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A GBT SURVEY FOR H I 21 CM ABSORPTION IN THE DISKS AND HALOS OF LOW-REDSHIFT GALAXIES¹

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ABSTRACT

We present a H I 21 cm absorption survey with the *Green Bank Telescope* (GBT) of galaxy-quasar pairs selected by combining galaxy data from the *Sloan Digital Sky Survey* (SDSS) and radio sources from the *Faint Images of the Radio Sky at Twenty-Centimeters* (FIRST) survey. Our sample consists of 23 sightlines through 15 low-redshift foreground galaxy - background quasar pairs with impact parameters ranging from 1.7 kpc up to 86.7 kpc. We detected one absorber in the GBT survey from the foreground dwarf galaxy, GQ1042+0747, at an impact parameter of 1.7 kpc and another possible absorber in our follow-up Very Large Array (VLA) imaging of the nearby foreground galaxy UGC 7408. Both of the absorbers are narrow (FWHM of 3.6 and 4.8 km s⁻¹), have sub Damped Lyman α column densities, and most likely originate in the disk gas of the foreground galaxies. We also detected H I emission from three foreground galaxies including UGC 7408. Although our sample contains both blue and red galaxies, the two H I absorbers as well as the H I emissions are associated with blue galaxies. We discuss the physical conditions in the 21 cm absorbers and some drawbacks of the large GBT beam for this type of survey.

Subject headings: galaxies: abundances — galaxies: ISM — quasars: absorption lines

1. INTRODUCTION

Inflow and outflow of gas between the intergalactic medium (IGM) and the disks of galaxies is a crucial aspect of galaxy evolution. In the last decade our theoretical understanding on how galaxies acquire gas has advanced considerably. Traditionally, galaxies were believed to accrete via the “hot” mode (White & Rees 1978; White & Frenk 1991), where accreting gas shock-heats to the virial temperature and then all halo gas within a “cooling radius” cools to form stars. Recent hydrodynamic cosmological simulations have discovered that gas accretion also occurs in the so called “cold” mode where the gas is able to cool as it descends into galactic potentials and thus never approaches the virial temperature (see Birnboim & Dekel 2003; Kereš et al.

2005). Unlike the hot-mode, the cold-mode accretion occurs through long filaments connecting the extended disks and halos of galaxies to their surrounding intergalactic medium (IGM). Although initial results by Kereš et al. (2005) showed that the cold mode accretion is dominant in galaxies with lower masses ($< 10^{10.3} M_{\odot}$) at redshift $z = 0$, recent simulations by Kereš et al. (2009) have shown that cold-mode accretion also persists even in massive galaxies with hot envelopes. In addition to filamentary accretion of cool gas from the intergalactic medium, cool gas is also expected to condense out of the hot galactic halos. Maller & Bullock (2004) argued that small clouds of cool gas condense out of hot gaseous halos through a multiphase process due to thermal instabilities. These halo clouds are predicted to have a characteristic size of ≈ 1 kpc and are expected to be another continuous source of cool gas for galaxies.

On the other hand, a substantial amount of material is ejected from galactic disks into their halos via galactic winds (see the review by Veilleux et al. 2005). Evidence of outflows has been detected in low-ionization metal absorption lines of Na I (e.g. Schwartz & Martin 2004;

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Rupke et al. 2005a,b) and Mg II (Tremonti et al. 2007; Rubin et al. 2010). Rubin et al. (2010) found that the strength of the outflow correlates with the absolute star-formation rate (SFR) and claimed that the mass outflow rates were of the order of the SFRs for their sample. Assuming a negligible halo drag, Rupke et al. (2005b) found that only 5-10% of the neutral material in starburst-driven winds escape into the IGM. In other words, the wind speeds are not larger than the galactic escape velocity for most galaxies and thus the outflowing material is primarily retained in the halos of galaxies. According to the Galactic Fountain model (Shapiro & Field 1976), the baryonic material deposited by the winds can eventually cool down and rain back into the galaxies. This model and subsequent work by Bregman (1980) predicts the existence of high-velocity clouds (HVCs) as a by-product of the cooling of the hot gas.

Besides inflow and outflow signatures, there is growing evidence of the extended nature of the galactic disks. Mg II absorbers, which trace cool low-ionization gas, have been believed to be associated with extended disks (Wagoner 1967; Bowen et al. 1995; Charlton & Churchill 1996). Observations of sightlines probing foreground galaxies at impact parameters up to $75 h^{-1}$ kpc with inclination angles between 40° to 75° have confirmed Mg II absorbers exhibiting disk-like rotation (Steidel et al. 2002; Kacprzak et al. 2010). Steidel et al. (2002) found that four of five background QSOs that lie to one side of a foreground galaxy exhibit Mg II absorption from the latter that are shifted in the same sense from the systemic redshift of the foreground galaxy. They concluded that the kinematics of their sample could be understood in terms of a disk-like rotation combined with a lagging halo. This was confirmed in a recent study by Kacprzak et al. (2010), where most of the Mg II absorption systems in their sample fully reside to one side of the galaxy systemic velocity. They also found a close coupling between galaxy inclination and Mg II absorption velocity spread. While it is clear that there is cool gas at these distances (> 100 kpc for $h = 0.7 \text{ km s}^{-1} \text{ Mpc}^{-1}$), it is unclear how much baryonic mass exists in these systems. Part of the problem is that these observations do not probe the hydrogen directly, and the coupling of neutral gas and metals is not well understood.

To properly characterize the total baryonic content in the extended disks as well as gaseous halos, H I observations must be made at a wide range of distances from the galaxies. Absorption spectroscopy is a powerful tool for probing low-column density systems with high sensitivity independent of the redshift of the absorbing material. Unfortunately, the Lyman α transition for $z < 1.7$ galaxies occurs in the ultraviolet (UV) band and thus ground-based observations of these systems are not possible. However, the H I 21 cm hyperfine transition for low- z systems can be observed with existing radio facilities. Various authors (e.g., Kanekar & Chengalur 1997; Lane & Briggs 2001; Vermeulen et al. 2003; Darling et al. 2004; Keeney et al. 2005; Gupta et al. 2007, 2009, 2010) have probed neutral gas in extragalactic sources at various redshifts using 21 cm absorption. In their recent study of the radio galaxy Centaurus A, Struve et al. (2010) discovered an H I 21 cm absorbing cloud of column density

$\sim 10^{20} \text{ cm}^{-2}$ (assuming $T_s = 100 \text{ K}$ and filling factor 1) in a region beyond the optical extent of the galaxy, at a distance of 5.2 kpc from the galactic nucleus using the northeastern radio lobe of the galaxy itself as a background source. While there is a substantial amount of literature on the using H I absorption toward radio lobes to probe the immediate vicinity of AGN (e.g. Saikia & Jamrozy 2009), this example shows how radio galaxies can also be used to probe the ISM of galaxies far from the nuclear region."

Another advantage of radio absorption studies is that radio sources are often extended and thus unlike UV or optical absorption spectroscopy, which is a single pencil beam study, absorption in the radio band provides a unique opportunity to probe multiple sight lines through the ISM of foreground galaxies simultaneously, and thus is an excellent way to map the small-scale structures in extragalactic ISMs. Such small-scale structures down to AU-scales were detected in the Milky Way as early as in 1976 by Dieter et al. (1976) and later confirmed by Diamond et al. (1989). Over the last few years, multi-epoch mapping of the extended radio emission of quasars by Brogan et al. (2005) and Lazio et al. (2009) have confirmed the small-scale structures in the ISM of the Milky Way detected via 21 cm absorption. These authors report that the covering fraction of these structures are roughly 10% which implies that their volume filling fraction is only about 1%. Similar spatial variations in the ISM of galaxies outside the Milky Way have not yet been detected. One way to constrain the sizes and physical nature of the cold clouds, including their spin temperature and covering fraction, is to conduct a survey of H I absorption against multiple bright background sources that probe a single galaxy. Such a survey would provide unique information on the (1) structure of the cold gas clouds in a galaxy, (2) variation of cloud properties as a function of distance from the galaxy/galactic disk; and (3) the cold gas distribution as a function of global properties (such as color, stellar mass) of galaxies.

In this paper, we investigate the existence and properties of cold gas in the disks and halos of low-redshift galaxies ($z < 0.3$). We seek to constrain the frequency and strength of the H I distribution as function of the distance from the center of a galaxy. A similar study was carried out by Gupta et al. (2010) with 5 quasar-galaxy pairs using the *Giant Metre Radio Telescope* (GMRT). The most important difference between their observational set up and ours is in the spectral and spatial resolution. The two studies are highly complementary. On one hand, the study conducted by Gupta et al. has a high spatial resolution which enables them to detect absorption more efficiently. On the other hand, our study achieved a high spectral resolution, which is crucial in identifying narrow H I features. The 21 cm absorbing gas is expected to be cold and thus the line widths can be extremely narrow. Such features were observed in the Milky Way ($\sim 1 \text{ km s}^{-1}$ by Knapp & Verschuur 1972), the Large Magellanic Cloud (LMC) (1.4 km s^{-1} by Dickey et al. 1994), and Small Magellanic Cloud (SMC) (1.2 km s^{-1} by Dickey et al. 2000). Unlike the extended sample used by Gupta et al. (2010) in their analysis, we use an unbiased sample, which was selected solely on the basis of the impact parameter without any consideration

to absorption detection through other line transitions.

We also wish to investigate the H I distribution in terms of galaxy luminosity, mass, and color. This goal is motivated by the fact that the intensity of accretion and outflow processes are expected to vary from galaxy to galaxy. For instance, blue star-forming galaxies are expected to have a significant reserve of gas compared to red galaxies. Similarly, cold-mode accretion is predicted to be predominant in galaxies of baryonic mass $< 10^{10.3} M_{\odot}$. Secondly, we intend to study the properties of H I clouds, such as their size, filling factor, and temperature, in various regions of an individual galaxy. We begin by presenting our sample and elaborating on the selection procedure, followed by our H I 21 cm observations in section 2. In section 3 we present our results and discuss their implications in section 4. Finally, we summarize our findings and conclusions in section 5. Throughout the paper we use $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.3$, and $\Omega_{\Lambda} = 0.7$.

2. OBSERVATIONS

2.1. Sample Selection

Our sample was selected in a two-stage procedure. First, a set of background QSO sightlines close to foreground galaxies were identified by cross-correlating galaxy positions with quasar positions from objects observed by the *Sloan Digital Sky Survey* (SDSS). The SDSS has redshift information for some of the galaxies, while others have photometric redshift determinations. Hence to identify quasar-galaxy pairs at various impact parameters from the entire SDSS catalog, we used one of the four criteria listed below.

1. For SDSS galaxies with spectroscopic redshifts between $0.005 < z_{\text{gal}} \leq 0.2$, we searched for SDSS QSOs that had sightlines passing within $20''$ of a galaxy, or, within $1.2 \times r_P(r)$, where $r_P(r)$ is the r -band Petrosian radius of a galaxy in arcsec. This latter selection was useful for finding low redshift galaxies where the QSO-galaxy impact parameter was $> 20''$ but where the sightline still passed close to the optical extent of a galaxy. Similarly, we also searched for QSO sightlines from galaxies separated by $< 30''$ but where the galaxy had $r_P(r) > 8.0$, in order to find very low redshift galaxies (with, therefore, large $r_P(r)$ on the sky) and where larger angular separations still translated to relatively small physical impact parameters. In order to remove galaxies whose absorption features might be confused with the AGN activity of the QSO, we excluded galaxies whose redshifts were within $\Delta z = 0.1$ of the QSO's redshift.
2. We searched for SDSS QSO sightlines that passed within $30''$ of SDSS galaxies that had *no* measured redshifts, with the intention of obtaining redshift information using available telescopes. We discuss these data in Section 2.2.
3. We also searched for SDSS QSOs within $1.5 D_{25}$ of all galaxies in the *Third Reference Catalogue of Bright Galaxies* (RC3) (de Vaucouleurs et al. 1991), where D_{25} is the major isophotal diameter of a galaxy measured to a surface brightness level of $\mu_B = 25.0$ square mags.

4. For probing the smallest impact parameters, we chose foreground galaxies that were identified by detecting emission lines superimposed on the background QSO spectrum (Quashnock et al. 2008; Borthakur et al. 2010).

Based on the above criteria, we selected about a thousand QSO-galaxy pairs. In the second stage of selecting suitable QSO-galaxy pairs, we extracted the radio flux densities associated with the optical QSOs from the *Faint Images of the Radio Sky at Twenty-Centimeters* (FIRST) Survey (Becker et al. 1995). We chose a limiting 1.4 GHz flux density value of 25 mJy for the background quasar, corresponding to 3σ optical depth of 0.35 for 2 hours of on-source integration for our observational settings described in the next section. Based on the observable redshift range and excluding pairs that were being observed by (Gupta et al. 2010) with the GMRT, this flux density cut resulted in a sample of 18 galaxy-quasar pairs.

We carried out observations of the 18 galaxy-pairs with the *Green Bank Telescope* (GBT). We encountered severe radio frequency interference (RFI) for three of the pairs, thus reducing our sample to 15 galaxy-quasar pairs. Some of the background quasars are extended radio sources with multiple peaks at the resolution of the FIRST Survey, and therefore allow us to probe different regions of the foreground galaxies. The GBT spectrum for such a pair would have contributions from each of the peaks seen in FIRST image as the GBT beam encompasses all the peaks of the extended quasars. Although, this would cause confusion in the case of a detection, however, this is particularly advantageous for nondetections. More stringent limits can be estimated for each of the background continuum peaks for nondetections. 6 of our 15 galaxy-quasar pairs have extended quasars and these GBT nondetections essentially probe 14 sightlines. Hence, we present our data for the 23 sightlines through 15 galaxy-quasar pairs.

Details on the sightlines along with their 1.4 GHz radio flux densities are presented in Table 1. Radio contours from the FIRST survey overlaid on the composite SDSS r , g , and i -band images for our sample are presented in Figure 1. The QSO is centered in each image and the foreground galaxy is marked by a white arrow. The QSO name is printed at the top left-hand side of the image whereas the sightline numbers used in Table 1 are reproduced near the galaxy or background radio source. Including multiple radio peaks, our sample covers a wide range of impact parameters from 1.7 to 86.7 kpc as well as a wide variety of galaxy properties (see next section). The sightline toward the background QSO SDSS J104257.58+074850.5 allows us to probe the inner stellar disk of the foreground dwarf spiral galaxy GQ1042+0747 at an impact parameter of 1.7 kpc (see Borthakur et al. (2010) for details). Besides a wide continuous range of impact parameters, we also cover a wide range of galaxy properties including luminosity, mass, and color.

2.2. Optical Properties and Galaxy redshifts

For foreground galaxies without redshift information from SDSS, we conducted spectroscopic observations from the Apache Point Observatory (APO) to obtain their redshifts. The APO redshifts were measured either by fitting Gaussian profiles to emission lines or from

TABLE 1
A DETAILED DESCRIPTION OF OUR SAMPLE WITH INFORMATION ON THE BACKGROUND OPTICAL QSOs, THEIR RADIO COUNTERPARTS, AND THE FOREGROUND GALAXIES.

SL	QSO _{opt}	QSO _{radio}	S _{1.4 GHz} (mJy)	Galaxy	z _{gal}	Telescope	m _r	L _r (L _*)	M _{stellar} (Log[M _☉])	Color	ρ (kpc)
1	J010644.15-103410.5	010644.124-103410.55	265.90	J010643.94-103419.3	0.1460	APO	18.68	0.571	10.0	Blue	23.4
2	J074841.77+173456.6	074841.773+173456.82	41.97	J074842.58+173450.6	0.0528	SDSS	16.30	0.595	10.6	Red	13.5
3	J074841.77+173456.6	074841.786+173512.20	21.94	J074842.58+173450.6	0.0528	SDSS	16.30	0.595	10.6	Red	25.1
4	J074841.77+173456.6	074842.084+173443.37	16.08	J074842.58+173450.6	0.0528	SDSS	16.30	0.595	10.6	Red	10.4
5	J082153.82+503120.4	082153.833+503120.57	53.94	J082153.76+503125.7	0.1835 ^a	APO	19.60	0.418	10.0	Blue	16.0
6	J084914.27+275729.7	084914.282+275729.90	53.94	J084912.42+275740.4	0.1948	APO	18.67	1.193	10.5	Blue	86.7
7	J091011.01+463617.8	091011.016+463617.87	163.10	J091010.55+463633.3	0.0998	APO	16.94	1.286	10.6	Blue	29.8
8	J102258.41+123429.7	102258.415+123426.26	93.98	J102257.92+123439.1	0.1253	APO	18.21	0.743	10.7	Red	33.1
9	J102258.41+123429.7	102258.552+123439.94	24.67	J102257.92+123439.1	0.1253	APO	18.21	0.743	10.7	Red	20.9
10	J104257.58+074850.5	104257.598+074850.60	381.57	GQ1042+0747 ^c	0.0332	SDSS	18.05	0.048	9.1	Blue	1.7
11	J110736.61+090114.8	110736.607+090114.72	36.97	J110736.88+090113.6	0.1050	APO	17.82	0.683	10.7	Red	8.1
12	J122106.87+454852.1 ^c	122105.480+454838.80	53.69	J122115.22+454843.2	0.0015	SDSS	14.72	0.003	7.4	Blue	3.3
13	J122106.87+454852.1 ^c	122106.854+454852.16	21.46	J122115.22+454843.2	0.0015	SDSS	14.72	0.003	7.4	Blue	2.8
14	J122106.87+454852.1 ^c	122107.811+454908.02	12.82	J122115.22+454843.2	0.0015	SDSS	14.72	0.003	7.4	Blue	2.6
15	J125248.28+474043.7	125247.588+474042.81	78.72	J125249.20+474105.9	0.0321	APO	13.92	1.861	11.1	Red	18.1
16	J125248.28+474043.7	125249.326+474042.19	45.06	J125249.20+474105.9	0.0321	APO	13.92	1.861	11.1	Red	15.2
17	J132534.55+495341.9	132534.240+495348.47	26.63	J132535.13+495324.4	0.0478	SDSS	16.27	0.497	10.5	Red	23.9
18	J132534.55+495341.9	132534.565+495342.26	15.87	J132535.13+495324.4	0.0478	SDSS	16.27	0.497	10.5	Red	17.5
19	J132840.57+622137.0	132840.599+622136.65	144.10	J132839.89+622136.0	0.0423	APO	18.76	0.038	9.0	Blue	4.2
20	J134528.74+034719.6	134528.765+034720.09	70.21	J134525.30+034823.8	0.0325	SDSS	14.22	1.431	10.9	Red	53.3
21	J141630.67+372136.8	141631.039+372203.01	30.71	J141629.25+372120.4	0.0341	APO	17.19	0.103	9.3	Blue	32.4
22	J141630.67+372136.8	141630.672+372137.09	30.53	J141629.25+372120.4	0.0341	APO	17.19	0.103	9.3	Blue	16.2
23	J160658.30+271705.5	160658.315+271705.86	141.01	J160659.13+271642.6	0.0462	APO	17.82	0.111	9.3	Blue	23.3
-	J082153.82+503120.4	082153.833+503120.57	53.94	J082152.85+503128.4	0.1694 ^b	APO	18.53	1.066	10.8	Red	35.3
-	J084914.27+275729.7	084914.282+275729.90	53.95	J084913.80+275735.9	0.2354	APO	18.69	1.886	10.9	Blue	32.8
-	J115839.90+625428.0	115836.882+625415.06	401.20	J115840.67+625423.7	0.2596	SDSS	18.49	2.887	11.1	Blue	109.7
-	J115839.90+625428.0	115837.852+625420.28	180.07	J115840.67+625423.7	0.2596	SDSS	18.49	2.887	11.1	Blue	78.6
-	J115839.90+625428.0	115842.780+625442.26	69.64	J115840.67+625423.7	0.2596	SDSS	18.49	2.887	11.1	Blue	94.5

^a Gupta et al. 2010 in press

^b Identified through emission lines on the QSO spectrum (see Quashnock et al. 2008; Borthakur et al. 2010)

^c Also known as UGC 7408.

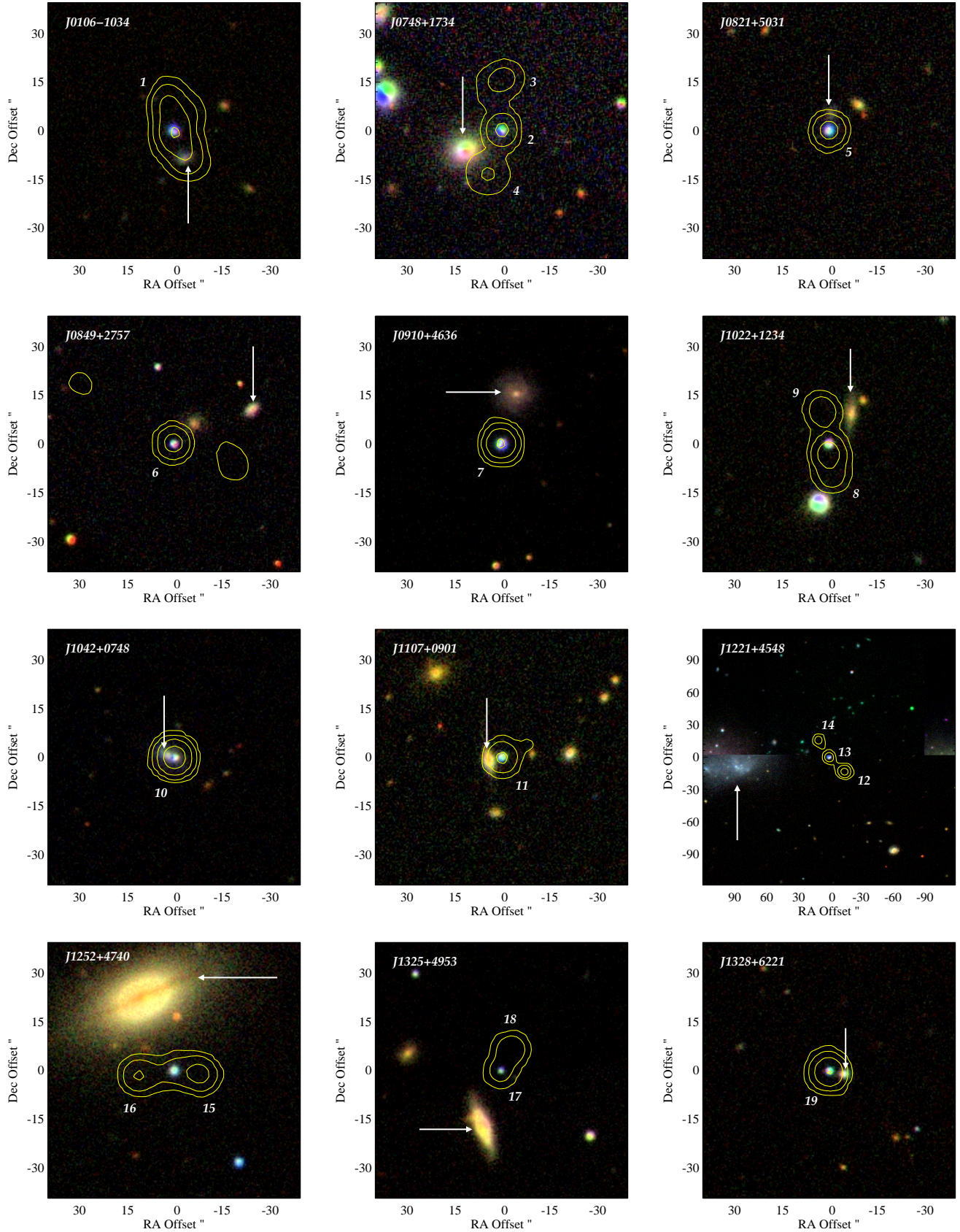


FIG. 1.— SDSS composite image of *i*-, *r*-, and *g*-band images for our sample of 15 galaxy-quasar pairs. The quasar is at the center of the image and its name is printed at the top left side of each plot. The yellow contours shows the 1.4 GHz VLA B-configuration continuum emission from the FIRST survey at levels 1, 5, 25, and 125 mJy/beam. The foreground galaxies are marked with a white arrow and each of the sightlines is numbered. Due to the large extent of the background quasars in 1.4 GHz some of our systems covered multiple sightlines. Details of our sample are presented in table 1.

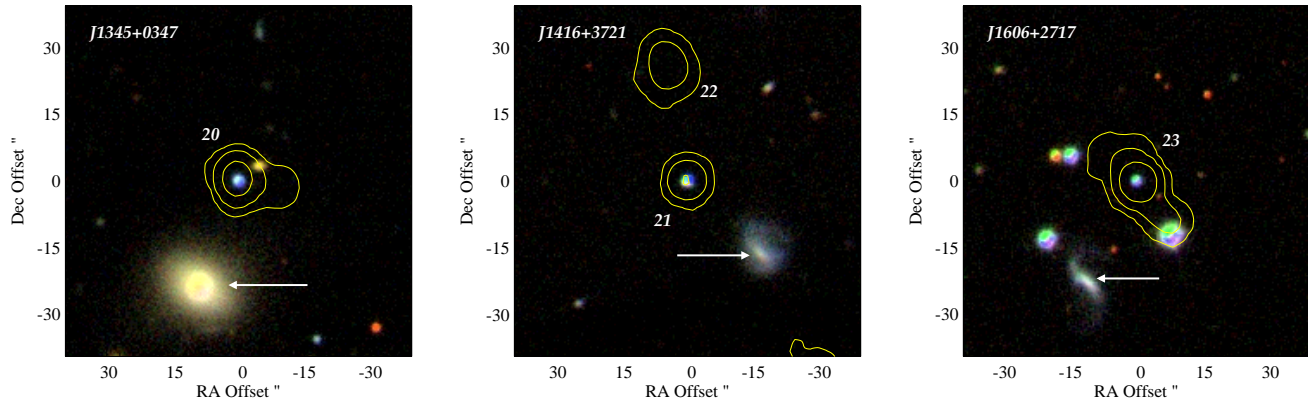


FIG. 1.— Continued.

TABLE 2
APO REDSHIFTS FOR FOREGROUND GALAXIES USING EMISSION
AND ABSORPTION LINES.

Foreground Galaxy	$z_{em-line}$	$z_{abs-line}^a$	$\delta z_{abs-line}^a$
J010643.94-103419.3	0.1460	-	-
J084913.80+275735.9	0.2356	0.2354	0.0004
J084912.42+275740.4	0.1941	0.1948	0.0008
J091010.55+463633.3	0.0998	0.0998	0.0004
J102257.92+123439.1	0.1254	0.1253	0.0004
J110736.88+090113.6	-	0.1050	0.0003
J132839.89+622136.0	0.0420	0.0423	0.0005
J141629.25+372120.4	0.0341	-	-

^a Estimated using IRAF fxcor routine.

the stellar absorption line measurements made using the Image Reduction and Analysis Facility (IRAF) “fxcor” routine, and cross-correlating the blue spectra with that of a radial velocity standard. The redshifts are presented in the sixth column of Table 1 with more details in Table 2. With the exception of two galaxies, the remaining APO redshifts are measured using absorption-lines. The r-band magnitude and luminosity for the foreground galaxies are presented in the eighth and ninth columns of Table 1. The luminosities, in units of L^* , were estimated using $M_r^* = -21.2$ (Table 2 of Blanton et al. 2003) and were k-corrected for r-band absolute magnitude at $z = 0.0$. The k-corrections were derived using the k-correct algorithm by Blanton & Roweis (2008) that uses SDSS photometric u, g, r, i, and z-band data. Due to the unavailability of SDSS photometric data for GQ1042+0747 we have used g, r, and i-band photometric values from Borthakur et al. (2010) to estimate the k-correction for this galaxy. We also present stellar masses for the foreground galaxies, which were estimated using Equation (1) from McIntosh et al. (2008) based on the k-corrected magnitudes.

From the galaxy color-magnitude diagram that relates galaxy luminosity and color one can classify galaxies as part of the red sequence or the blue cloud. For our sample, we follow the empirical results from van den Bosch et al. (2008, equation A6), which uses the parameter $^{0.0}A$ defined as

$$^{0.0}A = 0.0 (g - r) - 0.10 (\log[M_*] - 10.0) \quad (1)$$

where $^{0.0}(g-r)$ is based on g and r band magnitudes that have been k-corrected to $z = 0.0$ and M_* is the stellar mass of the galaxy. The two populations are found to have an overlap in terms of the $^{0.0}A$ value intersecting

at $^{0.0}A = 0.675$ (see Figure A1 of van den Bosch et al. 2008). Hence we use this value to distinguish the two populations and we classify galaxies with $^{0.0}A < 0.675$ as belonging to the blue cloud and ones with $^{0.0}A > 0.675$ as red sequence. In total, our sample consists of nine blue cloud and six red sequence galaxies.

2.3. Radio Observations

In order to search for cold gas in the outer disks and halos of galaxies, we carried out GBT observations of our sample followed by higher spatial resolution Very Large Array (VLA) and Very Large Baseline Array (VLBA) observations. Details of the VLA and VLBA observations of the pair SDSS J104257.58+074850.5-GQ1042+0747 have been presented in Borthakur et al. (2010) and are not presented in this section.

2.3.1. 21 cm Spectroscopy with the GBT

We conducted GBT observations of the H I 21 cm and OH 18 cm transitions for 18 galaxy-quasar pairs (Table 1) under project IDs GBT06B-052, GBT08A-082, and GBT09C-039 in August 2006, March through July 2008, and October and November 2009 respectively. We used the dual polarization L-band system with 9-level sampling. The intermediate frequency (IF) was set to yield a channel width of 1.56 kHz (0.33 km s^{-1}) using 8192 channels over a total bandwidth of 12.5 MHz. Observations were made in the standard position-switching scheme by cycling through the ON-OFF sequence, then dwelling for 300 s at each position. Each 300 s scan consisted of thirty 10 s integrations to minimize data corruption by sporadic RFI. The OFF position was chosen $+20'$ offset in Right Ascension from each of the sightlines. This was done to track the possible presence of a source in the OFF position. During each of the observing sessions, we used one of the three standard flux calibrators - 3C 48 (16.5 Jy), 3C 147 (22.5 Jy), and 3C 286 (15.0 Jy) - for pointing and estimating antenna gain. Local pointing corrections (LPC) were performed using the observing procedure AutoPeak and the corrections were then automatically applied to the data. This should result in a pointing accuracy of $3''$. Similarly, flux calibrations were performed by applying antenna gain to data for each session separately (values are presented in Table 3). The corrections also include attenuation due to opacity of the air (around 1%).

The data were analyzed using the NRAO software GBTIDL. Most of the sightlines were observed for almost 2 hours on-source. However, due to the corruption of the data by RFI the noise levels achieved for each of the sightlines span a wide range. Severe RFI corrupted entire datasets for sightlines toward the background radio QSOs 082153+503120 at $z = 0.1640$, 084914+275729 at $z=0.2357$, and 115836+625415 at $z=0.2596$ rendering the data unusable. Besides RFI, the GBT data also suffered from standing waves for some of the brighter sources, which limited our H I emission detection capabilities, especially for H I features with $\Delta V \geq 300$ km s $^{-1}$. The details of the noise properties and GBT measurements are presented in Table 3. No OH emission was detected in any of our sources and rms noise levels are presented in the above mentioned table.

2.3.2. Very Large Array Observations of J122115+454838

In order to understand the effect of the large GBT beam size on the survey, we imaged the H I distribution in the foreground galaxy J122115+454838 (UGC 7408) with the VLA under program AY190 on 28 June and 16 July 2008. The observations were carried out in the VLA D-configuration for 3.5 hours of on-source integration. We used dual polarization using 2 intermediate frequency (IF) correlator modes to achieve a total bandwidth of 1.5625 MHz (330 km s $^{-1}$) with a 12.2 kHz resolution after Hanning smoothing. This correlator setup allowed us to cover the entire H I line-width with a velocity resolution of 2.6 km s $^{-1}$. The data were calibrated following the standard VLA calibration and imaging procedures using NRAO’s Astronomical Image Processing Software (AIPS). The data were imaged using both natural and uniform weighting schemes. The natural weighting was applied to enhance the fainter emission features whereas uniform weighting was applied to achieve a higher spatial resolution for the purpose of extracting a spectrum with minimum contamination from the H I emitting region. The synthesized beam produced using natural weighting was $67.51'' \times 45.31''$ (Position angle = -86.75°) and uniform weighting was $52.17'' \times 38.11''$ (Position angle = -80.56°). The absolute uncertainty in the resulting flux density scaling is about 10%, and this is the formal uncertainty we quote for all physical parameters derived from the flux density. The average rms noise achieved, after combining the two polarizations, is 1.07 mJy/beam/channel.

3. RESULTS

We detected H I emission from three foreground galaxies: J122115.22+454843.2 (UGC 7408), J141629.25+372120.4, and J160659.13+271642.6, thus suggesting the presence of gas-rich galaxies in our sample. Figure 2 shows the smoothed GBT H I spectra for these galaxies. The spectra are presented in terms of the optical velocity of the galaxies, which show very good correspondence with the H I distribution. The H I masses presented in Table 3 were estimated using the relationship $M(\text{HI}) = 2.36 \times 10^5 D^2 (S\Delta V) M_\odot$, where D is the luminosity distance to the galaxy in megaparsec (Mpc) and the $S\Delta V$ is the velocity integrated H I flux density in Jy km s $^{-1}$. All the three galaxies are “blue” (as defined in Section 2.2) although they vary in luminosity and stellar mass. While J122115.22+454843.2 (UGC7408)

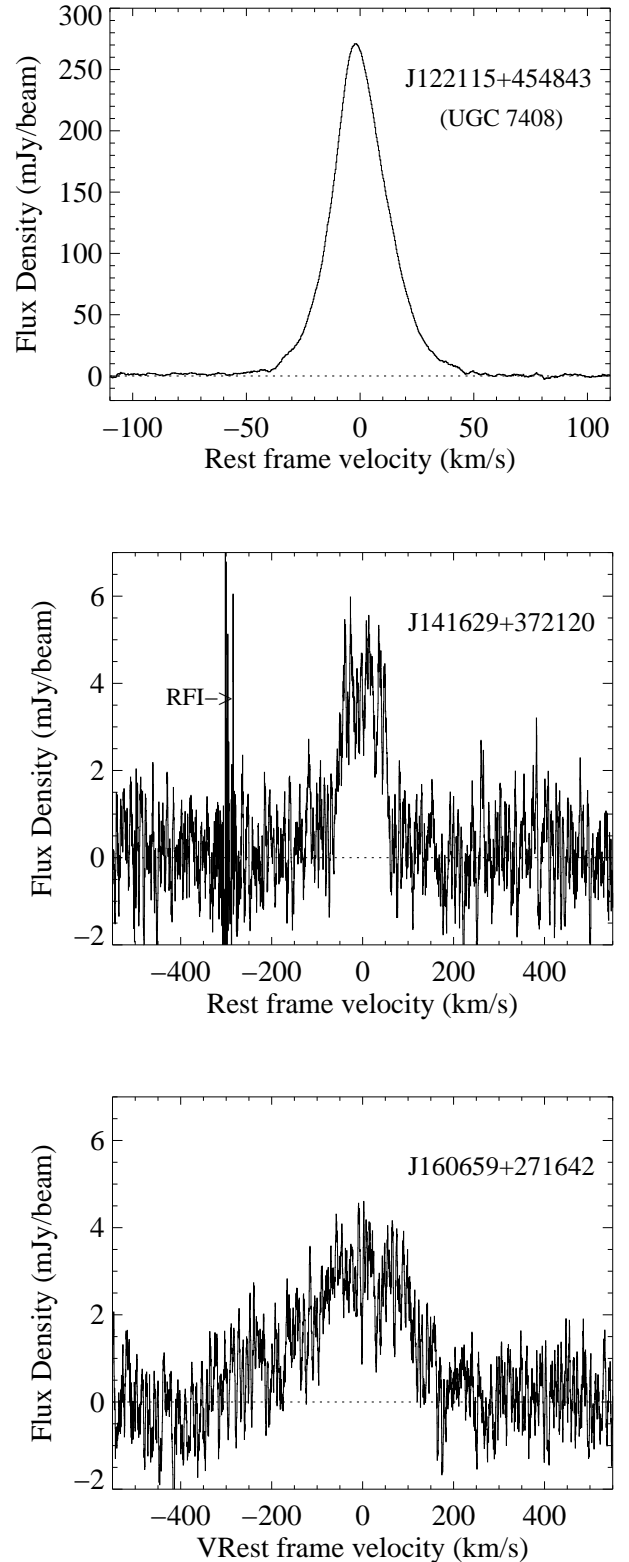


FIG. 2.— H I spectra for foreground galaxies UGC 7408 (top), J141629+372120 (center), and J160659+271642 (bottom). The spectra are presented at the rest frame of the optical galaxies to show the match between the optical and H I velocity distribution. All three spectra have been smoothed to 3 km s $^{-1}$ (i.e. 10 pixels). The H I mass associated with each of these sources are presented in Table 3.

TABLE 3
GBT 21 CM OBSERVATIONS AND MEASUREMENTS OF H I OPTICAL DEPTH AND MASS ASSOCIATED WITH THE FOREGROUND GALAXIES.

SL	QSO _{radio}	S _{L,AGHz}	Gain	σ_{HI}	τ^a	σ_{OH}	Foreground Galaxy	Distance	HI Mass
1	010644.124+103410.55	265.90	1.68	3.65	< 0.02	2.02	J010643.94-103419.3	692.4	-
2	074841.773+173456.82	41.97	1.72	2.45	< 0.10	1.63	J074842.58+173450.6	235.2	-
3	074841.786+173512.20	21.94	1.72	2.45	< 0.19	1.63	J074842.58+173450.6	235.2	-
4	074842.084+173443.37	16.08	1.72	2.45	< 0.26	1.63	J074842.58+173450.6	235.2	-
5	082153.833+503120.57	53.94	1.70	4.96	< 0.16	2.41	J082153.76+503125.7	890.5	-
6	084914.282+275729.90	53.94	1.72	6.36	< 0.20	1.98	J084912.42+275740.4	951.7	-
7	091011.016+463617.87	163.10	1.71	1.31	< 0.01	RFI	J091010.55+463633.3	459.3	-
8	102258.415+123426.26	93.98	1.70	1.86	< 0.03	1.88	J102257.92+123439.1	586.4	-
9	102258.552+123439.94	24.67	1.70	1.86	< 0.13	1.88	J102257.92+123439.1	586.4	-
10	104257.598+074850.60	381.57	1.65	4.2	0.04962	RFI	GQ1042+0747 ^b	145.9	-
11	110736.607+090114.72	36.97	1.82	3.81	< 0.18	RFI	J110736.88+090113.6	484.9	-
12	122105.480+454838.80	53.69	1.74	3.49	< 0.11	1.90	J122115.22+454843.2 ^c	8.4	1.4
13	122106.854+454852.16	21.46	1.74	3.49	< 0.28	1.90	J122115.22+454843.2 ^c	8.4	1.4
14	122107.811+454908.02	12.82	1.74	3.49	< 0.47	1.90	J122115.22+454843.2 ^c	8.4	1.4
15	125247.588+474042.81	78.72	1.86	2.64	< 0.06	RFI	J125249.20+474105.9	140.9	-
16	125249.326+474042.19	45.06	1.86	2.64	< 0.10	RFI	J125249.20+474105.9	140.9	-
17	132534.240+495348.47	26.63	1.64	5.90	< 0.38	RFI	J132535.13+495324.4	212.2	-
18	132534.565+495342.26	15.87	1.64	5.90	< 0.64	RFI	J132535.13+495324.4	212.2	-
19	132840.599+622136.65	144.10	1.78	3.02	< 0.04	2.13	J132839.89+622136.0	187.0	-
20	134528.765+034720.09	70.21	1.91	3.12	< 0.08	RFI	J134525.30+034823.8	142.7	-
21	141631.039+372203.01	30.71	1.87	2.75	< 0.16	RFI	J141629.25+372120.4	149.9	23.2
22	141630.672+372137.09	30.53	1.87	2.75	< 0.16	RFI	J141629.25+372120.4	149.9	23.2
23	160658.315+271705.86	141.01	1.76	2.34	< 0.03	RFI	J160659.13+271642.6	204.8	81.3
-	082153.833+503120.57	53.94	RFI	RFI	RFI	1.88	J082152.85+503128.4	815.1	-
-	084914.282+275729.90	53.95	RFI	RFI	RFI	2.40	J084913.80+275735.9	1177.1	-
-	115836.882+625415.06	401.20	RFI	RFI	RFI	2.20	J115840.67+625423.7	1315.4	-
-	115837.852+625420.28	180.07	RFI	RFI	RFI	2.20	J115840.67+625423.7	1315.4	-
-	115842.780+625442.26	69.64	RFI	RFI	RFI	2.20	J115840.67+625423.7	1315.4	-

^a Integrated optical depth binned over 1 km s⁻¹ bins i.e 3 pixels. The limiting values show 3 σ 1 optical depth.

^b Identified through emission lines on the QSO spectrum (see Quashnock et al. 2008; Borthakur et al. 2010)

^c Also known as UGC 7408

is a gas rich dwarf galaxy of luminosity $0.003 L_*$ and H I mass of $1.4 \times 10^8 M_\odot$, J141629.25+372120.4 and J160659.13+271642.6 are $\sim 0.1L_*$ galaxies and have H I masses of $2.3 \times 10^9 M_\odot$ and $8.1 \times 10^9 M_\odot$, respectively.

The VLA D-configuration image of H I emission associated with UGC 7408 (green contours), overplotted on an SDSS r-band image, is presented in the left panel of Figure 3. The FIRST image (beam size = $5''.40 \times 5''.40$) of the background QSO J122106.87+454852.1 with its lobes is shown in blue in (enlarged version in the right panel) and its D-configuration continuum image (beam size = $52''.17 \times 38''.11$) is shown with white contours. The sightlines (sightline (SL)# 12, 13, and 14) to the background continuum source probes the galaxy at impact parameter of 3.2, 2.8, and 2.6 kpc. The H I distribution in this galaxy is extended, extending 2-3 times beyond the optical size of the galaxy. The H I is asymmetric and extends $5'.6$ to the southeast of the galaxy. *The Westerbork observations of neutral Hydrogen in Irregular and SPiral galaxies* (WHISP Survey; Swaters et al. 2002) also imaged this galaxy. The subsequent analysis by Swaters et al. (2009) found that its H I distribution is clumpy and lacking ordered rotation.

As we observed UGC 7408 with both the VLA and GBT, we were able to compare the HI emission from the galaxy that we obtained from each telescope. Figure 4 shows a good match between the spectra from the two telescopes. The VLA recovered integrated flux of 7.4 Jy km s^{-1} (removing negative pixels) as compared to 8.1 Jy km s^{-1} of GBT. This confirms that the VLA data

recovered more than 91% of the H I flux density detected by the GBT. The H I spectral profile shows a deviation from a Gaussian profile at the peak of the emission and in the high velocity wing between 480 and 500 km s⁻¹. The VLA spectrum obtained from a region of equal size as the synthesized beam in the uniform weighting case (the black rectangle in the left panel of Figure 3) over the quasar reveals two emission peaks with a dip at 442.2 km s⁻¹. This could be a sign of absorption and is discussed in detail in section 4.2.

In the other two galaxies with H I emission (J141629.25+372120.4 and J160659.13+271642.6), the expected H I diameter estimated using the relationship shown in Figure 7 of (Swaters et al. 2002) are 30.7 and 60.2 kpc, respectively. Thus, the sightlines, SL# 22 and 23 are expected to pierce through the H I distribution in these systems. Although we detect dips in the GBT spectra, it is unclear if there is real absorption or are overlapping individual emission features mimicking absorption. Hence, we consider these two cases as non-detections. A detailed discussion of possible beam effects and how the GBT spectra in these two cases may be a result of the beam being filled-in by emission is presented in Section 4.2.

Besides the H I emitting galaxies, we detected one unambiguous H I absorption features at 21 cm in the foreground galaxy GQ1042+0747 (SL#10), a result we discussed in detail in Borthakur et al. (2010). GQ1042+0747 is a dwarf spiral galaxy of luminosity $0.048 L_*$. The sightline pierces through the optical disk in this moderately star-forming galaxy. We estimated the

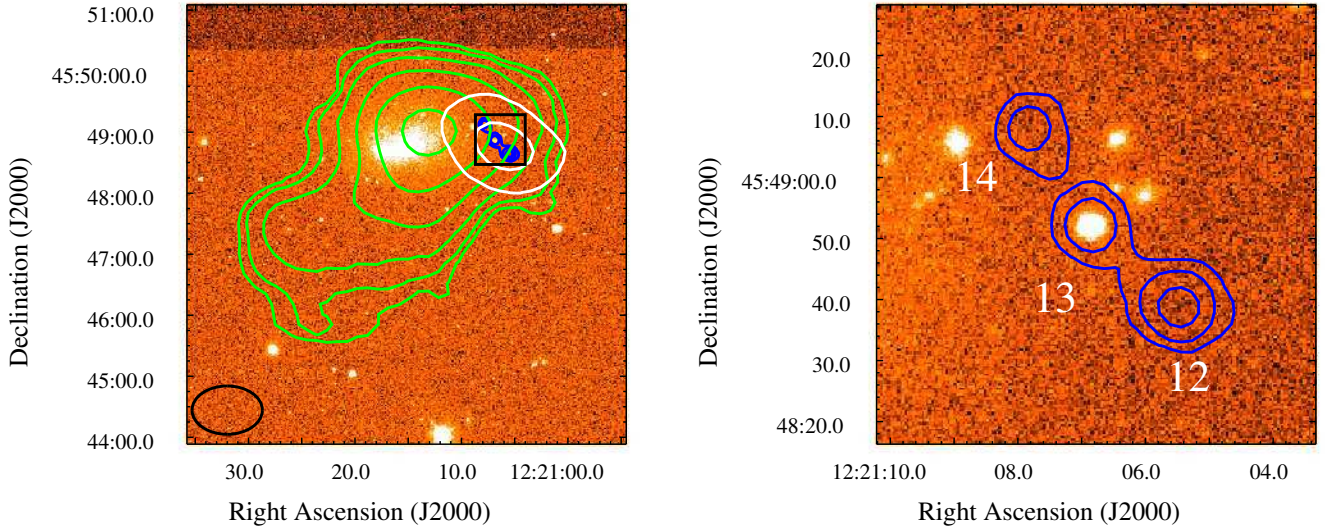


FIG. 3.— LEFT: VLA D-configuration H I 21 cm column density contour image of UGC 7408 overlaid as green contours on an SDSS-r band image. The contours levels are at $1.8, 3.6, 7.2, 14.4, 28.8, \& 57.6 \times 10^{19} \times \text{cm}^{-2}$ respectively. The beam size associated with the H I data obtained using natural weighting is shown at the bottom left corner ($67''.51 \times 45''.31$). The blue contours represent the VLA FIRST survey image of the background quasar J122106.87+454852.1 at 1, 5, and 25 mJy/beam and the white contours represent the VLA D-configuration 21 cm continuum image obtained using uniform weighting at 5 and 50 mJy/beam (Position angle = -86.75°). The higher resolution FIRST survey images shows three peaks from the background source, which are blended together into a single source in the D-configuration image. The black square represents the region corresponding to the red spectrum in Figure 4, which is roughly the size of the beam for the uniformly weighted data. RIGHT: A magnified version of the left hand panel showing the background quasar as imaged by the FIRST survey. The brightest background source at 1.4 GHz is the southern component of the quasar. The three peaks of the background quasar mark the sightlines 12, 13, and 14.

kinetic temperature of the cloud to be ≤ 283 K based on the absorption line FWHM of 3.6 km s^{-1} . Assuming the spin temperature to be same as the kinetic temperature, we estimated the column density of the absorber to be $N(\text{H I}) \leq 9.6 \times 10^{19} \text{ cm}^{-2}$. A detailed analysis of the properties of the absorber, the foreground galaxy, and its environment is presented in Borthakur et al. (2010). We did not detect any H I absorber in our GBT data from the three H I emitting regions. Part of the problem may be the large GBT beam size, which can include a substantial portion of H I emitting regions that may fill up the absorption. This possibility is also discussed in detail in section 4.2. Table 3 presents the 3σ limiting optical depths for each of the sightlines along with the flux density sensitivity achieved. Assuming $T_{\text{spin}} = 100$ K for the absorbing gas with $f_{\text{cov}} = 1$ and a line-width of $\sim 5 \text{ km s}^{-1}$, the corresponding limiting column densities would range from $0.09 - 5.83 \times 10^{20} \text{ cm}^{-2}$, with a median of $1.46 \times 10^{20} \text{ cm}^{-2}$.

4. DISCUSSION

4.1. Covering Fraction

The covering fraction of cold gas in galactic halos has been a matter of debate. While low-ionization gas in the Lyman α and Mg II transition have been detected in the halos of various galaxies, H I 21 cm transitions have not been explored extensively. For instance, Lanzetta et al. (1995) found that a significant portion of the Lyman α absorbers (1/3 to 2/3) were associated with normal, luminous galaxies with Lyman α absorbing halos of size $160 h^{-1} \text{ kpc}$. They found an anti correlation between the strength of the absorption (equivalent width) and galaxy impact parameter, which was later confirmed by other authors (e.g. Tripp et al. 1998; Chen et al. 1998, 2001, but see also Bowen et al. 2002; Wakker & Savage

2009). In this study, our aim is to estimate the frequency of occurrence of 21 cm absorbers in the halos and extended disks of galaxies. Figure 5 shows our findings in terms of two plots. The left plot shows the variation of measured optical depth as a function of impact parameter and the plot on the right shows the same quantity as a function of the ratio of impact parameter over the r-band Petrosian radius of the foreground galaxies. While the left plot shows the strength of the absorption or limiting strength versus the rest frame projected distance, the plot on the right shows the same quantity as a function of the optical size of the foreground galaxy. In addition, the optical colors of the foreground galaxies are denoted using respective colored symbols, and the luminosity with symbol-size, as described in the legend. Our single 21 cm absorber detection towards, SDSS J104257.58+074850.5, is shown as a filled circle while all the other non-detections are shown with open circles with arrows pointing down. We do not include the limiting optical depth values for the Gupta et al. (2010) sample as their channel width is more than an order of magnitude higher than ours. In addition, Gupta et al. (2010) presented their limiting optical depth for equivalent velocity resolution of 10 km s^{-1} . In contrast, we present our optical depth limits for a resolution of 1 km s^{-1} . The high resolution guarantees that such a study will detect absorbers similar to the one in GQ1042+0747 or for H I at a temperature of ~ 50 K.

The nondetection of cold H I can be understood in terms of three possibilities: (1) the nonexistence of cold gas along the QSO line of sight, (2) a low covering fraction of the background QSO by the absorber or high spin temperature of the absorber, and/or (3) confusing effects due to the large beam size of the GBT (FWHM of $9''.1$). The possible existence of cold gas clouds with Damped

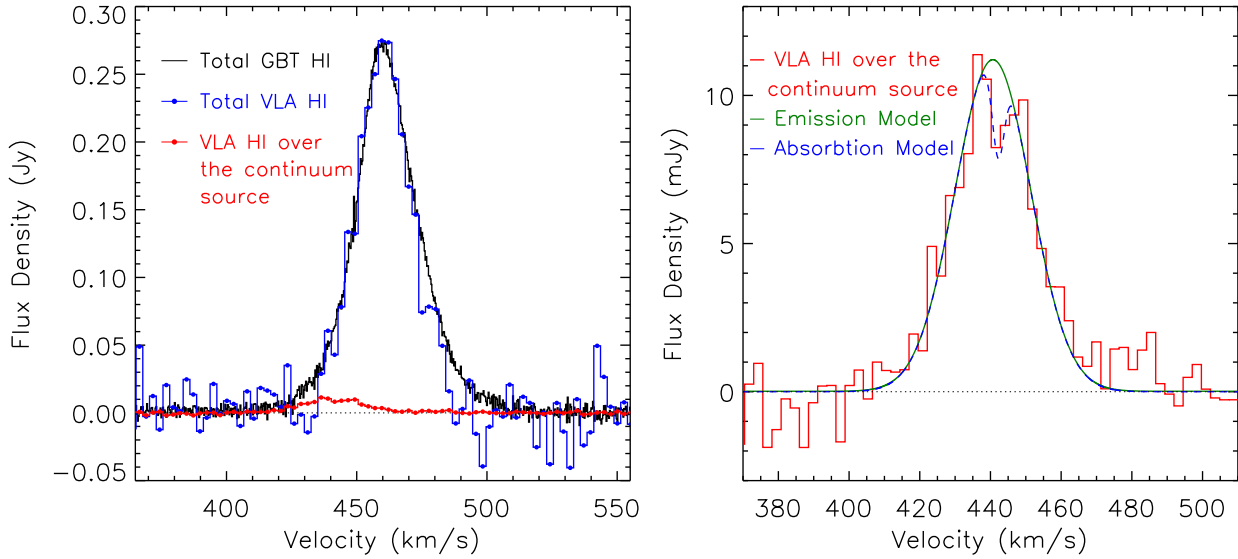


FIG. 4.— LEFT: A comparison plot showing the GBT HI spectrum from a region of size $9.1'$ (equivalent to the GBT beam at 21 cm) toward UGC 7408 in black and the VLA spectrum, obtained using natural weighting, encompassing the entire HI associated with UGC 7408 in blue. The good match between the two spectra proves that the VLA recovered almost 100% of the HI associated with this galaxy. The VLA HI spectrum from the region marked by the black rectangle (Figure 3, right panel) of size $40'' \times 50''$ is shown in red in the comparison plot. The spectrum was obtained from data reduced using uniform weighting, which resulted in a better spatial resolution. RIGHT: A magnified version of the red spectrum showing two emission peaks and a dip in between at 442.2 km s^{-1} . Interestingly, the stronger peak at 436.3 km s^{-1} is toward the low-velocity end of the profile unlike the global HI profile of the galaxy at that velocity. The green solid line and the blue dotted line represent Gaussian fits to the emission and the absorption profile respectively. The absorption is narrow with FWHM of 4.75 km s^{-1} corresponding to a kinetic temperature of $\leq 493 \text{ K}$.

Lyman α (DLA) column density of $2 \times 10^{20} \text{ cm}^{-2}$ in the region surveyed can be estimated using the cross section implied by the dN/dz of DLA systems (see e.g., section 6.1 in Schaye et al. 2007), where the covering fraction of such clouds is described as

$$f_{cov}^{DLA} \approx 7.3 \times 10^9 \left(\frac{dN}{dz} \right) \left(\frac{n_g}{10^{-2} \text{ Mpc}^{-3}} \right)^{-1} \left(\frac{r}{\text{pc}} \right)^{-2}, \quad (2)$$

where dN/dz is the rate of incidence of DLA systems at $z = 0$, n_g is the comoving number density of galaxies affiliated with the absorbers, and r is the radius of the absorber cross section. dN/dz for 21 cm column densities comparable to those of DLAs in the nearby universe was computed by various authors from large HI surveys like the *Arecibo Dual-Beam Survey* (ADBS) and *HI Parkes All Sky Survey* (HIPASS). Rosenberg & Schneider (2003) found $dN/dz = 5.3 (\pm 1.3) \times 10^{-2}$ from the ADBS whereas Ryan-Weber et al. (2003) and Zwaan et al. (2005) found $dN/dz = 5.8 (\pm 0.6) \times 10^{-2}$ and $4.5 (\pm 0.6) \times 10^{-2}$ respectively from the HIPASS. For our calculation, we adopt $dN/dz = 5.3 \times 10^{-2}$ from Rosenberg & Schneider, which is also the median of the three values cited above. This value is consistent with the DLA incidence per unit redshift $n_{DLA}(z) = (0.044 \pm 0.005) \times (1+z)^{1.27 \pm 0.11}$ for $0 < z < 5$ as described by Rao et al. (2006) based on their HST survey of SDSS quasars with intervening Mg II-Fe II absorption systems. The number density of galaxies within a magnitude range can be estimated from the Schechter luminosity function (Schechter 1976). For our analysis, we use the Schechter function fit of $\phi_* = 1.49 \times 10^{-2} h^3 \text{ Mpc}^{-3}$, $M_* - 5 \log_{10} h = -20.44$, and $\alpha = -1.05$, estimated by Blanton et al. (2003) for SDSS r-band data k-corrected to $z=0.1$. For the range of

luminosities in our sample, $n_g \sim 2 \times 10^{-2} \text{ Mpc}^{-3}$. Based on the Equation 2, the dN/dz statistics require $f_{cov}^{DLA} = 1$ for cross section of 14 kpc or less. The covering fraction drops as the cross section of the galaxy increases. For our sample, at a cross section radius of 55 kpc, f_{cov}^{DLA} drops to ≈ 0.08 , and at 90 kpc it further drops to ≈ 0.02 .

The presence of detectable 21 cm absorption not only depends on f_{cov}^{DLA} and the column density of the gas but also on the spin temperature, T_s , as well as the fraction of the background quasar covered by the absorber. This fraction, which we will refer to as f hereafter, which is different from f_{cov}^{DLA} . For an HI absorption feature the column density can be estimated using the following equation from Rohlfs & Wilson (2000):

$$N(\text{HI}) = 1.823 \times 10^{18} \frac{T_s}{f} \int \tau_{21}(v) dv \text{ cm}^{-2}, \quad (3)$$

where $\tau_{21}(v)$ is the 21 cm optical depth as a function of velocity in km s^{-1} . In cold clouds collisional excitation is the primary source of excitation for the hydrogen atom emitting the 21 cm line. Therefore, the above equation can be simplified by assume the spin temperature to be same as the kinetic temperature, T_k , which in turn can be related to the full-width at half maximum (FWHM) of the absorption feature by

$$T_k \leq 21.855 (\Delta v)^2 \quad (4)$$

where Δv is the FWHM velocity in km s^{-1} . Therefore, Equation 3 can be rewritten as

$$N(\text{HI}) = 0.390 \times 10^{18} \frac{T_s^{3/2}}{f} \tau_{avg} \text{ cm}^{-2} \quad (5)$$

where τ_{avg} represents the average optical depth such that $\tau_{avg} \Delta v = \int \tau_{21}(v) dv$. Figure 6 shows the variation of HI column density as a function of optical

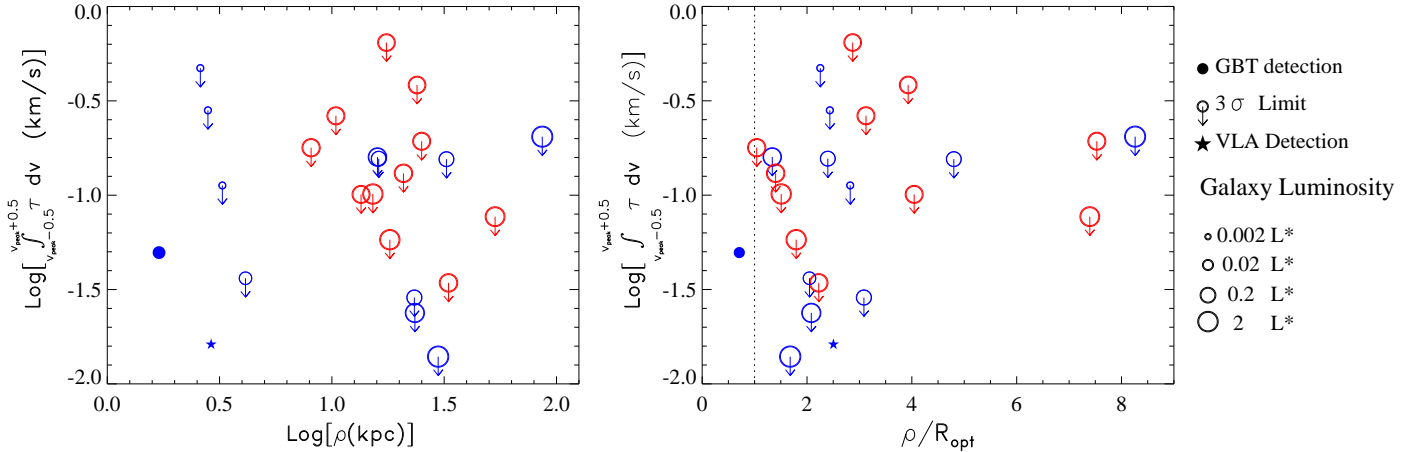


FIG. 5.— LEFT: Impact parameter, ρ , versus the GBT integrated optical depth over channels corresponding to 1 km s^{-1} (i.e. 3 channels), $\int_{v_{\text{peak}}-0.5}^{v_{\text{peak}}+0.5} \tau dv$, for our sample. The filled circles indicate 21 cm absorbers and open circles represents the 3σ limiting values. The colors show color classification of the galaxy as discussed in Section 2.2. The peak optical depth per km s^{-1} for the absorber detected in the VLA data of UGC 7408 is shown as a filled star. RIGHT: Ratio of impact parameter to optical Petrosian r-band radius, ρ/R_{opt} , versus the integrated optical depth. The dotted line represents $\rho/R_{\text{opt}} = 1$. The sightline probing GQ1042+0747 is the only sightline to probe the optical disk. This is also our only unambiguous 21 cm absorption detection with the GBT.

depth, assuming $f = 1$. The contours represent logarithms of H I column density and are labeled accordingly. A column density of $2 \times 10^{20} \text{ cm}^{-2}$, corresponding to a DLA system, is shown with a solid line with filled circles overplotted and labeled 20.3. The filled cyan circle shows the absorber associated with GQ1042+0747. The dashed line represents the maximum spin temperature measured in the Milky Way (Braun & Walterbos 1992), although the majority of their observations for the Milky Way and M31 have T_s below 300 K. The absorber associated with GQ1042+0747 is consistent with these observations. However, higher values of spin temperature have been commonly seen at higher redshifts and even in low-redshift absorbers associated with non-spiral galaxies (Lane et al. 1998; Chengalur & Kanekar 2000; Kanekar & Chengalur 2001, 2005). Our survey aims to detect cooler clouds with DLA equivalent column densities up to 1100 K (if $f = 1$). It is clear from Figure 6 that even for our best limiting optical depth value of 0.0139, if $T_s > 1100 \text{ K}$ then the corresponding H I column density would have to be well over the DLA limit to be detectable as a 21 cm absorber. On the other hand, we can be certain that a non-detection confirms that there are no cold clouds with $T_s < 100 \text{ K}$ (assuming $f = 1$) and column density equivalent to a DLA in our entire sample, with the exception of our noisiest sightline (SL# 18 with $\tau_{\text{avg}} < 0.64$).

4.2. Cold gas in the Dwarf Galaxy UGC 7408

The GBT spectrum towards 122106.854+454852.16 does not reveal any absorption from UGC 7408 despite the fact that there is H I emission in the region in front of the quasar as seen in the VLA D-configuration image (see Figure 3). The H I column density in the foreground region within the two lobes is between 1.44 and $2.88 \times 10^{20} \text{ cm}^{-2}$ at the resolution of VLA D-configuration. In this section, we investigate the possible causes that may be responsible for the lack of absorption features in the GBT spectrum and what that may mean for the rest of our sample. The nondetection of cold

gas in absorption, despite the good sensitivity achieved, could be due to any of the three following underlying causes. For an H I emitting galaxy such as UGC 7408, it is highly possible that the absorption is filled in by emission from the rest of the galaxy. This also extends to the other two H I emitting foreground galaxies in our sample. The second possibility is that the H I is mostly warm and does not absorb the background continuum. The third possibility is that the ISM in this galaxy is highly patchy and $N(\text{H I})$ is much lower along the QSO sightlines than the average inferred at larger scales from the emission map.

Ideally, 21 cm absorption studies require high angular resolution to prevent absorption from being filled in by emission from the rest of the system. Dickey et al. (2000) suggests a resolution of $< 10''$ for extragalactic sources. Unfortunately, the high sensitivity and spectral resolution achieved by the GBT comes at the cost of spatial resolution. On the other hand, our VLA D-configuration beam is much smaller than the GBT beam and this provides us with the opportunity to search for absorption from the region on the line of sight to the background quasar. The right panel of Figure 4 shows the VLA spectrum (in red) from the region marked by the black rectangle in Figure 3, which is also the size of beam for the uniform weighted data. The spectrum exhibits peaks at 436.3 and 449.3 km s^{-1} with a dip in between at 442.2 km s^{-1} . The fact that the absorption dip sits right at the peak of the local emission suggests that the dip is an absorption feature. Furthermore, the ratio of the strength of the two peaks in the red spectrum is opposite to that of the global H I profile of UGC 7408 at that velocity. This further suggests that we are spatially and kinematically resolving the underlying H I emission. Even with the VLA D-configuration beam, we are still primarily detecting the emission from the foreground region. For an H I cloud and a background quasar contained within a single beam, the observed line temperature, T_L , can be expressed as Equation (12.22)

of Rohlfs & Wilson (2000)

$$T_L = (f_{cl} T_s - f_0 f_C T_C)(1 - e^{-\tau}), \quad (6)$$

where, f_{cl} and f_C are beam filling factors of the H I cloud and continuum source, respectively, f_0 is the fraction of the continuum source covered by the H I cloud, and T_C is the brightness temperature of the background source. From the FIRST image of background quasar, J122106.854+454852.16, we find that $f_C < 0.33$ and $T_C = 25$ K for the VLA in D-configuration, and thus we expect the emission to be significantly stronger than absorption even for very cold clouds ($T_s \sim 25$ K).

In order to get a proper estimate of the H I column density associated with this dip, it is necessary to separate the emission profile from the absorption. By assuming the emission from the region covered by the VLA beam is a Gaussian, we fit the emission profile (excluding the dip). The green solid line in the right panel of Figure 4 shows the emission profile. Similarly the absorption profile is also fitted with a Gaussian profile yielding a FWHM=4.75 km s⁻¹ centered at 442.2 km s⁻¹. The column density associated with this profile can be estimated using equation 3. If we integrate this profile and divide by the total background flux density of 67 mJy (background continuum flux density of 56 mJy + background H I emission flux density of 11 mJy), we get $N(\text{HI}) = 3.8 \times 10^{19} \frac{T_s}{(100 \text{ K})} / f \text{ cm}^{-2}$. The kinetic temperature associated with this profile of FWHM = 4.75 km s⁻¹ is ≤ 493 K and if $T_s = T_k$, then the column density associated with this absorber can be as high as $1.86 \times 10^{20} \text{ cm}^{-2}$ (assuming $f = 1$). The VLA D-configuration data may also suffer from emission filling-in the absorption as the beam encompasses a much larger region than the background continuum peaks. Higher resolution VLA A-configuration and/or VLBA data are needed to resolve the structures in the background quasar and to map the H I only at the region in front of the continuum source. These observations will be crucial in understanding the small-scale structure of the ISM of UGC 7408.

In the absence of higher resolution data, it cannot be proven that two peaks and the dip in the VLA spectra represents a signature of absorption. The feature may be part of the emission profile itself. If the feature were not an absorber, then it would mean that the H I is at a higher temperature than that seen in the Cold Neutral Medium (CNM) of the Milky Way. Spin temperatures of around 1000 K or higher have been seen at low-redshifts ($z < 0.5$) in non-spiral systems (Kanekar & Chengalur 2001). Dickey et al. (2000) found that in the SMC the fraction of cool-phase H I is much lower than the overall H I abundance. The ISM properties of dwarf galaxies are different from bigger galaxies like the Milky Way and this may be the reason for the low cold-H I abundances in their ISM. Dwarf galaxies have lower abundances of metals and thus, cooling from fine-structure lines (especially from carbon and oxygen) in these galaxies is much less predominant than brighter galaxies (see Dickey et al. 2000 for details). In addition, the typical dust to gas ratios in dwarf galaxies are much lower than in the Milky Way and thus dwarf galaxies are inefficient in confining the UV-ionizing radiation to individual H II regions. Therefore, UV radiation can effectively heat up the cooler gas in dwarf galaxies to a larger radii outside

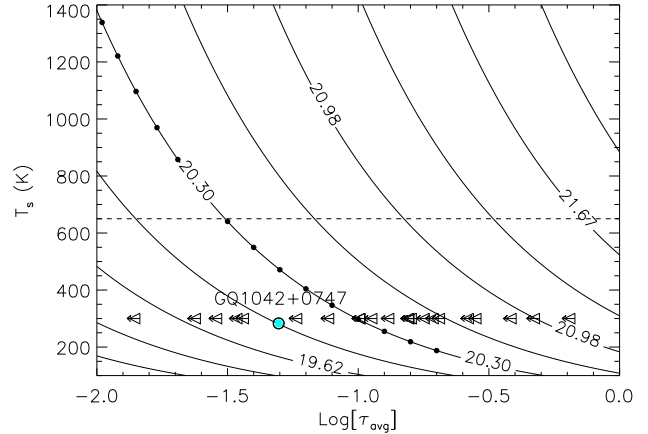


FIG. 6.— Variation of H I column density as a function of average optical depth, τ_{avg} and spin temperature, T_s for absorbers with covering fraction, $f = 1$. The contour levels mark column density in logarithmic scale. The contour at $\log N(\text{HI}) = 20.3$, which is indicated by a solid line with filled circles, represents the DLA limit of $N(\text{HI}) = 2 \times 10^{20} \text{ cm}^{-2}$. The horizontal dashed line represents the maximum T_s detected in our Galaxy within the range of τ values shown in the plot (see Fig 2a of Braun & Walterbos 1992). The maximum T_s for M31 is around 1000 K, although the majority of the data for both M31 and the Milky Way lie below 300 K. Our 21 cm absorber detected in the foreground galaxy GQ1042+0747, is plotted with a cyan circle. This is consistent with absorbers detected in the Milky Way. The GBT 3σ limiting optical depths for the rest of our foreground galaxies that show no 21 cm absorption are plotted at $T_s = 300$ K with arrow headed triangles and span a wide range of column densities.

the stellar disks. Hence, due to increased heating and reduced cooling, it is very likely that most of the H I in dwarf galaxies is at a much higher temperature.

Another possibility for the nondetection of H I absorption is the lack of H I in the foreground region. This may be a result of a clumpy ISM, which when observed with VLA D-configuration resolution get smoothed out. Puche et al. (1992) found holes in the H I distribution of Holmberg II, a much-studied star-forming dwarf galaxy, with diameters ranging from ~ 100 to 1700 pc. Similarly, Begum et al. (2008) also found gaps and/or holes in the H I distribution for dwarf galaxies such as DDO43. Maps from the WHISP survey also show that galaxies have similar patches of low intensity H I emission even in the inner regions. In the case of UGC 7408, since we have three sightlines through the ISM at impact parameters 2.6, 2.8, and 3.3 kpc, according to the last possibility, the ISM needs to be extremely clumpy. Also it requires the clumps to have very high column densities so that the average column density in emission when smoothed to VLA D-configuration resolution is around $2 \times 10^{20} \text{ cm}^{-2}$. In conclusion, we emphasize the need for high spatial resolution VLA A-configuration imaging. These observations will not only provide information on the clumpiness of the ISM in dwarf galaxies but also will shed light on the spin temperature of the H I in such systems. Furthermore, the background QSO is a UV bright source and can be observed with the *Cosmic Origin Spectrograph* (COS) on the *Hubble Space Telescope* (HST). The spectroscopic UV data would provide complementary information which will allow us to learn more about the multiphase ISM including the warmer phases of neutral gas in UGC 7408.

4.3. Connection of 21 cm Absorbers to ISM of Galaxies

Our only unambiguous GBT detection comes from the foreground galaxy GQ1042+0747, which is intercepted by a QSO sightline that has the smallest impact parameter of all the QSO-galaxy pairs in our sample, and which pierces the stellar disk of the galaxy. As discussed in Borthakur et al. (2010), the absorber is very narrow and has a column density of $N(\text{H I}) \leq 9.6 \times 10^{19} \text{ cm}^{-2}$, i.e., lower than a DLA. It is evident from our sample that we do not find cold clouds with column densities of the order of DLAs or more, associated with the halos of galaxies ubiquitously. The second possible absorber, associated with UGC 7408, probably arises in the bound ISM of the dwarf galaxy.

Other studies have found 21 cm absorbers and some of them have been associated with tidally stripped matter. For example, Stocke et al. (1991) found a 21 cm H I absorber associated with a tidal feature of NGC 3067. Cool gas associated with galactic outflows has been observed in the halos of galaxies probed by various metal-line transitions. Veilleux et al. (2005) observed such outflows in nearby starburst galaxies and Bland-Hawthorn & Cohen (2003) and Keeney et al. (2006) found indications of a bound outflow in the central region of the Milky Way. However, the nondetection of cold gas in the halos of galaxies in our sample suggests that either the amounts of H I in these systems are not high enough or $T_s > 1100$ for the outflowing material. Since most of the measurements of outflowing material come from metal-line transitions, the temperature of the gas may be well above our limit to detect 21 cm H I absorption even for our best-case sensitivity scenario. In addition, the radio background quasars are often extended and thus a clumpy structure of out-flowing material can have a relatively low covering fraction, f , thus significantly increasing the column density associated with our limiting optical depth.

The same argument can also be extended to accreting material or cold clouds condensing out of the halos. The above conclusions hold irrespective of the metallicity of the cold clouds, and therefore provide information about accreting and outflowing material with the same accuracy. Out of 15 foreground galaxies in our sample, three of them have masses $< 10^{10.3} M_\odot$ with no detectable H I emission, and sightlines toward the background quasars through the halo of the galaxies. However we did not find any 21 cm absorber in their halos. This is not surprising as “cold” accretion could easily have spin temperatures that makes it difficult to detect in 21 cm absorption. We also do not find ubiquitous 21 cm absorption in the halos of luminous galaxies unlike Lyman α studies.

5. SUMMARY

The process of gas flow into galaxies and subsequent condensation into cold gas has been a topic of various observational and numerical studies. In this paper, we investigated the presence of cold gas in the halos and extended disks of low-redshift galaxies. We find the occurrence of H I 21 cm absorption is rare (1/12) as we move outwards from the optical disk of the galaxies for our sample. We exclude the three, H I-emitting foreground galaxies as nondetections as any absorption along the background lines of sight may have been filled in by emission from H I in the ISM of the galaxy itself. The

only case where we unambiguously discovered atomic gas in absorption was in the stellar disk of the dwarf galaxy GQ1042+0747 at an impact parameter of 1.7 kpc from the QSO sightline SDSS J104257.58+074850.5. For systems with $\rho < 7.5$ kpc, which was reported as the median impact parameter for H I systems corresponding to DLA column densities by Zwaan et al. (2005), we detected H I in one out of two galaxies. Our detection rates for sightlines with $\rho < 15$ kpc do not match the results from Gupta et al. (2010). This may be partly because of the inclusion of pairs selected due to the presence of Ca II and Na I absorption (sightlines from Carilli & van Gorkom 1992) in the Gupta et al. (2010) sample.

On the other hand, we detected H I in emission in three foreground galaxies including the dwarf galaxy UGC 7408. The H I masses range from 0.9 to $81.3 \times 10^8 M_\odot$. The VLA H I imaging of UGC 7408 shows that its H I envelope extends to 2-3 times the optical size of the galaxy in the form of an ellipsoid extending in the southeast direction. Interestingly, the background continuum sources (SL# 14, 13, and 12) at impact parameters of 2.6, 2.8, and 3.3 kpc from the optical center of the galaxy is well within the H I distribution seen in VLA imaging. We did find a dip in the H I emission at the position of the background quasar. The dip is centered at 442.2 km s^{-1} and has a FWHM of 4.75 km s^{-1} corresponding to a kinetic temperature of $\leq 493 \text{ K}$. Assuming the spin temperature to be same as the kinetic temperature, we estimate the column density associated with this dip to be $1.86 \times 10^{20} \text{ cm}^{-2}$. However, higher resolution imaging is required to confirm this detection, because the dip may be a result of the emission profile from the foreground ISM of UGC 7408 and not an absorption signature. This may suggest that the H I seen in the VLA imaging may have too high a temperature to detect 21 cm absorption or the H I distribution in this galaxy is highly patchy with no H I structures at the position of the background continuum source.

Our study was designed to be most sensitive to cold gas with temperatures observed in the CNM of the Milky Way. However in a few cases our sensitivities are high enough to look for DLAs at temperatures above 500 K. For instance, along the sightline where we achieved the best 3σ limiting optical depth of 0.01, we were sensitive to DLA systems even at $>1000 \text{ K}$. On the other hand, for low temperature systems, this means that we could detect absorbers with column densities as low as $5.6 \times 10^{18} \text{ cm}^{-2}$ if $T_s = 100 \text{ K}$. Therefore, we can conclude that in most cases, there is no H I with DLA column densities at temperatures similar to that of CNM in Milky Way in the halos and extended disks of galaxies. Although our data are insufficient to confirm the inflow and/or outflow of neutral gas through the galactic halos, our nondetection of H I in the halos does hint that the process of condensation of warmer gas into H I may be occurring at regions closer to the optical disk of galaxies. Our present understanding of the cold gas distribution in the halos and extended disks of galaxies is quite limited. To extend our knowledge, high-resolution H I imaging of galaxy-quasar pairs as well as complementary information about warmer phases of gas through UV observations is crucial. In the future, the Expanded VLA (EVLA) with broad

bandwidth and the COS with high sensitivity will be able to probe systems through observations that would yield information on the physical nature and distribution of various gaseous species in the halos and extended disks of galaxies.

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Facilities: ARC (), GBT (), Sloan (), VLA ()

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