) are recognized as one of the most important coolants in the warm ISM. In particular, the emission from singly ionized carbon (i.e. [C II],  $\lambda_{rest} = 158 \mu\text{m}$ ) is known to trace warm ( $\geq 200$  K) and dense ( $n_{cr} = 3 \times 10^3 \text{ cm}^{-3}$ ) photo-dissociation regions (PDRs; Kaufman et al. 1999), directly identifying the gas surrounding active star forming regions. The [C II] line is the dominant coolant in the Galactic ISM, accounting for 0.1 – 1 % of the FIR luminosity (Hollenbach & Tielens 1997). The brightness of the line makes it an attractive tracer of star formation in galaxies, and a relation holds between [C II] luminosity and the massive star formation rate derived from H $\alpha$  emission for galaxies that have  $L_{\text{FIR}} < 3 \times 10^{10} L_{\odot}$  (Boselli et al. 2002). In the FIR luminous galaxies (ULIRGs), the [C II]-to-FIR flux ratio is an order of magnitude smaller than in normal star forming galaxies (Luhman et al. 1998, 2003). The reason for the deficiency in the [C II] line luminosity is not well understood, but several possible scenarios have been offered (Luhman et al. 1998; Malhotra et al. 2001; Pierini, Leech & Völk 2003; Luhman et al. 2003).

Despite the large line luminosity expected, detecting [C II] line emission from luminous high redshift sources using a single-dish submillimeter telescope has proved difficult (e.g. Isaak et al. 1994; van der Werf 1998; Bolatto et al. 2004; Marsden et al. 2005). The low [C II]-to-FIR ratio suggests that the physical properties in high- $z$  sources resemble those of the local ULIRG population (e.g. Blain et al. 2002). Detection of the [C II]

line in the  $z = 6.4$  QSO J1148+5251 has been reported recently by Maiolino et al. (2005). Here we report a first detection of the [C II] line emission from the  $z = 4.7$  hyper-luminous ( $\sim 10^{13} L_{\odot}$ ) QSO BR1202-0725 observed using the Submillimeter Array (SMA; Ho, Moran & Lo 2004)<sup>5</sup>.

BR1202-0725 is an optically bright QSO ( $M_B = -28.5$ ) with a  $2''.3$  Lyman  $\alpha$  extension toward the northwest (Hu, McMahon & Egami 1996; Petitjean et al. 1996; Fontana et al. 1998; Ohyama, Taniguchi & Shioya 2004). Radio continuum emission was detected at 1.4, 4.9 and 43 GHz (Yun et al. 2000; Carilli et al. 2002). The emission at 1.4 GHz was resolved into two components; one at the location of the QSO and the other  $\sim 4''$  northwest. This apparent companion is also seen in millimeter continuum emission (Omont et al. 1996; Guilloteau et al. 1999). Abundant CO emission ( $M_{H_2} \geq 10^{10} M_{\odot}$ ) was detected toward both the optical QSO and the northwestern companion (Omont et al. 1996; Ohta et al. 1996; Guilloteau 2001; Carilli et al. 2002). The CO-derived redshifts are  $z = 4.6947$  and  $4.6916$  for the QSO and the northwestern companion, respectively (Omont et al. 1996).

We adopt  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ,  $\Omega_m = 0.3$ ,  $\Omega_{\Lambda} = 0.7$ , which gives  $D_L = 43.27 \text{ Gpc}$  ( $1'' = 6.5 \text{ kpc}$ ) for the luminosity distance to BR1202-0725.

### 2. OBSERVATION AND DATA REDUCTION

A total of 5 tracks (Table 1) were obtained toward BR1202-0725, where the total on-source integration time was 19.6 hours. The SIS receivers were tuned to a redshifted [C II] frequency of 333.969 GHz ( $900 \mu\text{m}$ ) (LSB). The SMA correlator had 2 GHz total bandwidth with 0.8 MHz ( $0.7 \text{ km s}^{-1}$ ) spectral resolution. The data were calibrated using the IDL based SMA calibration tool MIR. The quasars 3C279 and 3C273 were observed

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to calibrate the antenna based time dependent gain, and Callisto, 3C279, 3C273, and 1924-292 were used to calibrate the bandpass. The flux scaling is derived assuming Callisto has a total flux density of 15 Jy at 334 GHz, and we estimate a 20% uncertainty in the derived fluxes.

Imaging and analysis of the visibility data were carried out using MIRIAD. The angular resolution achieved by adopting ROBUST = 2 was  $3''.4 \times 2''.7$  (P.A. =  $-8^\circ$ ). The continuum image constructed using the line-free data has an RMS noise of 3.0 mJy. Although the varying S/N across the the SMA bandpass makes it difficult to characterize the overall noise properties, the RMS noise in the [C II] spectrum calculated using several line-free regions of the data cube is estimated to be 7.5 mJy.

### 3. RESULTS

Figure 1 (*left*) presents the 900  $\mu\text{m}$  continuum image of BR1202-0725. The coordinates of the two components are  $\alpha$  (J2000) =  $12^{\text{h}}05^{\text{m}}22.98^{\text{s}}$  and  $\delta$  (J2000) =  $-7^\circ42'30''.0$  for BR1202N, and  $\alpha$  (J2000) =  $12^{\text{h}}05^{\text{m}}23.12^{\text{s}}$  and  $\delta$  (J2000) =  $-7^\circ42'32''.5$  for BR1202S with uncertainties of  $\pm 0''.2$ . The peak of the northern component (BR1202N) is coincident with the CO (2–1) peak (Carilli et al. 2002) within the  $0''.2$  astrometric accuracy. The peak of the southern component (BR1202S) is displaced to the north from the CO (2–1) and the radio continuum peaks (Carilli et al. 2002) by  $\sim 0''.5$ . The location of BR1202S is consistent with the optical/NIR QSO reported by Hu, McMahon & Egami (1996). The derived flux densities from the line-free region of the spectrum are  $S_{900} = 27 \pm 4$  mJy (BR1202N) and  $32 \pm 4$  mJy (BR1202S). BR1202S emits  $54 \pm 9\%$  of the total 900 $\mu\text{m}$  flux, which is slightly lower than the earlier observations at 1.3 mm in which  $65 \pm 10\%$  of the total flux was found in BR1202S (Omont et al. 1996).

A spectrum at the peak of the 900 $\mu\text{m}$  continuum in BR1202N is shown in Figure 1 (*middle*). A similar spectrum toward BR1202S is also shown in Figure 1 (*right*). Continuum subtraction for both spectra was performed by fitting a zeroth order baseline in the line-free channels of the visibilities. The BR1202N spectrum shows a 22 mJy peak at  $z \sim 4.691$  with a line profile that spans the redshift range of  $z = 4.687 - 4.696$ . No significant detection is seen in the BR1202S spectrum. The FWHM of the [C II] line is  $240 \pm 50$  km  $\text{s}^{-1}$ , which is narrower than the CO (5–4) profile (FWHM =  $350 \pm 60$  km  $\text{s}^{-1}$ ) obtained by Omont et al. (1996). The formal significance of the [C II] line detection over the entire line-width shown in Figure 1 is  $4\sigma$ , assuming Gaussian statistics. The true significance of this feature is likely to be higher because noise is non-uniform across the bandpass. The excellent agreement between the [C II] line profile and the published CO (5–4) line profile further supports the robustness of our [CII] line detection. By integrating the emission in the velocity range -230 to 130 km  $\text{s}^{-1}$ , we estimate the [C II] line flux of  $S_{\text{[CII]}} = (6.8 \pm 1.1)$  Jy km  $\text{s}^{-1}$  which translates to a [C II] luminosity of  $L_{\text{[CII]}} = (4.5 \pm 0.7) \times 10^9 L_\odot$ . Using a FIR luminosity of  $L_{\text{FIR}} = (1.2 \pm 0.2) \times 10^{13} L_\odot$  (see §4.2), we obtain  $L_{\text{[CII]}}/L_{\text{FIR}} = (3.8 \pm 0.9) \times 10^{-4}$  (0.04%) in BR1202N. In addition, assuming  $\Delta v = 500$  km  $\text{s}^{-1}$  and  $L_{\text{FIR}} = 2.6 \times 10^{13} L_\odot$  (see §4.2), we obtain a  $3\sigma$  upper limit of  $L_{\text{[CII]}}/L_{\text{FIR}} < 2.8 \times 10^{-4}$  (0.03%) in BR1202S.

## 4. DISCUSSION

### 4.1. Properties of the [C II] Line Emission

The exceedingly low  $L_{\text{[CII]}}/L_{\text{FIR}}$  ratios for both BR1202N and BR1202S are consistent with the similarly low ratios found in local ULIRGs where the average was found to be  $L_{\text{[CII]}}/L_{\text{FIR}} = 0.04\%$  (Luhman et al. 1998, 2003). This is in stark contrast to the values found in the local galaxies by Malhotra et al. (2001) in which  $\sim 70\%$  of the sources show  $L_{\text{[CII]}}/L_{\text{FIR}} > 0.2\%$ , with the majority in the range 0.1 – 1%. Past studies have reported  $L_{\text{[CII]}}/L_{\text{FIR}} \sim 0.02\%$  for the  $z = 6.42$  QSO SDSS J1148+5251 (Maiolino et al. 2005), and  $L_{\text{[CII]}}/L_{\text{FIR}} < 0.4\%$  in a  $z = 4.926$  QSO CL 1358+62 (Marsden et al. 2005). Our new measurement suggests that the [C II]-to-FIR ratio (0.04%) in the high- $z$  hyperluminous source BR1202-0725 is comparable to local ULIRGs. Table 2 summarizes these results.

The integrated [C II] line intensity is known to be a tracer of star formation activity in normal galaxies. We derive the star formation rate (SFR) of 2900  $M_\odot \text{ yr}^{-1}$  using the calibration of Maiolino et al. (2005). This is 1 – 2 orders of magnitude higher than the SFRs derived from optical recombination lines observed near BR1202N (i.e. 10 – 230  $M_\odot \text{ yr}^{-1}$ ; Hu, McMahon & Egami 1996; Petitjean et al. 1996; Ohta et al. 2000; Ohya, Taniguchi & Shioya 2004), but consistent with the SFR derived from the SED fit to the millimeter/submillimeter continuum measurements (i.e. SFR = 2000  $M_\odot \text{ yr}^{-1}$  in BR1202N; see §4.2).

Another important diagnostic of the physical properties is the [C II]-CO flux ratio. Early studies suggest the close spatial and kinematical association between the [C II] and CO (1–0) emissions in nearby galaxies, suggesting that this ratio is a good tracer of star formation activity that is independent of the beam filling factor (Crawford et al. 1985; Stacey et al. 1991). Assuming the luminosity ratio between the CO (5–4) and CO (1–0) lines in BR1202N is similar to that of the Cloverleaf (Barvainis et al. 1994, 1997), the CO (1–0) luminosity is estimated to be  $L_{\text{CO(1-0)}} \sim 10^6 L_\odot$ , yielding  $L_{\text{[CII]}}/L_{\text{CO(1-0)}} \sim 4500$ . This value is highly uncertain since the CO (5–4) – CO (1–0) ratio may be different in BR1202-0725. Nevertheless, this is a factor of 2.5 larger than the average values of ULIRGs (i.e.  $L_{\text{[CII]}}/L_{\text{CO(1-0)}} \sim 1700$ ; van der Werf 1998), but comparable to the average values of normal galaxies (i.e.  $L_{\text{[CII]}}/L_{\text{CO(1-0)}} = 4400 - 6300$ ; Crawford et al. 1985; Stacey et al. 1991). The higher observed ratio may suggest a factor of few higher  $G_0/n$  ratio (where  $G_0$  and  $n$  are the far UV flux and the cloud density respectively) in BR1202N compared to the local ULIRGs (see Figure 9 of Kaufman et al. 1999).

### 4.2. Properties of the Continuum Emission

The spatially resolved submillimeter images allow us to investigate the properties of the individual components in BR1202-0725. The radio to FIR SED of BR1202-0725 is shown in Figure 2, where the individual flux densities are given for BR1202N and BR1202S for the 1.4 GHz (Carilli et al. 2002), 1.3 mm (Omont et al. 1996) and 900 $\mu\text{m}$  observations (see figure caption for the adopted parameters). The SEDs of the two sources are ana-

lyzed using the model by Yun & Carilli (2002), yielding  $L_{\text{FIR}} = 1.2 \times 10^{13}$  and  $2.6 \times 10^{13} L_{\odot}$ , and  $\text{SFR} = 2000$  and  $4500 M_{\odot} \text{ yr}^{-1}$  for BR1202N and BR1202S respectively. The discrepancy between the SFR derived by us and Yun & Carilli (2002) ( $2300 M_{\odot} \text{ yr}^{-1}$  for the entire BR1202-0725) is due to the difference in the adopted cosmology. Note that while the millimeter/submillimeter continuum flux densities of BR1202N and BR1202S are comparable (§3), the radio continuum flux density of BR1202N is 3 times larger. This suggests a presence of a highly obscured AGN in BR1202N (see §4.3), or a strong effect of a highly energetic jet from BR1202S (Klamer et al. 2004).

Assuming a dust absorption coefficient of  $\kappa_d = 1.5 \text{ m}^2 \text{ kg}^{-1}$  at  $160 \mu\text{m}$  (Hildebrand 1983; Draine & Lee 1984), the dust mass of the individual components is estimated to be  $M_d$  (BR1202N) =  $(5.4 \pm 0.8) \times 10^8 M_{\odot}$  and  $M_d$  (BR1202S) =  $(9.2 \pm 1.1) \times 10^8 M_{\odot}$ . Using the  $\text{H}_2$  mass derived from the CO (5-4) emission (Omont et al. 1996), we find  $M_{\text{H}_2}/M_d \sim 50 \pm 10$  for both components. This ratio is an order of magnitude lower than the mean value for LIRGs/ULIRGs ( $M_{\text{H}_2}/M_d = 540 \pm 290$ ; Sanders, Scoville & Soifer 1991), nearby spiral galaxies<sup>6</sup> ( $M_{\text{H}_2}/M_d \sim 500$ ; see Fig. 2b of Devereux & Young 1990) and a sample of *Spitzer*  $160 \mu\text{m}$  selected galaxies ( $M_{\text{H}_2}/M_d = 330$ ; Yun et al. 2006). These gas-to-dust ratios are highly uncertain because dust emissivity, temperature, and CO-to- $\text{H}_2$  conversion factor are all poorly constrained. If the low ratio for BR1202-0725 is correct, however, it may indicate that the  $\text{H}_2$  mass is significantly underestimated in BR1202-0725, and/or the warm component of dust dominates the far infrared emission, suggesting that the entire ISM is involved in the starburst activity.

#### 4.3. BR1202-0725: A Colliding Galaxy System at $z = 4.7$

Our analysis of previously unpublished archival *Chandra* data (ObsID 3025) clearly shows that BR1202S is an X-ray source and gives 26.4 net counts in 9.64 ksec. The derived flux from BR1202S is  $f_{0.5-2\text{keV}} = 6 \times 10^{-15} \text{ erg cm}^{-2} \text{ s}^{-1}$  which yields a luminosity of  $L_{0.5-2\text{keV}} = 3 \times 10^{45} \text{ erg s}^{-1}$  (assuming  $\Gamma = 2.0$  and Galactic foreground  $N_H = 3.35 \times 10^{20} \text{ cm}^{-2}$ ; Stark et al. 1992). X-ray emission is also seen within  $1''$  of the position of BR1202N. With just 2 counts

<sup>6</sup> The CO to  $\text{H}_2$  conversion factor used by Omont et al. (1996), Sanders, Scoville & Soifer (1991) and Devereux & Young (1990) are all consistent to within 20%. In addition, the rest frame  $160 \mu\text{m}$

in a single pixel, this detection seems marginal, but the local background of 0.03 counts/pixel implies that the likelihood of a random occurrence is only 0.04%. The derived flux from BR1202N is  $f_{0.5-2\text{keV}} = 4.8_{-4.3}^{+6.5} \times 10^{-16} \text{ erg cm}^{-2} \text{ s}^{-1}$  (95% confidence; Regener 1951), implying  $L_{0.5-2\text{keV}} \sim 3 \times 10^{44} \text{ erg s}^{-1}$ .

If the X-ray emission from BR1202N is real, its high X-ray luminosity suggests that BR1202N also hosts a luminous AGN. In addition, our new [C II] line detection, large dust emission, and the presence of abundant molecular gas suggest that BR1202N is forming stars at a high rate. BR1202S and BR1202N thus form a system of two massive galaxies undergoing an interaction or a merger, separated by a projected distance of  $\sim 25$  kpc, observed when the universe was only 1.2 Gyrs old. The presence of a highly obscured AGN (and an optically bright QSO) during a massive merger is consistent with the scenario predicted in recent galaxy collision simulations that investigate the evolution of a central black hole (e.g. Hopkins et al. 2005).

A high degree of clustering near massive galaxies is expected in the theoretical models of large scale structure formation (e.g. White & Rees 1978). This is supported by observations such as the higher surface density of submillimeter sources in the fields surrounding high- $z$  radio galaxies (Stevens et al. 2003). Although an earlier study by Giallongo et al. (1998) found little evidence for clustering of galaxies near BR1202-0725 with  $R' \leq 25$ , two massive galaxies in proximity of each other makes BR1202-0725 a particularly interesting test ground for the hierarchical scenarios, soon after the reionization.

## 5. SUMMARY

We present a  $3''$  resolution  $900 \mu\text{m}$  continuum image and a detection of the redshifted [C II] line emission from the  $z = 4.7$  QSO BR 1202-0725 obtained using the SMA. This is one of the first detections of the [C II] line from a high redshift source. The [C II] line is associated with BR1202N. The low [C II]-FIR ratio of  $\sim 3.8 \times 10^{-4}$  is similar to local ULIRGs. X-ray emission is clearly detected from BR1202S, and at 99.6% confidence from BR1202N, suggesting that BR1202-0725 is the first example of a pair of AGN hosts at  $z \sim 4.7$ .

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emission from BR1202-0725 traces the same warm dust (and hence the same  $M_d$ ) as measured by *IRAS* (Yun et al. 2006).

## REFERENCES

- Barvainis, R., Tacconi, L., Antonucci, R., Alloin, D., & Coleman, P. 1994, *Nature*, 371, 586  
 Barvainis, R., Maloney, P., Antonucci, R., & Alloin, D., 1997, *ApJ*, 484, 695  
 Blain, A. W., Smail, I., Ivison, R. J., Kneib, J. -P. & Frayer, D. T. 2002, *PhR*, 369, 111  
 Bolatto, A. D., Di Francesco, J. & Willott, C. J. 2004, *ApJL*, 606, L101  
 Boselli, A., Gavazzi, G., Lequeux, J., & Pierini, D. 2002, *A&A*, 385, 454  
 Carilli, C. L. et al. 2002, *AJ*, 123, 1838  
 Crawford, M. K., Genzel, R., Townes, C. H., & Watson, D. M. 1985, *ApJ*, 291, 75  
 Devereux, N. A. & Young, J. S. 1990, *ApJ*, 359, 42  
 Draine, B. T. & Lee, H. M. 1984, *ApJ*, 285, 89  
 Fontana, A., D'Odorico, S., Giallongo, E., Cristiani, S., Monnet, G., & Petitjean, P. 1998, *AJ*, 115, 1225  
 Giallongo, E., D'Odorico, S., Fontana, A., Cristiani, S., Egami, E., Hu, E., & McMahon, R. G., *AJ*, 115, 2169  
 Guilloteau, S., Omont, A., Cox, P., McMahon, R. G. & Petitjean, P. 1999, *A&A*, 349, 363  
 Guilloteau, S. 2001, in *ASP Conf. Ser. 235, Science with the Atacama Large Millimeter Array*, ed. A. Wootten (San Francisco:ASP)  
 Hildebrand, R. H. 1983, *QJRAS*, 24, 267  
 Ho, P. T. P., Moran, J. M. & Lo, Kwok Yung 2004, *ApJ*, 616, 1  
 Hopkins, P. F., Hernquist, L., Cox, T. J., Di Matteo, T., Martini, P., Robertson, B., & Springel, V. 2005, *ApJ*, 630, 705

- Hu, E. M., McMahon, R. G. & Egami, E. 1996, *ApJL*, 459, L53  
Hollenbach, D. J. & Tielens, G. G. M., 1997, *ARAA*, 35, 179  
Isaak, K. G., McMahon, R. G., Hills, R. E. & Withington, S. 1994, *MNRAS*, 269, L28  
Kaufman, M. J., Wolfire, M. G., Hollenbach, D. J. & Luhman, M. L., 1999, *ApJ*, 527, 795  
Klamer, I. J., Ekers, R. D., Sadler, E. M. & Hunstead, R. W. 2004, *ApJL*, 612, L97  
Luhman, M. L. et al. 1998, *ApJL*, 504, L11  
Luhman, M. L. et al. 2003, *ApJ*, 594, 758  
Maiolino, R., Cox, P., Caselli, P., Beelen, A., Bertoldi, F. et al. 2005, *A&A*, 480, L51  
Marsden, G., Borys, C., Chapman, S. C., Halpern, M. & Scott, D. 2005, *MNRAS*, 359, 43  
Malhotra et al. 2001, *ApJ*, 561, 766  
Ohyama, Y., Taniguchi, Y. & Shioya, Y. 2004, *AJ*, 128, 2704  
Ohta, K., Yamada, T., Nakanishi, K., Kohno, K., Akiyama, M. & Kawabe, R. 1996, *Nature*, 382, 426  
Ohta, K. et al. 2000, *PASJ*, 52, 557  
Omont, A., Petitjean, P., Guilloteau, S., McMahon, R. G., Solomon, P. M. & Pecontal, E. 1996, *Nature*, 382, 428  
Petitjean, P., Pecontal, E., Valls-Gabaud, D. & Charlot, S. 1996, *Nature*, 380, 411  
Pierini, D., Leech, K. J., & Völk, H. J. 2003, *A&A*, 397, 871  
Regener, V. H. 1951, *PhRv*, 84, 161  
Sanders, D. B., Scoville, N. Z., & Soifer, B. T. 1991, *ApJ*, 370, 158  
Stacey, G. J., Geis, N., Genzel, R., Lugten, J. B., Poglitsch, A., Sternberg, A., & Townes, C. H. 1991, *ApJ*, 373, 444  
Stevens, J. A. et al. 2003, *Nature*, 425, 264  
Stark, A. A., Gammie, C. F., Wilson, R. W., Bally, J., Linke, R. A., Heiles, C. & Hurwitz, M. 1992, *ApJS*, 79, 77  
van der Werf, P. P. 1998, in *ASP Conf. Ser. 156, Highly Redshifted Radio Lines*, ed. C. L. Carilli, S. J. E. Radford, K. M. Menten, & G. I. Langston (San Francisco: ASP), 91.  
White, S. D. M. & Rees, M. J. 1978, *MNRAS*, 183, 341  
Yun, M. S., Carilli, C. L., Kawabe, R., Tutui, Y., Kohno, K. & Ohta, K. 2000, *ApJ*, 528, 171  
Yun, M. S., & Carilli, C. L. 2002, *ApJ*, 568, 88  
Yun, M. S., et al. 2006, in prep.

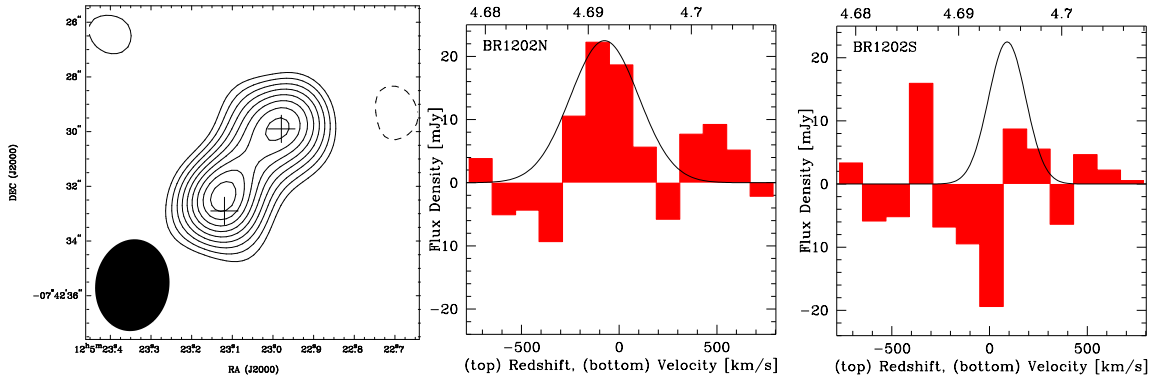


FIG. 1.— (*Left*) 900 $\mu\text{m}$  continuum map of BR1202-0725 made by averaging line free channels in the LSB. The crosses mark the locations of the CO (2–1) peaks from Carilli et al. (2002). The contour levels are  $-2, 3, 4, 5, 6, 7, 8, 9, 10, 11\sigma$  ( $1\sigma = 3.0 \text{ mJy beam}^{-1}$ ). The beam is shown in the lower left corner. The SMA line profile of BR1202N (*middle*) and BR1202S (*right*) obtained at the peak pixels using  $120 \text{ km s}^{-1}$  velocity averaging. Velocity of  $0 \text{ km s}^{-1}$  corresponds to 334 GHz. The solid lines are schematic representations of the central velocities and linewidths of the CO (5–4) line observed in BR1202N and BR1202S by Omont et al. (1996), where the peaks have been scaled to match the [C II] line intensity in BR1202N.

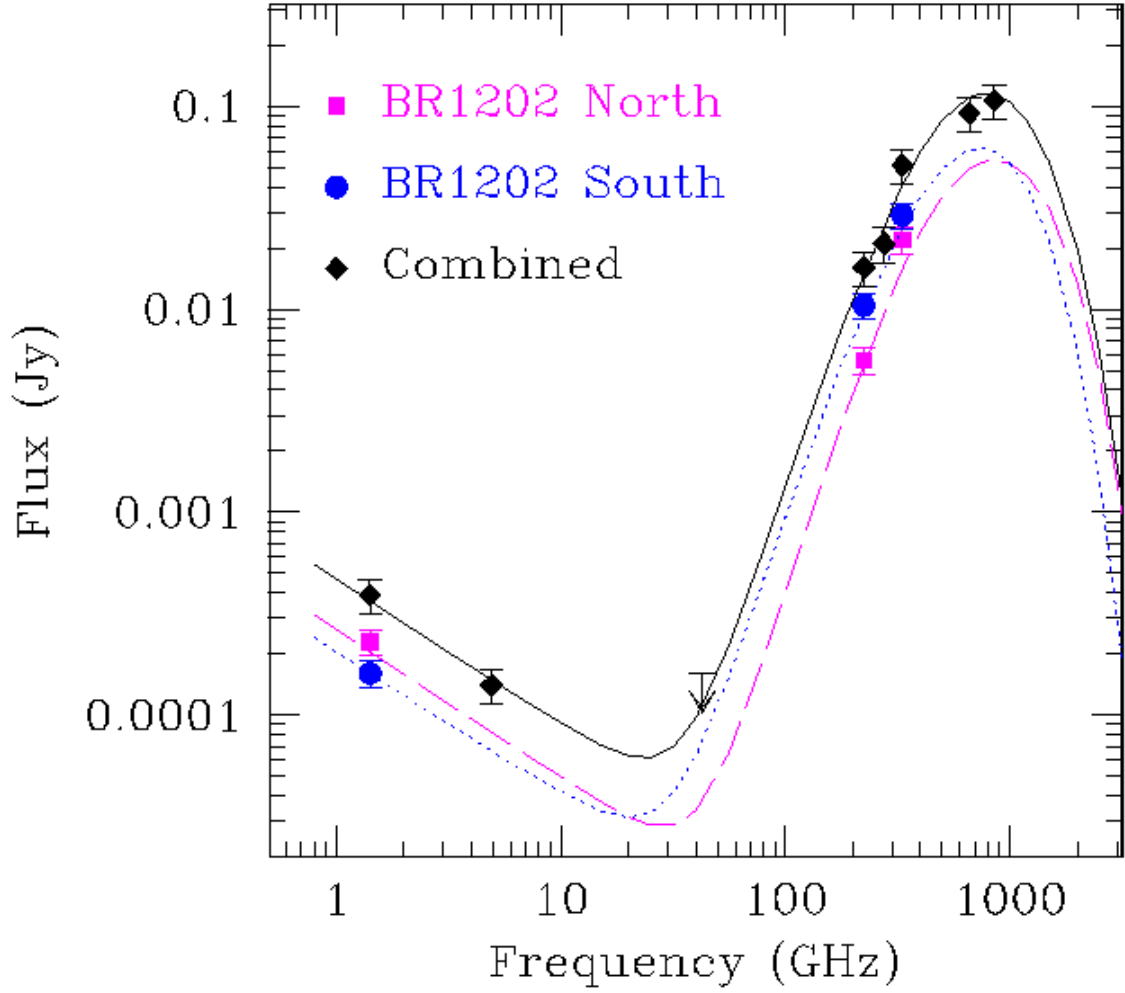


FIG. 2.— The radio to FIR SED of BR1202-0725. Two values are given when BR1202N and BR1202S are spatially resolved. The solid dark line represents the fit to the SED of the entire BR1202-0725 system, while the SEDs are not unique fits for BR1202N and BR1202S. The parameters used for the SEDs are  $T_d = 80\text{K}$ ,  $\alpha = 0.75$ ,  $\beta = 1.75$ ,  $f_{nth} = 6.0$  and  $\text{SFR} = 2000 \text{ M}_\odot \text{ yr}^{-1}$  for BR1202N and  $T_d = 65\text{K}$ ,  $\alpha = 0.75$ ,  $\beta = 1.50$ ,  $f_{nth} = 2.0$  and  $\text{SFR} = 4500 \text{ M}_\odot \text{ yr}^{-1}$  for BR1202S, where  $\alpha$ ,  $\beta$ , and  $f_{nth}$  are the synchrotron spectral index, dust emissivity index, and the normalization factor to account for the non-thermal synchrotron emission (see Yun & Carilli 2002, for details).



TABLE 1  
SUMMARY OF SMA OBSERVATIONS

Date	$N_{ants}^1$	$t_{source}^2$	$T_{sys}^3$	RMS phase <sup>4</sup>
02/22/2005	5	5.3	240	5
04/05/2005	6	5.2	310	15
03/21/2006	7	1.3	300	14
04/19/2006	8	4.5	250	6
05/11/2006	5	3.3	300	11

<sup>1</sup>Number of SMA antennas used.

<sup>2</sup>On source integration time in hours.

<sup>3</sup>The average DSB system temperature in Kelvins.

<sup>4</sup>The average RMS phase (in degrees) taken over all baselines for each track. These are calculated toward 3C279 for the first two tracks and 3C273 for the three tracks taken in 2006.

TABLE 2  
SUMMARY OF LUMINOSITY RATIOS

Source	$L_{\text{FIR}} (L_{\odot})$	$L_{[\text{CII}]} / L_{\text{FIR}}$	Ref.
Normal Galaxies	$10^7 - 10^{11}$	0.1 – 1%	1
ULIRGs	$\sim 10^{12}$	0.01 – 0.1%	2
High- $z$ Sources			
BR1202N ( $z = 4.7$ )	$1.2 \times 10^{13}$	0.04%	3
BR1202S ( $z = 4.7$ )	$2.6 \times 10^{13}$	$< 0.03\%$	3
SDSS J1148+5251 ( $z = 6.4$ )	$1.2 \times 10^{13}$	0.02%	4
CL 1358+62 ( $z = 4.9$ )	$2.4 \times 10^{12}$	$< 0.4\%$	5

REFERENCES. — (1) Malhotra et al. (2001); (2) Luhman et al. (2003); (3) this work; (4) Maiolino et al. (2005); (5) Marsden et al. (2005).